

Searching for swept-up hydrogen and helium in the late-time spectra of 11 nearby Type Ia supernovae

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ABSTRACT

The direct detection of a stellar system that explodes as a Type Ia supernova (SN Ia) has not yet been successful. Various indirect methods have been used to investigate SN Ia progenitor systems but none have produced conclusive results. A prediction of single-degenerate models is that H- (or He-) rich material from the envelope of the companion star should be swept up by the SN ejecta in the explosion. Seven SNe Ia have been analysed to date looking for signs of H-rich material in their late-time spectra and none were detected. We present results from new late-time spectra of 11 SNe Ia obtained at the Very Large Telescope using XShooter and FORS2. We present the tentative detection of H α emission for SN 2013ct, corresponding to $\sim 0.007 M_{\odot}$ of stripped/ablated companion star material (under the assumptions of the spectral modelling). This mass is significantly lower than expected for single-degenerate scenarios, suggesting that $>0.1 M_{\odot}$ of H-rich is present but not observed. We do not detect H α emission in the other 10 SNe Ia. This brings the total sample of normal SNe Ia with non-detections (<0.001 – $0.058 M_{\odot}$) of H-rich material to 17 events. The simplest explanation for these non-detections is that these objects did not result from the explosion of a CO white dwarf accreting matter from a H-rich companion star via Roche lobe overflow or symbiotic channels. However, further spectral modelling is needed to confirm this. We also find no evidence of He-emission features, but models with He-rich companion stars are not available to place mass limits.

Key words: line: profiles – supernovae: general

1 INTRODUCTION

The stellar systems and explosion mechanisms that produce Type Ia supernovae (SNe Ia) are still under debate. There are two main competing scenarios: the double-degenerate scenario (DD; Iben & Tutukov 1984; Webbink 1984) where the companion star to the exploding white dwarf is another white dwarf, and the single-degenerate scenario (SD; Whelan & Iben 1973) where the companion star is a non-degenerate star such as a main-sequence (MS), giant or sub-giant. The ‘double-detonation’ explosion mechanism has also seen a recent resurgence in popularity. In this scenario, the explosion results from the detonation of He on the white dwarf surface (accreted from either a degenerate or non-degenerate companion star), which triggers a subsequent C detonation in the core that unbinds the star (Nomoto 1982; Livne & Glasner 1990; Bildsten et al. 2007; Shen & Bildsten 2014).

Recent work has highlighted that it is likely that more than one of these channels contributes to produce ‘normal’ SNe Ia (Gilfanov & Bogdán 2010; Li et al. 2011; Nugent et al. 2011; Sternberg et al. 2011; Dilday et al. 2012; Schaefer & Pagnotta 2012; Maguire et al. 2013; Wang et al. 2013; Maoz et al. 2014; Scalzo et al. 2014). However, clear observational diagnostics of the progenitor channel for individual objects have still not been determined.

The presence of circumstellar material (CSM) around SNe Ia was originally suggested as a key observational diagnostic of the SD scenario (expected not to be present in DD scenarios). CSM signatures have been detected in a number of SNe Ia through studies of time-varying Na I D absorption lines in maximum-light spectra (Patat et al. 2007; Blondin et al. 2009; Simon et al. 2009). The presence of CSM has also been inferred through statistical studies that have shown that ~ 20 per cent of SNe Ia show signs of CSM (Sternberg et al. 2011; Maguire et al. 2013) and its presence is more common in more luminous SNe Ia (Maguire et al. 2013).

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However, recent results have also shown at least qualitatively that the observed Na I D absorption profiles could also be produced in merger scenarios and/or where the companion star is a second white dwarf (Raskin & Kasen 2013; Shen et al. 2013; Soker et al. 2013).

A unusual class of thermonuclear SNe, dubbed ‘SNe Ia-CSM’ (Silverman et al. 2013a), display strong signatures of H (H α emission with typical widths of ~ 2000 km s $^{-1}$) in their spectra at early times, consistent with interaction with a H-rich CSM: SN 2002ic (Hamuy et al. 2003; Deng et al. 2004; Kotak et al. 2004; Wang et al. 2004; Wood-Vasey et al. 2004), SN 2005gj (Aldering et al. 2006; Prieto et al. 2007), SN 2008J (Fox et al. 2011; Taddia et al. 2012), and PTF11kx (Dilday et al. 2012). A sample including new objects discovered by the Palomar Transient Factory is discussed in Silverman et al. (2013a).

SNe Ia-CSM are suggested to result from a progenitor channel involving a non-degenerate companion star such as a symbiotic system (e.g. Hamuy et al. 2003; Dilday et al. 2012). Some of these objects, such as PTF11kx, have been observed at late times (>200 d post maximum), where their H emission due to interaction with CSM is found to be still strong (Silverman et al. 2013a,b). Since these SNe appear very unusual also at early times and perhaps represent a distinct class of SNe Ia, the astrophysical community does not consider them to be ‘normal’ SNe Ia, i.e. SNe Ia that would be used in cosmological studies.

Another observational signature of a non-degenerate companion star would be shock interaction from the SN ejecta impacting on the companion star, which was calculated by Kasen (2010). This interaction would take the form of an increased luminosity in the very early-time (a few days after explosion) light curves of SNe Ia but would be dependent on the SN viewing angle. Signatures of this were searched for in a number of optical-wavelength studies but no clear detection was made (Hayden et al. 2010; Bianco et al. 2011; Ganeshalingam et al. 2011; Tucker 2011; Brown et al. 2012; Zheng et al. 2013; Olling et al. 2015). Recent early photometric observations of a SN Ia, iPTF14atg, have shown evidence of ultra-violet emission soon after explosion, consistent with the predictions of interaction with a non-degenerate companion star (Cao et al. 2015). However, iPTF14atg was not a normal SN Ia; it displayed unusual photometric and spectroscopic properties being most similar to the peculiar class of sub-luminous ‘SN 2002es-like’ SNe Ia (Ganeshalingam et al. 2012).

Given these inconclusive results on the progenitor configuration of ‘normal’ SNe Ia from CSM and shock-interaction studies to date, complementary tracers of progenitor scenarios are very important. One particular area of interest is the detection of hydrogen in late-time ($\gtrsim 200$ d) SN Ia spectra – H features are expected to be present only in the SD scenario, where H-rich material is removed from a non-degenerate companion star. A key benefit of this nebular-phase probe is that the SN ejecta is expected to be optically thin at these late phases, and therefore the detection of H-rich material does not depend on orientation angle. This H emission is distinct from that seen in SNe Ia-CSM at late times, where the broader H emission is due to interaction with H-rich CSM.

The ablation (heating) or stripping (momentum transfer) of H-rich material from a non-degenerate companion star through the impact of the SN Ia ejecta on the companion was discussed by Wheeler et al. (1975) and Chugai (1986). A number of early numerical studies were performed looking at the interaction between the SN ejecta and a non-degenerate companion star (Fryxell & Arnett 1981; Taam & Fryxell 1984; Livne et al. 1992). The predicted properties of the H- and He-rich material have since been constrained in greater detail in a number of independent

studies (Marietta et al. 2000; Meng et al. 2007; Pakmor et al. 2008; Pan et al. 2010, 2012; Liu et al. 2012, 2013).

The amount of unbound material present in the system after explosion, as well as its velocity distribution, is found to depend on the properties of the companion star. Assuming the case of Roche lobe overflow (RLOF), where the binary separation is typically $\sim 3R_*$ (where R_* is the radius of the companion star), the latest three-dimensional simulations of MS, red giant (RG) and He-star companions find unbound masses of ~ 0.1 – 0.2 , 0.6 and 0.02 – $0.06 M_\odot$, respectively (Liu et al. 2012, 2013; Pan et al. 2012). The smallest amount of material is removed in the WD+He star scenario since although the initial binary separation is small (a key parameter determining the amount of mass removed through stripping and ablation), a He star is more compact and has a higher binding energy than an MS or RG companion. This makes the removal of its envelope in the explosion more difficult. In the case of the WD+He star scenario, the material is expected to be He- instead of H-rich. The peak velocities of the unbound material are also predicted to depend on the companion-star properties. In the simulations of Pan et al. (2010, 2012), the peak velocities of the unbound material are ~ 550 , 660 , and 955 km s $^{-1}$ for WD+MS, WD+RG and WD+He star systems, respectively. The Liu et al. (2012, 2013) velocity predictions are in good agreement with these values.

Alternative mass-transfer mechanisms such as WD+RG systems that transfer mass to the white dwarf through a stellar wind (symbiotic systems) can have binary separations larger than those permitted for RLOF systems. Some recurrent nova systems such as RS Ophiuchi and T Coronae Borealis can have separations of up to $\sim 5R_*$ (Bode 1987), greater than the allowed separation for RLOF. However, $>0.5 M_\odot$ of material is still expected to be removed from the companion star in this scenario (Pan et al. 2012).

Observational studies of seven nearby ‘normal’ SNe Ia have been performed to look for low-velocity (<1000 km s $^{-1}$) H-features in their late-time spectra (Mattila et al. 2005; Leonard 2007; Shappee et al. 2013; Lundqvist et al. 2013, 2015). No positive detections have been made. Three of the late-time spectra studied were of low-resolution (~ 700 km s $^{-1}$), comparable to the width of the expected H features: SN 2001el (Mattila et al. 2005), SN 1998bu, SN 2000cx (Lundqvist et al. 2013). The spectra used by Leonard (2007) for SN 2005am and SN 2005cf, and by Shappee et al. (2013) for SN 2011fe, were significantly higher resolution (~ 150 km s $^{-1}$). A spectrum obtained for SN 2014J, the closest SN Ia for several decades, had a resolution of ~ 300 km s $^{-1}$ (Lundqvist et al. 2015).

The tightest limits on the H flux present have been obtained for the nearby SN 2011fe (Shappee & Stanek 2011), corresponding to an upper limit of $0.001 M_\odot$ of swept-up material if a linear extrapolation of the H fraction of the Mattila et al. (2005) spectral synthesis model holds (Shappee et al. 2013). Possible signatures of H at near-infrared (NIR) wavelengths were investigated for SN 1998bu using the Pa α and Pa β lines, but less stringent mass limits of $\sim 0.5 M_\odot$ were placed (Lundqvist et al. 2013).

A search for signatures of a He-rich companion star in the late-time spectra of SNe 2011fe and 2014J was carried out (Lundqvist et al. 2015). Based on the results of Lundqvist et al. (2013), it is expected that line fluxes of O and Ca produced in He-rich material would be stronger than those of He. Therefore, the presence of [Ca II] and [O I] lines with widths of <1000 km s $^{-1}$ was investigated by Lundqvist et al. (2015) – they detected no features at the expected wavelengths and placed upper He-mass limits from the [O I] 6300 Å line of <0.002 and $<0.005 M_\odot$ for SNe 2011fe and 2014J, respectively. For SN 2011fe, they suggested that this limit

Table 1. Spectral observations of the SNe Ia in our sample.

SN Name	Observation ^a MJD	Phase ^b scaled to (d)	Date ^b scaled to	MJD ^b scaled to	Instrument	Grism	Exposure ^c time (s)	References for time of max.
SN 2009ig	55485.3	405	20101016	55485.3	FORS2	300V, 300I	2400, 2400	Foley et al. (2012a)
SN 2010gp	55681.4, 55682.3, 55683.4	277	20110501	55682	FORS2	300V, 300I	8100, 8100	Miluzio et al. (2013)
SN 2011ek	56210.2, 56211.2	421	20121010	56210	FORS2	300V, 300I	5400, 5400	Maguire et al. (2012)
SN 2011iv	56210.4, 56212.3, 56223.3	318	20121023	56223	FORS2	300V, 300I	1800, 900	Foley et al. (2012b)
SN 2012cg	56420.2	339	20130508	56420.2	XSH	–	4800	Silverman et al. (2012)
SN 2012cg	56423.2, 56424.1	343	20130512	56424	FORS2	300V, 300I	600, 600	Silverman et al. (2012)
SN 2012cu	56424.2, 56445.1	344	20130602	56445	FORS2	300V, 300I	3600, 3600	Marion et al. (2012)
SN 2012fr	56600.3	357	20131104	56600.3	XSH	–	6000	Childress et al. (2013)
SN 2012ht	56718.2, 56741.2, 56742.2	433	20140312	56728	XSH	–	12600	Maguire et al. (2013)
SN 2013aa	56696.4, 56714.4	360	20140216	56704	XSH	–	3800	Maguire et al. (2013)
SN 2013cs	56741.3, 56742.3	303	20140325	56741	XSH	–	11732	Howerton et al. (2013)
SN 2013ct	56645.1	229	20131219	56645.1	XSH	–	1900	Parker et al. (2013)

^aMJD = Modified Julian date. The observation MJD refers to the mid-point of each observation block taken for the SN.

^bThe spectra have been calibrated to match the photometry at this phase, date and MJD.

^cExposure times for FORS2 are for the 300V, 300I grisms, respectively.

is sufficient to rule out all He-rich companion star models from Pan et al. (2012) and Liu et al. (2013), including for larger-than-RLOF binary separations.

If the model predictions are correct, then the non-detection mass limits of the literature sample are sufficient to rule out MS or RG companion stars in RLOF or accreting mass via a stellar wind for these seven SNe Ia. However, the sample size is small and it is unclear if these SNe Ia are representative of the full SN Ia population. Therefore, a study looking at a significantly larger sample of nebular phase spectra of SNe Ia is needed.

In this paper, we present low- and medium-resolution nebular-phase spectroscopic data of 11 nearby SNe Ia obtained using the European Southern Observatory (ESO) Very Large Telescope (VLT) with the XShooter spectrograph (Vernet et al. 2011) and with the FOCAL Reducer and low dispersion Spectrograph (FORS2; Appenzeller et al. 1998). In Section 2, we detail the observations and data reduction steps employed. The analysis and search for the predicted $H\alpha$ lines, as well as for $Pa\alpha$ and $Pa\beta$ lines at NIR wavelengths are described in Section 3. In the same section, we also discuss the results of a search for signatures of He-rich material measured using the available He I, as well as [Ca II] and [O I], lines. The discussion and conclusions are presented in Section 4. Throughout this paper we assume a Hubble constant, $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$.

2 OBSERVATIONS AND DATA REDUCTION

Nebular-phase spectra ($\gtrsim 230$ d) of 11 SNe Ia are used in this study: six obtained using VLT+XShooter and six obtained using VLT+FORS2. One object, SN 2012cg was observed with both instruments and serves as a check of the analysis and methods. We discuss the data reduction for the two instruments separately.

2.1 XShooter spectra

XShooter is an echelle spectrograph with three arms (UVB, VIS and NIR) covering the wavelength range of $\sim 3000\text{--}25000 \text{ \AA}$. The spectral format of XShooter is fixed but the resolution can be adjusted using different slit widths. For our data, we used slit widths

of 0.8, 0.9 and 0.9 arcsec in the UVB, VIS and NIR arms, corresponding to resolutions of $R \sim 6200, 8800$ and 5300 , respectively. Details of the spectral observations are given in Table 1.

The spectra were reduced using the REFLEX pipeline (version 2.6) with the XShooter module (version 2.5.0), producing flux-calibrated one-dimensional spectra in each of the three arms (Modigliani et al. 2010; Freudling et al. 2013). Photometric conditions and seeing of < 0.8 arcsec were required before execution of the observations could begin. Due to the relatively high spectral resolution of XShooter ($\sim 35 \text{ km s}^{-1}$), host galaxy features were easily identified and removed in the reduction process. For one SN in our sample, SN 2013ct, due to the potential detection of a feature consistent with $H\alpha$ emission (see Section 3.1.3 for more details), we also performed an additional extraction of the spectra using a custom-written pipeline. The two-dimensional data products from the REFLEX pipeline were used as input, but with sky-subtraction, extraction and flux-calibration performed order-by-order on the unresampled REFLEX images. This was to check carefully the host galaxy feature removal and relative flux calibration, and confirm that this was not causing the tentative emission feature that was seen.

We scale the flux of our spectra to photometry taken close in time to the spectroscopic observations to obtain absolute flux-calibrated spectra. We obtained acquisition images using r' and i' band filters, similar to the filters used by the Sloan Digital Sky Survey (York et al. 2000), and in the Johnson I band filter. The SN magnitudes were determined using point spread function (PSF) photometry. The zero-points of the images were obtained using aperture photometry of comparison stars in the images, which were then calibrated by comparison to either catalogue magnitudes from the SDSS Data Release 10 (Ahn et al. 2014) for SNe 2012cg and 2012ht, or using the XShooter acquisition and guide camera zero-points for SNe 2012fr, 2013ct, 2013cs, 2013aa. XShooter zero-points were not available for r' . Therefore, for the three SNe (2013ct, 2013cs, and 2013aa) that were observed in the r' band, a transformation, using the relative flux throughputs of the SN spectra in the relevant filters, was applied to convert r' to Johnson R for which an XShooter zero-point was available.

The flux-calibrated spectra were corrected to rest-frame wavelengths using the heliocentric velocities given in Table 2. These val-

Table 2. SN host galaxy information.

Name	Host galaxy	Vel. _{helio} ^a (km s ⁻¹)	Distance ^b (Mpc)	E(B-V) ^c (mag)
SN 2009ig	NGC 1015	2629±6	34.5±7.0	0.029
SN 2010gp	NGC 6240	7339±9	105.3±7.4	0.067
SN 2011ek	NGC 0918	1507±3	17.0±1.6	0.312
SN 2011iv	NGC 1404	1947±4	20.9±1.2	0.010
SN 2012cg	NGC 4424	461±4	15.2±1.9	0.20 ^d
SN 2012cu	NGC 4772	1040±5	43.8±7.3	0.914 ^d
SN 2012fr	NGC 1365	1636±1	19.1±0.9	0.018
SN 2012ht	NGC 3447a	1066±1	20.1±1.5	0.026
SN 2013aa	NGC 5643	1199±2	18.1±3.6	0.169
SN 2013cs	ESO 576-017	2771±1	44.1±3.1	0.082
SN 2013ct	NGC 0428	1152±2	12.1±1.2	0.025

^aHeliocentric velocities from NED.^bDistances calculated from cosmic microwave background redshifts from NED apart from redshift-independent distances for SN 2009ig (Tully 1988), SN 2011ek (Tully et al. 2013), SN 2011iv (Tully et al. 2013), SN 2012cg (Cortés et al. 2008), SN 2012cu (Tully et al. 2009), SN 2012fr (Tully et al. 2013), SN 2013aa (Tully 1988), and SN 2013ct (Tully et al. 2013).^cGalactic E(B-V) values from Schlafly & Finkbeiner (2011).^dHost galaxy extinction of E(B-V) = 0.18 mag was found for SN 2012cg by Silverman et al. (2012). For SN 2012cu, an additional host galaxy extinction of E(B-V) = 0.89 mag was determined using the results of Amanullah et al. (2015) corrected to an R_V of 3.1. The values quoted are the combined Galactic and host galaxy E(B-V) values.

ues were obtained from the NASA/IPAC Extragalactic Database¹ (NED) galaxy spectra, apart from SN 2012cg, which following Maguire et al. (2013) used the stellar velocity of Cortés et al. (2006) at the SN position.

Following Leonard (2007), the optical regions of the XShooter spectra were rebinned to 3 Å. The NIR portions of the spectra were rebinned to 6 Å to increase the signal-to-noise ratio. A host-galaxy origin for any potential ~1000 km s⁻¹ Hα emission can be confidently ruled out since it would only contribute to the central spectral bin at the binned XSH dispersion.

2.2 FORS2 spectra

For the FORS2 observations, a combination of grisms 300V and 300I (+OG590) and a 1 arcsec slit were used, yielding a resolution $R \sim 400$ –1000 over a wavelength range of 3500–10500 Å. After bias subtraction and flat-fielding an optimal, variance-weighted extraction of the spectra was performed with IRAF² task APALL. The dispersion solution was established using arc-lamp exposures and cross-checked against night-sky emission lines. Spectrophotometric standard stars, usually observed during the same nights as the SNe, were used to perform a relative flux calibration of the spectra and to correct for telluric absorptions.

To calibrate to an absolute flux scale and correct for slit losses, the spectra were scaled to match FORS2 *BVRI* photometry of the SNe obtained close in time to the spectral observations (see Table 1). The SN magnitudes in the *BVRI* frames were measured us-

ing PSF photometry, and calibrated using the FORS2 zero-points and extinction coefficients provided by the ESO data-quality monitoring³.

As a result of the significantly lower spectral resolution of the FORS2 spectra (~450 km s⁻¹) compared to the XShooter spectra, the host-galaxy subtraction was more complex. For nine of the SNe Ia⁴ observed with FORS2, emission from a H II-region was visible in the two-dimensional spectra either at the SN location or very nearby. Therefore, a clean background subtraction could not be guaranteed for these objects and we remove these SNe from further discussion since the uncertain background subtraction could mask Hα emission from the SN itself. However, we note that these SNe Ia could have displayed narrow H features that went undetected. Therefore, the number of SNe Ia with narrow H features could be higher than measured in our sample. The six remaining SNe Ia observed with FORS2, and that have a clean background subtraction, are detailed in Table 1. For these spectra, the dispersion was left at its native value of 3.25 Å.

3 ANALYSIS

We have searched for the presence of material that could have been removed from a non-degenerate companion star (via ablation and/or stripping) in a new late-time SN Ia spectral sample, as well as used model predictions to constrain quantitatively the presence or absence of solar abundance material.

3.1 Searching for signatures of H

We searched for the presence of narrow (<1000 km s⁻¹) lines of H (Hα, Paα, Paβ) in the spectra. To do this an underlying continuum must first be defined. The underlying continuum of the spectra in these regions was fit using a second-order Savitzky-Golay smoothing polynomial (Press et al. 1992), with the search wavelength region of ±22 Å of the H features excluded from the smoothing. This was to avoid biasing the continuum towards higher values if a H-emission feature was present. The smoothing scales used were significantly larger than the predicted widths of the narrow features, with widths of 80–140 Å in the optical and ~200 Å in the NIR. Further analysis and the associated uncertainties on estimating the continuum for the Hα region is discussed in Section 3.1.2.

The rebinned spectra and continuum fits at the position of Hα for our SN Ia sample are shown in Fig. 1. The Paβ (12822 Å) regions of the XShooter spectra are shown in Fig. 2. The position of any potential Paα (18750 Å) feature falls in the middle of a telluric band and detection of SN flux is very difficult. Therefore, the only NIR H line we discuss is the Paβ feature. We did not detect any strong unambiguous Hα or Paβ features in our sample when inspecting the appropriate wavelength regions. However, this does not exclude less prominent emission features being present. A quantitative estimate of the presence of Hα and associated non-detection limits is described in Section 3.1.4.

We have estimated the potential effect of telluric features on detection of narrow features in the Hα and Paβ spectral regions. The closest significant telluric feature (transmission <0.9) to the wavelength of Hα is a feature at ~6519 Å identified using the ESO

¹ <http://nedwww.ipac.caltech.edu/>² IRAF is distributed by the National Optical Astronomy Observatory, which is operated by the Association of Universities for Research in Astronomy under cooperative agreement with the National Science Foundation.³ <http://www.eso.org/observing/dfo/quality/FORS2/qc/zeropoints/zeropoints.html>⁴ SNe 2009le, 2010ev, 2010hg, 2010kg, 2011K, 2011ae, 2011at, 2011im, 2011jh

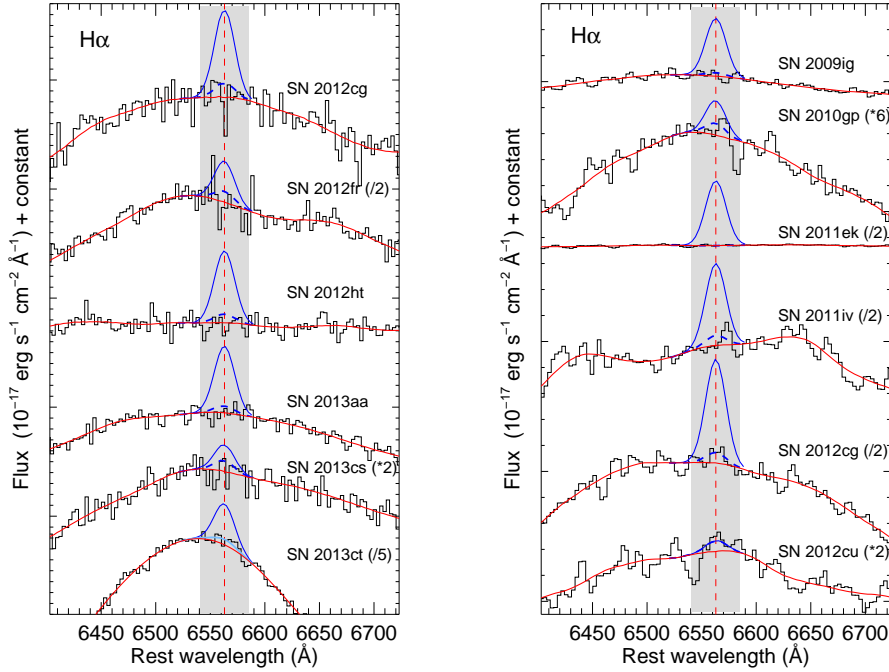


Figure 1. Left-hand panel: the $H\alpha$ spectral region of the SNe Ia with XShooter spectra at phases greater than +200 d. The black solid lines are the binned spectra (binned dispersion of 3 Å) and the red solid lines are the calculated continuum fits as described in Section 3. The red vertical dashed line marks the rest position of $H\alpha$ (6563 Å). The grey shaded region of $\sim 1000 \text{ km s}^{-1}$ ($\pm 22 \text{ Å}$) is where $H\alpha$ emission is expected based on modelling efforts. The expected $H\alpha$ emission from $0.05 M_{\odot}$ of material at +380 d as predicted by the modelling of Mattila et al. (2005) is shown as a blue solid line for each SN and has been corrected for the distance and extinction towards each SN. The estimated 3σ detection limit is shown for each SN as a blue dashed line. For presentation purposes, some spectra have been multiplied or divided by the constants specified in parentheses after the SN name. The detection of $H\alpha$ for SN 2013ct is shown as a light blue solid line. Right-hand panel: the $H\alpha$ spectral region for the SNe Ia in our sample with FORS2 spectra (binned dispersion of 3.25 Å) is shown.

SKYCALC Sky Model Calculator. For the SN in our sample with the lowest heliocentric redshift, SN 2012cg, this feature would lie at a rest frame wavelength of $\sim 6508 \text{ Å}$. This is well outside the region of interest of $\pm 22 \text{ Å}$ of the $H\alpha$ feature at 6563 Å and the feature is weak enough not to be visible in our spectra. Therefore, we conclude that this has negligible effect on our results. We have also investigated potential telluric features in the $\text{Pa}\beta$ wavelength region and find no significant telluric features within the observed wavelength region.

3.1.1 Model predictions of stripped/ablated material

Based on the calculations of Marietta et al. (2000), Mattila et al. (2005) used a one-dimensional nebular-phase spectral code (see Kozma & Fransson 1998; Sollerman et al. 2004; Kozma et al. 2005, for more details) to give a quantitative estimate of the $H\alpha$ emission expected from different amounts of solar-abundance material stripped or ablated from a companion star. This solar abundance material was located in the inner 1000 km s^{-1} of the ejecta using the W7 explosion model (Nomoto et al. 1984; Thielemann et al. 1986). The peak luminosity of the $H\alpha$ line for $0.05 M_{\odot}$ of solar-abundance material was found to be $\sim 3.36 \times 10^{35} \text{ erg s}^{-1} \text{ Å}^{-1}$ (Mattila et al. 2005; Leonard 2007).

These model $H\alpha$ luminosities were calculated at +380 d and the $H\alpha$ emission is expected to be time-dependent (Mattila et al. 2005). However, Leonard (2007) noted that since the optical depth to gamma rays should be higher at earlier times, the $H\alpha$ emission

should be stronger at epochs earlier than +380 d. Therefore, the estimated $H\alpha$ flux at +380 d can be considered a lower limit for earlier spectra as long as the condition that the ejecta are transparent is fulfilled ($\geq 200 \text{ d}$) and we can observe this low-velocity material. Sources of additional uncertainty in the modelling are discussed in Section 4.5.

For three SNe Ia in our sample (SNe 2009ig, 2011ek, and 2012ht), the spectra were obtained at epochs greater than +380 d (between +405 and +433 d). Given the small difference in the epoch studied in the model and the observed data, the model fluxes are assumed to be still applicable at these slightly later phases.

We have calculated the expected $H\alpha$ emission for this luminosity for each SN in our sample, using the distance and Galactic extinction values [corrected to the wavelength of $H\alpha$ using the Fitzpatrick (1999) extinction curve with $R_V = 3.1$] in Table 2. Host galaxy extinction corrections were not applied for most SNe Ia in our sample since we found no evidence for significant host galaxy extinction⁵. This was based on an analysis of their maximum-light spectra, absolute magnitudes at peak, and locations within their host galaxy.

The expected integrated $H\alpha$ fluxes for $0.05 M_{\odot}$ of H-rich material, assuming a Gaussian profile with a full width at half-

⁵ Based on the analysis of Silverman et al. (2012), we used a combined Galactic and host galaxy extinction of $E(B-V) = 0.20 \text{ mag}$ for SN 2012cg. For SN 2012cu, a combined extinction of $E(B-V) = 0.914 \text{ mag}$ was used (Amanullah et al. 2015).

Table 3. Measured and derived quantities for H α non-detection limits

SN	Phase (d)	H α flux ($0.05 M_{\odot}$) ^a (10^{-17} erg s $^{-1}$ cm $^{-2}$)	H α flux limit (3σ) ^b (10^{-17} erg s $^{-1}$ cm $^{-2}$)	Mass limit (3σ) ^c (M_{\odot})
SN2009ig	405	5.18 ± 1.60	0.33	0.003
SN2010gp	277	0.51 ± 0.07	0.18	0.018
SN2011ek	421	11.59 ± 1.91	0.18	0.001
SN2011iv	318	14.70 ± 1.55	1.87	0.006
SN2012cg	339	18.57 ± 3.90	3.09	0.008
SN2012cg	343	18.57 ± 3.90	1.98	0.005
SN2012cu	344	0.48 ± 0.13	0.56	0.058
SN2012fr	357	17.30 ± 1.52	4.79	0.013
SN2012ht	433	15.36 ± 2.06	2.05	0.007
SN2013aa	360	14.19 ± 4.32	1.46	0.005
SN2013cs	303	2.83 ± 0.36	1.18	0.021
SN2013ct	229	42.46 ± 7.32	$5.66 \pm 0.98^*$	$0.007 \pm 0.001^*$

^aExpected integrated line flux of H α (with FWHM of 1000 km s^{-1}) from $0.05 M_{\odot}$ of solar metallicity material according to the models of Mattila et al. (2005).

^bEstimated 3σ upper limit on the integrated flux of H α (with FWHM of 1000 km s^{-1}) estimated as described in Section 3.1.4.

^cDerived 3σ upper limits to the amount of solar metallicity material that could remain undetected.

*For SN 2013ct, the values quoted are the derived flux and mass for the detected feature at the position of H α (with best-fitting FWHM of 850 km s^{-1}).

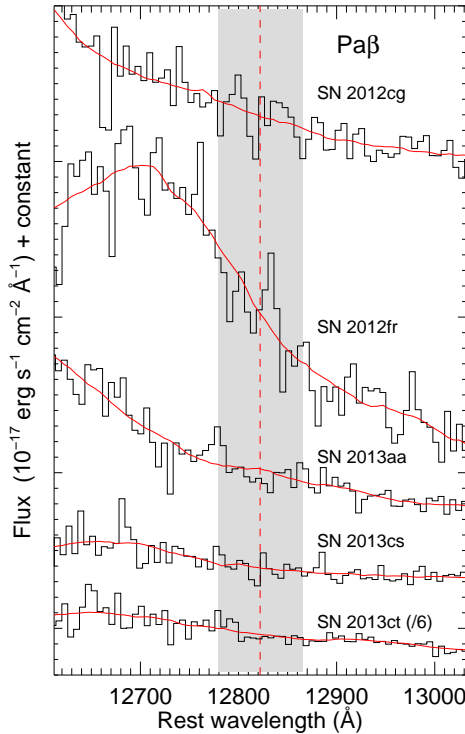


Figure 2. The Pa β (12822 Å) region is shown with the red vertical dashed line marking its rest position. The black solid lines are the binned spectra and the red solid lines are the smoothed continuum fits as described in Section 3. The grey shaded region correspond to $\pm 1000 \text{ km s}^{-1}$ of the Pa β rest wavelength. No features consistent with narrow Pa β emission are identified.

maximum (FWHM) of 1000 km s^{-1} , and a central wavelength of 6563 Å are given in Table 3. The chosen value of the FWHM is consistent with the previous studies, as well as the model constraint of Mattila et al. (2005) that the H-rich material is located at veloc-

ities below $\sim 1000 \text{ km s}^{-1}$. The expected H α line profiles for $0.05 M_{\odot}$ of solar abundance material are plotted in Fig. 1 for each SN in our sample.

3.1.2 Estimation of contributions from narrow H α

To estimate the amount of H α that could be present, as well as set limits on any non-detections, we made model spectra combining the continuum fit with a Gaussian to represent a potential H α emission feature. The Gaussian was set to have a fixed FWHM of 1000 km s^{-1} and fixed central wavelength of 6563 Å , but a varying peak flux. A likelihood function was calculated as a function of varying peak flux in steps of $0.001 M_{\odot}$ within the range corresponding to $\pm 0.05 M_{\odot}$ (negative values are included to sample a complete peak-flux distribution). Mattila et al. (2005) and Lundqvist et al. (2013) calculated the expected flux at the position of H α for differing input masses of stripped/ablated H-rich companion star material ($0.01, 0.05, 0.1, 0.5 M_{\odot}$), and found a linear relation between H α flux and mass. Therefore, we use a linear scaling between H α flux and mass in our analysis.

The largest uncertainty in the likelihood calculation is the estimation of the continuum fit. Therefore, to estimate the associated uncertainty we have calculated a number of continuum fits using the Savitzky-Golay smoothing polynomial with widths varying from 80 to 140 Å . This range was chosen so has to be significantly larger than the width of the feature we are searching for ($\pm 22 \text{ Å}$) but not so big that the underlying continuum features are not well fit. A value of 100 Å was used in the analysis of Leonard (2007). The mean of these continuum fits is then used as input to the likelihood calculation.

The sigma entering the likelihood was calculated from the rms scatter per wavelength bin from the continuum fits (an estimation of the uncertainty in the continuum definition) combined with the scatter around the normalized continuum (an estimation of the spectral noise).

3.1.3 Tentative detection of $H\alpha$ in SN 2013ct

For one SN in our sample, SN 2013ct, we have identified a weak but broad feature at the position of $H\alpha$, consistent with the stripping/ablation model predictions of $H\alpha$ emission with velocities of ~ 600 – 1000 km s^{-1} . By varying the width of the feature, we found the strongest significance feature for an FWHM of 850 km s^{-1} . Using this FWHM and the method detailed in Section 3.1.2, we found that the $H\alpha$ flux of SN 2013ct corresponded to ~ 0.005 M_{\odot} of stripped or ablated H-rich material.

Due to the strong underlying continuum for SN 2013ct and the potential detection of a narrow $H\alpha$ feature, we investigated further fits to the underlying continuum. We expanded the continuum fitting models to include a Gaussian fit to the underlying continuum. The parameters of the Gaussian were set to be variable. The region of the spectrum included in the Gaussian was varied between ± 70 and ± 130 \AA of the centre of the underlying profile (6540 \AA). The region ± 1000 km s^{-1} from the rest wavelength of $H\alpha$ was again excluded from the continuum fitting.

The mean continuum of the combined Savitzky–Golay and Gaussian fitting was used as the continuum in the likelihood calculation. The best-fitting for the $H\alpha$ feature for SN 2013ct using this expanded continuum fitting is shown in Fig. 3. This $H\alpha$ flux corresponds to $\sim 0.007 \pm 0.001$ M_{\odot} of stripped or ablated H-rich material at a 3.7σ significance.

To help determine the robustness of our $H\alpha$ detection for SN 2013ct, we performed two additional tests. First, we searched for similar features at different rest wavelengths by varying the wavelength of the expected feature within ± 120 \AA of the detected $H\alpha$ position. We used the same analysis method as for the search at the position of $H\alpha$, including excluding the search region from the continuum fitting. Secondly, we tested if a feature of a similar significance was detected on top of another broad nebular emission feature (the 5800 \AA feature) in the SN 2013ct spectrum. This was to determine if residuals in the fitting of the continuum for the broad emission features would result in a similar feature to that seen at the position of $H\alpha$. In both cases, no significant detection ($\geq 3\sigma$) is associated within any feature outside of the $H\alpha$ search region. However, we caution that the significance of the $H\alpha$ detection is under the assumptions of the fitting to an unknown underlying continuum, as well as the predictions of the models of the position and velocity distribution of H-rich material.

If our measurements of the SNe are assumed to be independent, then as the sample size increases, there is an increased probability of finding a 3σ detection by chance. We used the correction of Šidák (1967) to estimate the increased significance needed to claim a 3σ detection for one object out of the eight SNe Ia in our combined new data and literature sample for which the mass detection limit at the position of $H\alpha$ is ≤ 0.007 M_{\odot} (equivalent to the SN 2013ct detection). We found that a 3.7σ detection in any one spectrum is equivalent to a 3.1σ detection if we make eight comparisons, which we use as our significance value for this detection.

3.1.4 $H\alpha$ detection limits

The rest of the SNe Ia in our sample showed no obvious $H\alpha$ emission. Using the same method as for SN 2013ct, we determined that the Gaussian fits at the position of $H\alpha$ for the rest of the sample are consistent with zero within the 3σ uncertainties. We placed 3σ upper limits on the strength of features that could have remained undetected at the position of $H\alpha$.

The line profiles representing these flux limits are shown in

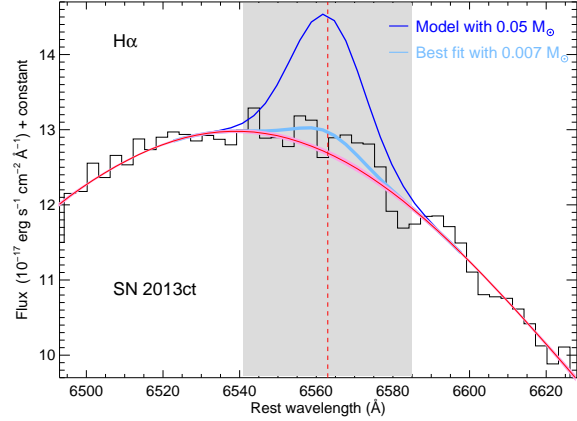


Figure 3. The $H\alpha$ region of the SN 2013ct spectrum at 229 d past maximum. The black solid line is the binned spectrum with the mean continuum fit over plotted as a red solid line. The pink shaded region is the standard deviation of the continuum fits. The red vertical dashed line marks the rest position of $H\alpha$ (6563 \AA). The grey shaded region of 1000 km s^{-1} (± 22 \AA) is where $H\alpha$ emission is expected based on modelling efforts. The expected $H\alpha$ emission from 0.05 M_{\odot} of material at $+380$ d as predicted by the modelling of Mattila et al. (2005) is shown as a blue solid line. The light blue solid line is the best-fitting detection of ~ 0.007 M_{\odot} of H with an FWHM of 850 km s^{-1} .

Fig. 1, and values for the integrated $H\alpha$ flux and corresponding mass limits are given in Table 3. The 3σ limiting mass range for the SN sample is 0.001 – 0.058 M_{\odot} . The bright, nearby SN 2012cg was observed with both XShooter and FORS2 and we obtain a consistent mass limit (< 0.010 M_{\odot}) for both spectra.

Since these limits are strictly statistical, Lundqvist et al. (2015) estimated how systematic uncertainties from our lack of knowledge of the underlying spectral continuum could affect the measured $F(3\sigma)$ values of $H\alpha$ emission for SNe 2011fe and 2014J. They suggested, using an inspection by eye, that the limits on the masses of stripped/ablated material could be 2–3 greater than their measured values.

As discussed in Section 3.1, to minimize the uncertainty in the selection of the continuum in our sample, we have excluded the region of ± 22 \AA with respect to $H\alpha$ when smoothing the spectra to avoid biasing the continuum fit of the search wavelength region. We have also tested our continuum fits by checking for similar strength features outside the $H\alpha$ search region, as well as on top of the broad 5800 \AA nebular feature, and no detection was made. Therefore, we have chosen to use our original $F(3\sigma)$ values calculated from the probability distributions to measure $M(3\sigma)$ and these are the values we quote in Table 3.

3.2 Searching for signatures of He

The companion star to the exploding white dwarf could also potentially be a He-rich star instead of a H-rich star, which could result in the presence of narrow (< 1000 km s^{-1}) He lines in late-time SN Ia spectra. Less material is expected to become unbound than in the H-rich case; just 0.02 – 0.06 M_{\odot} is predicted by the models (Pan et al. 2012; Liu et al. 2013).

To investigate the presence of He emission at late times, we focus first on the He I 5876 , 10830 , and 20587 \AA features in the XShooter VIS and NIR spectra. The wavelength regions for the He I lines are shown in Fig. 4. The He I 5876 and 10830 \AA spectral regions are found to be relatively free of telluric absorption features.

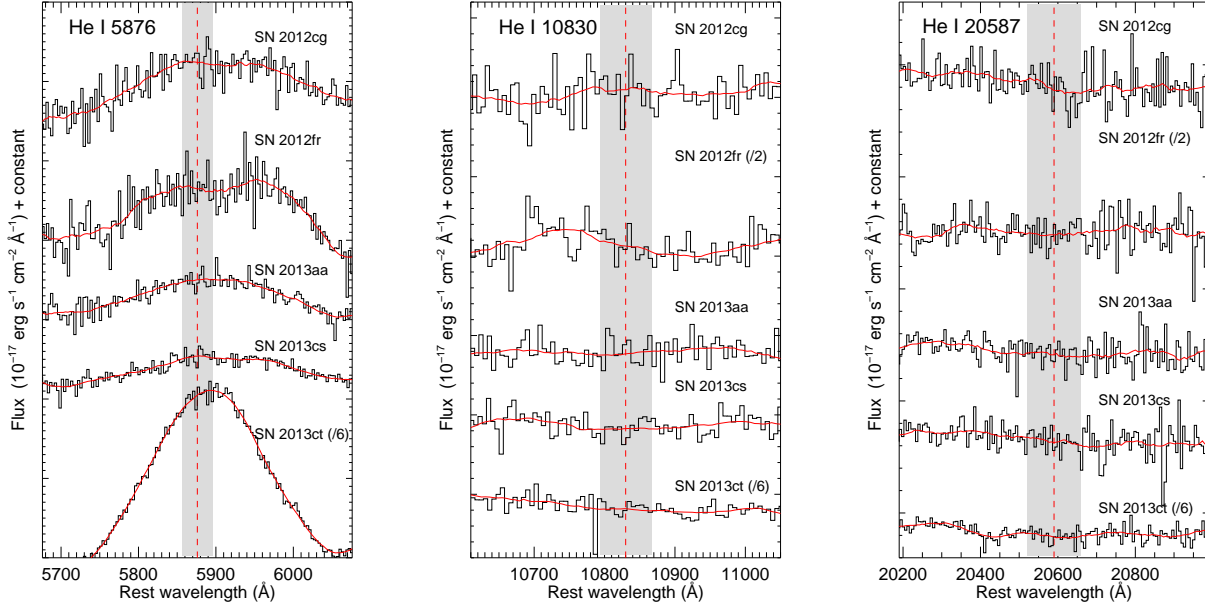


Figure 4. The He I spectral regions for five SNe Ia in our sample at phases greater than +220 d is shown. The black solid lines are the binned spectra and the red solid lines are the smoothed continuum fits as described in Section 3. The red vertical dashed lines mark the rest position of He I lines and the grey shaded regions correspond to ± 1000 km s $^{-1}$. No He I emission is identified in our spectra. We note that the He I 20590 Å region is moderately affected by telluric absorption features and therefore, our results are based predominantly on the He I 5876 and He I 10830 Å features.

However, the He I 20590 Å region is moderately affected. For completeness, we show this region in Fig. 4 but do not base our results on this region alone. We did not have model predictions of the peak luminosity for these He features but from Fig. 4 it is clear that there are no strong He I signatures detected in our SN Ia sample within the velocity range predicted by Pan et al. (2012).

Lundqvist et al. (2013) detailed how [O I] and [Ca II] emission features are promising probes of stripped/ablated He-rich material. We have also searched for [O I] 6300, 6364 Å and [Ca II] 7291, 7324 Å emission features with velocities of ± 1000 km s $^{-1}$ of the rest wavelength in our XShooter and FORS2 spectra but did not find any features consistent with the qualitative model predictions. However, further analysis, when robust spectral modelling predictions of peak luminosities are available, will be necessary to put flux (and mass) limits on these He, Ca and O non-detections.

4 DISCUSSION

Using the expected H α luminosities from a time-dependent nebular synthesis calculation for solar abundance material confined within 1000 km s $^{-1}$ (Mattila et al. 2005), we have quantified the presence or absence of H-rich companion star material in 11 SNe Ia for which late-time spectra were obtained. For 10 SNe Ia in the sample, we did not find evidence of H-rich companion star material in their late-time spectra. Under the assumptions of the modelling, we can place limits of <0.001 – $0.058 M_{\odot}$ on the mass of H-rich material that could remain undetected. For one SN in our sample, SN 2013ct, we have made a tentative detection (3.1σ) of emission at the position of H α corresponding to $0.007 \pm 0.001 M_{\odot}$ of H-rich material. Combining these new data with previous samples (Mattila et al. 2005; Leonard 2007; Lundqvist et al. 2013, 2015; Shappee et al. 2013), this means one potential detection of

H α emission out of 18 ‘normal’ SNe Ia that have been studied at late times to look for signatures of H-rich material.

Emission at the wavelength of Pa β was also investigated for five SNe Ia in the sample with suitable NIR spectral coverage and no clear emission was identified. However, as noted in Lundqvist et al. (2013), the mass limits obtained from Pa β are less constraining than for H α .

We do not detect strong signatures of He I 5876, 10830, 20587 Å features within a velocity range of ~ 1000 km s $^{-1}$. We also do not detect [O I] or [Ca II] emission features, which were suggested by Lundqvist et al. (2013) to be promising tracers of He-rich material. Lundqvist et al. (2015) estimated very rough limits on the He mass present for SNe 2011fe and 2014J by adjusting the H mass limits for the He to H number density ratio assuming solar metallicity. For the tightest mass limits for SN 2011fe in Lundqvist et al. (2015), all He-rich companion stars from the models of Pan et al. (2012) and Liu et al. (2012, 2013) can be ruled out based on the non-detection of [O I] emission. However, since the He-mass predictions obtained are simple extrapolations from models of H-rich solar abundance material placed in the inner ~ 1000 km s $^{-1}$ of the SN ejecta and the He-rich material is expected to be at higher velocity than this, they should be taken as very rough upper limits.

4.1 First detection of H α in the late-time spectrum of a normal SN Ia?

The detection of a potential H α feature consistent in velocity (<1000 km s $^{-1}$) with the models of Pan et al. (2012) and Liu et al. (2012) is suggested for SN 2013ct. Using the models of Mattila et al. (2005), this is estimated to be equivalent to $0.007 \pm 0.001 M_{\odot}$ of solar abundance material being present at velocities below 1000 km s $^{-1}$. However, the models of Pan et al. (2012) and Liu et al. (2012) have shown that if the companion star

to the exploding white dwarf was an MS star in any realistic binary scenario, at least $0.1\text{--}0.2\text{ M}_{\odot}$ of material should be removed from the companion star after explosion. The companion star can not be artificially moved to greater separations to reduce the amount of stripped/ablated material because then the necessary condition of RLOF would not be fulfilled. For an RG companion, $\sim 0.6\text{ M}_{\odot}$ of material is predicted to be stripped/ablated from the companion in RLOF. For systems transferring mass via a stellar wind, $>0.5\text{ M}_{\odot}$ of material would be removed from the companion star (Pan et al. 2012).

Therefore, in order to explain an emission feature corresponding to $\sim 0.007\text{ M}_{\odot}$ of H-rich material, an additional $\sim 0.1\text{--}0.6\text{ M}_{\odot}$ of material would have to be present but not visible as a narrow H α emission within 1000 km s^{-1} . A tail of material extending to higher velocities could hide some material that would not contribute significantly to the observed narrow H α feature. However, the models predict that the amount of stripped or ablated material present in a high-velocity tail is expected to be low (Pan et al. 2012; Liu et al. 2012).

Another possibility is that the majority of the H-rich companion star material is not sufficiently powered by radioactive heating to produce H α emission. This may be because there is little SN ejecta remaining below 1000 km s^{-1} at these times (Liu et al. 2012). Therefore, while the presence of a narrow H α feature suggests a H-rich non-degenerate companion star to the white dwarf for SN 2013ct, further constraints on the companion star properties can not currently be placed. The uncertainties of the model and observations are detailed further in Section 4.5.

4.2 Are H-rich SD systems sub-dominant for producing normal SNe Ia?

Unfortunately, little early-time data are available for SN 2013ct to determine its detailed properties. The SN was discovered on 2013 May 10 but not spectroscopically classified until 2013 May 22, where the best near-infrared spectroscopic match was a normal SN Ia approximately 20 d after maximum light (Marion et al. 2013). Therefore, while this spectrum appears similar to other normal SNe Ia, subtle differences that may have been apparent with a larger data set are not quantifiable.

Although the first tentative detection of H α emission in a late-time spectrum of a ‘normal’ SN Ia is very interesting, the absence of H α emission in the late-time spectra of 17 other SNe Ia is also worthy of discussion. Using the models of Mattila et al. (2005), stripped/ablated mass limits in the range $0.001\text{--}0.058\text{ M}_{\odot}$ were determined for these 17 SNe Ia. If we assume that the stripping/ablation and spectral model calculations are correct, then these limits are sufficient to rule out all MS and RG companions transferring mass via RLOF or a stellar wind at all plausible separations.

However, the H α detection equivalent to $\sim 0.007\text{ M}_{\odot}$ of H-rich material for SN 2013ct potentially weakens these constraints, since it suggests that $\sim 0.1\text{--}0.6\text{ M}_{\odot}$ of H-rich companion star material was hidden for this SN. Therefore, H-rich material may have also been present in the other 17 SNe Ia but also hidden.

The spin-up/spin-down scenario of Justham (2011) and Di Stefano et al. (2011) could reduce the amount of H-rich material stripped from a non-degenerate companion star at the time of explosion. This involves a white dwarf becoming spun-up by the mass accreted from its companion star, resulting in a stable white dwarf above the Chandrasekhar mass. This gives time for the companion star to evolve and contract before the SN explosion. Therefore, at the time of explosion, the companion star is much smaller and

more tightly bound, significantly reducing the amount of material removed by the impact of the SN ejecta. However, some fine-tuning is necessary to produce this scenario and as the sample of SNe Ia without H emission grows, it is becoming increasingly unlikely that this can explain all the observed non-detections.

It has also been suggested that the broad underlying emission feature at the position of H α in our sample (usually attributed to [Fe II] 6559 Å emission) may have a contribution from a broad H α emission component (Mazzali et al. 2015). However, in this case, the H-rich material would have to be present at much higher velocities than those currently predicted by modelling efforts (Liu et al. 2012; Pan et al. 2012). Further studies are necessary to determine if this is physically plausible.

However, it is likely that the sub-class of SNe Ia, SNe Ia-CSM, do result from an SD progenitor channel. The most likely scenario is that of a symbiotic system involving an RG or asymptotic giant branch star (Hamuy et al. 2003; Dilday et al. 2012). This scenario would produce the necessary CSM to explain the interaction features seen in their spectra such as broad H emission extending till late times (Silverman et al. 2013b). If SNe Ia-CSM and at least some so-called ‘normal’ SNe Ia originate from a SD scenario, then it may not be surprising to identify weaker H features, such as that detected for SN 2013ct, in some SNe Ia (coming from stripped/ablated companion material instead of pre-explosion mass loss).

4.3 He-rich companion stars

Another solution for producing ‘normal’ SNe Ia through the SD channel and avoiding H contamination is to invoke a He- instead of H-rich companion star. The latest simulations of the interaction between a He-rich companion star and the SN ejecta predict stripped/ablated masses of $0.02\text{--}0.06\text{ M}_{\odot}$ (Pan et al. 2012; Liu et al. 2013). Lundqvist et al. (2015) put limits on the mass of He-rich material present in SNe 2011fe and 2014J of <0.002 and $<0.005\text{ M}_{\odot}$, respectively. Under the assumption of RLOF, these limits rule out the He-rich companion star models of Pan et al. (2012) and Liu et al. (2013) for both SNe Ia.

We have presented a qualitative discussion of the non-detection of He emission in our late-time spectral sample of 11 SNe Ia – we do not find obvious features that could be attributed to He I, [Ca II] or [O I] emission. However, as cautioned in Lundqvist et al. (2015), these limits are based on a rough correction from H- to He-rich material, and have not been modelled explicitly. Therefore, we await more detailed spectral modelling to place limits on the presence of He-rich material from a companion star swept-up in the SN ejecta.

Companion stars with He-rich outer layers are also present in the ‘double-detonation’ scenario, where a thin layer of He on the surface of the primary white dwarf is responsible for the first detonation (Nomoto 1982; Shen & Bildsten 2014). This He material is accreted from a He star, He WD, a CO white dwarf with a thin layer of He on its surface, or it may be already present on the primary white dwarf surface. At the time of the subsequent detonation of the core, the He on the surface is expected to have velocities of $>20000\text{ km s}^{-1}$. Therefore, this material is not expected to be visible at the low velocities studied here, and the non-detection of He features in the late-time spectra does not place constraints on the presence of high-velocity He-rich material in the ‘double-detonation’ scenario. Detection of this high-velocity He is also unlikely in early-time observations because of insufficient heating of this material so far out in the ejecta to cause He emission lines to be observed.

4.4 Sample selection

The SNe Ia in our late-time spectral sample were selected for observation based on their proximity ($z \lesssim 0.025$) and visibility at Paranal at >200 d after maximum light. If an SN Ia was significantly sub-luminous at maximum light then it would not have been scheduled for observations because it would be deemed too faint at >200 d after maximum light. This would bias our sample towards more luminous SNe Ia. However, the SNe Ia in our sample span a wide range of host galaxy types from early- to late-type galaxies, suggesting a spread in SN luminosity (Hamuy et al. 1995, 1996, 2000; Riess et al. 1999; Sullivan et al. 2006). SN 2011iv which occurred in an elliptical galaxy was a low-luminosity SN Ia (Foley et al. 2012b). Therefore, we conclude that our sample is not significantly biased towards more luminous events, and covers a range of ‘normal’ SN Ia luminosities. There is also no reason to expect that less-luminous SNe Ia are more likely to have H or He features present in their spectra. In fact, recent work suggested the opposite – that it is likely that more luminous SNe Ia occur preferentially through SD channels (Silverman et al. 2013a; Maguire et al. 2014; Leloudas et al. 2015).

4.5 Observational and model uncertainties

In this section, we discuss the uncertainties and limitations of the observations and modelling that could cause the results and interpretation of our analysis to be less constraining. A number of independent simulations have been made of the amount of stripped/ablated material, and its velocity distribution, that is expected to be removed from a non-degenerate companion star after explosion (Marietta et al. 2000; Meng et al. 2007; Pakmor et al. 2008; Pan et al. 2010, 2012; Liu et al. 2012, 2013). There is now reasonable agreement among the different groups for these masses and velocity distributions for different companion star setups. However, only one analysis has been done to determine the spectral line strengths associated with different mass and velocity values (Mattila et al. 2005; Lundqvist et al. 2013). This study was based on the earliest impact simulations of Marietta et al. (2000). However, the amount and velocity distributions of the stripped/ablated material have not been updated drastically in more recent modelling, and this should not significantly affect the spectral modelling results.

The largest uncertainty in the estimation of the presence of H- or He-emission features in the observed spectra is the calculation of the underlying continuum. We have used a second-degree Savitzky–Golay smoothing polynomial (Press et al. 1992) to fit the underlying continuum, as was used in previous late-time narrow H α searches. This gives a good fit to the underlying continua in our sample, even in the case of broad underlying emission features. For SN 2013ct, we expanded our analysis of the continuum fitting to include also broad Gaussian fits to the underlying spectral feature, and again found a significant detection. However, given the tentative nature of the detection of H α emission in SN 2013ct, we could not completely exclude the possibility that the continuum fits results in a residual consistent with the detected H α emission. To determine how likely this is, we performed tests, detailed in Section 3.1.3, looking for similar strength features at different wavelengths around H α and also on top of a different broad emission feature at ~ 5800 Å. In neither case, was a similar feature found.

With regard to the spectral synthesis models, the addition of the H-rich material potentially swept-up from a non-degenerate companion star in the models of Mattila et al. (2005)

and Lundqvist et al. (2013) is somewhat ad hoc – it is made by adding varying amounts of solar metallicity material with velocities <1000 km s $^{-1}$ in a W7 density model, artificially increasing the density in the innermost region. It is uncertain if this is consistent with the density structure obtained from the three-dimensional modelling of the impact of the SN ejecta on the companion star.

In particular, a major source of uncertainty in the spectral synthesis modelling is whether the low-velocity H-rich material is sufficiently powered by radioactive heating to produce H α emission. This depends on the location of the H-rich material relative to the radioactive material; if they are not co-located then H-rich companion star material may be present in the ejecta but not observable. Mattila et al. (2005) performed the spectral line strength calculations at an epoch of +380 d. In the range of 150–300 d, the SN envelope is expected to become transparent to gamma-rays and enter a positron-dominated phase (Seitenzahl et al. 2009). However, this is not expected to be the case for the central high-density regions, where the H is located. In these high-density regions, the optical depth to gamma-rays is likely to be high enough to power the H lines (Mattila et al. 2005), but more detailed modelling need to be carried out to confirm this. Indeed, the tentative detection of low-velocity H emission in SN 2013ct, corresponding to $\sim 0.007 M_{\odot}$ of H-rich material, suggests that >0.1 – $0.6 M_{\odot}$ of stripped/ablated material must be present but not observable.

The one-dimensional models of Mattila et al. (2005) and Lundqvist et al. (2013) also assume spherical symmetry. However, the impact simulations show that the stripping and ablation of material from the companion star is not symmetric; the material is predominantly confined to the downstream region behind the companion star (Marietta et al. 2000; Pakmor et al. 2008; Liu et al. 2012, 2013; Pan et al. 2012). Since the ejecta are assumed to be optically thin at the late phases studied here ($\gtrsim 200$ d), the detection of the swept-up material is not viewing angle dependent. However, the predicted asymmetry of the unbound material could affect the shape and wavelength of the observed line profiles. As discussed in Lundqvist et al. (2015), the model used by Mattila et al. (2005) and Lundqvist et al. (2013) also included only a limited number of elements, ionization states and atomic levels. No macroscopic mixing of the companion material was included in the spectral modelling either, which could affect the predicted shape and flux of the emission lines but is not expected to be a dominant source of uncertainty.

Therefore, given the discussed uncertainties and limitations, future modelling, using the three-dimensional simulated ejecta structure of Pan et al. (2012) and Liu et al. (2012, 2013) as input to a multi-dimensional radiative transfer calculation for computing expected spectral fluxes, is of vital importance to confirm the model predictions, and hence observational mass limits.

5 CONCLUSIONS

We have presented a search for the presence of H- and He-rich material stripped or ablated from a non-degenerate companion star in new late-time spectra of 11 SNe Ia, obtained at the VLT+XShooter and the VLT+FORIS2. The observed fluxes (or limits) at the position of H α have been converted to masses using the spectral synthesis modelling described in Mattila et al. (2005) and Lundqvist et al. (2013). Our main results are as follows.

- (i) We find evidence at the 3.1σ level of H α emission with a best-fitting FWHM of ~ 850 km s $^{-1}$ for one SN Ia in our sample, SN 2013ct. This corresponds to $0.007 \pm 0.001 M_{\odot}$ of H-rich material stripped/ablated from a non-degenerate companion star.

(ii) This mass is much lower than expected for MS+WD or RG+WD progenitor systems, suggesting at least $0.1 M_{\odot}$ of H-rich material is present in SN 2013ct but not observed as narrow H α emission.

(iii) We find no evidence of H emission (H α , Pa β) in the late-time spectra of 10 other SNe Ia, bringing the total sample with no H emission detected to 17 SNe Ia.

(iv) Upper limits on the stripped/ablated mass of solar abundance material of $0.001\text{--}0.058 M_{\odot}$ are placed for these SNe Ia.

(v) These upper mass limits of H-rich solar abundance material are inconsistent with MS or RG companion stars transferring mass via RLOF or wind-driven accretion (under the assumptions of current modelling).

(vi) No signatures of He-rich material in the form of He I, [O II] or [Ca II] emission lines are identified. However, spectral modelling of the expected flux of lines from He-rich material is not available.

While future observational studies will increase the sample size of SNe Ia with the necessary late-time observations, and perhaps identify narrow H α emission in more objects, major future improvements are also likely to come from the next generation of spectral-synthesis modelling, allowing us to confirm (or adjust) these mass limits and determine if the SD scenario for producing the majority of SNe Ia is really in jeopardy.

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REFERENCES

- Ahn C. P. et al., 2014, *ApJS*, 211, 17
 Aldering G. et al., 2006, *ApJ*, 650, 510
 Amanullah R. et al., 2015, *MNRAS*, 453, 3300
 Appenzeller I. et al., 1998, *The Messenger*, 94, 1
 Bianco F. B. et al., 2011, *ApJ*, 741, 20
 Bildsten L., Shen K. J., Weinberg N. N., Nelemans G., 2007, *ApJ*, 662, L95
 Blondin S., Prieto J. L., Patat F., Challis P., Hicken M., Kirshner R. P., Matheson T., Modjaz M., 2009, *ApJ*, 693, 207
 Bode M. F., 1987, *RS Ophiuchi (1985) and the recurrent nova phenomenon*. Proc. Manchester Conf. VNU Science Press, UT.
 Brown P. J., Dawson K. S., Harris D. W., Olmstead M., Milne P., Roming P. W. A., 2012, *ApJ*, 749, 18
 Cao Y. et al., 2015, *Nature*, 521, 328
 Childress M. J. et al., 2013, *ApJ*, 770, 29
 Chugai N. N., 1986, *SvA*, 30, 563
 Cortés J. R., Kenney J. D. P., Hardy E., 2006, *AJ*, 131, 747
 Cortés J. R., Kenney J. D. P., Hardy E., 2008, *ApJ*, 683, 78
 Deng J. et al., 2004, *ApJ*, 605, L37
 Di Stefano R., Voss R., Claeys J. S. W., 2011, *ApJ*, 738, L1
 Dilday B. et al., 2012, *Science*, 337, 942
 Fitzpatrick E. L., 1999, *PASP*, 111, 63
 Foley R. J. et al., 2012a, *ApJ*, 744, 38
 Foley R. J. et al., 2012b, *ApJ*, 753, L5
 Fox O. D. et al., 2011, *ApJ*, 741, 7
 Freudling W., Romaniello M., Bramich D. M., Ballester P., Forchi V., García-Dabó C. E., Moehler S., Neeser M. J., 2013, *A&A*, 559, A96
 Fryxell B. A., Arnett W. D., 1981, *ApJ*, 243, 994
 Ganeshalingam M., Li W., Filippenko A. V., 2011, *MNRAS*, 416, 2607
 Ganeshalingam M. et al., 2012, *ApJ*, 751, 142
 Gilfanov M., Bogdán Á., 2010, *Nature*, 463, 924
 Hamuy M., Phillips M. M., Maza J., Suntzeff N. B., Schommer R. A., Aviles R., 1995, *AJ*, 109, 1
 Hamuy M. et al., 2003, *Nature*, 424, 651
 Hamuy M., Phillips M. M., Suntzeff N. B., Schommer R. A., Maza J., Aviles R., 1996, *AJ*, 112, 2391
 Hamuy M., Trager S. C., Pinto P. A., Phillips M. M., Schommer R. A., Ivanov V., Suntzeff N. B., 2000, *AJ*, 120, 1479
 Hayden B. T. et al., 2010, *ApJ*, 722, 1691
 Howerton S. et al., 2013, *Central Bureau Electronic Telegrams*, 3533, 1
 Iben I., Tutukov A. V., 1984, *ApJS*, 54, 335
 Justham S., 2011, *ApJ*, 730, L34
 Kasen D., 2010, *ApJ*, 708, 1025
 Kotak R., Meikle W. P. S., Adamson A., Leggett S. K., 2004, *MNRAS*, 354, L13
 Kozma C., Fransson C., 1998, *ApJ*, 496, 946
 Kozma C., Fransson C., Hillebrandt W., Travaglio C., Sollerman J., Reinecke M., Röpke F. K., Spyromilio J., 2005, *A&A*, 437, 983
 Leloudas G. et al., 2015, *A&A*, 574, A61
 Leonard D. C., 2007, *ApJ*, 670, 1275
 Li W. et al., 2011, *Nature*, 480, 348
 Liu Z. W., Pakmor R., Röpke F. K., Edelmann P., Wang B., Kromer M., Hillebrandt W., Han Z. W., 2012, *A&A*, 548, A2
 Liu Z.-W. et al., 2013, *ApJ*, 774, 37
 Livne E., Glasner A. S., 1990, *ApJ*, 361, 244
 Livne E., Tuchman Y., Wheeler J. C., 1992, *ApJ*, 399, 665
 Lundqvist P. et al., 2013, *MNRAS*, 435, 329
 Lundqvist P. et al., 2015, *A&A*, 577, A39
 Maguire K. et al., 2012, *MNRAS*, 426, 2359
 Maguire K. et al., 2014, *MNRAS*, 444, 3258
 Maguire K. et al., 2013, *MNRAS*, 436, 222
 Maoz D., Mannucci F., Nelemans G., 2014, *ARA&A*, 52, 107
 Marietta E., Burrows A., Fryxell B., 2000, *ApJS*, 128, 615
 Marion G. H., Hsiao E. Y., Morrell N., Parker S., 2013, *The Astronomer’s Telegram*, 5081, 1
 Marion G. H., Milisavljevic D., Rines K., Wilhelmly S., 2012, *Central Bureau Electronic Telegrams*, 3146, 2
 Mattila S., Lundqvist P., Sollerman J., Kozma C., Baron E., Fransson C., Leibundgut B., Nomoto K., 2005, *A&A*, 443, 649
 Mazzali P. A. et al., 2015, *MNRAS*, 450, 2631
 Meng X., Chen X., Han Z., 2007, *PASJ*, 59, 835
 Miluzio M. et al., 2013, *A&A*, 554, A127
 Modigliani A. et al., 2010, in *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*, Vol. 7737, *Society of Photo-Optical Instrumentation Engineers (SPIE) Conference Series*

- Nomoto K., 1982, *ApJ*, 253, 798
- Nomoto K., Thielemann F.-K., Yokoi K., 1984, *ApJ*, 286, 644
- Nugent P. E. et al., 2011, *Nature*, 480, 344
- Olling R. P. et al., 2015, *Nature*, 521, 332
- Pakmor R., Röpke F. K., Weiss A., Hillebrandt W., 2008, *A&A*, 489, 943
- Pan K.-C., Ricker P. M., Taam R. E., 2010, *ApJ*, 715, 78
- Pan K.-C., Ricker P. M., Taam R. E., 2012, *ApJ*, 750, 151
- Parker P., Maury A., Hsiao E. Y., 2013, *Central Bureau Electronic Telegrams*, 3539, 1
- Patat F. et al., 2007, *Science*, 317, 924
- Press W. H., Teukolsky S. A., Vetterling W. T., Flannery B. P., 1992, *Numerical recipes in FORTRAN. The art of scientific computing*. Cambridge: University Press, —c1992, 2nd ed.
- Prieto J. L. et al., 2007, *arXiv:0706.4088*
- Raskin C., Kasen D., 2013, *ApJ*, 772, 1
- Riess A. G. et al., 1999, *AJ*, 117, 707
- Scalzo R. A., Ruiter A. J., Sim S. A., 2014, *MNRAS*, 445, 2535
- Schaefer B. E., Pagnotta A., 2012, *Nature*, 481, 164
- Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Seitenzahl I. R., Taubenberger S., Sim S. A., 2009, *MNRAS*, 400, 531
- Shappee B. J., Stanek K. Z., 2011, *ApJ*, 733, 124
- Shappee B. J., Stanek K. Z., Pogge R. W., Garnavich P. M., 2013, *ApJ*, 762, L5
- Shen K. J., Bildsten L., 2014, *ApJ*, 785, 61
- Shen K. J., Guillochon J., Foley R. J., 2013, *ApJ*, 770, L35
- Šidák Z. K., 1967, *Journal of the American Statistical Association*, 62, 626
- Silverman J. M. et al., 2012, *ApJ*, 756, L7
- Silverman J. M. et al., 2013a, *ApJS*, 207, 3
- Silverman J. M. et al., 2013b, *ApJ*, 772, 125
- Simon J. D. et al., 2009, *ApJ*, 702, 1157
- Soker N., Kashi A., García-Berro E., Torres S., Camacho J., 2013, *MNRAS*, 431, 1541
- Sollerman J. et al., 2004, *A&A*, 428, 555
- Sternberg A. et al., 2011, *Science*, 333, 856
- Sullivan M. et al., 2006, *ApJ*, 648, 868
- Taam R. E., Fryxell B. A., 1984, *ApJ*, 279, 166
- Taddia F. et al., 2012, *A&A*, 545, L7
- Thielemann F.-K., Nomoto K., Yokoi K., 1986, *A&A*, 158, 17
- Tucker B. E., 2011, *ApJS*, 335, 223
- Tully R. B., 1988, *Nearby galaxies catalog*
- Tully R. B. et al., 2013, *AJ*, 146, 86
- Tully R. B., Rizzi L., Shaya E. J., Courtois H. M., Makarov D. I., Jacobs B. A., 2009, *AJ*, 138, 323
- Vernet J. et al., 2011, *A&A*, 536, A105
- Wang L., Baade D., Höflich P., Wheeler J. C., Kawabata K., Nomoto K., 2004, *ApJ*, 604, L53
- Wang X., Wang L., Filippenko A. V., Zhang T., Zhao X., 2013, *Science*, 340, 170
- Webbink R. F., 1984, *ApJ*, 277, 355
- Wheeler J. C., Lecar M., McKee C. F., 1975, *ApJ*, 200, 145
- Whelan J., Iben I. J., 1973, *ApJ*, 186, 1007
- Wood-Vasey W. M., Wang L., Aldering G., 2004, *ApJ*, 616, 339
- York D. G. et al., 2000, *AJ*, 120, 1579
- Zheng W. et al., 2013, *ApJ*, 778, L15