

A joint *NICER* and *XMM-Newton* view of the “Magnificent”
thermally emitting X-ray Isolated Neutron Star RX J1605.3+3249

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ABSTRACT

Thermally emitting X-ray isolated neutron stars represent excellent targets for testing cooling surface emission and atmosphere models, which are used to infer physical parameters of the neutron star. Among the seven known members of this class, RX J1605.3+3249 is the only one that still lacks confirmation of its spin period. Here we analyze *NICER* and *XMM-Newton* observations of RX J1605.3+3249, in order to address its timing and spectral behavior. Contrary to a previous tentative detection, but in agreement with the recent work by Pires et al. (2019), we find no significant pulsation with pulsed fraction higher than 1.3% (3σ) for periods above 150 ms. We also find a limit of 2.6% for periods above 2 ms, despite searches in different energy bands. The X-ray spectrum can be fit by either a double-blackbody model or by a single-temperature magnetized atmosphere model, both modified by a Gaussian absorption line at ~ 0.44 keV. The origin of the absorption feature as a proton cyclotron line or as an atomic transition in the neutron star atmosphere is discussed. The predictions of the best-fit X-ray models extended to IR, optical and UV bands are compared with archival data. Our results are interpreted in the framework of a fallback disk scenario.

Keywords: X-rays: stars – stars: neutron, magnetars, atmospheres, individual (RX J1605.3+3249)

1. INTRODUCTION

Isolated Neutron Stars (INSs) are a class of neutron stars (NSs) with no evidence for any stellar companion and an undetected or extremely weak radio counterpart (Kaspi et al. 2006). Among these is the group of thermally emitting INSs discovered by the ROSAT satellite, seven in number, and hence dubbed the “Magnificent Seven” (M7, Haberl 2007; Kaplan 2008; Turolla 2009), otherwise called thermally emitting X-ray Isolated Neutron Stars (XINS, Potekhin et al. 2015).

XINSs are thought to be members of the nearby OB associations of the Gould Belt (Walter 2001; Popov et al. 2003; Motch et al. 2003, 2005, 2009), located within a few hundreds of pc from the Sun (Posselt et al. 2007; van Kerkwijk & Kaplan 2007). Their relatively low X-ray luminosity ($\sim 10^{31-32}$ erg s⁻¹) is consistent with cooling NSs of age ~ 0.5 Myr (Page et al. 2004), in rough agreement with estimates from their kinematic ages (Page et al. 2006; Haberl 2007; Motch et al. 2009). Their X-ray spectra are soft, thermal in nature, and usually fitted by blackbody components with temperatures in the range kT $\sim 40 - 100$ eV (e.g., van Kerkwijk & Kaplan 2007; Haberl 2007). Atmospheric emission models have also been proposed to fit

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the thermal emission from INSSs (see Potekhin et al. 2016 and references therein). However, low-magnetic-field atmosphere models (Gänsicke et al. 2002; Zavlin & Pavlov 2002) fail to reproduce the spectra of M7 members. On the other hand, for the highly-magnetized atmosphere models, high-metallicity (e.g., iron) models predict a large number of absorption lines that are not observed, while pure hydrogen models reproduce spectral continua that are similar overall to the observed ones, and imply an effective temperature considerably lower than that suggested by blackbody models (Pavlov et al. 1996; Ho et al. 2007). Finally, while blackbody models underachieve the observed flux at optical wavelengths (Pavlov et al. 1996; Burwitz et al. 2003), atmospheric models overestimate it in some cases.

Timing studies of the M7 members in X-rays reveal spin periods in the range 3 – 17s and relatively high magnetic fields, $B \sim 10^{13}$ G (Kaplan et al. 2011; Haberl 2007; Hambaryan et al. 2017). Further commonality is observed among most M7 members, whose X-ray spectra show broad absorption features at energies $\sim 0.2 - 0.8$ keV (Haberl et al. 2003; Zane et al. 2005; Haberl 2007; van Kerkwijk & Kaplan 2007; Schwope et al. 2007). Such features are usually attributed either to electron/proton cyclotron resonant scattering features (eCRSF), proton cyclotron absorption features (pCF), or to bound-bound/bound-free transitions in atoms of strongly magnetized NS hydrogen atmospheres (Zavlin & Pavlov 2002; van Kerkwijk & Kaplan 2007).

The only member of the M7 group that still lacks a coherent timing solution is RX J1605.3+3249 (J1605 hereafter, Schwope et al. 1999). A possible candidate spin period of 6.9s was tentatively proposed by Haberl (2007) but not confirmed in later observations. Another possible candidate spin period of 3.4s (and a spin period derivative of $\dot{P} \sim 1.6 \times 10^{-12}$ s s $^{-1}$) was proposed for J1605 by Pires et al. (2014), which however was significant only at a low confidence level ($\sim 4\sigma$).

The source distance has been analyzed in several works. Posselt et al. (2007) find two solutions for the distance, namely 390 or 325pc. On the other hand, Motch et al. (1999) consider closer distance values, as low as ~ 100 pc, while Motch et al. (2005) link the source with the Sco OB2 association within the Gould Belt, at a mean distance of 120 – 140pc. Tetzlaff et al. (2012a) argue that J1605 was probably born in the Octans association from a supernova at ≈ 100 pc and calculates the current distance of the NS as 300 – 400pc.

A spectral feature at ~ 0.45 keV was first discovered by van Kerkwijk et al. (2004). Haberl (2007) found two additional absorption lines in the spectrum of J1605 obtained with *XMM-Newton*, at energies of 0.59 and 0.78 keV, and consistent with energies in a 2:3:4 ratio. A narrow absorption fea-

ture at 0.58keV is also found by van Kerkwijk et al. (2004) and Hohle et al. (2012) using the high energy-resolution instrument (RGS) on-board *XMM-Newton*, with a width of 3.3eV. Further analysis of *XMM-Newton* data by Pires et al. (2014) also finds significant absorption lines in J1605 at slightly different energies than previous works, that is 0.44, 0.58, and 0.83keV, while Pires et al. (2019) find no evidence of other absorption features than that at ~ 0.4 eV.

In this work we report the results of recent observations of J1605 performed with *XMM-Newton* (*XMM*) plus unpublished observations performed with the *Neutron Star Interior Composition Explorer* (*NICER*). We perform timing and spectral analysis in order to address open questions about J1605’s main properties. Combining both *NICER* and *XMM*, our observations do not show evidence of pulsations in the X-ray light curve of J1605. Moreover, our long *NICER* exposures allow to perform the most sensitive pulse search down to 2 ms to date. The measured spectra are well fitted by a double-blackbody model or by a magnetized hydrogen atmospheric emission model, and we also confirm the presence of a broad absorption feature at ~ 0.45 keV. Our timing and spectral results are consistent with the most recent work by Pires et al. (2019). We integrate archival IR/Optical/UV data in our study, and interpret our findings in the context of the most recent XINS emission scenarios and the incidence of emission and absorption of their surrounding medium.

2. OBSERVATIONS AND DATA REDUCTION

In this work, we analyzed available *NICER* data for J1605, and only the most recent *XMM* data. The *NICER* and *XMM* observation log is reported in Table 1.

2.1. *NICER*

The *NICER* X-ray Timing Instrument (XTI, Gendreau et al. 2016) is an array of 56 co-aligned X-ray concentrator optics, each associated with a silicon drift detector sensitive in the 0.2 – 12keV band (Prigozhin et al. 2012). A single concentrator consists of 24 nested grazing-incidence gold-coated aluminum foil mirrors, parabolically shaped with a common focal length. To date, 52 detectors are operating, providing a peak effective area of 1900cm 2 at ~ 1.5 keV, with an energy resolution of ~ 100 eV and a photon time-tagging resolution of ~ 100 ns. With an effective area at 1 keV of about 2000cm 2 (that is about 2 times that of *XMM*-pn and several times that of *Chandra*/ACIS), *NICER* perfectly matches the needs to analyze soft, thermal emission from NSs.

NICER observed J1605 for a total of 32 segments between 2017 July 19 (Obs ID: 1032020101) and 2018 April 8 (Obs ID: 103202032), collecting a total of ~ 165 ks of unfiltered exposure. The data were reduced using the software HEASOFT version 6.23 and NICERDAS version 2018-03-

01_V004¹. Because of the low intrinsic source flux and the relatively high spectroscopic sensitivity needed for the purposes of the present work, accurate background removal is important. Good Time Intervals (GTIs) were first created applying standard filtering (e.g., removing events detected during South Atlantic Anomaly passages). Then, further filtering was applied to remove high particle-radiation intervals associated with the Earth’s auroral zones, i.e., “the polar horns”, by applying a cut on the cutoff rigidity with $COR_SAX > 4.0$. In addition, detectors flagged as “hot” by the data analysis software were removed for each observation. Then, GTIs were separated into times when *NICER* was exposed to direct sunlight (orbit day), and times when the satellite was within the Earth shadow (orbit night). This procedure deals with different background components separately, such as the optical loading prominent only during day orbits at energies below ~ 0.35 keV. In this way low energy data are free from artificial structures, although at the expense of a considerable amount of exposure time. This does not represent the standard procedure for *NICER* data analysis but ensures a low background level in the energy range $\sim 0.3 - 0.5$ keV, where an absorption feature is expected (see Section 1). Finally, a flat count rate cut (at ~ 10 c/s) to the resulting 0.25–12 keV light curve was applied to remove possible remaining background flaring events. The resulting final exposures are 58 ks and 16 ks for day and night orbits, respectively. The source spectra were grouped using the *GRPPHA* tool to have a minimum of 25 counts per bin. The most recent response files provided by the *NICER* instrumental calibration team were used².

Background spectra were created from data acquired from one of seven “blank sky” targets based on the *Rossi X-Ray Timing Explorer (RXTE)* background fields (Jahoda et al. 2006). Among the seven available fields, we selected the target with the smallest angular separation from J1605 (BKGD #8, $\Delta\theta \sim 45^\circ$). All observations of the background field were reduced as described above. The resulting final exposures are 10 ks and 15 ks for day and night background spectra, respectively.

Source counts are 1.5×10^5 and 2.2×10^5 , corresponding to a mean count rate of 5.9 and 3.6 s^{-1} for night and day spectra, respectively, and accounting for about 94% of the total.

2.2. XMM-Newton

XMM observed J1605 a total of four times between 2015 July 21 (Obs ID: 0764460201) and 2016 February 10 (Obs ID: 0764460501). In this work, we only use data from the *XMM* European Photon Imaging Camera (EPIC) cameras—

Table 1. Observations log of RX J1605.3+3249

Telescope	Obs ID	Start Time	Exposure ^a
		[UTC]	[s]
<i>NICER</i>	1032020101	2017-07-19T22:46:40	115
<i>NICER</i>	1032020102	2017-07-20T04:57:20	2620
<i>NICER</i>	1032020103	2017-07-21T01:01:14	4545
<i>NICER</i>	1032020104	2017-07-22T00:12:38	6038
<i>NICER</i>	1032020105	2017-07-23T00:53:19	5924
<i>NICER</i>	1032020106	2017-07-24T00:04:00	3184
<i>NICER</i>	1032020107	2017-10-25T20:00:40	1111
<i>NICER</i>	1032020108	2017-10-26T00:43:15	6324
<i>NICER</i>	1032020109	2017-10-26T23:52:55	7141
<i>NICER</i>	1032020110	2017-10-28T00:30:41	9187
<i>NICER</i>	1032020111	2017-10-29T01:12:41	8565
<i>NICER</i>	1032020112	2017-10-30T00:21:42	10246
<i>NICER</i>	1032020113	2017-12-01T02:03:10	2529
<i>NICER</i>	1032020114	2017-12-02T02:50:35	2367
<i>NICER</i>	1032020115	2017-12-03T00:27:35	10105
<i>NICER</i>	1032020116	2017-12-04T01:09:16	3535
<i>NICER</i>	1032020117	2017-12-05T00:10:11	1587
<i>NICER</i>	1032020118	2017-12-06T02:29:17	6060
<i>NICER</i>	1032020119	2017-12-07T01:43:20	9660
<i>NICER</i>	1032020120	2017-12-08T00:50:43	16545
<i>NICER</i>	1032020121	2017-12-08T23:58:00	14728
<i>NICER</i>	1032020122	2017-12-10T00:41:16	10856
<i>NICER</i>	1032020123	2017-12-10T23:55:30	4672
<i>NICER</i>	1032020124	2017-12-19T00:44:31	2681
<i>NICER</i>	1032020125	2017-12-19T23:47:40	7175
<i>NICER</i>	1032020126	2017-12-21T02:02:20	2700
<i>NICER</i>	1032020127	2017-12-23T00:32:19	2944
<i>NICER</i>	1032020128	2017-12-24T01:01:53	822
<i>NICER</i>	1032020129	2017-12-26T19:56:44	143
<i>NICER</i>	1032020130	2017-12-27T23:17:00	574
<i>NICER</i>	1032020131	2018-04-01T05:57:20	186
<i>NICER</i>	1032020132	2018-04-08T03:25:20	308
<i>XMM</i>	0764460201	2015-07-21T20:19:26	121353
<i>XMM</i>	0764460301	2015-08-20T18:07:20	68000
<i>XMM</i>	0764460401	2015-08-20T18:07:20	73000
<i>XMM</i>	0764460501	2016-02-10T22:36:56	62900

^a Unfiltered times.

i.e., the PN CCD (Strüder et al. 2001) and the MOS CCDs (Turner et al. 2001)—both sensitive in the 0.15–12 keV range and offering spectral resolution $E/\Delta E \sim 20 - 50$ at 6.5 keV (see Pires et al. 2019, for analysis of RGS data). However, as J1605 has a soft spectrum, to avoid background contamination we restricted our analysis to the range 0.2–1.2 keV. The PN and MOS cameras were set in Full and Large Window modes with thin filters, respectively. The total unfiltered exposure was ~ 325 ks. Data reduction was performed using the Science Analysis System (SAS) software *xmmsas_20170719_1539-16.1.0*, with the latest available calibration files. Step-by-step reduction was

¹ https://heasarc.gsfc.nasa.gov/docs/nicer/nicer_analysis.html

² Response Matrix File and Ancillary Response File version 1.02.

performed following the official SAS Science Threads³. To remove high background flaring activity for EPIC cameras, single event (`PATTERN==0`), high energy light curves were extracted for each camera; then, a count-rate threshold was chosen corresponding to the low and steady background. Applying such a threshold to the light curves resulted in the selection of Good Time Intervals. For the pn-camera, single and double events were selected (`pattern≤4`), while single, double, triple, and quadruple events were accepted for the MOS cameras (`pattern≤12`). Background circular regions of size 60'' to 100'' were defined off-source, on the same chip as the target for the PN camera, while on a different yet close and largely source-free chip for MOS, and used to generate background spectra. PN and MOS spectra were extracted separately for each observation.

Because of the relatively high source flux, the thin filter, and the Full/Large Window observing mode, the observations are affected by pile-up at a few percent level in both PN and MOS cameras. Even though the observed count rate is within “tolerant” levels according to Jethwa et al. (2015), we opted for a conservative approach in order to ensure a confident energy redistribution of the recorded events, thus helping in constraining spectral parameters and spectral features. To minimize the resulting spectral distortion, we excised the core of the Point Spread Function (PSF) in each observation and extracted counts in an annulus centered on the source, retaining only the lower count-rate wings of the PSF. To optimize the excising radius, we excluded progressively larger radii of the PSF core (up to a radius of about 15''), testing their impact on the pile-up reduction using the `epatplot` task⁴ until pile-up effects were negligible. This resulted in a loss of about 10 and 20% of the original MOS and PN camera exposures, respectively.

We then used the SAS task `epicspeccombine` to combine spectra in order to improve statistics. However, we notice that according to the SAS team the task can be used only to merge spectra and response files that have been generated in the same PI channel interval, but also that selecting a spectrum with a non-standard PI range results in wrong response matrices and therefore unreliable spectra⁵. For this reason, we merged all the different cameras’ spectra separately, ending up with one merged PN spectrum and one merged MOS spectrum. The resulting final exposures are 285 and 280 ks for MOS2 and MOS1 and 225 ks for PN. Spectral bins have been grouped to have a minimum of 25 counts per spectral bin while the maximum oversample of the instrumental energy resolution was fixed to a factor of 3.

3. ANALYSIS

3.1. Pulsation Searches

To search for a periodic pulsed signal associated with the rotation of the NS, the event detection times were first translated to the Solar System barycenter with the `barycorr` tool in HEASOFT for *NICER* and the `barycen` task in SAS for *XMM*. For this purpose we adopted the DE405 Solar System ephemeris and the position of J1605 derived from the *XMM* EPIC pn imaging data from the longest exposure (Obs ID 0764460201), RA=16:05:18.48, Dec=+32:49:21.0.

The periodicity searches were conducted using the PRESTO Fourier-domain pulsar search software package (Ransom et al. 2002). The acceleration search technique implemented in PRESTO allowed us to coherently search long time series (up to several months) by considering a wide range of frequency drifts caused by a range of possible rotational spin-down values of the neutron star. For *XMM*, only the EPIC PN data were used for this analysis, due to the significantly greater sensitivity and better time resolution ($\Delta t = 73.4$ ms) of this instrument compared to the MOS cameras ($\Delta t = 0.9$ s). In addition, just a subset of the *NICER* observations were used for the pulsation search—in particular, the observations taken in 2017 July 19–24 (22.4 ks of unfiltered exposure), 2017 October 25–30 (42.4 ks unfiltered exposure), and 2017 December 1–27 (82.7 ks unfiltered exposure), as they provide the most compact set of deep exposures, which is desirable for sensitive coherent pulsation searches. These *NICER* subsets from July, October, and December data was first searched separately and also combined to perform a coherent search.

The *XMM* events were binned at the intrinsic EPIC pn 73.4 ms detector time resolution, while the events from *NICER* (which has an absolute time resolution of ~ 100 ns) were binned at a time resolution of 0.977 ms (1024 Hz) for the separate searches of the 2017 July, October, and December subsets. To coherently search the combined October and December *NICER* event lists, we used a 0.0625 s binning, and to search the July–December data set, we used a time binning of 0.25 s.

We first conducted periodicity searches over the 0.3–1.2 keV band. However, since J1605 may exhibit a multi-temperature thermal spectrum as suggested in previous works (e.g., Pires et al. 2014) and shown in Section 3.2, it is possible that pulsations may only arise from the hotter and smaller regions on the stellar surface. Alternatively, the pulsations of the cool and hot emission regions may be significantly out of phase such that pulsations are strongly suppressed when integrated over the full 0.3–1.2 keV band (see, e.g., Gotthelf & Halpern 2009, for the curious case of the Puppis A pulsar). To account for this scenario, we conducted additional searches for periodicity restricted to events

³ <https://www.cosmos.esa.int/web/xmm-newton/sas-threads>

⁴ <https://www.cosmos.esa.int/web/xmm-newton/sas-thread-epatplot>.

⁵ <https://www.cosmos.esa.int/web/xmm-newton/sas-thread-epic-merging>

in the 0.3–0.5 keV and 0.8–1.2 keV bands, where the cool and hot components dominate, respectively.

The XINS RX J0720.4–3125 (Borghese et al. 2015) and RX J1308.6+2127 (Borghese et al. 2017) are known to exhibit narrow absorption features that only appear over a fraction of their rotation periods. In principle, the same could be the case for J1605 such that pronounced pulsations at the neutron star rotation period only occur in the absorption line. Motivated by this prospect, we conducted searches using only events in the energy range around the absorption feature apparent in the X-ray spectrum, 0.4–0.5 keV.

No statistically significant ($\geq 4\sigma$, as determined by the Z_n^2 test; see Buccheri et al. 1983) periodic signals are found in either the *NICER* or *XMM* data sets for any choice of energy band. From the *XMM* data we can set a 3σ upper limit of 1.3% on the pulsed fraction over the 0.3–1.2 keV band, assuming a sinusoidal pulse, for spin periods greater than 0.1468 s, comparable to what Pires et al. (2019) find. For the *NICER* data in the 0.3–1.2 keV band, after accounting for the additional number of trials from the acceleration search, we obtain a pulsed fraction limit of $<2.6\%$ for spin periods greater than 1.95 ms from the 2017 December observations, a $<1.6\%$ limit for periods greater than 0.125 s by combining the 2017 October and December observations. Including the *NICER* exposures from 2017 July 19–24 data does not result in a significant improvement in sensitivity to pulsations due to the 3 month gap and shorter exposure relative to the October and December data, yielding a $<1.6\%$ limit for periods >0.5 s. The $<2.6\%$ limit (at 3σ) for short periods (>2 ms) is significantly stronger than the 3.2–5.0% limits (at 2σ) for periods >1.2 ms obtained by van Kerkwijk et al. (2004) using *XMM* pn fast timing data (see in particular their Table 3).

The absence of pulsations in substantially deeper exposures indicates that the period reported in Pires et al. (2014) based on a 60 ks *XMM* exposure was spurious, in agreement with Pires et al. (2019).

3.2. Spectral analysis

Spectra were fitted using the XSPEC 12.10.0 package (Arnaud 1996). We fitted *NICER* (day and night) and *XMM* (pn and MOS) spectra simultaneously, allowing for a cross-normalization factor among the different spectra. The energy band for fitting was limited to 0.2–1.2 keV, above which the background dominates.

Photoelectric absorption model and elemental abundances were set according to Wilms et al. (2000) (tbabs in XSPEC) to account for photoelectric absorption by neutral interstellar matter (or column density N_H), and assuming model-relative (wilm) solar abundances. During our spectral analysis we allowed the column density parameter to vary (although we also explored the case where the absorption column density was kept fixed to the Galactic value, $2.4 \times 10^{20} \text{ cm}^{-2}$). This

Table 2. Best-fit results of J1605 spectral analysis with a double-blackbody model and two atmospheric models, NSA and NSMAXG. All reported errors are at 90% c.l.

	Double-BB	NSA ($g = 2.43$)	NSMAXG
N_H [10^{20} cm^{-2}]	$4.6^{+1.4}_{-1.4}$	$3.4^{+0.2}_{-0.2}$	$5.6^{+0.3}_{-0.3}$
kT_{cool} [eV]	63^{+7}_{-6}	–	–
kT_{hot} [eV]	119^{+6}_{-4}	–	–
$\log T_{\text{eff}}$ [K]	–	$5.737^{+0.005}_{-0.005}$	$5.729^{+0.016}_{-0.022}$
M_{NS} [M_{\odot}]	–	1.4 (fixed)	$2.04^{+0.19}_{-0.49}$
R_{NS} [km]	–	10 km (fixed)	$15.6^{+0.62}_{-0.79}$
B [10^{13} G]	–	1 (fixed)	1 (fixed)
K_{cold}	$3.3^{+8.0}_{-2.2} \text{ E}+5$	–	–
K_{hot}	$1.9^{+0.8}_{-0.8} \text{ E}+3$	–	–
Distance ^a [kpc]	$0.174^{+0.127}_{-0.079}$	$0.092^{+0.005}_{-0.005}$	0.1 (fixed)
K_{atmos}	–	$1.19^{+0.12}_{-0.12} \text{ E}-4$	1.0 (fixed)
E_{abs} [keV]	$0.435^{+0.013}_{-0.006}$	$0.452^{+0.003}_{-0.003}$	$0.445^{+0.003}_{-0.003}$
σ_{abs} [keV]	$0.110^{+0.010}_{-0.011}$	$0.087^{+0.004}_{-0.005}$	$0.092^{+0.005}_{-0.004}$
τ_{abs}	$0.24^{+0.13}_{-0.08}$	$0.098^{+0.009}_{-0.008}$	$0.133^{+0.011}_{-0.016}$
C_{MOS}^b	$0.849^{+0.006}_{-0.005}$	$0.849^{+0.006}_{-0.005}$	$0.857^{+0.006}_{-0.006}$
C_{Night}^c	$0.848^{+0.013}_{-0.013}$	$0.848^{+0.013}_{-0.013}$	$0.858^{+0.013}_{-0.013}$
C_{Day}^d	$0.725^{+0.005}_{-0.005}$	$0.725^{+0.005}_{-0.005}$	$0.733^{+0.005}_{-0.005}$
Flux ^e	$1.10^{+0.42}_{-0.27}$	$1.08^{+0.04}_{-0.03}$	$1.16^{+0.40}_{-0.34}$
$\chi^2_{\text{red}}/\text{d.o.f.}$	1.12/483	1.15/485	1.14/484

Parameters indicated with K_i represent the normalization value of the *cold* and *hot* component for the double-blackbody model, and that of the *atmospheric* NSA and NSMAXG models. ^a Distance D of the source calculated as proportional, or equal to, $1/K_x^2$ for the double-BB and NSA model, respectively (see text), or reported as a parameter of the model for the NSMAXG model. ^{b,c,d} Cross-normalization factors for MOS, *NICER* night and day spectra, respectively. PN cross-normalization factor was kept fixed to unity. ^e Unabsorbed flux calculated for the continuum component(s) in the 0.3–1.2 keV band and reported in units of $10^{-11} \text{ erg cm}^{-2} \text{ s}^{-1}$. Flux values with estimated errors were derived using the cflux model from XSPEC.

resulted in values larger than the Galactic N_H value to the source ($\sim 3-5 \times 10^{20} \text{ cm}^{-2}$), as well as larger than the value found in previous works (e.g., van Kerkwijk & Kaplan 2007; Pires et al. 2014). This is similar to what Pires et al. (2019) also found, i.e. nominal values of the N_H parameter generally higher than Galactic (that is, in the range $2.5-5.3 \times 10^{20} \text{ cm}^{-2}$). Moreover, as discussed in Viganò et al. (2014), the single-blackbody model used in previous works to fit the X-ray data of J1605 (see Section 1) tends to underestimate N_H by 20–30% with respect to the actual value. However, the higher than Galactic N_H value in the case of J1605 is related to the co-variance with the parameters of the broad absorption line at $\sim 0.4 \text{ keV}$ and to our ignorance of the true continuum spectral shape. We also verified that employing different absorption models (e.g., wabs in XSPEC, see Morrison & McCammon 1983) and abundances (e.g., angr in XSPEC, see Anders & Grevesse 1989) returns consistent N_H values to those obtained using the tbabs model.

The continuum of J1605 is known to be fitted by an absorbed single- or double-blackbody model modi-

fied by at least one Gaussian absorption line at 0.45 keV (van Kerkwijk et al. 2004; Pires et al. 2014) and we tested both models to fit our data. We used the `bbodyrad` model (instead of the simpler `bbody`) from XSPEC because it allows to link the normalization K of the blackbody component to the emitting surface area through the relation $K = R_{\text{km}}^2/D_{10\text{kpc}}^2$, where R_{km} is the radius of the emitting surface in units of km and $D_{10\text{kpc}}$ is the distance from the source in units of 10kpc. A single-blackbody model returns a reduced $\chi^2 > 5$ (for 488 d.o.f.). Our data are best fit by an absorbed double-blackbody model with an absorption feature at ~ 0.43 keV (see Table 2). The absorption feature has been modeled with a Gaussian profile (the multiplicative `gabs` component in XSPEC):

$$F(E) = \exp \left\{ -\frac{\tau_{\text{abs}}}{\sqrt{2\pi}\sigma_{\text{abs}}} \exp \left(-\frac{(E - E_{\text{abs}})^2}{2\sigma_{\text{abs}}^2} \right) \right\}, \quad (1)$$

where E_{abs} is the line centroid energy, τ_{abs} and σ_{abs} are the optical depth and the width of the line, respectively. The second blackbody component and the absorption feature improve the reduced χ^2 to a value of 2.4 and 1.1, respectively. The blackbody temperatures and absorption line energy are in general agreement with findings from Pires et al. (2019). On the contrary, previous works (e.g., Haberl 2007; Pires et al. 2014) require additional absorption features at ~ 0.58 and 0.8 keV. However, we note that the observation of spectral features in the spectra of XINS can be model-dependent, and can also depend on the exact details of the analysis. Moreover, previous results based on *XMM* observations of J1605 are now superseded by Pires et al. (2019) analysis and data set. On the other hand, we notice that results from Haberl (2007) are obtained using a single-temperature blackbody model and can not therefore be compared to that used in the present work. Moreover, the fit obtained by Haberl (2007) is only marginally acceptable, with a reduced $\chi_{\text{red}}^2 \sim 1.4$, and structured residuals resulting from the best-fit model.

The absorption feature found in previous works at ~ 0.58 keV (van Kerkwijk et al. 2004; Hohle et al. 2012; Pires et al. 2014) is narrow (Gaussian width in the range $\sim 3 - 16$ eV) and needs the resolution power of instruments like RGS on board *XMM* to be resolved, which is beyond the scope of the present work. However, we note that Pires et al. (2019) also found no evidence of the narrow absorption feature at ~ 0.58 keV in the most recent *XMM*-RGS observations.

Finally, we note that Schwope et al. (2007) find that the energy of the two absorption features observed in the XINS RBS 1223 spectrum considerably changes among different observations, with the lowest energy absorption feature changing from 0.39 keV in 2003 to 0.20 keV in 2005 and the

higher energy absorption feature going from undetected to 0.73 keV, respectively (while roughly harmonic in the rest of the observations). Therefore, the centroid energy of the detected features may vary on time scales of years (a somewhat analogous phenomenon is observed in accreting NSs, e.g. Her X-1 Staubert et al. 2016, although the accretion process in those sources is likely responsible for the long-term variation of the cyclotron line energy). More recent works (Hambaryan et al. 2011; Borghese et al. 2017) also find inconsistent line energies for RBS 1223, although the inconsistency might be due to the difference in the analysis approach.

More complex, physical models have also been tested, such as those representing the X-ray spectrum emitted from the atmosphere of a NS. Various NS atmospheric models have been proposed in the literature, and tested in the present work, e.g., the magnetic/non-magnetic versions of the fully ionized hydrogen atmosphere model (NSA in XSPEC, Zavlin et al. 1996; Pavlov et al. 1995), the non-magnetic hydrogen atmosphere model with variable surface gravitational acceleration (NSAGRAV in XSPEC, Zavlin et al. 1996), and the weakly/strongly magnetized versions of the partially ionized atmospheric model that allows for variable surface gravitational acceleration and is composed of hydrogen (H) or heavier elements (e.g., carbon, oxygen, iron, NSX and NSMAXG model in XSPEC, Ho & Heinke 2009; Ho 2014). Among all tested atmospheric models, the only ones that returned an acceptable fit were (1) the NSA model in the magnetic case ($B=10^{13}$ G); and (2) the NSMAXG model with $B=10^{13}$ G, the latter with the distance and normalization values kept fixed (see Table 2). None of the other tested combinations of the above mentioned atmospheric models returned a statistically acceptable fit or physically meaningful values of the model parameters, and will not be discussed further. Moreover, we stress that the source distance commonly adopted in the literature (~ 350 pc; see, e.g., Posselt et al. 2007) is nonetheless uncertain and cannot be readily accommodated by our data, while a possible distance of ~ 150 pc (see Section 1) is consistent with our data.

Similarly to the double-blackbody model, the atmospheric model fits also require a Gaussian absorption line at ~ 0.45 keV. Finally, because the NSA model is developed for a standard gravitational acceleration $g = 2.43 \times 10^{14} \text{ cm s}^{-2}$, corresponding to standard NS mass $M_{\text{NS}} = 1.4 M_{\odot}$ and radius $R_{\text{NS}} = 10$ km, these two parameters were kept fixed during the fitting procedure⁶.

Analogous spectral continuum models and spectral features have been employed by Pires et al. (2019) to fit *XMM* data of J1605. Their spectral results are overall consistent with those found in the present work. However, their best-

⁶ <https://heasarc.gsfc.nasa.gov/xanadu/xspec/manual/node196.html>

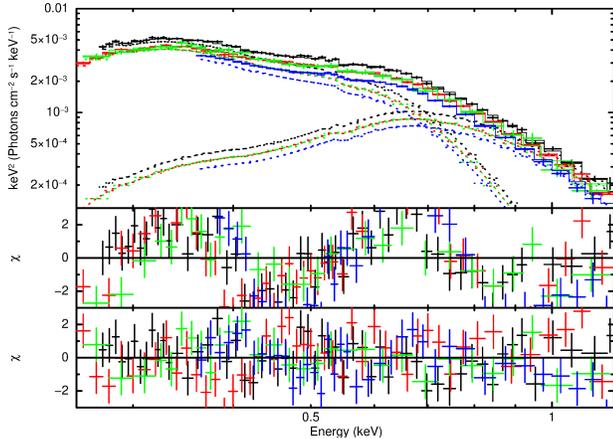


Figure 1. *NICER* and *XMM* spectra of J1605 fitted with an absorbed double-blackbody continuum plus a Gaussian absorption line. *Top panel:* Data and folded model of the *XMM* (PN and MOS – black and red, respectively) and *NICER* (night and day – green and blue, respectively) spectra of J1605. Dotted lines represent the cold and hot components (peaking at softer and higher energy, respectively). *Central panel:* residuals for the double-blackbody model without the Gaussian absorption component. *Bottom panel:* residuals for the best-fit model. Spectra and residuals have been rebinned for plotting purpose.

fit spectral models of EPIC data find an absorption feature at 385 eV, a significantly different energy than that found here, while the energy of the absorption feature found by the analysis of *XMM*-RGS spectra is consistent with our results.

4. DISCUSSION

4.1. A double-blackbody model for the continuum emission

The spectral continuum of XINS is broadly consistent with a single-temperature blackbody component ($kT_{bb} \sim 40\text{--}100\text{ eV}$) slightly modified by interstellar absorption ($N_{\text{H}} \sim 10^{20}\text{ cm}^{-2}$). This result is generally interpreted as thermal emission from the NS surface. However, with the availability of high signal-to-noise spectra, deviations from that simple model emerge (see, e.g., [Zane et al. 2011](#) and references therein). Phase-averaged as well as phase-resolved spectra from XINS generally are better fitted by a combination of two or three blackbody components, physically interpreted as coming from different regions of the star surface. Such deviations from a purely single-temperature component are expected as the result of an inhomogeneous temperature distribution across the NS surface, expected from theoretical arguments and stemming from the presence of strong magnetic fields causing, e.g., anisotropic thermal conductivity and non-spherically symmetric magnetic field dissipation, leading to the presence of hot spots ([Page et al. 2007](#); [Pons et al. 2009](#)). In particular, the origin of the double-blackbody continuum is generally attributed to two thermally emitting spots: a smaller hot spot, usually associated with the

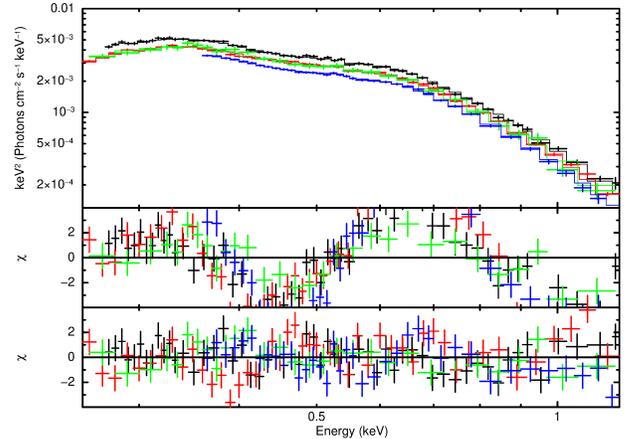


Figure 2. *NICER* and *XMM* spectra of J1605 fitted with an absorbed magnetized hydrogen atmospheric (*NSA*) model plus a Gaussian absorption line. *Top panel:* Data and folded model of the *XMM* (PN and MOS – black and red, respectively) and *NICER* (night and day – green and blue, respectively) spectra of J1605. *Central panel:* residuals for the pure *NSA* model without the Gaussian component. *Bottom panel:* residuals for the best-fit model. Spectra and residuals have been rebinned for plotting purpose.

magnetic poles, and a larger cool spot as wide as the NS itself and due to the cooling surface.

Our analysis finds that a double-blackbody model fits well the continuum emission from J1605 (see Figure 1). The temperatures of the cold and hot regions are 63 and 119 eV, respectively, in general agreement with previous work on J1605 and other XINS. Assuming a 10km-radius NS for the cold component (see Section 4.3) we get a distance of $\sim 174^{+127}_{-80}\text{ pc}$, while the normalization of the hot component, given a nominal distance of 174 pc returns a radius of the emitting hot spot equal to $\sim 0.76^{+0.18}_{-0.18}\text{ km}$. The case of the double-blackbody model with the column density value fixed to the Galactic value ($N_{\text{H}} = 2.4 \times 10^{20}\text{ cm}^{-2}$) has also been investigated. This model fits well the X-ray data ($\chi^2_{\text{red}}/d.o.f. = 1.14/484$) and returns a distance of $393 \pm 18\text{ pc}$, more in line with previous works ([Tetzlaff et al. 2012b](#), and references therein, see also Sect. 1), while both the hot and cold blackbody components show about 10 eV hotter temperature values.

Finally, to investigate deviations from a pure, single-blackbody continuum as found in [Motch et al. \(2005\)](#) for J1605, we compared that with the broadband spectral energy distribution obtained from our double-blackbody model. Our results are presented in Fig. 3. Below $\sim 0.5\text{ keV}$, the flux predicted by the double-blackbody model is about 10 times higher than that predicted by the single-blackbody, while the hot blackbody component returns a spectrum that is roughly comparable to that of the single-blackbody model of [Motch et al. \(2005\)](#), considering the relatively small difference between the temperature values (119 and 99 eV, respec-

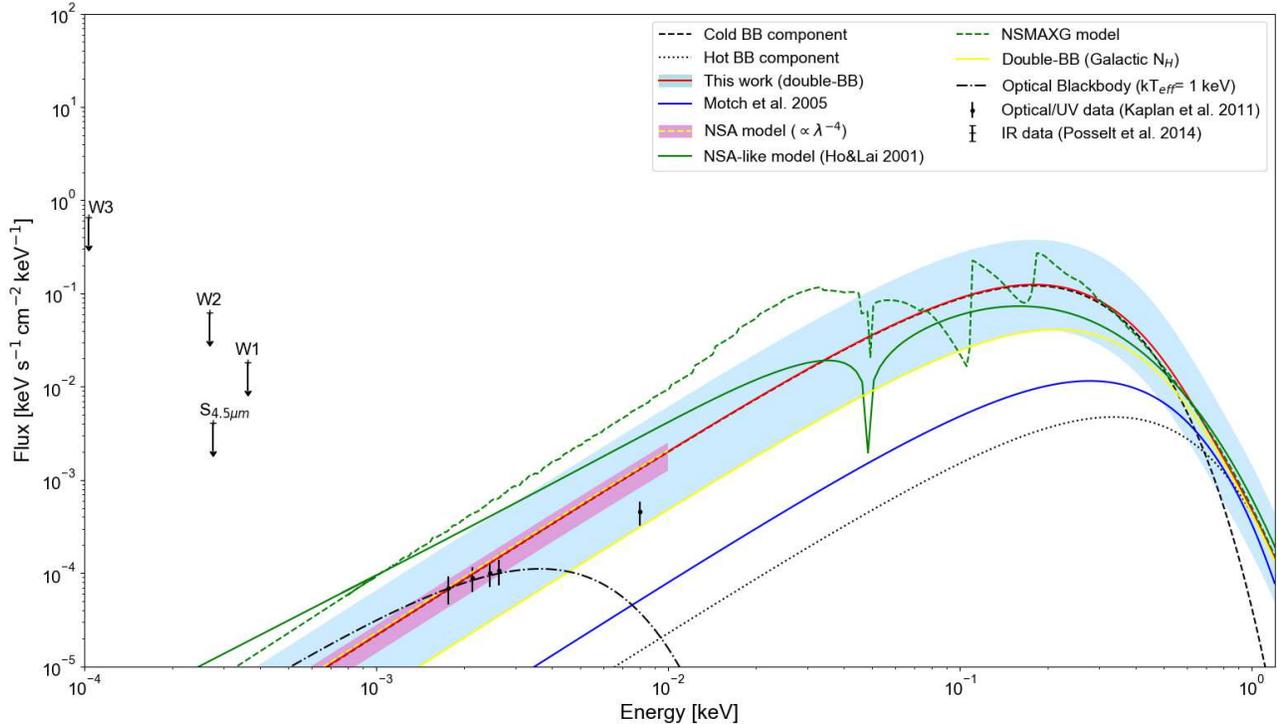


Figure 3. Broadband spectral energy distribution of J1605. Black dotted and dashed lines are the hot and cold (unabsorbed) blackbody components, respectively, obtained from the fit of X-ray data (see Table 2). The red continuous line represents the sum of the two blackbody components (uncertainty at 90% c.l. shown as the cyan shaded region). The yellow solid line represents the best-fit double-blackbody model obtained with N_{H} fixed to the Galactic value to the source. For comparison, the best-fit single-blackbody model obtained by Motch et al. (2005) is also shown (continuous blue line). IR upper limits are shown as black arrows. Optical/UV data are shown as black points with error bars, while the black dash-dotted line represents blackbody emission at $kT_{\text{eff}} = 1$ keV (see text and Ertan et al. 2017). The best-fit NSA-like model (Ho & Lai 2001) and NSMAXG model are also reported for comparison (continuous and dashed green lines, respectively). The low-energy tail predicted by the NSA model corrected for a color factor of 2.5 is shown (yellow dashed line), including propagated uncertainty (plum shaded region) in the relevant energy band.

tively). However, the cold component brings an important contribution to the soft X-ray band ($\lesssim 0.5$ keV), modifying the emerging spectrum at softer X-ray energies.

An important consequence deriving from the inclusion of a cold component is the model prediction at lower wavelengths, namely in the IR/optical/UV band. Fig. 3 compares the prediction of our model with Hubble Space Telescope (*HST*)/*Subaru* optical/UV data from Kaplan et al. (2003); Motch et al. (2005); Kaplan et al. (2011) as well as *WISE* and *Spitzer* IR data (Posselt et al. 2014). Optical/UV magnitudes have been converted to flux according to the standard STD MAG conversion (Kaplan et al. 2003, 2011), and corrected to account for optical extinction following the empirical relation described in Foight et al. (2016), $N_{\text{H}} = (2.87 \pm 0.12) \times 10^{21} \text{ A}_V \text{ cm}^{-2}$ (which, however, suffers from considerable scatter at low N_{H} values such as those derived for J1605). The model can fit optical data from the B to the R band without requiring additional components or conditions, such as a power law component or a thin hydrogen atmosphere (Motch et al. 2003, 2005; Ho et al. 2007). This is a direct result of the inclusion of a second, colder blackbody

component in the fit. However, as illustrated by Kaplan et al. (2011), J1605’s optical/UV data show a trend that is less steep than a $\propto \lambda^{-4}$, blackbody-like function. As a result, the nominal double-blackbody model fits the optical data but results in an overestimation of the UV flux, consistent with it only within the large uncertainty (see Figure 3).

4.2. The fallback disk scenario

As outlined in Sect. 4.1 and Fig. 3, the best-fit double-blackbody model is subject to further considerations. If the physical interpretation of the double-blackbody model is plausible, and only the nominal predicted values are considered, then a mechanism must be at work to suppress the inferred flux in the UV band. A possible mechanism responsible for the suppression of the UV flux in XINS is due to potential fallback disks or dusty belts surrounding the compact object which are, at least in some cases, indicated as a distinct possibility (Perna et al. 2000; Posselt et al. 2014, 2018). When present, dust grains around the NS are heated by the UV radiation, for which the grains behave as nearly perfect absorbers, thus reradiating the absorbed flux in the

infrared band. However, the double-blackbody model predicts a flux that is ~ 3 times larger than that observed in UV (see Fig. 3), which requires absorption values of a few times 10^{21} cm^{-2} , that is about 10 times the Galactic value to the source, and many times the amount observed in our best-fit model (see Table 2). Assuming that the fallback disk is distributed as a torus with inner and outer radii of 10^{10} and 10^{16} cm, respectively (see Posselt et al. 2014), a column density of $3 \times 10^{21} \text{ cm}^{-2}$ corresponds to a torus mass of about $30 M_{\oplus}$. The IR emission from such a massive torus would therefore be intense, e.g. according to Eq. 1 in Posselt et al. (2014), of the order of 10^2 mJy at $160 \mu\text{m}$ and for a source distance of 173 pc. However, Posselt et al. (2014) found no significant infrared counterpart at the position of J1605, a result that puts upper limits on the IR emission from this source, $< 12.2 \text{ mJy}$ at $160 \mu\text{m}$ (that is, approximately $18 \text{ keV s}^{-1} \text{ cm}^{-2} \text{ keV}^{-1}$ at $8 \times 10^{-6} \text{ keV}$, see Fig. 3 for comparison) and on the possible disk mass, $< 2.2 M_{\oplus}$. Therefore, a dusty fallback disk surrounding the NS seems unsuitable to explain the UV suppression implied by the double-blackbody model. Alternatively, the best-fit double blackbody model with the column density value fixed to the Galactic value can be considered. Fig. 3 shows that this model is consistent with the UV data point (at $\sim 10^{-2} \text{ keV}$), but not with the optical data. This result can be interpreted in terms of the model employed by Ertan et al. (2017), where the optical flux is emitted mainly from the inner rim of the fallback disk in the form of a blackbody spectrum. In Fig. 3 we show a blackbody model with $kT_{\text{eff}} \sim 1.3 \text{ keV}$ and optical flux $F_{\text{opt}} \sim 10^{-3} F_X$, where F_X is the X-ray flux. These values are typical of XINS (Ertan et al. 2017) and result to fit well the optical data leaving the UV flux almost unaffected.

4.3. An atmospheric model for the continuum emission

Besides the double-blackbody model, X-ray spectra from J1605 analyzed in this work are equally well fit by specific configurations of the atmospheric NSA and NSMAXG models (see Table 2). In our analysis, both models assume constant magnetic field strength and temperature across the NS surface. While one could expect magnetic field (and temperature) variation across the emission region, the non-detection of pulsations (see Section 3.1) supports our assumption.

The NSA model provides the effective unredshifted temperature of the NS surface and can be fitted for a few (fixed) values of the magnetic field, while providing the source distance as a free parameter. Our data are well fitted by the NSA model with $B = 10^{13} \text{ G}$. Similar to the double-blackbody model employed in Section 4.1, the normalization factor K_{atmos} of the NSA model is linked to the distance d (in units of parsec) through the relation $K_{\text{NSA}} = 1/d^2$, thus resulting in a distance of 92_{-5}^{+5} pc , consistent with the results from Section 4.1 and those hinted by Motch et al. (1999,

2005). The surface temperature returned by this model is $\log T_{\text{eff}} = 5.737 \pm 0.005 \text{ K}$ ($47.0 \pm 0.5 \text{ eV}$).

A constrained configuration of the NSMAXG model has also been found to fit the data. The best-fit NSMAXG model consists of a magnetized ($B = 10^{13} \text{ G}$) NS at a fixed distance of 100 pc and a normalization value, equal to the ratio of the emitting region compared to the NS radius $(R_{\text{em}}/R_{\text{NS}})^2$ fixed to unity. This model returns a similar surface temperature as that of the NSA model, $\log T_{\text{eff}} = 5.729_{-0.022}^{+0.016} \text{ K}$ ($46.2_{-2.3}^{+1.7} \text{ eV}$), and a ‘‘scaled-up’’ version of the standard NS, with $M_{\text{NS}} = 2.04 M_{\odot}$ and $R_{\text{NS}} = 15.6 \text{ km}$ (corresponding to a gravitational redshift of $z_g = 0.28$).

In analogy with Section 4.1, the NSA and NSMAXG models also have been tested with the absorption column density fixed to the Galactic value. However, these models do not fit the data well, leaving enhanced wave-like residuals ($\chi_{\text{red}}^2 = 1.4$), especially in the softest ($< 0.35 \text{ keV}$) part of the spectrum, and will not be further discussed.

Although atmospheric models have been found to satisfactorily fit the X-ray emission of XINS in previous work as well, the temperature values so obtained generally overestimate the observed optical flux by a factor of $\sim 10-100$ (see, e.g., Pavlov et al. 1996; Pons et al. 2009; Motch et al. 2003; Burwitz et al. 2003). For this, we compare the NSA model predictions with the optical/UV data provided in Motch et al. (2005); Kaplan et al. (2003, 2011), as well as with the double-blackbody model (see Section 4.1). In fact, at energies softer than the X-ray domain, the blackbody and the NSA models both follow a power-law, $F(\lambda) \propto \lambda^{-4}$ (see also Ho et al. 2008). However, the NSA model spectra implemented in XSPEC only extends down to 0.05 keV. Furthermore, at $B = 10^{13} \text{ G}$, the proton cyclotron spectral feature occurs at 0.063 keV and significantly distorts the continuum spectrum near this energy (see Figure 3). This prevents extending the NSA model to optical wavelengths in the form of a power-law. Thus, in order to illustrate a fully ionized hydrogen atmosphere spectrum at optical wavelengths, we compute and show an analogous spectrum (i.e., a NSA-like model) using the method described in Ho & Lai (2001) and values obtained from the NSA best-fit (see Table 2). As shown in Fig. 3, the NSA-like model is consistent at optical wavelengths with the double-blackbody model within its relatively large uncertainty. This is still noticeable, since other XINS show high discrepancy between the two models (as reported above). On the other hand, the best-fit NSMAXG model shows a discrepancy from the nominal double-blackbody at optical wavelengths by a factor of almost 5. Note that the NSMAXG spectrum shown in Figure 3 takes into account the dense-plasma effect described by Ho et al. (2003), which occurs when photons of frequency below the local plasma frequency have their propagation hindered; this effect causes the spectrum to deviate from a λ^{-4} behavior but the pre-

cise nature of the deviation is uncertain, and therefore the spectrum shown here is not definitive.

Finally, in Figure 3 we also show the low-energy tail predicted by the NSA model ($\propto \lambda^{-4}$) approximated as a blackbody spectrum (dashed yellow line) whose temperature is equal to that found by the best-fit model of the NSA component ($\log T_{\text{eff}} = 5.737\text{K}$, see Table 2) and corrected by a color factor of 2.5⁷. Despite the fact that color factors found by those authors and employed here are computed for non-magnetic models, it is remarkable how well the resulting spectrum meets the double-blackbody model at optical/UV wavelengths. The NSA model so obtained, similarly to the nominal double-blackbody model, is consistent with the optical data but not with the UV data, leading to a possible similar interpretation (see Section 4.1).

4.4. A pCF origin for the absorption feature

Spectral absorption features are commonly observed among XINS. These features are usually attributed to electron or proton cyclotron interactions, and/or to electronic transitions in partially ionized or condensed atmospheres (see Section 4.5). However, absorption features have been also shown to result spuriously as a consequence of fitting, e.g., multi-temperature blackbody emission models (Viganò et al. 2014).

A cyclotron resonant feature is expected in NSs with high ($B \geq 10^{12}\text{G}$) magnetic fields, where the electron/proton motion perpendicular to the magnetic field lines is quantized in discrete Landau levels, and so are the energies corresponding to those levels, thus resulting in resonant scattering of photons at those energies. The energy of the fundamental line is

$$E_{\text{CRSF}} \approx \frac{11.6}{(1+z_g)} \frac{m_e}{m_x} B_{12} \text{ keV} \quad (2)$$

where B_{12} is the magnetic field in units of 10^{12}G , m_e and m_x are the mass of the electron and that of the particle responsible for the photon scattering, respectively, and z_g the gravitational redshift (~ 0.3 for standard NS mass and radius).

In our analysis, the absorption feature in the spectrum of J1605 is found at a nominal energy of 0.432–0.451 keV (depending on the best-fit model, see Table 2). Assuming canonical values for the NS mass and radius ($M = 1.4M_{\odot}$ and $R = 10\text{km}$) and an average value of the absorption line of 0.44 keV, the resulting magnetic field is 9.0×10^{13} and $4.9 \times 10^{10}\text{G}$ for the proton and electron features, respectively.

Timing studies point out that all other pulsating XINS harbor a magnetic field of the order of 10^{13}G , an order of mag-

nitude that is consistent with the results from our fit of the atmospheric model and that inferred by the pCF (see Table 2 and Section 4.3). Moreover, we notice that in the case of pCF, the feature is expected to be narrower, with line widths of the order of hundreds of eV (Nishimura 2003, and references therein), contrary to widths of the order of keV for eCRSF, although line strength suppression by vacuum polarization in high magnetic fields plays a role in both features (see, e.g., Ho & Lai 2003). We observe a line width of $\sim 110\text{eV}$, in agreement with expectations for a pCF and comparable to the width of Gaussian absorption lines in other XINS whose independent measurements of the magnetic field favor the pCF interpretation (see, e.g., Cropper et al. 2004). Therefore, if the origin of the absorption feature is to be ascribed to cyclotron resonance, our analysis tends to favor the proton rather than the electron as the particle responsible for the scattering/absorption process.

Previous work claimed other absorption features in the spectrum of J1605 that have not been detected in the present work (Haberl 2007; Pires et al. 2014). These features had centroid energies with a 2:3:4 ratio, and were therefore interpreted as the result of harmonic cyclotron features. In Section 3.2 we gave technical reasons that might explain the different results presented in this work. Here we point out a further physical reason that highlights the difficulty to observe proton harmonic cyclotron resonant features in the spectra of highly-magnetic, thermally-emitting INs. In fact, in such physical conditions the strength of each pCF harmonic would scale with the feature centroid energy E as $\sim E/m_x c^2$ with respect to the fundamental (where m_x is the mass of the particle responsible for the scattering). Accordingly, pCFs would result in progressively weaker lines (van Kerkwijk & Kaplan 2007; Schwope et al. 2007; Potekhin 2010), contrary to the observations. However, we note that quantum effects can make electron cyclotron harmonics stronger than the simple $\sim E/m_x c^2$ scaling (Suleimanov et al. 2010, 2012), and therefore make them possible to be observed. Furthermore, Pires et al. (2019) also found no evidence of the previously reported absorption features.

Finally, as noted by Viganò et al. (2014), surface temperature inhomogeneities can mimic absorption lines, at least in the spectra of some XINS. Those authors find that a pure blackbody model can result in a spurious absorption feature at $\sim 0.45\text{keV}$ which, instead, can be accounted for by a synthetic model composed of several surface temperature distributions. However, the employment of a double-blackbody model should ensure our analysis to be free from spurious detections at $\sim 0.45\text{keV}$ because this energy value lies close to the Wien peak of one of the blackbody components, thus enhancing the robustness of our results. Nonetheless, a double-blackbody model also can result in spurious absorption features around the energy where the flux from the cold and hot

⁷ This value of the color factor is obtained by Zavlin et al. (1996) for an atmospheric spectrum with the same temperature and composition (pure hydrogen) of the NSA model considered here (see Figure 5 of their work).

components is comparable. For J1605, our analysis reveals that comparable flux contributions in the double-blackbody model are found between 0.7–0.8 keV (see Figure 1), consistent with the centroid energy of the broad absorption line reported in previous works on J1605 (see Section 1). Therefore, we conclude that the claimed broad absorption line at ~ 0.8 keV in J1605 is most likely model-dependent.

4.5. An atmospheric origin for the absorption feature

An alternative explanation for the absorption features observed among XINS spectra considers atomic transitions in a hydrogen atmosphere (Lai 2001; Medin & Lai 2006; van Kerkwijk & Kaplan 2007; Medin et al. 2008; Potekhin 2014). Absorption features resulting from magnetized atmospheres have been considered in, e.g., Sanwal et al. (2002); van Kerkwijk et al. (2004); Suleimanov et al. (2012). With a gravitational redshift of $z_g = 0.3$ (see Section 4.3), and for a mean value of the absorption feature of 0.44 keV, the unredshifted energy of the feature goes up to 0.57 keV. Considering a partially ionized hydrogen atmosphere (see Section 4.3), such a value does not correspond to expected energies except possibly a transition from ground state to first excited tightly bound state ($\nu = 0, s = 0 \rightarrow 1$, where ν, s are the principal and magnetic quantum numbers, respectively) and only if $B > 10^{13}$ G (see e.g., Lai 2001; Medin & Lai 2006; van Kerkwijk & Kaplan 2007; Ho et al. 2008). Such a high magnetic field is also required for the observed feature to be associated with spectral features due to a condensed surface spectrum (see e.g., Potekhin et al. 2012; Hambaryan et al. 2017). However, for a magnetic field strength as high as that implied by the pCF (9.0×10^{13} G, see Section 4.4), absorption features may be washed out due to the vacuum resonance mode conversion (Lai & Ho 2003; Ho & Lai 2003; Potekhin et al. 2004; van Adelsberg & Lai 2006; van Kerkwijk & Kaplan 2007; Potekhin et al. 2012; Potekhin 2014).

Even though our best-fit atmosphere models are composed of hydrogen, we notice that in Medin et al. (2008) there is an expected bound-free transition at 564.9 eV in a helium atmosphere with a magnetic field of 10^{13} G. This energy is consistent with the unredshifted feature’s nominal energy found by the best-fit double-blackbody model, that is 0.565 keV.

4.6. Lack of pulsations

A periodic X-ray signal associated with the rotation period of J1605 remains undetected. Non-detection of pulsation for this source is also reported by Pires et al. (2019). The upper limit on the pulsed fraction of $\sim 1.3\%$ we have derived is comparable to the very low modulation fraction of 1.2% observed for another XINS, RX J1856.5–3754 (Tiengo & Mereghetti 2007). Such a low level of pulsations indicates that the variations of the apparent thermal

flux caused by the changing view of the surface of the neutron star as it rotates are modest. This could be a consequence of a close alignment between the observer’s line of sight and the spin axis of the NS. However, for the case of double-blackbody emission, Pires et al. (2019) argue that the likelihood that we do not see pulsations from the source due to the particularly unfavorable viewing geometry is small, $\sim 2\%$.

Moreover, or alternatively, the temperature contrast across the stellar surface may be low, or the heat distribution may be approximately symmetric about the rotation axis, which when combined with relativistic light bending (which acts to suppress the amplitude of thermal pulsations) can result in a low pulsed fraction. At least in the case of the double-blackbody model (see Section 4.1), the temperature contrast between the hot spot and the cooling NS is large, favoring the symmetrical heat distribution interpretation.

On the other hand, the fit of atmospheric models hints to a uniform temperature across the entire surface of the NS (see Section 4.3), which would lead to isotropic emission, thus to the observed lack of pulsation. Although the above represents a tempting scenario, we note that an inhomogeneous temperature distribution is expected across the NS surface due to the presence of strong magnetic fields (see Ho 2007; Page et al. 2007; Pons et al. 2009; Hambaryan et al. 2017, and present work). Nonetheless, other theoretical works (e.g., Potekhin et al. 2003; Kaminker et al. 2006) do not expect strong deviations of the temperature distribution from spherical symmetry, even in the case of strong magnetic fields, e.g., if a “patched” multipole geometry of the magnetic field is present (see Pérez-Azorín et al. 2006).

5. SUMMARY

We have performed X-ray spectral and timing analyses of the XINS RX J1605.3+3249 observed with NICER and XMM. Our main results can be summarized as follows:

- In agreement with Pires et al. (2019), we found no evidence of pulsation for J1605 with pulsed fraction greater than 1.3% (3σ) for periods above 0.15 s. With NICER, we find no pulsations with pulsed fraction greater than 2.6% for periods above 2 ms. As such, J1605 remains the only XINS member that exhibits no pulsation. This may be due to either geometrical effects or to isotropic atmospheric emission.
- The X-ray spectrum of J1605 is equally well fitted by a double-blackbody model and by two magnetic atmospheric models: the NSA/NSA-like model and the NSMAXG model. Double-blackbody and NSA-like models predict consistent flux values at optical wavelengths. The former also fits optical archival data. A

color-corrected version of the NSA model also fits the optical data.

- Those models that fit optical data show a UV excess that, if due to absorption by material surrounding the NS, would require a column density of a few times 10^{21} cm^{-2} . However IR archival data are difficult to reconcile with this scenario. If the best-fit double-blackbody model with Galactic N_{H} is considered, UV data are also well fitted, while optical data can be accounted as blackbody emission at $kT_{\text{eff}} = 1 \text{ keV}$ from a fallback disk.
- The X-ray spectrum of J1605 shows an absorption feature at $\sim 0.44 \text{ keV}$, consistent with prior analyses but without the harmonically related features previously claimed. Our analysis favors the interpretation of this feature as a proton cyclotron resonant feature, implying a magnetic field strength of $9 \times 10^{13} \text{ G}$. Contribution to this feature may come from atomic hydrogen transition from the ground state to the first excited tightly bound state.

The lack of a well established value of the distance to J1605 prevents us from inferring additional physical characteristics of this source. It is therefore of key importance

to perform further observations in order to measure J1605's distance conclusively.

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