

Discovery of Thermonuclear Type-I X-ray Bursts from the X-ray binary MAXI J1807+132

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ABSTRACT

MAXI J1807+132 is a low-mass X-ray binary (LMXB) first detected in outburst in 2017. Observations during the 2017 outburst did not allow for an unambiguous identification of the nature of the compact object. MAXI J1807+132 was detected in outburst again in 2019 and was monitored regularly with *NICER*. In this paper we report on five days of observations during which we detected three thermonuclear (Type-I) X-ray bursts, identifying the system as a neutron star LMXB. Time-resolved spectroscopy of the three Type-I bursts revealed typical characteristics expected for these phenomena. All three Type-I bursts show slow rises and long decays, indicative of mixed H/He fuel. We find no strong evidence that any of the Type-I bursts reached the Eddington Luminosity; however, under the assumption that the brightest X-ray burst underwent photospheric radius expansion, we estimate a < 12.4 kpc upper limit for the distance. We searched for burst oscillations during the Type-I bursts from MAXI J1807+132 and found none ($< 10\%$ amplitude upper limit at 95% confidence level). Finally, we found that the brightest Type-I burst shows a ~ 1.6 sec pause during the rise. This pause is similar to one recently found with *NICER* in a bright Type-I burst from the accreting millisecond X-ray pulsar SAX J1808.4–3658. The fact that Type-I bursts from both sources can show this type of pause suggests that the origin of the pauses is independent of the composition of the burning fuel, the peak luminosity of the Type-I bursts, or whether the NS is an X-ray pulsar.

Key words: stars: neutron – stars: individual (MAXI J1807+132) – X-rays: binaries – X-rays: bursts

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1 INTRODUCTION

Low-mass X-ray binaries (LMXBs) consist of either a neutron star (NS) or a black hole (BH) primary accreting material from a low mass ($\lesssim 1 M_{\odot}$) companion star through Roche lobe overflow. Gas flowing from the companion star forms an accretion disc around the primary. As the gas in the disc spirals closer to the compact object, gravitational potential energy is released in the form of X-rays. A LMXB spends most of its life in a quiescent (very low or non-accreting) state, but when in outburst, it can reach X-ray luminosities of $L_X \simeq 10^{34-38} \text{ erg s}^{-1}$ (see, e.g., [Tauris & van den Heuvel 2006](#), for a review).

It is challenging to differentiate between LMXBs hosting a BH or NS. Dynamical measurements of a compact object's mass can be used to help identify its nature, however, the sensitivity of the observations needed to make these estimations is difficult to acquire (e.g., [Casares & Jonker 2014](#)). The detection of coherent pulsations confirms a NS identification, as these are associated with the spin frequency of the NS (e.g., [Patruno & Watts 2012](#)). Similarly, thermonuclear burning requires a solid surface for the fuel to settle on. A BH has no solid surface, but an event horizon instead, hence the detection of thermonuclear burning also secures a NS identification. There are additional observables that can be used to identify the nature of the compact object, although the identification is based on empirical evidence that certain phenomenology has only been observed in either a BH or a NS so far. For example, quasi-periodic oscillations (QPOs) with frequencies higher than 500 Hz have only been seen in X-ray light curves from observations of NS systems (see, e.g., [van der Klis 2006](#), for a review). However, the observation and characterisation of QPOs, and the power-spectral broadband noise, with frequencies below a few hundred Hertz are not always conclusive (e.g., [Klein-Wolt & van der Klis 2008](#)). Multi-wavelength observations have been also used to distinguish between BHs and NSs. For example, NSs can be ~ 30 times fainter in the radio band when compared with BHs observed at similar X-ray luminosities. However, recent works have shown that there is a population of radio-faint BHs that have similar radio luminosities to NSs ([Tetarenko et al. 2018](#)).

Thermonuclear burning in a NS atmosphere can manifest as stable, unstable, or marginally stable (e.g., [Galloway & Keek 2017](#)). Here, we concentrate on unstable thermonuclear (Type-I) X-ray bursts (hereafter referred to as “X-ray bursts”). X-ray bursts appear as a sudden increase in X-ray emission over timescales of seconds. They occur when pure or mixed material - Hydrogen, Helium, and sometimes Carbon - accreted onto the NS surface reaches a critical density and temperature which allows for runaway thermonuclear burning (see, e.g., [Lewin et al. 1993](#); [Strohmayer & Bildsten 2003](#), for reviews). During this process, the X-ray flux increases rapidly ($\lesssim 1 - 10$ s) and is generally followed by an exponential-like decay over tens of seconds as the NS atmosphere cools.

When a Type-I X-ray burst reaches the Eddington limit, L_{Edd} , it results in photospheric radius expansion (PRE, see, e.g., [Tawara et al. 1984](#); [Lewin et al. 1984](#)), where the outward radiation pressure exceeds the inward gravitational force. Since black body luminosity scales as $L \propto R^2 T^4$, when the X-ray burst reaches the Eddington limit an increase in

Table 1. Details of the *NICER* observations analysed in this paper. Quoted exposure times are as after processing.

#	ObsID	Start Time (MJD)	Date (DD-MM-YY)	Exposure (sec)	X-ray Burst
1	2200840122	58783.04946	27-10-2019	15838	-
2	2200840123	58784.01681	28-10-2019	12557	Yes
3	2200840124	58785.04902	29-10-2019	5839	Yes
4	2200840125	58786.27471	30-10-2019	4480	Yes
5	2200840126	58787.04898	31-10-2019	4850	-

radius and decrease in temperature can be seen in time-resolved spectral analysis of PRE X-ray bursts. An X-ray burst's luminosity remains roughly constant at L_{Edd} during PRE, thus PRE X-ray bursts can be used as empirical standard candles ([van Paradijs 1978](#); [Kuulkers et al. 2003](#)).

MAXI J1807+132 (hereafter MAXI J1807) was first discovered by the nova-alert system of the Monitor of All-sky X-ray Image Gas Slit Camera (*MAXI/GSC*; [Matsuoka et al. 2009](#); [Mihara et al. 2011](#)) during its 2017 outburst ([Negoro et al. 2017](#)). This detection was followed up with The Neil Gehrels Swift Observatory (*Swift*; [Gehrels 2004](#)) X-Ray Telescope (XRT; [Burrows et al. 2003](#)) observations, spectral analysis of which suggested a LMXB with a NS primary ([Shidatsu et al. 2017](#)). Further X-ray studies with *XMM-Newton*'s European Photon Imaging Camera (EPIC; [Strüder et al. 2001](#)) and ground-based optical telescopes supported the NS identification, although the possibility of a BH primary could not be ruled out ([Jiménez-Ibarra et al. 2019](#)).

After roughly 2 years in quiescence, a new outburst of MAXI J1807 was detected on 10 September 2019 (MJD 58736) by the *MAXI/GSC* nova-alert system ([Shidatsu et al. 2019](#)). Subsequently, the Neutron Star Interior Composition Explorer (*NICER*; [Gendreau & Arzoumanian 2017](#)) monitored the source from 16 September (MJD 58742), and observed the system on a regular basis. The outburst of MAXI J1807 was characterised by flaring events lasting days to weeks ([Rapisarda et al. 2019](#)). Two X-ray bursts were detected by *NICER* on 28 and 29 October, and preliminary analysis of these two events reported by [Arzoumanian et al. \(2019\)](#) confirmed MAXI J1807 as a neutron star LMXB. In this paper we report on the detection of a third X-ray burst with *NICER*, and present a detailed analysis of all three Type-I bursts.

2 OBSERVATIONS AND DATA ANALYSIS

NICER is an X-ray telescope on board the International Space Station. It was launched in 2017 with instrumentation specifically designed for the study of NSs. *NICER*'s X-ray Timing Instrument (XTI) operates in the 0.2–12 keV energy band, providing high timing and spectral resolution of 100 ns and $6 < E/\Delta E < 80$ from 0.5 keV to 8 keV respectively. With 52 active detectors, *NICER* provides an effective area of 1900 cm² at 1.5 keV.

NICER observed MAXI J1807 between 16 September and 26 November 2019, generating a total of 47 observation IDs (ObsIDs). We searched all available data for X-ray

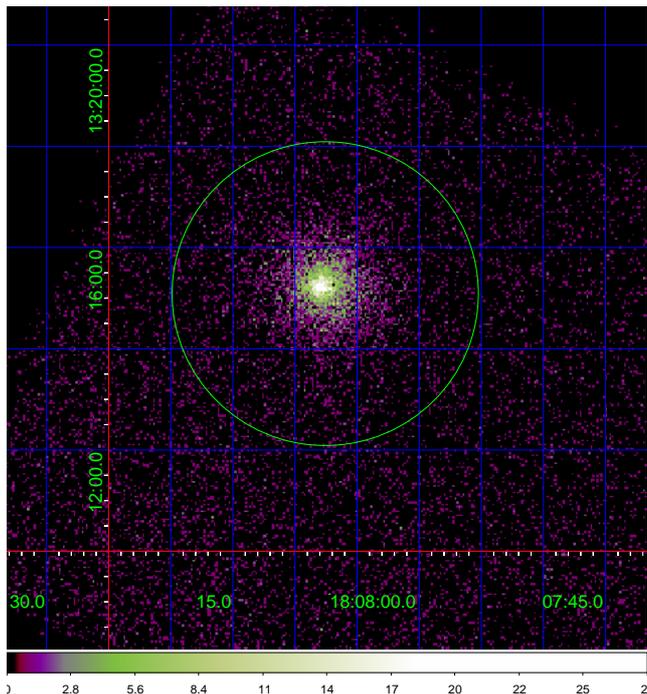


Figure 1. A 3–78 keV *NuSTAR* image of the *NICER* field-of-view, 26 September 2019 (ObsID 90501342002, exposure time of 22 ks). The *NICER* field-of-view for our observations is denoted by a green circle of 3 arcmin radius.

bursts; here we report on the five observations around the time of the detection of three X-ray bursts (see Table 1). The observations correspond to ObsIDs 2200840122–26, performed between 27 October (MJD 58783) and 31 October (MJD 58787), where X-ray bursts are observed in ObsIDs 2200840123–25. The detailed analysis of the spectral and variability evolution of the full outburst will be presented elsewhere.

To process the data, we used *HEASOFT* v6.26.1 and *NICERDAS* v6 (HEASARC 2014), and applied standard filtering criteria, i.e., pointing offset $< 54''$, Earth limb elevation angle $> 15^\circ$, and bright Earth limb angle $> 30^\circ$. Because the observations occurred during an epoch of high optical loading, the parameter “*underonly_range*” was increased to 0 – 400 when reprocessing the data using the *nicer12* tool. The first X-ray burst occurred during a nominal South Atlantic Anomaly (SAA) passage. As such, “*nicersaafilt*” was set to “no” on the second observation analysed, in order to include the data taken during the SAA passage. To test if the data were affected by the apparent passage through the SAA, we extracted the 13 – 15 keV lightcurve to look for the presence of high-energy background flares (see Bult et al. 2018); we found none. The total good exposure after processing was 43.6 ksec.

We determined the background contribution of our observations from *NICER* observations of a Rossi X-ray Timing Explorer (*RXTE*; Bradt et al. 1993; Jahoda et al. 2006) blank-sky region ($\sim 1\text{--}2\text{ cts s}^{-1}$ from *RXTE*-6).

Table 2. The black body temperatures, photon indices, 0.5–10 keV unabsorbed fluxes, and reduced χ^2 obtained from modelling the pre-burst emissions of MAXI J1807, with all errors quoted at a 90% confidence interval.

Burst	kT_{bb} (keV)	Photon Index	Unabsorbed Flux ($10^{-10}\text{ erg cm}^{-2}\text{ s}^{-1}$)	χ^2_{ν} (χ^2/dof)
1	0.10 ± 0.01	2.1 ± 0.2	0.54 ± 0.03	205.8/177
2	0.19 ± 0.01	2.0 ± 0.2	1.75 ± 0.05	263.2/233
3	0.11 ± 0.02	2.6 ± 0.3	1.14 ± 0.04	214.8/178

2.1 Light Curves and Hardness Ratios

To study the X-ray bursts in the context of MAXI J1807’s outburst, we constructed a 0.3–10 keV long-term light curve using 25 s bins. We extracted individual X-ray burst light curves in the 0.3–10 keV energy band using 0.1 s bins.

We define the start of each X-ray burst as the first 0.1 s bin which has an intensity 1σ above the persistent emission (calculated as the median count rate of the exposure) when scanning backwards in time from the X-ray burst peak. We define the end of an X-ray burst as the initial time of a 10 s bin that is consistent within 1σ of the median count rate of the last 100 s of the data segment containing the X-ray burst.

The hardness ratio was defined as the count rate in the 1–10 keV energy band divided by the count rate in the 0.3–1 keV energy band. Light curves for the hardness ratios were extracted with 0.2 s binning.

2.2 Spectral Analysis

2.2.1 Persistent Emission

Before examining the X-ray burst emissions, the pre-burst (persistent) spectra of MAXI J1807 were first explored. For this, we extracted the spectra of the persistent emission from exposures of 509 s, 395 s, and 110 s before the first, second, and third X-ray bursts, respectively. The 0.5–10 keV energy spectrum was then successfully fitted with an absorbed black body plus power-law model in *Xspec* v12.10.1:

$$\text{tbabs} \text{ (bbodyrad + powerlaw) },$$

where *tbabs* takes into account the effect of the interstellar absorption (Wilms et al. 2000). This simple model fits the data well. As the burst flux is generally much brighter than the persistent emission, our results remain unaffected if we use more complex models for the persistent emission. We found a column density of $N_H = (1.3 \pm 0.9) \times 10^{21}\text{ cm}^{-2}$ before the first X-ray burst and chose the same value to fit the persistent emission spectra before the second and third X-ray bursts. We note that our results were the same within error by (i) using the column density as explained above, (ii) using column densities as estimated from the persistent emission before each burst, or (iii) using an average of the three column densities.

The black body temperatures, photon indices, 0.5–10 keV unabsorbed fluxes, and reduced χ^2 obtained from modelling the pre-burst emissions are reported in Table 2. In all cases, we obtained a reduced χ^2 of ~ 1.1 . The errors are quoted for a 90% confidence interval.

2.2.2 Time-Resolved Spectroscopy

We performed time-resolved spectroscopy on all three X-ray bursts. Each X-ray burst was divided into time bins containing a minimum of 500 counts. We fitted energy spectra to each time bin using the variable persistent flux method (f_a -method; see, e.g., Worpel et al. 2013, 2015).

The f_a -method describes the X-ray burst thermal emission with a black body component and any possible excess is accounted for by scaling the pre-burst emission. This model is usually constructed by first fixing the pre-burst spectral components and then adding a black body component for the X-ray burst emission. A multiplicative factor f_a is then applied on the persistent part, which allows us to account for the effect of X-ray burst emission on the accretion processes. In `Xspec`, this model reads as:

$$\text{tbabs} \left(\text{bbodyrad} + f_a (\text{bbodyrad} + \text{powerlaw}) \right),$$

where `tbabs` takes into account the interstellar absorption and `bbodyrad` the black body emission from the X-ray burst. In our study, the model for the persistent emission (`bbodyrad + powerlaw`) is fixed to the respective values we found in Section 2.2.1.

The f_a -method has already been used for time-resolved spectroscopy of X-ray bursts using *NICER* data (see, e.g., Keek et al. 2018; Jaisawal et al. 2019). Although the f_a -method has been shown to be a good way to parameterise the effects on the accretion disc, it is likely that f_a should be a function of energy rather than a constant (see, e.g., Keek et al. 2014; Degenaar et al. 2018). In our analysis, given the relatively low count rates during MAXI J1807's X-ray bursts, we assumed an f_a that does not depend on energy.

3 RESULTS AND DISCUSSION

3.1 X-ray Imaging Observations

NICER is a non-imaging instrument. Here we investigate whether the X-ray bursts could originate from a different source than MAXI J1807.

MAXI J1807 was observed with *NuSTAR* (Harrison et al. 2013) on 26 September 2019, i.e., about 35 days before the first X-ray burst we detected with *NICER*. In Figure 1 we show the 3–78 keV *NuSTAR* image including the *NICER* field-of-view (a circle with 3 arcmin radius, corresponding to a 30 arcmin² field-of-view; Gendreau & Arzoumanian 2017). Only one source was significantly detected, with coordinates consistent with that of MAXI J1807.

There are also 28 photon counting (PC) mode *Swift*/XRT observations taken over the period 26 March–30 May 2017 (there are no *Swift*/XRT imaging observations during 2019). We created a 0.3–10 keV image of the *NICER* field-of-view by integrating the 28 images (total exposure time of 24 ks). The image was reduced using the *Swift*/XRT data product generator provided by the University of Leicester¹. Only one source was significantly detected, with coordinates consistent with that of MAXI J1807. Given the evi-

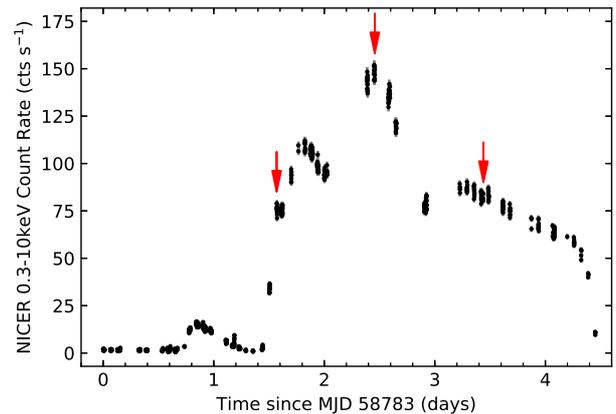


Figure 2. *NICER* long-term light curve of MAXI J1807 with 25 s time resolution in the 0.3–10 keV energy band. This section of the outburst contains the three thermonuclear X-ray bursts detected in October 2019. The X-ray burst data have been removed for clarity, and their onsets marked by arrows.

dence provided by *NuSTAR* and *Swift*/XRT, we concluded that the X-ray bursts originate from MAXI J1807.

3.2 Outburst Evolution and Occurrence of X-ray bursts

Figure 2 shows the outburst evolution of MAXI J1807 from 27 to 31 October 2019 (MJD 58783–58787). The occurrence of the X-ray bursts are marked by arrows. On 27 October ($t \approx 0.7$ days), the light curve shows a flare-like feature, where the count rate rises from ~ 2 cts s⁻¹ to ~ 15 cts s⁻¹ and falls back to the original count rate within half a day. Roughly 3 hrs later, we observe the onset of a larger flare with multiple peaks. The count rate increases to ~ 70 cts s⁻¹ over approximately 3 hrs, after which the first X-ray burst occurs on MJD 58784.61727 ($t \approx 1.6$ days in Figure 2). The flare reaches a first peak at ~ 110 cts s⁻¹ roughly 4.5 hrs after the X-ray burst, and drops to ~ 95 cts s⁻¹ in the following 4 hrs. There is an additional peak almost 10 hrs later during which the second X-ray burst occurs on MJD 58785.50558 ($t \approx 2.5$ days in Figure 2) at a persistent count rate of ~ 150 cts s⁻¹. The persistent count rate then drops to ~ 80 cts s⁻¹ over approximately 10 hrs before a final rise to ~ 90 cts s⁻¹, after which the flare begins to decay. At this stage in the outburst the third X-ray burst occurs on MJD 58786.48803 ($t \approx 3.5$ days in Figure 2) at a persistent count rate of ~ 80 cts s⁻¹. There is a sudden decrease in flux just less than a day after the third X-ray burst, where we observe the flux decreasing from ~ 60 cts s⁻¹ to ~ 10 cts s⁻¹ in roughly 4 hrs.

We observed the three X-ray bursts over a period of three days. The waiting time between the first and second X-ray bursts was 21.3 hrs, while between the second and third X-ray bursts was 23.6 hrs. We note that the data-gaps prevent us to understand whether there were other X-ray bursts than those we detected.

3.3 Type-I X-ray Burst Light Curves

Figure 3 shows light curves and hardness ratios of the three X-ray bursts exhibited by MAXI J1807. The first, second,

¹ https://www.swift.ac.uk/user_objects/index.php

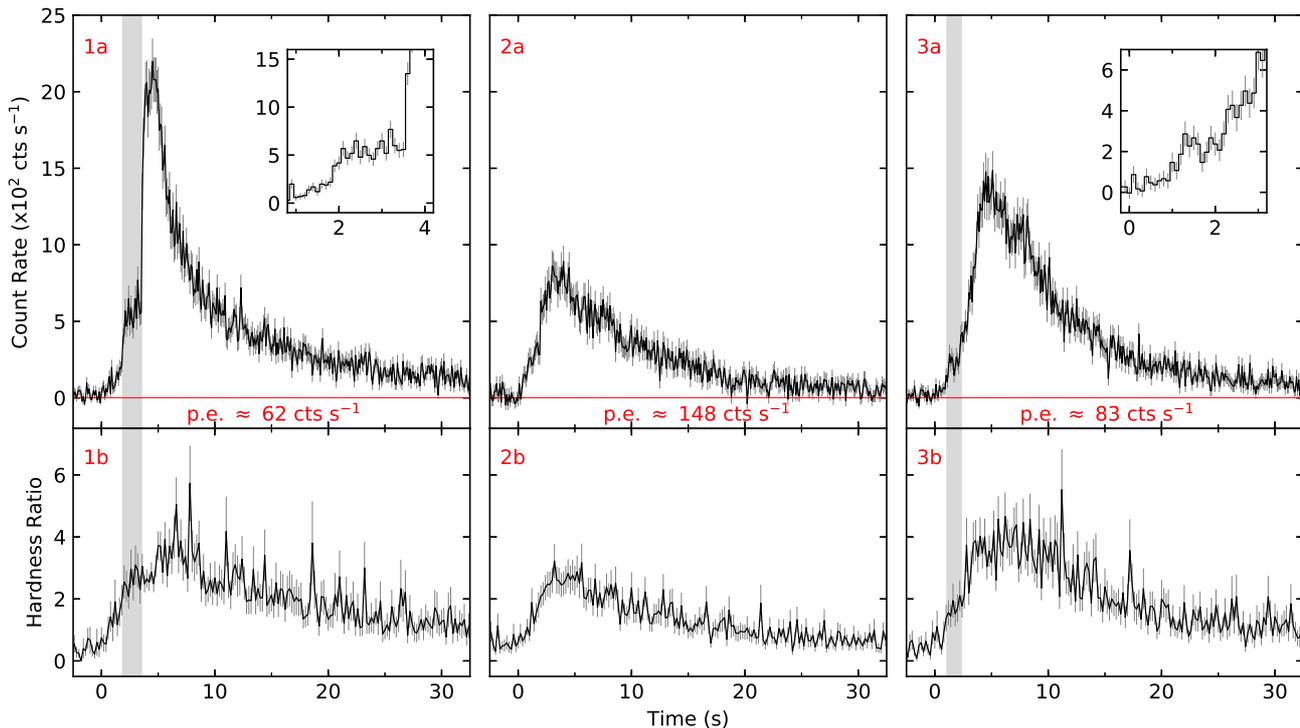


Figure 3. Top panels: light curves of the three X-ray bursts from MAXI J1807 with 0.1 s time resolution in the 0.3–10 keV energy band. We found the persistent emission levels by taking the mean count rate for 100 s of data before the start of each X-ray burst, and subtracted them from the light curves. The persistent emission count rates are denoted on the panels as “p.e.”. The insets in panels 1a and 3a show rise features in detail, which are also shaded in the main panels. Bottom panels: hardness ratios (1–10 keV / 0.3–1 keV count rate) with 0.2 s time resolution.

and third X-ray bursts will hereafter be referred to as B1, B2, and B3, respectively. At the burst onset, $t_{B1} = 0$, B1 count rate increases from the persistent rate of $\sim 62 \text{ cts s}^{-1}$ to $\sim 500 \text{ cts s}^{-1}$ over approximately 2 s. At this point, the burst exhibits a “pause” lasting ~ 1.6 s (shaded region and inset in Figure 3, panel 1a). Following the pause, the burst reaches its peak in roughly 0.2 s. The count rate remains approximately constant during this peak at $\sim 2250 \text{ cts s}^{-1}$ between $t_{B1} = 3.8 - 5.1$ s. After the peak, the count rate decreases for roughly 133 s to a persistent rate of $\sim 70 \text{ cts s}^{-1}$, i.e., at a flux $\sim 13\%$ higher than the persistent emission before the burst onset. The overall burst duration is ~ 137 s.

The rise of B2 is in two parts. In the initial rise, starting at $t_{B2} = 0$, the count rate increases from the persistent rate of $\sim 148 \text{ cts s}^{-1}$ to $\sim 460 \text{ cts s}^{-1}$ over approximately 2 s. The count rate then rises suddenly and reaches the peak of the burst in less than one second. This peak has an average count rate of $\sim 920 \text{ cts s}^{-1}$ between $t_{B2} = 2.6 - 4.6$ s. As the data segment containing the burst ends less than 100 s after the burst peak, the end of the burst is ill-defined while the count rate is still 8% above the pre-burst level (160 cts s^{-1} , the median count rate of the last 20 s of the data segment containing B2).

B3 exhibits interesting features in both its rise and decay. At $t_{B3} = 0$, B3’s count rate increases for roughly 1.5 s from the persistent rate of $\sim 83 \text{ cts s}^{-1}$ to a potential pause similar to that we observed in B1, but at $\sim 290 \text{ cts s}^{-1}$ (marked by the shaded region and inset in Figure 3, panel 3a). After the potential pause, the count rate continues to

rise reaching a peak average count rate of $\sim 1417 \text{ cts s}^{-1}$ between $t_{B3} = 4.0 - 5.3$ s. During the decay there is indication of a double peak at $t \approx 7.5$ s lasting roughly 2 s; however the double peak is not statistically significant given the large error bars. After this, the count rate decreases over approximately 78 s to a post-burst persistent count rate of $\sim 100 \text{ cts s}^{-1}$, which is $\sim 20\%$ higher than the pre-burst persistent rate. The overall burst duration is ~ 83 s.

The hardness ratios of all three X-ray bursts track similar profiles to the light curves, increasing through the burst rise and decreasing through the decay. The hard band (1–10 keV) dominates after less than 1 s in each case, and increases to peaks of ~ 4 in B1 and B3, and ~ 3 in B2. There is a plateau in the hardness ratio that starts during the pause in B1 and appears to continue after it. There is also evidence of a plateau during the potential pause found in B3, after which the ratio increases; however in this case the plateau is not as well constrained as in B1.

We measure different average count rates for the persistent emission before and after each X-ray burst. Given our data set, we are unable to understand if this is due to an intrinsic change in the flux of the persistent emission, or whether it is the effect of long decay tails which we cannot differentiate from the continuum (see, e.g., [in’t Zand et al. 2017](#)). Out of the three X-ray bursts, B1 has the highest peak count rate, whilst B2 has the lowest. All three X-ray bursts have a rise time of ~ 4 s and exhibit long decay tails (>1 minute). A slow rise and long decay is indicative of H-rich fuel at the moment of ignition, which is likely the re-

sult of accretion of a mixed H/He fuel (Galloway et al. 2008; Galloway & Keek 2017; Schatz et al. 2001).

3.4 X-ray Burst Spectral Characteristics

In Figure 4 we show the best-fit parameters for the time-resolved spectroscopy of the X-ray bursts. The bolometric unabsorbed flux (F_{bol}) of all three X-ray bursts approximately follows the light curve contours. The peak fluxes are reported in Table 3. The black body temperature (kT_{bb}) also follows the contours of the burst profiles in the light curves, increasing during the burst rise and decreasing during the decay. The evolution of kT_{bb} is typical of the heating and cooling processes of the NS atmosphere during a Type-I X-ray burst (see, e.g., Lewin et al. 1993, for a review), thus strongly suggesting that these features are indeed thermonuclear in origin. The peak temperatures achieved by each of MAXI J1807's X-ray bursts (~ 1 keV, reported in Table 3) are lower than what has typically been seen in previous works for other sources ($\gtrsim 2$ keV). We did not find any evidence that the low temperatures are due to any instrumental effects or due to our spectral modelling. Indeed, previous results on other sources using *NICER* data and the f_a -method have shown that X-ray bursts can reach temperatures in the 2–3 keV range (see, e.g., Keek et al. 2018). Therefore, we conclude that the low temperatures we measure are intrinsic to the X-ray bursts detected in this source.

Thermonuclear X-ray bursts can be used to estimate the distance to the system if they exhibit photospheric radius expansion. The brightest of the three X-ray bursts, B1, does not show strong evidence, if any, of PRE. However, under the assumption that B1 reached the Eddington luminosity (L_{Edd}), MAXI J1807 contains a $1.4 M_{\odot}$ NS and $L_{\text{Edd}} = 2 \times 10^{38} \text{ erg s}^{-1}$ (as expected for H-rich fuel, see, e.g., Lewin et al. 1993), we found a distance upper limit of 12.4 kpc.

We see differences in the evolution of the black body radius, calculated assuming a distance of 12.4 kpc, during the three X-ray bursts. For B1, it remains approximately constant at ~ 28.5 km. The radius during B2 approximately follows the contour of the light curve, peaking at ~ 28.4 km. In the case of B3, the radius increases throughout the rise in count rate, and then remains roughly constant at ~ 28.5 km after the burst peak. The radii reported here are larger than what might be expected (see, e.g., Galloway et al. 2008). This is due to the fact that they are calculated using the 12.4 kpc upper limit for the distance. If MAXI J1807 were at, for example, half this distance, the peak radius would be roughly 18.5 km, which, whilst it is still large, is more consistent with the expected radius during X-ray bursts.

The parameter f_a describes the photospheric evolution of the X-ray burst as well as its effect on the accretion processes. For B2, f_a remains at ~ 1.9 throughout the burst. For B1 and B3, f_a increases and peaks at the same time as the burst, and decreases afterwards. This can be the result of real changes in the accretion disc or due to the choice of model. Testing whether the accretion disc varies significantly is difficult to assess (e.g., Degenaar et al. 2018), particularly with the low count rates in our data.

In Table 3, we report the X-ray burst fluences and time scales τ (= fluence / peak F_{bol}). We calculated the fluences of B1, B2, and B3 over the first 31 s, 25 s, and 42 s of data,

Table 3. The best fit parameters obtained from time-resolved spectroscopy for peak black body temperatures and bolometric fluxes achieved by each X-ray burst from MAXI J1807, and calculated estimates for X-ray burst parameters fluence and τ .

Burst	Peak kT_{bb} (keV)	Peak F_{bol} ($10^{-8} \text{ erg s}^{-1} \text{ cm}^{-2}$)	Fluence ($10^{-8} \text{ erg cm}^{-2}$)	τ (s)
1	$1.21^{+0.14}_{-0.12}$	$1.01^{+0.17}_{-0.15}$	2.30 ± 0.29	2.29 ± 0.07
2	0.9 ± 0.1	0.26 ± 0.03	1.36 ± 0.11	5.30 ± 0.13
3	0.9 ± 0.07	$0.56^{+0.06}_{-0.05}$	3.17 ± 0.27	5.70 ± 0.06

respectively (during these time-intervals the fluence was well constrained). The fluences are in the (~ 2 to ~ 3.5) $\times 10^{-8} \text{ erg cm}^{-2}$ range, which is the lower end of typical fluence distributions found in other sources (Galloway et al. 2008; Lyu et al. 2015; Boirin et al. 2007). However, given the uncertainties in our calculations of X-ray burst duration and the intensity of the continuum after the burst, our fluences are probably underestimated by 10–15%. We found that τ was in the ~ 2 to ~ 6 s range. These values are also in the lower end of the distribution usually found in other sources (Galloway et al. (2008, 2020)).

Calculating an accurate value for the X-ray burst parameter α (ratio of the integrated persistent flux to the burst fluence, see, e.g., Galloway et al. 2008) is not possible without a well-constrained recurrence time. However, as an exercise, we calculate a rough estimate for α with certain caveats. We only detected X-ray bursts from MAXI J1807's 2019 outburst during a ~ 1.9 day window. During this window, *NICER* covered $\sim 8\%$ of the time. If we assume that only $\sim 8\%$ of the X-ray bursts were detected in this time frame, this implies a total of ~ 36 X-ray bursts with a recurrence time of ~ 1.27 hrs. Taking a pre-burst flux of $\sim 1 \times 10^{-10} \text{ erg s}^{-1} \text{ cm}^{-2}$ and fluence of $\sim 2 \times 10^{-8} \text{ erg cm}^{-2}$, this gives an averaged $\alpha \approx 23$. These assumptions make the statistical and systematic uncertainties on this calculation rather large. Comparing this value with what we see in Galloway et al. (2008), our estimate for α is much smaller than what we would expect for our calculated values of τ (if compared with other sources, we expect $\alpha > 60$ for $\tau < 10$; see Figure 14, Galloway et al. 2008). The uncertainty in our calculations may be able to account for this tension, although it may also be an intrinsic property of the X-ray bursts in MAXI J1807 as the fluences we measured are smaller than average.

3.5 Search for Burst Oscillations

To search for burst oscillations, we constructed 0.8–8.0 keV X-ray burst light curves at 1/8192 s time resolution. Each light curve started 10 s prior to the burst onset, and had a duration of 50 s, such that the entire burst epoch was included. We then applied a $T = 2$ s, 4 s, and 8 s wide window to each light curve, which we moved across the respective burst in steps of $T/2$. For each combination of window size and position we constructed the power spectrum and searched between 100 Hz and 1000 Hz for the presence of a coherent burst oscillation signal in excess of the counting noise distribution (see Watts et al. 2005; and Watts 2012 for a review). No significant signals were observed to a 95%

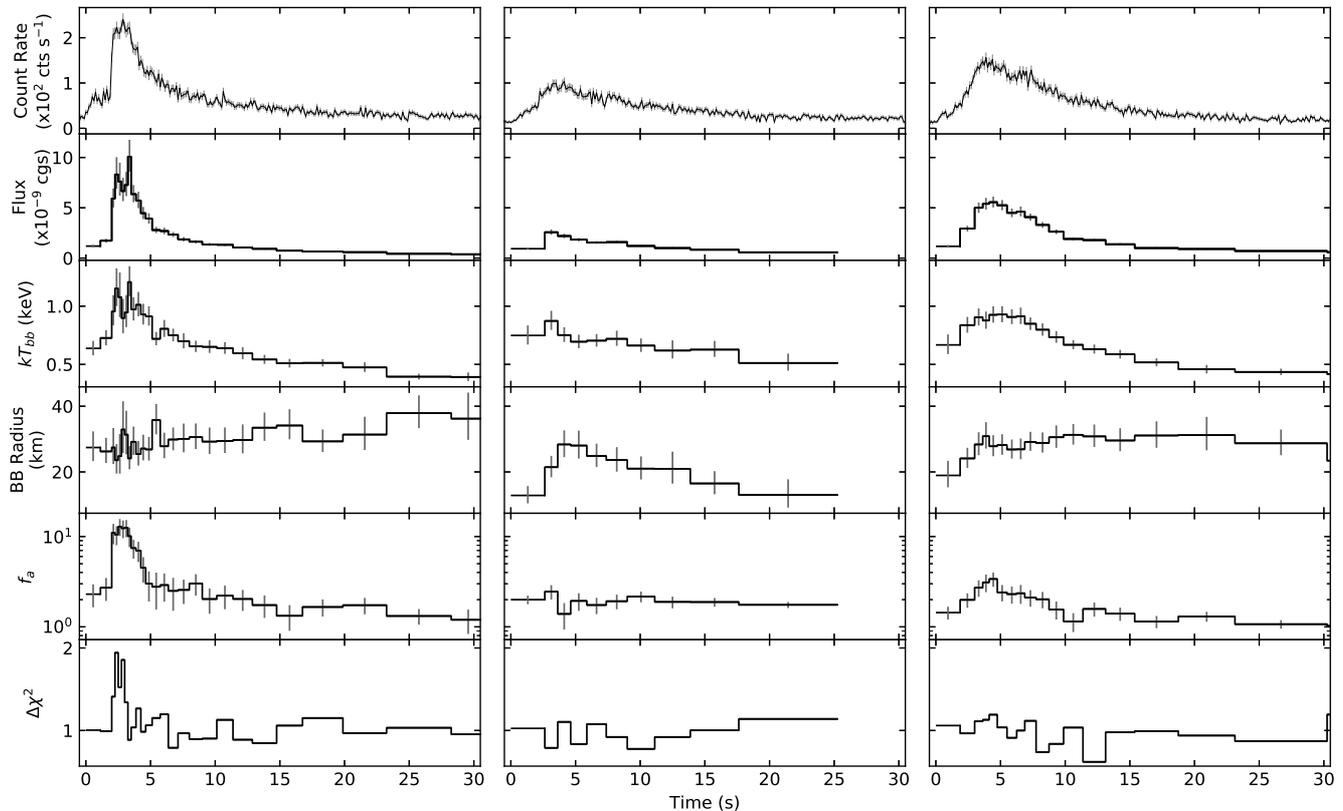


Figure 4. Time-resolved spectroscopy of the three X-ray bursts (B1, B2 and B3, left to right, respectively) using the variable f_a method. From top, we show the count rate, the estimated bolometric flux (in units of $\text{erg s}^{-1} \text{cm}^{-2}$), the black body temperature and radius (assuming a distance of 12.4 kpc), the variable scale factor f_a and the reduced χ^2 .

confidence upper limit of approximately 10% fractional amplitude at peak intensity.

3.6 Pauses and Comparison

Historically, the rises of thermonuclear Type-I X-ray bursts have generally been described as relatively fast and smooth. In the last two decades, thanks to the high time resolution and effective area of *RXTE*, we have learned that X-ray burst profiles can be very complex. These include slower or faster rises, two-step rises (like we see in panel 2a of Fig 3), concave and convex rises (Maurer & Watts 2008), and X-ray bursts with multiple (2 or 3) peaks. We note that, sometimes an X-ray burst with a double-peaked profile can be interpreted as two different events (e.g., Bhattacharyya & Strohmayer 2007).

X-ray bursts with multiple peaks in their bolometric luminosity profile are not uncommon (see, e.g., Galloway et al. 2008; Lewin et al. 1993; Watts & Maurer 2007, and references therein). Notably, the source 4U 1636–536 shows the largest sample of multi-peaked X-ray bursts with a variety of profiles, including X-ray bursts with two peaks that reach the same luminosity (e.g., Bhattacharyya & Strohmayer 2006a), X-ray bursts where the two peaks are not at the same luminosity (e.g., Bhattacharyya & Strohmayer 2006b; Maurer & Watts 2008), and even triple-peaked X-ray bursts (van Paradijs et al. 1986; Zhang et al. 2009). Double-peaked

X-ray bursts from this source have been interpreted in a number of different ways. For example, a second peak may be caused by multi-step thermonuclear energy release, or the ignition of fresh or leftover material (see, e.g., Jaisawal et al. 2019, and references therein). Bhattacharyya & Strohmayer (2006a,b) interpreted the double-peaked profile as the result of a temporary stalling in the burning front at the NS’s equator. This interpretation requires rare high-latitude ignition (Spitkovsky et al. 2002), explaining why double-peaked bursts are far less common than single-peaked bursts (e.g., Galloway et al. 2008, but also see discussions by Cooper & Narayan (2007); Watts & Maurer (2007); Maurer & Watts (2008)). However, triple-peaked bursts are difficult to explain solely by a temporary stalling in the burning front. Zhang et al. (2009) showed that observations of triple peaks pose particular challenges for the polar ignition mechanism required by the interpretation of Bhattacharyya & Strohmayer (2006a,b).

In this paper, we report on the discovery of a < 2 sec pause in the rise of one single-peaked burst in the LMXB MAXI J1807+132. Given the complex profiles that X-ray bursts may have (see above), it is important to understand whether pauses during an X-ray burst rise are a different manifestation of the multi-peak phenomena (i.e. what we observe as a pause is in reality a very short, weak peak of a double-peaked X-ray burst), or the pause is unrelated to the physical process that produces multi-peak X-ray bursts.

Bult et al. (2019) recently reported on a bright, He-fueled PRE Type-I X-ray burst from the accreting millisecond X-ray pulsar SAX J1808.4–3658. The light curve profile of the X-ray burst (hereafter S1) was complex: it showed a pause in the rise and a double peak unrelated to the PRE event. Additionally, S1 showed burst oscillations. Bult et al. (2019) found that the pause, the dip between the X-ray burst’s double peaks, and the onset of burst oscillations all occur at similar count rates, suggesting that the similar count rates are related to the Eddington limit of a hydrogen envelope.

The pause in the first burst we detected in MAXI J1807 (B1) is ~ 1.6 s in duration, and occurs at roughly 20% of the peak count rate. The pause in SAX J1808.4–3658’s S1 is shorter, lasting ~ 0.7 s, and occurs at roughly 40% of the peak count rate. We calculated the ratio between peak and pause fluxes to be $5.8^{+0.3}_{-0.2}$, whereas this ratio was calculated to be 1.68 ± 0.13 for S1. In both B1 and S1, the rise time before the pause is roughly the same duration as the pause, lasting ~ 1.6 s and ~ 0.6 s for B1 and S1 respectively. However, the rise time from the end of the pause to the burst peak in B1 is ~ 0.5 s, much shorter than that of ~ 3 s in S1. B1 and S1 also have different fuel compositions, with mixed H/He (H-rich) and He only, respectively.

Although we currently have a sample of only two X-ray bursts with confirmed pauses during the rise, we suggest that the pauses are not a shorter and weaker manifestation of the multi-peaked X-ray bursts seen in other sources. Our interpretation is supported by the fact that SAX J1808.4–3658’s X-ray burst (Bult et al. 2019) exhibits both a pause and a double-peaked profile, showing that the two phenomena can manifest in the same X-ray burst. In addition, the X-ray burst we report for MAXI J1807 shows a pause and a single-peaked profile. These two results together, and the lack of other reports on pauses in X-ray bursts studied with *RXTE*, suggest that there is no link between the detection of double-peaked profiles and the detection of a pause during the rise. However, our interpretation would be incorrect if the pauses *are* short and weak manifestations of a peak in the X-ray burst profile, meaning that the X-ray burst in SAX J1808.4–3658 can be considered a triple-peaked X-ray burst, and B1 in MAXI J1807 can be considered a double-peaked burst.

in’t Zand et al. (2003) reported on an unusual Type-I X-ray burst from GRS 1747-312² observed with *RXTE* whilst the system was apparently in quiescence. The 2–10 keV light curve profile exhibits a pause-like feature in the rise of the X-ray burst, while the full 2–60 keV light curve did not show a pause. The pause in the 2–10 keV light curve is considerably longer than those seen in S1 and B1 (~ 20 s), and, judging from Figure 2a of in’t Zand et al. (2003), the pause occurs at roughly 70% of the peak count rate. The X-ray burst also had an unusually strong PRE that was *not* at the start of the X-ray burst, as is expected. It is unclear whether this is a longer manifestation of the pause phenom-

ena we see in B1 and S1, but it shows that the presence of pause-like features can depend on the energy band used.

Independently of whether the pauses are related to the multi-peaked phenomena discussed above, our results, together with those presented by Bult et al. (2019), show that short pauses in the rise of X-ray burst light curves occur in at least two sources, suggesting that these types of pauses might be present in the X-ray light curves of Type-I bursts from other NS-LMXBs. We only know of two sources (the accreting millisecond X-ray pulsar SAX J1808.4–3658 and MAXI J1807), each with one reported X-ray burst with a clear pause during the rise. This small sample size prevents us from a detailed study on the relation between pauses and X-ray burst characteristics/properties, and naturally it is unclear whether the pauses seen in B1 and S1 are caused by the same mechanism. However, under the assumption that we are seeing manifestations of the same phenomena in B1 and S1, even with a sample of two we can conclude that pauses in the X-ray bursts are possible in both pulsating and non-pulsating systems, suggesting that the strength of the magnetic field might not play an important role, unless the lack of pulsations in MAXI J1807 is due to an alignment between rotation and magnetic axes (Lamb et al. 2009a,b). In addition, the occurrence of the pause does not appear to depend on the composition of the burning material either, or whether the X-ray burst reaches PRE.

To our knowledge, standard X-ray burst models do not predict pauses in the X-ray burst rise. However, we note that Keek et al. (2012) finds X-ray bursts with pauses in their rise when modelling X-ray bursts that occur shortly after a superburst. In the hot, He-rich atmospheres following a superburst, X-ray bursts are initially powered by the triple- α process. As the atmospheric temperature begins to rise due to these X-ray bursts, α -captures start playing a role powering the peaks of the X-ray bursts. It is during this transition that X-ray bursts with pauses in the rise occur, as it takes some time to ignite the second burning stage through α -captures. At the moment, it is unclear whether the model of Keek et al. (2012) could be adapted to explain our results, i.e. whether an X-ray burst can show pauses independently of the occurrence of a superburst to create a hot He-rich atmosphere.

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² GRS 1747-312 is a transient X-ray source located in the core of the 9.5 kpc distant globular cluster Terzan 6. Given *RXTE*/Proportional Counter Array’s 1° field-of-view, in’t Zand et al. (2003) could not discard the possibility that the X-ray burst originated in a different NS-LMXB

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DATA AVAILABILITY

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