

# On the Ca-strong 1991bg-like type Ia supernova 2016hnk: evidence for a Chandrasekhar-mass explosion

Lluís Galbany<sup>1,2,\*</sup>, Chris Ashall<sup>3</sup>, Peter Höflich<sup>3</sup>, Santiago González-Gaitán<sup>4</sup>, Stefan Taubenberger<sup>5,6</sup>, Maximilian Stritzinger<sup>7</sup>, Eric Y. Hsiao<sup>3</sup>, Paolo Mazzali<sup>8,6</sup>, Eddie Baron<sup>9</sup>, Stéphane Blondin<sup>10,11</sup>, Subhash Bose<sup>12</sup>, Mattia Bulla<sup>13</sup>, Jamison F. Burke<sup>14,15</sup>, Christopher R. Burns<sup>16</sup>, Régis Cartier<sup>17</sup>, Ping Chen<sup>12,18</sup>, Massimo Della Valle<sup>19,5</sup>, Tiara R. Diamond<sup>20</sup>, Claudia P. Gutiérrez<sup>21</sup>, Jussi Harmanen<sup>22</sup>, Daichi Hiramatsu<sup>14,15</sup>, T. W.-S. Holoien<sup>16</sup>, Griffin Hosseinzadeh<sup>23</sup>, D. Andrew Howell<sup>14,15</sup>, Yiwen Huang<sup>24</sup>, Cosimo Inserra<sup>25</sup>, Thomas de Jaeger<sup>26</sup>, Saurabh W. Jha<sup>27</sup>, Tuomas Kangas<sup>28</sup>, Markus Kromer<sup>29,30</sup>, Joseph D. Lyman<sup>31</sup>, Kate Maguire<sup>32,33</sup>, George Howie Marion<sup>34</sup>, Dan Milisavljevic<sup>35</sup>, Simon J. Prentice<sup>32</sup>, Alessandro Razza<sup>36,37</sup>, Thomas M Reynolds<sup>21</sup>, David J. Sand<sup>38</sup>, Benjamin J. Shappee<sup>39</sup>, Rohit Shekhar<sup>40,13</sup>, Stephen J. Smartt<sup>32</sup>, Keivan G. Stassun<sup>41</sup>, Mark Sullivan<sup>21</sup>, Stefano Valenti<sup>42</sup>, Steven Villanueva<sup>43,44</sup>, Xiaofeng Wang<sup>45</sup>, J. Craig Wheeler<sup>46</sup>, Qian Zhai<sup>47,48</sup>, Jujia Zhang<sup>47,48</sup>

(Affiliations can be found after the references)

Received April 5, 2019; accepted —

## ABSTRACT

**Aims.** We present a comprehensive dataset of optical and near-infrared photometry and spectroscopy of type Ia supernova (SN) 2016hnk, combined with integral field spectroscopy (IFS) of its host galaxy, MCG -01-06-070, and nearby environment. Our goal with this complete dataset is to understand the nature of this peculiar object.

**Methods.** Properties of the SN local environment are characterized by means of single stellar population synthesis applied to IFS observations taken two years after the SN exploded. We perform detailed analyses of SN photometric data by studying its peculiar light and color curves. SN 2016hnk spectra are compared to other 1991bg-like SNe Ia, 2002es-like SNe Ia, and Ca-rich transients. In addition, abundance stratification modelling is used to identify the various spectral features in the early phase spectral sequence and the dataset is also compared to a modified non-LTE model previously produced for the subluminous SN 1999by.

**Results.** SN 2016hnk is consistent with being a sub-luminous ( $M_B = -16.7$  mag,  $s_{BV} = 0.43 \pm 0.03$ ), highly reddened object. IFS of its host galaxy reveals both a significant amount of dust at the SN location, as well as residual star formation and a high proportion of old stellar populations in the local environment compared to other locations in the galaxy, which favours an old progenitor for SN 2016hnk. Inspection of a nebular spectrum obtained one year after maximum contains two narrow emission lines attributed to the forbidden [Ca II]  $\lambda\lambda 7291, 7324$  doublet with a Doppler shift of  $700 \text{ km s}^{-1}$ . Based on various observational diagnostics, we argue that the progenitor of SN 2016hnk was likely a near Chandrasekhar-mass ( $M_{Ch}$ ) carbon-oxygen white dwarf that produced  $0.108 M_{\odot}$  of  $^{56}\text{Ni}$ . Our modeling suggests that the narrow [Ca II] features observed in the nebular spectrum are associated with  $^{48}\text{Ca}$  from electron capture during the explosion, which is expected to occur only in white dwarfs that explode near or at the  $M_{Ch}$  limit.

**Key words.** supernovae: general – supernovae: individual: SN 2016hnk

## 1. Introduction

Type Ia supernovae (SNe Ia) are among the brightest transient objects in the Universe, typically reaching peak absolute magnitudes of  $\lesssim -19$  mag. They are also very homogeneous, with peak brightness varying  $\sim 1-2$  mag (e.g., Hamuy et al. 1996). However, through empirical calibrations, such as the relations between absolute peak magnitude and both the post-maximum brightness decline (e.g.  $\Delta m_{15}(B)$ , the parameter accounting for the magnitude decay after 15 days post-max in the  $B$  band; Phillips 1993) and the peak color (Riess et al. 1996; Tripp 1998), this scatter around peak brightness is reduced to  $\sim 0.1$  mag. Due to their standardizability, and because they can be discovered from the ground up to high redshifts ( $z \lesssim 1.0$ ), SNe Ia have been widely used in the last decades to put constraints on cosmological parameters and the Hubble constant (e.g. Howell 2011; Burns

et al. 2018), they were the key to the discovery of the accelerated expansion of the Universe (Riess et al. 1998; Perlmutter et al. 1999) and remain one of the key probes of dark energy (e.g. Betoule et al. 2014; DES Collaboration et al. 2018).

However, their success as cosmological probes is not matched by our understanding of their progenitor systems, the details on how they explode, and the underlying physical interpretation of the known empirical relations that enable them to provide precise cosmological distances. Further improvement requires a much better understanding of the physical model of the explosion, observational constraints on which kind of progenitors can produce SNe Ia, a better control of reddening, and a complete characterization of peculiar (non-standardizable) types.

The fact that SNe Ia are discovered in all types of galaxies, including passive galaxies with no signs of star formation, points to old stars being the most probable progenitors ( $>100$ s Myr;

\* E-mail: lgalbany@pitt.edu

Mannucci et al. 2006). At the same time, their uniform brightness suggests that their progenitor systems form a homogeneous family. The consensus picture is that a SN Ia results from a degenerate (held by electron degeneracy pressure) carbon-oxygen (C-O) white dwarf (WD) undergoing a thermonuclear runaway (Hoyle & Fowler 1960), and that they originate from close binary systems. However, the nature of the system and the explosion mechanism is still under debate. Potential progenitor systems may either consist of two WDs, called a double degenerate (DD) system, or a single WD with a non-degenerate companion, called a single degenerate (SD) system. Candidate progenitor systems have been inferred for both SD and DD systems. Although the lack of donor stars surviving the explosion may favor DD systems (Tucker et al. 2019), other examples would favor SD progenitors, such as the potential He donor star identified for the well-observed 2002cx-like type Ia SN 2012Z (McCully et al. 2014a; Stritzinger et al. 2015) and the massive asymptotic-giant-branch star inferred for SN 2002ic (Hamuy et al. 2003).

Within these progenitor scenarios four classes of explosion scenarios are most favored: (1) the dynamical merging of two C-O WDs in a binary system after losing angular momentum via gravitational radiation, where the thermonuclear explosion is triggered by the heat of the merging process (García-Berro & Lorén-Aguilar 2017); (2) the explosion of a C-O WD with a mass close to  $M_{\text{Ch}}$  triggered by compressional heating near the WD center, caused by accretion from the donor star, which may be a red giant, a main sequence star of mass less than  $7-8M_{\odot}$ , a He star, or the tidally disrupted WD from a DD system (Piersanti et al. 2004); (3) the explosion of a sub- $M_{\text{Ch}}$  C-O WD triggered by detonating a thin surface He layer on the WD (double detonations; Kromer et al. 2010; Shen & Moore 2014); and (4) direct collision of two WDs in tertiary systems, where the third star induces oscillations in the eccentricity of the other two increasing gravitational-wave losses and finishing in a head-on collision (e.g. Thompson 2011; Mazzali et al. 2018). The debate about dynamical mergers,  $M_{\text{Ch}}$  explosions, and double-detonations is ongoing, but all channels may contribute to the SNe Ia population.

The connection of these progenitor scenarios and explosion mechanisms with SN Ia subtypes is still not well-established. More extreme sub- or over-luminous SNe Ia may require progenitors deviating from those discussed above. Clarifying which progenitor configuration and explosion mechanism lead to SNe Ia with particular properties, and with what rates, is of utmost importance for a number of fields including stellar evolution and SN cosmology. The presence of these peculiar over- and sub-luminous SNe Ia, and other heterogeneous objects that do not follow the empirical relations (Taubenberger 2017), may challenge the picture of all SNe Ia coming from the same family of progenitors.

In particular, subluminescent SN 1991bg-like objects seem to break such relations. Photometrically, they are characterized by being fainter at peak and by having faster light curve rise and decline rates, putting them at the faint and fast extreme of the light curve width-luminosity relation ( $\Delta m_{15} \gtrsim 1.7$  mag). The maxima in the near-infrared (NIR) bands occur after the maximum in optical bands, but one of the main peculiarities is the lack of either a shoulder or a secondary maximum in the NIR light curves. The latter can be explained by the recombination of Fe III to Fe II occurring earlier: the Fe II line blanketing absorbs flux in the blue that is re-emitted at longer wavelengths, causing the first and second NIR peaks to merge to form a single, slightly delayed maximum (Kasen & Woosley 2007; Höflich et al. 2017). As a result, their colors are redder at maximum ( $(B - V)_{B_{\text{max}}} \gtrsim 0.5$  mag) and their  $B - V$  color curves peak earlier ( $\sim 15$  days from peak

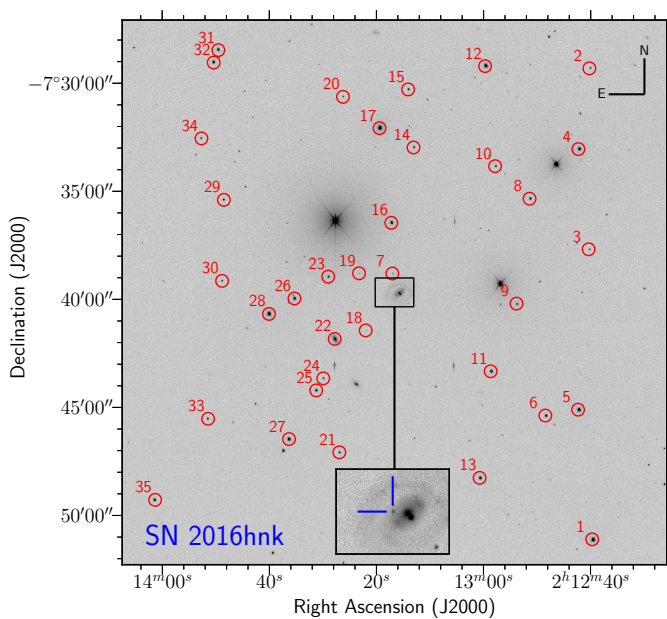
**Table 1.** Properties of SN 2016hnk and its host galaxy.

Parameters	Values	Refs
<b>SN 2016hnk:</b>		
RA (J2000)	$02^{\text{h}}13^{\text{m}}16^{\text{s}}.63$	(1)
	$33.319312^{\circ}$	(1)
DEC (J2000)	$-07^{\circ}39'40''.80$	(1)
	$-7.661332^{\circ}$	(1)
Date of explosion (JD)	$2,457,676.54 \pm 4.49$	This work
Date of maximum (JD)	$2,457,689.98 \pm 3.27$	This work
Projected offset	11.55 arcsec	This work
	3.71 kpc	This work
Deprojected offset	14.23 arcsec	This work
	4.57 kpc	This work
Milky Way extinction $E(B - V)_{\text{MW}}$	$0.0224 \pm 0.0008$ mag	(4)
Host galaxy extinction $E(B - V)_{\text{HG}}$	$0.45 \pm 0.06$ mag	This work
$\Delta m_{15}(B)_{\text{obs}}$	$1.297 \pm 0.071$ mag	This work
$\Delta m_{15}(B)_{\text{true}}$	$1.324 \pm 0.096$ mag	This work
$\Delta m_{15,s}(B)$	$1.803 \pm 0.200$ mag	This work
$s_{BV}$	$0.438 \pm 0.030$	This work
$(m, M)_{B_{\text{max}}}^{\text{corrected}}$	$(17.539, -16.656)$	This work
<b>Host galaxy:</b>		
Name	MCG -01-06-070	
Type	(R')SB(rs)a	(3)
RA (J2000)	$02^{\text{h}}13^{\text{m}}15^{\text{s}}.79$	(3)
	$33.315792^{\circ}$	(3)
DEC (J2000)	$-07^{\circ}39'42''.70$	(3)
	$-7.661861^{\circ}$	(3)
Recession velocity, $v$	$4780 \pm 45$ km s $^{-1}$	This work
Helio-centric redshift, $z_{\text{helio}}$	$0.01610 \pm 0.00015$	This work
Projection angle (N $\rightarrow$ E)	$308.84 \pm 4.09$ deg	This work
Inclination	$60.17 \pm 2.84$ deg	This work
Luminosity distance	$68.35 \pm 0.04$ Mpc	This work
$m_B$	$14.89 \pm 0.02$ mag	(5,6)
Stellar mass	$10.78^{+0.12}_{-0.21}$ dex	This work

Notes: (1) ATLAS survey; (2) Schlafly & Finkbeiner (2011); (3) NASA/IPAC Extragalactic Database (<http://ned.ipac.caltech.edu/>); (4) NASA/IPAC Infrared science archive (<https://irsa.ipac.caltech.edu/>); (5) SDSS DR14 (Abolfathi et al. 2018); (6) Jester et al. (2005).

maximum in  $B$ -band) compared to normal SNe Ia ( $\sim 30$  days). Spectroscopically, their main characteristics are the presence of prominent Ti II and strong O I  $\lambda 7744$  lines, indicating low ionization and temperature, deeper Si II  $\lambda 5972$  than in normal SNe Ia, and relatively low ejecta velocities with fast decline and no high-velocity features. To first order, the luminosity of a supernova is dependent on the mass of radioactive  $^{56}\text{Ni}$  synthesized in the explosion, where less luminous objects produce smaller amounts of  $^{56}\text{Ni}$  (Arnett 1982; Pinto & Eastman 2000; Mazzali et al. 2001), though there is some evidence that a variation of the total ejecta mass may also be required (Stritzinger et al. 2006; Scalzo et al. 2014; Blondin et al. 2017). Given their low luminosity, 1991bg-like SNe synthesize rather low amounts ( $\leq 0.1 M_{\odot}$ ) of radioactive  $^{56}\text{Ni}$  (Stritzinger et al. 2006). In addition, most 1991bg-like SNe Ia are found in elliptical galaxies with rather old stellar populations and no ongoing or recent star formation, although the well-observed SN 1999by occurred in the spiral Sa-type galaxy NGC 2841 (Höflich et al. 2002; Garnavich et al. 2004).

In this paper, we present a comprehensive dataset collected for the peculiar SN 2016hnk (see also Sell et al. 2017), which exploded in the Sb-type galaxy MCG -01-06-070. SN 2016hnk showed many similarities with 1991bg-like SNe Ia, including the lack of hydrogen features, the typical deep Si II  $\lambda 6355$  absorption, other typical SN Ia features with absorption velocities within  $6,000 < v < 10,000$  km s $^{-1}$ , and very strong calcium lines in the optical. With an ATLAS-c absolute magnitude at discovery of  $-16.1$ , it was also reported to be similar to the unusual Ca-rich transient PTF09dav (Pan et al. 2016). Our goal with this complete dataset is to understand the nature of this peculiar object. The paper is structured as follows. In Section 2 we describe the discovery and the photometric and spectroscopic follow up campaign. In Section 3, integral field spectroscopy of the SN host galaxy from the PISCO survey (Galbany et al. 2018) is used



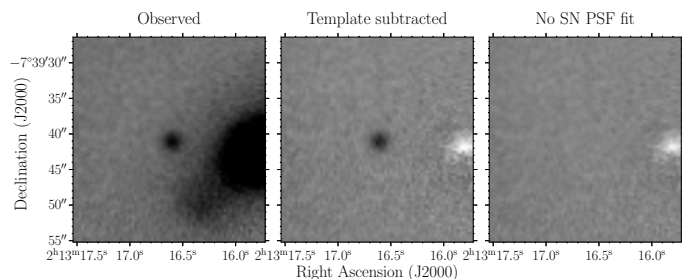
**Fig. 1.** Greyscale *B*-band image obtained with the Las Cumbres 1m telescope at Cerro Tololo Inter-American Observatory on November 3rd 2016. Orientation is North-up, East-left. SN 2016hnk is marked with two blue lines at the bottom inner panel, next to its host galaxy MCG -01-06-070. Red circles represent all 35 standard stars used for photometric calibration in the optical and near-infrared.

to study the SN local environment. Later we present the observational analysis in two sections: the photometric view in Section 4 and the spectroscopic view in Section 5, including an observational characterization within the 1991bg-like group. We make a comparison with theoretical models from the abundance stratification technique and detailed non-LTE modified models of SN 1999by in Section 6. Finally, we discuss the origin of this object in Section 7. Throughout the paper we employed a standard flat  $\Lambda$ CDM cosmology with  $\Omega_M=0.2865$  and  $H_0=69.32 \text{ km s}^{-1} \text{ Mpc}^{-1}$  (Hinshaw et al. 2013).

## 2. SN 2016hnk Observations and reduction

The discovery of SN 2016hnk (aka ATLAS16dpc) was reported by the Asteroid Terrestrial-impact Last Alert System (ATLAS; Tonry et al. 2018) using the ACAM1 instrument in the ATLAS Haleakala telescope on 2016 October 27.47 (JD=2457688.97) at  $\alpha = 02^{\text{h}}13^{\text{m}}16^{\text{s}}.63$ ,  $\delta = -07^{\circ}39'40''.80$  (J2000; see Fig. 1) and with a magnitude of 17.91 in the cyan-ATLAS filter. The object exploded in the type SB galaxy MCG -01-06-070 of the Morphological Catalogue of Galaxies (Vorontsov-Vel' Yaminov & Arhipova 1962), 11.5 arcsec East of the center, which at the redshift of the galaxy  $z=0.016$  corresponds to 3.71 kpc (see inner panel of Fig. 1). The main parameters of the supernova and host galaxy are listed in Table 1.

The last reported non-detection of the object is from 2016 October 19.53 (JD=2457681.03) in the orange-ATLAS filter ( $>18.66$  mag),  $\sim 8$  days before the first detection. Forced photometry of pre-discovery images which have been differenced with the all sky reference frame (see Stalder et al. 2017), confirmed this as an actual detection, pushing further the last non-detection to the previous image of the region on 2016 October 10.56 (JD=2,457,672.06) in the orange-ATLAS filter, which provides a constraint on the explosion epoch to within 9 days between the last non-detection and first detection. Throughout



**Fig. 2.** *g*-band imaging of SN 2016hnk and its host galaxy, before (left) and after (middle) subtracting the host galaxy template using hotpants integration in SNOoPY. This example shows how the subtracted image is clean enough to perform point-spread-function (PSF) fitting, as shown in the right panel where the SN PSF has been removed. Note the small residual of the galaxy core on the middle and right panels (white blob), however it does not affect SN PSF fitting.

the paper, we assume that the explosion epoch is the mid point between the last non-detection and the first detection with an uncertainty of half of this time, on JD=2457676.5  $\pm$  4.5, however, we base SN epochs on the date of peak brightness (see Section 4).

SN 2016hnk was initially classified by the NOT (Nordic Optical Telescope) Un-biased Transient Survey (NUTS; Cannizzaro et al. 2016a,b) on 2016 Nov 01.08 UT as a type I supernova. However, they did note it resembled the spectra of a 1991bg-like SN. The spectrum was relatively red, with strong lines dominated by the Ca NIR triplet, O I  $\lambda 7774$ , and Si II  $\lambda 6355$ , and with a Si II velocity of  $\sim 10,000 \text{ km s}^{-1}$ . A few hours later, on November 2nd at 20:16 UT, the Public ESO Spectroscopic Survey for Transient Objects (PESSTO; Smartt et al. 2015) confirmed the redshift and provided a definite classification as a type Ia SN spectroscopically similar to SN 1991bg (Dimitriadis et al. 2016). A third report a day later, on November 3rd at 04:24 UT, provided a more detailed classification (Pan et al. 2016). Although the report was cautious, it pointed out similarities between SN 2016hnk and the peculiar Ca-rich object PTF09dav (Sullivan et al. 2011; Kasliwal et al. 2012), from their peak magnitudes, similar colors, and red spectra with overall similar features. In particular, two deep absorptions at  $\sim 5350$  and  $\sim 5540 \text{ \AA}$  and another absorption at  $\sim 7120 \text{ \AA}$  (in the rest frame) were identified as Sc II and Ti II, respectively, by Sullivan et al. (2011). However, in Section 5 we will show that these features are actually formed by different elements removing any resemblance between both objects, other than the extremely strong calcium features and low luminosity. Finally, another spectrum was obtained with Hiltner MDM-OSMOS on Oct 29.29, however observations were not reported.

Given its peculiar nature, it was unusually faint and showed spectral features that clearly distinguish it from other 1991bg-like SNe, SN 2016hnk was recognized as a rare and potentially interesting object. A follow up campaign was started immediately the day after the PESSTO Telegram with the New Technology Telescope (NTT) and the Las Cumbres Observatory network, which ran for more than two months. All the spectra described above and those obtained during the PESSTO and Las Cumbres campaigns, together with extensive photometry from different sources, are presented in this work and described below.

## 2.1. Photometry

Most of the optical photometric coverage was obtained by the Las Cumbres Observatory 1.0 m telescope network (Brown et al. 2013) through PESSTO allocated time and a number of dedicated SN follow up programmes. In addition, supplementary data were obtained from: the 2.0 m Liverpool Telescope (LT; Steele et al. 2004); the twin 0.5 m ATLAS telescope system (Tonry et al. 2018); the 0.5 m DEdicated MONitor of EXotransits and Transients (DEMONEXT; Villanueva et al. 2018); the NUTS survey using ALFOSC at the NOT 2.56 m telescope; and pre-discovery data from the All-Sky Automated Survey for SuperNovae (ASAS-SN; Shappee et al. 2014). All data have been reduced following standard procedures with dedicated pipelines for each instrument, which consist of bias subtraction, trimming of the over scan region, and flat fielding.

Near-infrared (NIR) follow up broad-band imaging was obtained by both PESSTO and NUTS. PESSTO obtained NIR imaging with the 3.58 m NTT equipped with the Son of ISAAC (SOFI; Moorwood et al. 1998), while NUTS obtained data with NOTCAM on the NOT 2.56 m telescope. From these observations we were able to obtain six epochs of *JHK*-band photometry extending from +5 to +72 d from peak. NTT images were reduced using the PESSTO pipeline described in Smartt et al. (2015), while the NUTS images were reduced using a slightly modified version of the NOTCam Quicklook v2.5 reduction package<sup>1</sup>. In both cases, the reduction procedures consisted of flat-field, distortion, cross-talk, and illumination corrections, trimming, sky subtraction and stacking of the dithered images. After proper sky-subtraction, obtained through dithering, the individual science exposures were aligned and combined to produce a definitive stacked science image, to which we then applied accurate astrometry.

Given the location of the SN within the galaxy, we also obtained deep observations of the field in all bands and in all telescopes after more than a year from the discovery of the SN, to be used as a template for galaxy subtraction.

We performed photometry on the SN and up to 35 reference stars located within the field covered by these different instruments using the SuperNOva PhotometrY (SNOoPY<sup>2</sup>) package. First, each individual exposure is aligned to the USNO-B catalog (Monet et al. 2003), and a point-spread-function (PSF) with a width of the seeing is fitted to all reference stars shown in Figure 1 to calculate the zero points. We then subtracted the host galaxy templates using hotPants (Becker 2015), and performed PSF fitting on the SN in the subtracted image (see Figure 2 for an example). Finally, we calibrated all SN optical photometry against the Sloan Digital Sky Survey Data Release 15 (SDSS; *ugriz*; Aguado et al. 2019), the AAVSO Photometric All-Sky Survey (APASS; *BV*; Henden et al. 2015), and NIR photometry to the Two Micron All Sky Survey (2MASS; *JHK*; Skrutskie et al. 2006), accounting for corresponding airmass and color corrections in a standard manner.

ATLAS photometry in *c* and *o* bands, and ASASS-SN *V*-band photometry are presented in Tables C.1 and C.2, respectively. All other photometric observations and SN magnitudes are given in Tables C.3 (optical) and C.4 (NIR) in the AB system.

<sup>1</sup> <http://www.not.iac.es/instruments/notcam/guide/observe.html#reductions>

<sup>2</sup> SNOoPY is a package for SN photometry using PSF fitting and/or template subtraction developed by E. Cappellaro. A package description can be found at <http://sngroup.oapd.inaf.it/snoopy.html>.

## 2.2. Spectroscopy

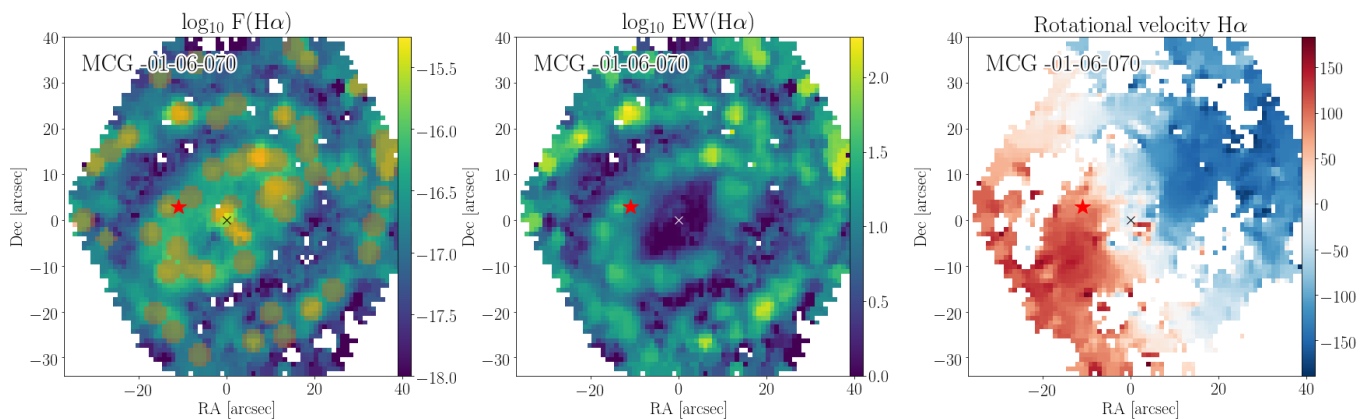
After the reported peculiarity of the object, we began an intensive optical spectroscopic follow-up campaign using several facilities. Most of the data (12 epochs, including 1 NIR spectrum) were obtained by the PESSTO collaboration with the NTT using the ESO Faint Object Spectrograph and Camera (EFOSC2; Buzoni et al. 1984) instrument and SOFI (for the NIR spectrum). A description of the reduction pipeline can be found in Smartt et al. (2015).

These observations are complemented with data from the following sources: the LT using the Spectrograph for the Rapid Acquisition of Transients (SPRAT), and reduced and calibrated in wavelength using the pipeline described in Piascik et al. (2014); the NOT using ALFOSC through the NUTS survey. These data were reduced in a standard manner with ALFOSCGUI; the Yunnan Faint Object Spectrograph and Camera (YFOSC) mounted on the Lijiang 2.4 m telescope (LJT; Zhang et al. 2014); the MDM 2.4 m Hiltner telescope at KPNO using OSMOS (Martini et al. 2011). The reductions follow normal procedures (Villanueva et al. 2016, 2018); the 6.5 m MMT telescope with the Blue Channel instrument (Schmidt et al. 1989). Standard procedures to bias-correct, flat-field, and flux calibrate the data using IRAF were followed; the 8.1 m Gemini South Telescope with the FLAMINGOS-2 spectrograph reduced following the prescriptions in Hsiao et al. (2019); and the two discovery spectra reported in the TNS and Astronomical Telegrams (ATELs) from the 4.1m Southern Astrophysical Research Telescope (SOAR) and the 11m Southern African Large Telescope (SALT) are also included here.

Before performing any analysis of the spectra, they have been first scaled to match the photometry following a procedure known as *color-matching* (Hsiao et al. 2007): we first fit a third order spline to both the interpolated observed photometry with gaussian processes at the epoch of each spectrum, and to the same synthetic bands from the observed spectrum; then, the observed spectrum is multiplied by a wavelength-dependent factor that matches the two splines. This procedure ensures that colors from the corrected spectra are the same as those from the photometric measurements, which we assume to be more reliable.

Finally, we obtained six spectra of the nebular phase at four different epochs from +346 to +377 d with X-shooter (Vernet et al. 2011) under the ESO programme 0100.D-0738(A). X-shooter provides coverage of the full near-UV (UVB), optical (VIS), and NIR wavelength range (300 to 2,500 nm) in a single exposure in three different channels. We used slits of 1''3 and 1''2 for the UVB and VIS arms, respectively. For the NIR arm, we opted for the 0.9'' JH slit to optimize the signal-to-noise (S/N) in the *J* and *H* bands while blocking the *K*-band region. With this setup, a relatively homogeneous slit transmission of ~85% is achieved throughout the wavelength range. Exposure times were fixed to 3060 sec in the UVB channel, 3000 sec in the VIS, and 3000 sec in the NIR. The six sky-subtracted (but not telluric-removed) observations were combined in a single spectrum to enhance the signal-to-noise ratio. The full spectral log is presented in Table C.5<sup>3</sup>.

<sup>3</sup> All spectroscopic observations are publicly available at <https://github.com/lgalbany/SN2016hnk>.



**Fig. 3.** Continuum-subtracted  $H\alpha$  emission (left),  $H\alpha$ EW (middle), and  $H\alpha$  velocity (right) maps of MCG -01-06-070.  $HII$  regions extracted with `HIIEXPLORER` are marked in brown polygons in the left panel. The red star and the black/white  $\times$  represent respectively the locations of SN 2016hnk and MCG -01-06-070 nucleus.

### 3. The host galaxy and local environment of SN 2016hnk

Integral field spectroscopy of the type SB galaxy MCG -01-06-070 was obtained on 2018 October 11 as part of the PMAS/Ppak Integral field Supernova hosts COmpilation (PISCO; Galbany et al. 2018).

Observations were performed with the Potsdam Multi Aperture Spectrograph (PMAS) in PPak mode (Roth et al. 2005) mounted to the 3.5 m telescope of the Centro Astronómico Hispano Alemán (CAHA) located at Calar Alto Observatory. PPak is a bundle of 331 fibers of 2.7'' diameter ordered in a hexagonal shape with a filling factor of the field-of-view (FoV) of 55% (Verheijen et al. 2004; Kelz et al. 2006). Three 900 s exposures were performed with the V500 grating of 500 lines  $\text{mm}^{-1}$ , which provides a spectral resolution of  $\sim 6 \text{ \AA}$  and covers the whole optical range from 3750 to 7300  $\text{\AA}$ . The second and third exposures were shifted by a few arcsec with respect to the first pointing to provide spectroscopic coverage of the full hexagonal  $\sim 1.3 \text{ arcmin}^2$  FoV with  $1'' \times 1''$  pixels, which corresponds to  $\sim 4000$  spectra. Reduction was done with the pipeline used for the Calar Alto Legacy Integral Field Area (CALIFA) survey Data Release 3 (v2.2; all details can be found in Sánchez et al. 2016).

The analysis was performed in a similar way as that presented in Galbany et al. 2014, 2016b. We used a modified version of STARLIGHT (Cid Fernandes et al. 2005; López Fernández et al. 2016, priv. comm.) to estimate the fractional contribution of different simple stellar populations (SSP) from the *Granada-Miles* base (González Delgado et al. 2015) with different ages and metallicities to the stellar continuum in the spectra, and adding dust effects as a foreground screen with a Fitzpatrick (1999) reddening law and  $R_V = 3.1$ . We then obtained pure gas emission spectra by subtracting the best SSP fit from the observed spectra and estimated the flux of the most prominent emission lines after correcting for dust attenuation from the Balmer decrement (assuming case B recombination, Osterbrock & Ferland 2006; a Fitzpatrick 1999 extinction law; and  $R_V = 3.1$ ). Line fluxes were used to estimate a number of galactic parameters, such as the ongoing star-formation rate (SFR; from the  $H\alpha$  emission line), the weight of young-to-old populations (from the  $H\alpha$  equivalent width;  $H\alpha$ EW), and oxygen abundances.

In this work we are mainly interested in studying the local properties of the SN 2016hnk environment and the integrated properties of its host. For the latter, we repeated the above proce-

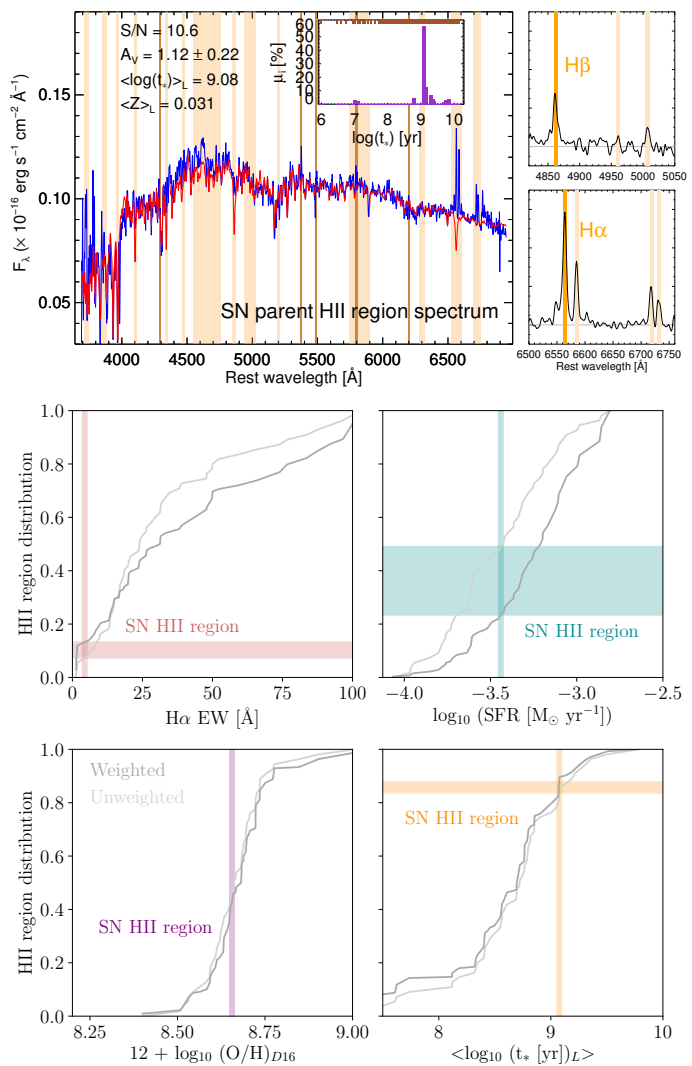
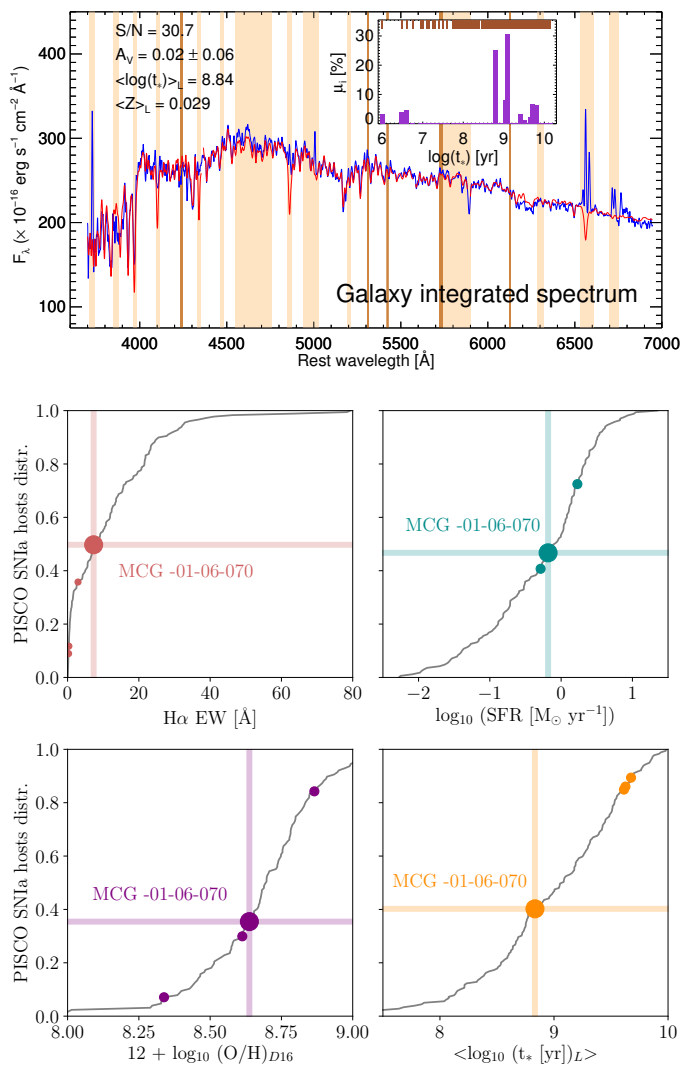
dures on an integrated spectrum of the host galaxy, summing up the spectra in all spaxels, and accounting for spatial covariance in the error budget. Wide-field IFS allows SN explosion site parameters to be compared not only to those of the overall host, but to all other stellar populations found within hosts, therefore exploiting the full capabilities of the data. Following Galbany et al. (2016a) we developed a method to characterize the SN environment comparing its properties to all other nebular clusters in the galaxy. Using our extinction-corrected  $H\alpha$  map we selected star-forming  $HII$  regions across the galaxy with a modified version of `HIIEXPLORER`<sup>4</sup> (Sánchez et al. 2012), a package that detects clumps of higher intensity in a map by aggregating adjacent pixels. This procedure selected 56  $HII$  clumps with an average radius of  $\sim 700$  pc. Once the  $HII$  regions were identified, the same analysis described above was performed on the extracted spectra.

Figure 3 shows two-dimensional maps of the continuum-subtracted  $H\alpha$  emission with the  $HII$  region segmentation performed with `HIIEXPLORER` on top (left), the  $H\alpha$  equivalent width (middle), and the rotational velocity field (right). In all panels the red star represents the SN position. SN 2016hnk occurred between two  $HII$  clumps, within the small disc created by the bar present in the host galaxy. One of these clumps has both high  $H\alpha$  brightness and  $H\alpha$ EW, pointing to a strong component of young stellar populations and ongoing star formation. However the SN is closer to the clump showing bright  $H\alpha$  emission but lower  $H\alpha$ EW, which means ongoing star formation in a location dominated by older non-ionizing stellar populations, as it is seen in the star-formation history of Figure 5. At the SN location the rotational velocity of the galaxy is  $82 \pm 6 \text{ km s}^{-1}$ , which we subtracted in all SN velocities measured in Section 5.

The resulting integrated spectrum of MCG -01-06-070 is presented in Figure 4 with the best STARLIGHT SSP fit. The recovered star-formation history is also shown in the inner plot. The stellar mass of the galaxy is  $4.84 \times 10^{10} M_{\odot}$ , and the average stellar age of the galaxy is 690 Myr, although we were able to recover three main peaks of formation: a very old component at around 6 Gyr, an intermediate burst at around 1 Gyr, and a young component of a few Myr. The best fit also provides an oversolar stellar metallicity of  $Z=0.03$  and a non-significant optical stellar extinction of  $A_V=0.02 \pm 0.06 \text{ mag}$ .

The lower panels of Figure 4, aim to compare the properties of SN 2016hnk and the other three 1991bg-like SNIa host galax-

<sup>4</sup> [http://www.caha.es/sanchez/HII\\_explorer/](http://www.caha.es/sanchez/HII_explorer/)



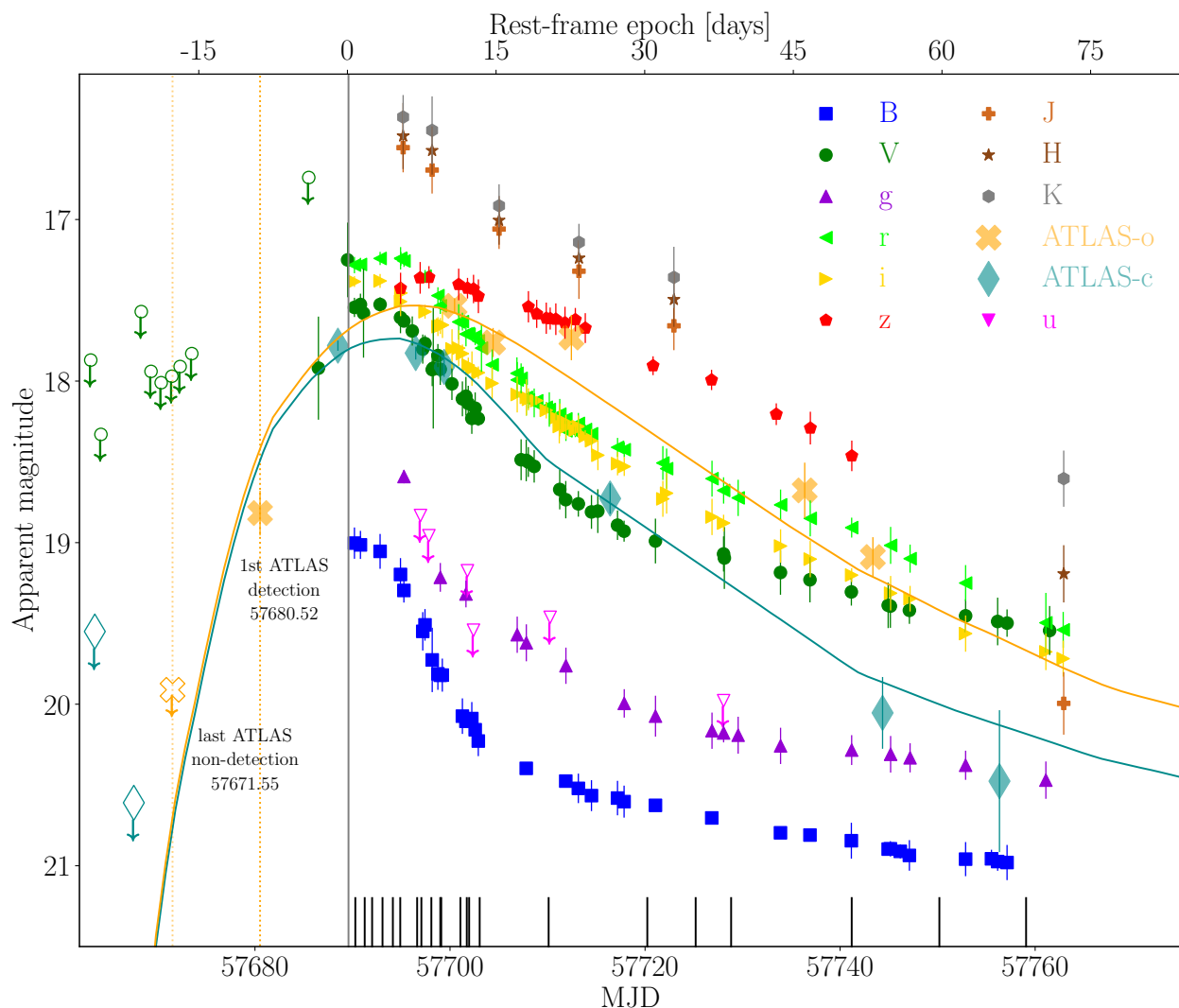
**Fig. 4.** Top: MCG -01-06-070 integrated spectrum (blue) and best simple stellar-population fit from STARLIGHT (red). Vertical strips (in beige) represent masked regions in the fit due to known emission lines (such as Balmer lines, oxygen, WR bumps, etc.) and regions with sky-lines (in brown). The inner panel shows the star-formation history of the spectra (in purple), where upper brown ticks represent the ages of the models used in the fit. Signal-to-noise ratio of the spectrum, optical extinction, average stellar age, and average metallicity are reported in the upper left corners. Lower rows: Distributions of the  $H\alpha$  equivalent width, star-formation rate, oxygen abundance, and average stellar age of all 180 SN Ia host galaxies in PISCO (Galbany et al. 2018). The large colored dots represent the position of MCG -01-06-070, and the three small dots represent other 1991bg-like SNIa hosts in PISCO.

ies with those of all other 180 SN Ia host galaxies in PISCO (Galbany et al. 2018; published sample plus observations up to March 2019 to be presented elsewhere). We present distributions of the  $H\alpha$ EW, SFR, oxygen abundance ( $12 + \log_{10} \text{O}/\text{H}$ ) in the Dopita et al. (2016) scale, and the average light-weighted stellar age. For MCG -01-06-070 we found an  $H\alpha$ EW of  $7.04 \pm 0.18$  Å, a SFR of  $0.649 \pm 0.234 M_\odot \text{ yr}^{-1}$ , and an oxygen abundance of  $8.64 \pm 0.11$  dex. The big colored dot represent the position of MCG -01-06-070 parameters in such distributions, and the vertical and horizontal strips show the actual parameter value of this galaxy and the rank within PISCO. The other three 1991bg-like SN Ia hosts are also represented with smaller colored dots. Overall, MCG -01-06-070 has no extreme proper-

**Fig. 5.** Top: Similar to Figure 4, the spectrum of the SN 2016hnk parent HII region (blue) and best simple stellar-population fit from STARLIGHT (red) are shown. The right panels show the stellar-subtracted spectra around the  $H\alpha$  and  $H\beta$  wavelengths. Lower rows: Distributions of the gas-based parameters of all 56 HII regions in MCG -01-06-070, weighted by the contribution of the region to the total SFR of the galaxy (dark lines) or unweighted (light lines). In this case, colored horizontal bands represent the band where the SN 2016hnk parent HII region lies within.

ties and lies within the 35-50% rank in all measured parameters within the PISCO sample. However, it has a younger average stellar population age and, equivalently, a higher  $H\alpha$ EW than the other 1991bg-like SN Ia hosts, which show less evidence of ongoing/recent star formation.

Regarding the local environment of the SN, in Figure 5 we present the spectrum of the closest HII region from the segmentation described above. We also show two panels at the top right centered at Balmer  $H\alpha$  and  $H\beta$  wavelengths, where the best stellar continuum fit has been subtracted from the observations leaving the pure emission spectra. In those, the O III, N II, and S II emission lines (in pale pink) and  $H\alpha$  and  $H\beta$  (in orange) have been highlighted. The extinction-corrected fluxes of these seven lines were used to calculate the  $H\alpha$ EW, the SFR, and the oxygen abundance, which are shown in the bottom panels together with the distributions of the same parameters measured from the spectra of all 56 HII regions found in our segmentation of the



**Fig. 6.** Optical and near-infrared light curves. A grey solid vertical line represents the epoch of peak brightness in the  $B$ -band obtained from a SNOOPy fit to the *cyan* and *orange* ATLAS light-curves, which are shown in orange and cyan solid lines. Ticks at the bottom represent epochs with optical spectra available. Dotted vertical lines mark the last non-detection (yellow) and first detection (orange) epochs. Open symbols with downward arrows correspond to ASAS-SN (circles), ATLAS (crosses/diamonds), and  $u$ -band (triangles) non-detections.

host galaxy. We included the distribution of the parameters as measured from the spectra, and once weighted by the contribution of the HII region to the total SFR (Lyman et al. 2018). For the HII region closest to SN 2016hnk we get a  $H\alpha$ EW of  $4.39 \pm 0.04 \text{ \AA}$ , a SFR of  $-3.44 \pm 0.34 M_{\odot} \text{ yr}^{-1}$ , an oxygen abundance of  $8.65 \pm 0.11 \text{ dex}$ , and an average stellar age of 1.2 Gyr, with a dominant peak ( $\sim 65\%$ ) at the average age and only a  $\sim 4\%$  contribution from young ( $< 15 \text{ Myr}$ ) populations. Interestingly, in this case we find that these parameter values deviate from the average properties of all HII regions: it is at the lower end of the  $H\alpha$ EW distribution indicating a residual contribution of young populations compared to other regions of the galaxy, and can also be seen in the stellar age distribution, in agreement with the expected old progenitors for 1991bg-like SNe Ia; conversely it is in an average star forming region (25 to 50% rank once weighted for the SFR) and has an average oxygen abundance ( $\sim 50$ th percentiles).

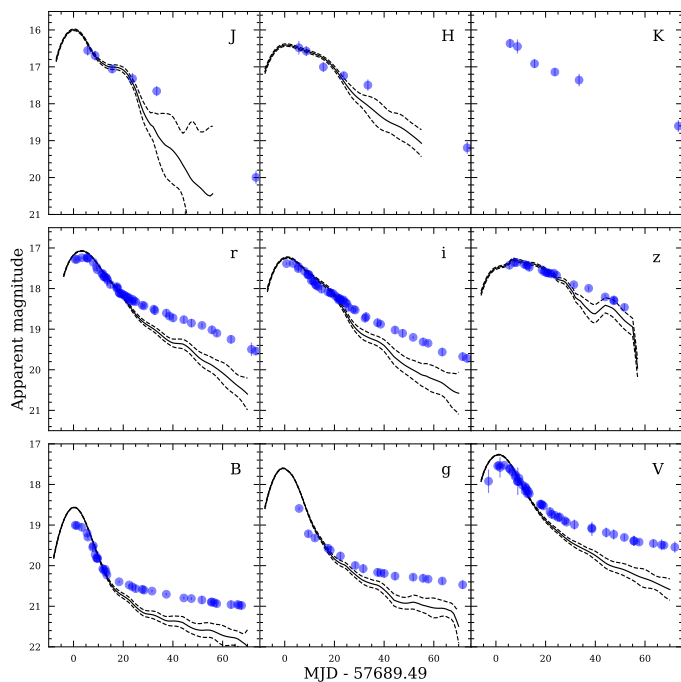
The  $H\alpha$  to  $H\beta$  flux ratio also provides an estimate of the color excess  $E(B - V)$  under certain assumptions (case B recombination; Osterbrock & Ferland 2006). We get a ratio of  $4.36 \pm 0.35$ , which compared to the theoretical value of 2.86 for

a non-extincted ratio and using a Fitzpatrick (1999) extinction law, gives an  $E(B - V)$  of  $0.363 \pm 0.049 \text{ mag}$ .

#### 4. Observational photometric properties

Final *ocBVugrizJHK* light curves are presented in Figure 6. SN 2016hnk photometry has rather sparse coverage before peak, since it was only observed in ATLAS *oc* bands and ASAS-SN  $V$ -band during the rise. However, from the shape of the light curves in all bands shown in Figure 6, we can infer that the follow up campaign started very close to the epoch of peak brightness ( $t_{\text{max}}$ ).

To determine  $t_{\text{max}}$  precisely, we used the SuperNovae in object-oriented Python (SNOOPy; Burns et al. 2011, 2014) fitter, which allows us to consider different extinction laws and amounts of reddening in our light curve fits. Also, SNOOPy makes studies involving color curves very flexible, such as the study of color laws, the construction of bolometric light curves, and estimates of the dimensionless  $s_{BV}$  parameter, which gives the stretch factor of the  $B - V$  color curve with respect to the typical behaviour of a SN Ia having a  $t_B^{\text{max}} - t_{B-V}^{\text{max}}$  of 30 days. The  $s_{BV}$  parameter provides a more physically significant inter-



**Fig. 7.** Multi-band light curves of SN 2016hnk (filled circles) compared with a 1991bg-like template (lines). Best-fit SNooPy parameters are  $R_V=2.1$  and a host-galaxy color excess of  $E(B - V)_{\text{host}}=0.55$  mag. Note that beyond +20 d the template fit does not succeed in capturing the SN brightness.

pretation of the light curve width than similar parameters, as it measures the time at which iron re-combines from Fe III to Fe II (e.g. Burns et al. 2014; Wygoda et al. 2018). So,  $s_{BV}$  may be providing more information when it comes to fast-decliners and peculiar types, such as SN 2016hnk. We fit the light curves of the three bands that have pre-maximum observations and the  $B$  band, with the 1991bg-like built-in model, and obtain a  $t_{\text{max}}$  of  $\text{MJD}=57689.5 \pm 3.3$  (ATLAS-*co* fits are included in Figure 6). This estimate constrains the SN rise time to  $17.4 \pm 5.5$  days, hence to the range from 11.9 to 22.9 days, consistent with most SNe Ia and 1991bg-like SNe, but larger than in Ca-rich transients, which have faster rise times of the order 9-15 days (Taubenberger 2017).

A SNooPy light curve fit using the Markov Chain Monte Carlo fitter and the 91bg model was simultaneously performed in all bands and to the whole temporal extent, fixing the  $t_{\text{max}}$  and also inputting the  $E(B - V)$  obtained from the analysis of the most nearby environment as a Gaussian prior. Our initial attempt provided a reasonable fit up to 15-20 days past maximum (See Figure 7), but the most remarkable mismatch was found at late times ( $>+20$  d) where SN 2016hnk showed an excess in brightness in all bands. Best-fit parameters were  $R_V=2.1 \pm 0.4$ ,  $E(B - V)=0.547 \pm 0.064$  mag, and  $s_{BV}=0.438 \pm 0.030$ .

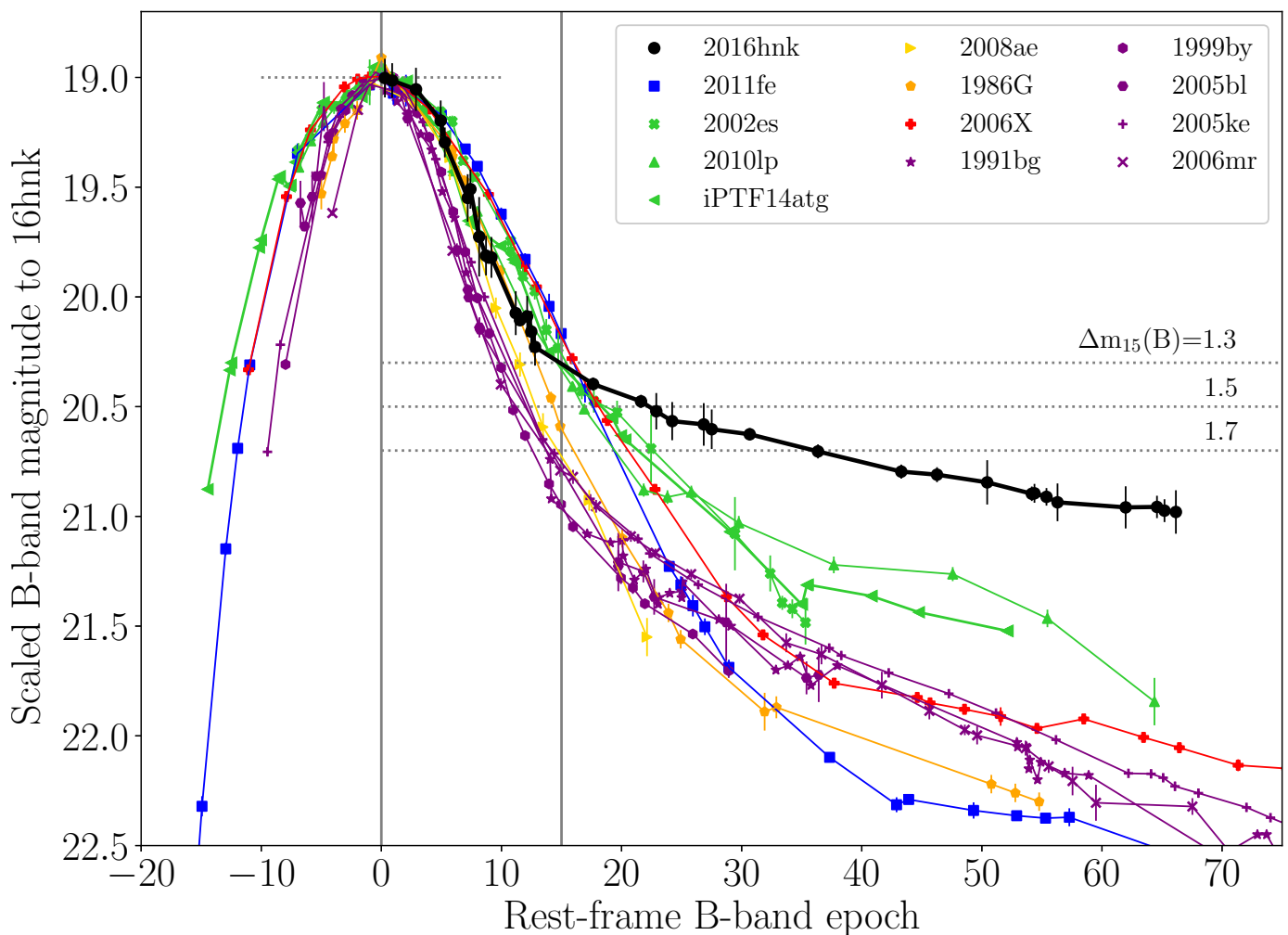
Figure 8 shows SN 2016hnk  $B$ -band compared to a variety of SNe Ia: the normal SN Ia 2011fe (Richmond & Smith 2012); three 2002es-like SNe Ia, 2002es (Ganeshalingam et al. 2012), 2010lp (Kromer et al. 2013), iPTF14atg (Cao et al. 2015); a 2002cx-like SN Ia, 2008ae (Krisciunas et al. 2017); the transitional SN 1986G (Phillips et al. 1987); the highly extinguished SN Ia 2006X (Wang et al. 2008; Krisciunas et al. 2017); and a sample of 1991bg-like SNe Ia including 1991bg (Filippenko et al. 1992; Leibundgut et al. 1993; Turatto et al. 1996; Krisciunas et al. 2004), 1999by (Ganeshalingam et al. 2010; Garnavich et al. 2004), 2005bl (Taubenberger et al. 2008; Friedman et al. 2015;

Hicken et al. 2009; Contreras et al. 2010), 2005ke (Friedman et al. 2015; Hicken et al. 2009; Contreras et al. 2010; Krisciunas et al. 2017), and 2006mr (Contreras et al. 2010; Stritzinger et al. 2010), the fastest SN Ia ever found. These 1991bg-like SNe will be used in the following sections as a base for comparison. All light curves have been restframed and shifted vertically to match the SN 2016hnk at peak brightness. For reference, two vertical lines correspond to epochs at peak and at +15 days, and three horizontal lines cross the +15 d line at the position where light curves with  $\Delta m_{15}(B)=1.3$  (similar to SN 2016hnk), 1.5 (intermediate), and 1.7 mag (approx. the minimum for 1991bg-like SN Ia) should cross. SN 2016hnk has a  $\Delta m_{15}(B)$  most similar to 2002es-like objects, but with two peculiarities: it shows a much narrower light curve before +15 d, and has a much shallower decay at later epoch. In fact, up to 12-13 days past max, it follows very closely the low-luminosity transitional SN Ia 1986G light curve, being only slightly wider than all other 1991bg-like SN Ia light-curves. Focusing on the late ( $>+15$  days) light curve, we clearly see that SN 2016hnk has the brightest tail of all objects in comparison. This figure clearly shows that  $\Delta m_{15}(B)$  is not capable of capturing all the diversity in SNe Ia and may mask intrinsic differences in light curve shapes, as it is seen in SNe 2002es and 2016hnk. We expand on the explanation of this light excess in Section 4.1.

As an independent check to increase the robustness of our findings, we used the Color-Decline-Magnitude intercept Calibration (CDMagic; Höflich et al. 2017) to estimate the host-galaxy reddening and preferred dust extinction law. First, following Goldhaber et al. (2001), we measured the stretch factor  $s$  that multiplies the temporal axis of the light curve providing the best match of the Hsiao et al. (2007) SN Ia template to the observations. In our case, we restricted the fit to the first 15 days, to avoid the dependence of the light curve on secondary parameters and the light excess detected for SN 2016hnk at later times, and obtained an  $s=0.59 \pm 0.04$ . The modified decline rate  $\Delta m_{15,s}(B)$  is then defined as the magnitude difference between the peak and the reference epoch, 15 days  $\times s$  (measured in the template), divided by the  $s$  factor  $[\Delta m_{15,s} = \frac{\Delta m_{15}(B)}{s}]$ , which in our case is  $\frac{\Delta m_{15}}{s} = \frac{1.06 \pm 0.13}{0.59 \pm 0.04} = 1.803 \pm 0.200$  mag. The measured value of  $\Delta m_{15,s}(B)$  determines the theoretical model from Höflich et al. (2017) that is used as a reference to correct the  $B$ - and  $V$ -band light curves, the  $B$  vs.  $(B - V)$  and  $V$  vs.  $(B - V)$  diagrams, and the  $(B - V)$  evolution of SN 2016hnk. In the CDMagic fit we fixed the  $t_B^{\text{max}}$  to the value previously obtained with SNooPy and the distance modulus measured from the assumed cosmology ( $\mu=34.195$  mag), and only left as free parameters  $E(B - V)$  and  $R_V$ . While the color excess is estimated from the shift between the uncorrected (as observed and Milky Way dust corrected) and the corrected points that best match the template, the  $R_V$  is measured from the direction of this shift in the color-magnitude diagrams. In Figure 9 we show, as an example, the absolute  $V$  and vs.  $B - V$  diagram, including a number of SNe Ia from the Carnegie Supernova Project I (CSP-I; Krisciunas et al. 2017) and the 1991bg-like SN Ia 2005ke (empty triangles) for reference, models from Höflich et al. (2017) with different stretch factor (green lines), and SN 2016hnk data before (blue) and after corrections (red). We obtained an  $E(B - V)$  of  $0.45 \pm 0.08$  mag, consistent with previous methods, and an  $R_V = 2.1 \pm 0.6$ , which we will be using hereafter, however, we note that our results are not significantly dependent on the particular choice of  $R_V$ .

The observed apparent  $B$ -band peak magnitude is  $19.003 \pm 0.088$  (at epoch +0.7 d). Using  $t_{\text{max}}$  as a reference, we estimated  $\Delta m_{15}(B)$  by interpolating the flux at 15 days with a 3rd order





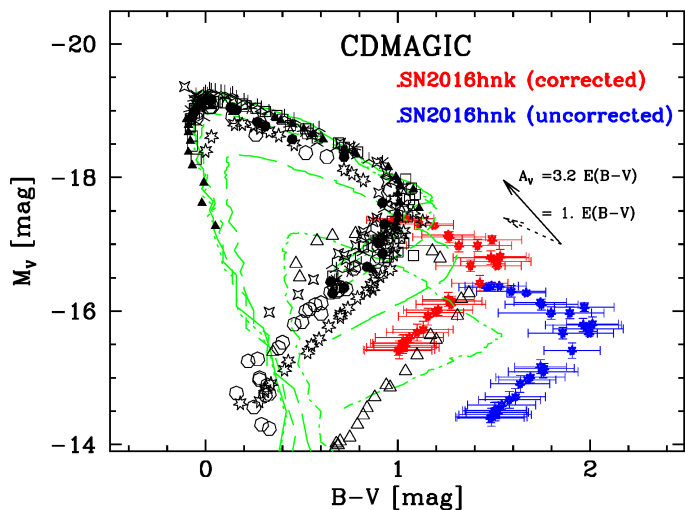
**Fig. 8.** *B*-band light curves of SN 2016hnk and a variety of SNe Ia of different subtypes: normal SN Ia (2011fe), 2002es-like (2002es, 2010lp, iPTF14atg), 2002cx-like (2008ae), transitional (1986G), and 1991bg-like (1991bg, 1999by, 2005bl, 2005ke, 2006mr). All data have been rest-framed and scaled to match the SN 2016hnk peak. Vertical lines represent epochs of the peak and +15 days, and horizontal lines mark the crossing point at 15 days of a SN with  $\Delta m_{15}(B)=1.3$ , 1.5, and 1.7 mag, respectively. SN 2016hnk has one of the lowest  $\Delta m_{15}(B)$  but is the object with the brightest tail in both bands.

spline and obtained a value of  $1.279 \pm 0.071$  mag. We then applied the correction provided by Phillips et al. (1999) to get the true  $\Delta m_{15}(B)_{\text{true}} = \Delta m_{15}(B)_{\text{obs}} + 0.1 \times E(B-V)$ , and found  $1.324 \pm 0.096$  mag. In order to find the absolute *B*-band peak magnitude, we corrected the observed value for Milky Way reddening of  $E(B-V) = 0.0224 \pm 0.0008$  mag from extinction maps of Schlafly & Finkbeiner (2011) and the average Milky Way value for  $R_V$  of 3.1, and the SN host galaxy dust extinction with the color excess of  $E(B-V) = 0.45$  mag and an  $R_V$  of 2.1 found above. We obtained an  $A_V^{\text{MW}} = 0.069$  mag and  $A_V^{\text{host}} = 0.945$  mag, which correspond to  $A_B^{\text{MW}} = 0.092$  mag and  $A_B^{\text{host}} = 1.395$  mag. Considering both corrections we get a corrected absolute *B* band magnitude of  $-16.656 \pm 0.162$ . This is at the low end of the luminosity distribution for SNe Ia (Ashall et al. 2016a), and confirms SN 2016hnk as a subluminal SN Ia.

This is clearly seen in Figure 10, where we show the absolute magnitude vs. light curve width relation for SN 2016hnk and SNe Ia from the CSP-I (Krisciunas et al. 2017). In the upper panel we used the  $\Delta m_{15}(B)$  parameter. While this parameter correlates strongly with the SN peak brightness, it is known to become degenerate for the fainter SNe Ia showing a branch

that follows the same relation as brighter SNe Ia, and a second branch that breaks the relation. In this plot, SN 2016hnk is an outlier, and the reason for this is discussed above and can be seen in Figure 8. It has a very low  $\Delta m_{15}(B)$  when compared to other subluminal SNe Ia, which usually have larger values ( $>1.7$  mag). In the lower panel we use the  $s_{BV}$  parameter instead. Burns et al. (2014) showed that this parameter keeps more information and breaks the degeneracy in the peak brightness vs. light curve width relation for the fainter objects. In this case, SN 2016hnk perfectly follows the relation.

In Figure 11 we present several SN 2016hnk color curves (upper panel), and the *B* – *V* color curve compared to the sample of 1991bg-like SNe Ia presented above, SNe 2011fe, and 2002es (bottom panel). The slope of the *B* – *V* color during the Lira law regime (between 35 and 80 days after maximum) provides an independent test on the nature of the reddening process in play (Lira et al. 1998; Phillips et al. 1999). Extinction by interstellar material (ISM) produces a shift in the *B* – *V* color curve to higher (red) values, without significant changes on the decline rate, whereas sufficiently close dust ( $\lesssim 10$  pc, either circumstellar material –CSM– or ISM) can also affect the decline rate due to an evolving  $E(B-V)$  (Förster et al. 2013; Bulla et al. 2018b).

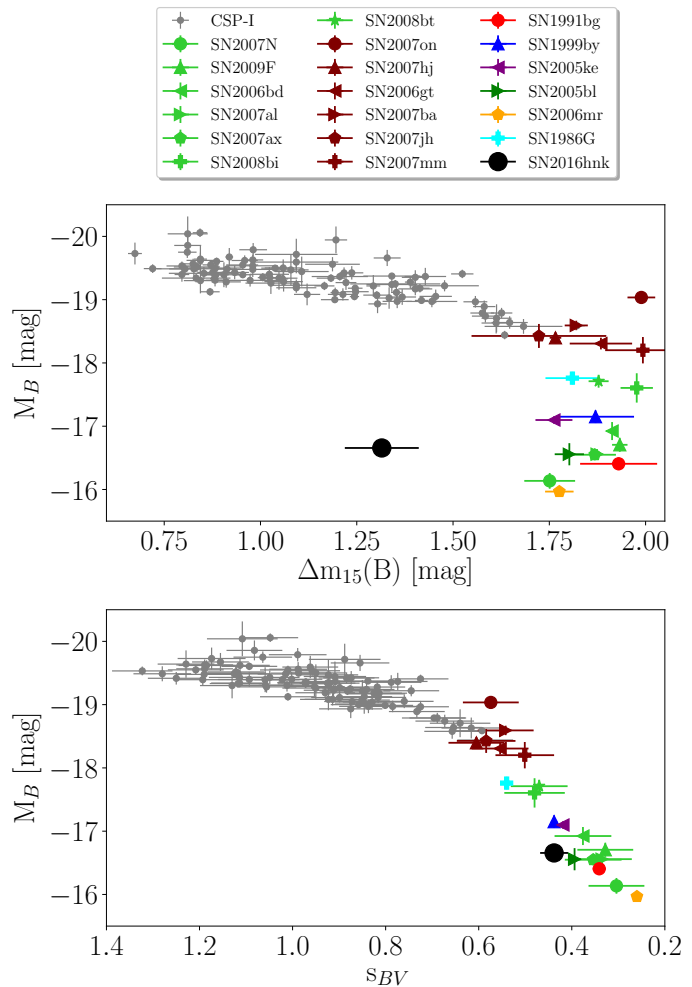


**Fig. 9.** Representative panel for the CDMagic method, where a simultaneous fit in the  $B$ - and  $V$ -band light curves, the  $B$  vs.  $(B - V)$  and  $V$  vs.  $(B - V)$  diagrams, and the  $(B - V)$  evolution, is performed. For more details we refer the reader to e.g. Figure 13 in Höflich et al. 2017. Here we show the absolute  $V$ -band vs.  $(B - V)$  diagram populated with eight normal SNe Ia from the CSP-I (Krisciunas et al. 2017; SNe 2004ef, 2005eo, 2005el, 2005iq, 2005ki, 2005M, 2006bh, 2006D; in different symbols), the SN Ia 1991bg-like 2005ke (empty triangles), and the uncorrected (blue) and corrected (red) SN 2016hnk datapoints. Models from Höflich et al. 2017 are shown in green for reference, and the two arrows are the reddening vectors representing the direction to which the datapoints are shifted when affected by  $R_V=3.2$  and  $R_V=1.0$ . SN 2016hnk corrected datapoints lie on top of the other subluminous SN Ia 2005ke, and from the shift we estimate an  $R_V=2.1\pm 0.6$ .

It can clearly be seen that the observed  $B - V$  curve is well above (redder) those other subluminous SNe at all epochs. While other 1991bg-like objects have  $B - V$  colors at maximum light around 0.4–0.8 mag, SN 2016hnk has 1.5 mag. Up to +20 days, the color curve is consistently half a magnitude larger when compared to other 1991bg-like SNe Ia. This is in agreement with the  $E(B - V)$  values found previously in the CDMagic and SNOOPY analyses, and confirms a large amount of reddening. On the other hand, the post maximum slope is clearly shallower, which is opposite of what is expected for objects affected by large amounts of extinction (Förster et al. 2013). We fit a straight line between epochs (+30,+70) and found a slope of  $-0.005\pm 0.001$  mag day $^{-1}$ , which is actually in the low regime than is expected for objects with insignificant reddening, and is in conflict with the expectations for CSM scattering. However, this may also be caused by either difference in the central density of the progenitor at the explosion (Gall et al. 2018b) or by differences in the transmission filters, which may cause differences in the color curve slope (Höflich et al. 2017).

#### 4.1. Late light curve excess

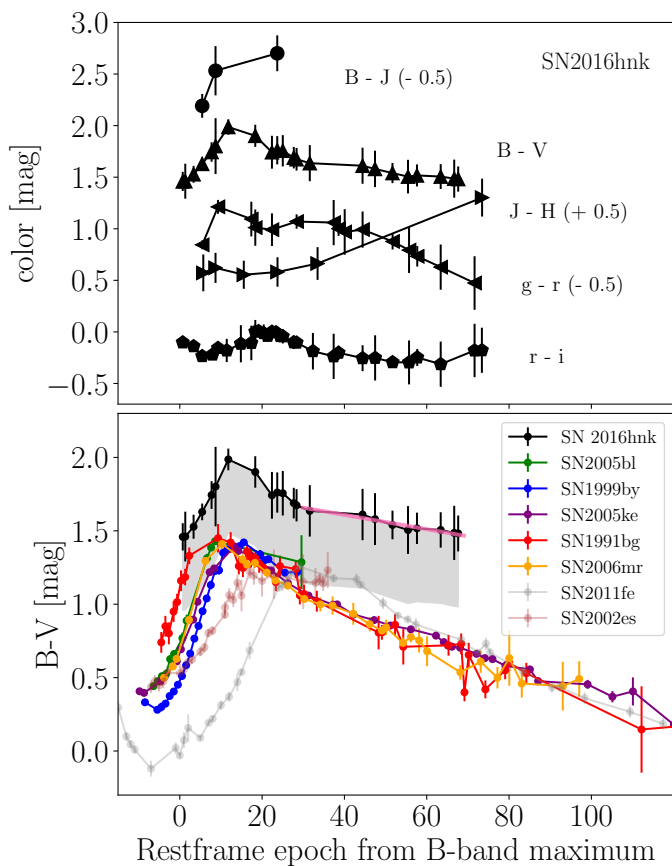
Figure 8 shows that SN 2016hnk presents the shallowest brightness decline at later epochs ( $\geq 12$  days) compared with a variety of thermonuclear events. To explain this light excess we performed a number of tests confirming that: (a) SN 2016hnk is similar to SN 1991bg when similar host-galaxy reddening with low extinction law is artificially applied to SN 1991bg, although it is still redder and no Na I D lines are detected at any epoch in its spectra; and (b) when a simple light-echo model is removed from the observed SN 2016hnk light curve, a SN 1991bg-like model



**Fig. 10.** The luminosity width relation as a function of  $\Delta m_{15}(B)$  (top) and  $s_{BV}$  (bottom) for a selection of CSP-I SNe Ia from Krisciunas et al. (2017), with the addition of SN 2016hnk (black circle). Note that SN 2016hnk sits off the  $\Delta m_{15}(B) - M_B$  relationship because of the degeneracy in  $\Delta m_{15}(B)$  for fast declining SNe Ia.

fits reasonably well to the observations at late epochs, however no light-echo (spectrum at maximum) features are detected in the late-time spectra. Details of those analysis are left for Appendix A, here we concentrate on our most plausible explanation.

At late epochs, the slow decline in the SN 2016hnk light curves suggests an additional energy source beyond  $^{56}\text{Co}$  to  $^{56}\text{Fe}$  radioactive decay. The most likely mechanism for this particular light curve shape is multiple scattering with either a circumstellar or interstellar dust cloud intervening in the line-of-sight, which would in turn cause the total-to-selective extinction ratio to be steeper (lower  $R_V$ ; Goobar 2008). Amanullah & Goobar (2011) presented predictions of how scattering by dust affects the SN light curve shape depending on the distance to the interfering cloud and its thickness (opacity). Qualitatively, the light curve suffers some broadening before 15 days, and it produces a light excess after 15 days (as observed in e.g. SN 2006X), which is more pronounced with thicker dust clouds (prolonged scattering of SN photons). We see similarities between simulated light curves shown in Figure 3 of Amanullah & Goobar (2011) for the  $10^{18}$  cm ( $\sim 0.1$  pc) case, with a broadening and a plateau at latter times, and the SN 2016hnk  $B$ -band light curve when compared to all other 1991bg-like objects. Importantly, the net effect of this multiple scattering would be a reduction of the  $\Delta m_{15}(B)$  param-



**Fig. 11.** SN 2016hnk color curves (top panel), and the  $B-V$  color curve of SN 2016hnk, the 1991bg-like comparison sample, and SNe 2002es and 2011fe for reference. The shaded region projects the SN 2016hnk color curve down by 0.5 mag, on top of the bulk of other 1991bg-like SNe Ia. In pink we show the linear fit to the  $B-V$  curve from 30 to 70 days past  $B$  band maximum.

eter (of up to  $\sim 0.4$  mag) for similar  $E(B-V)$  as in SN 2016hnk, as we have also shown in Figure 8.

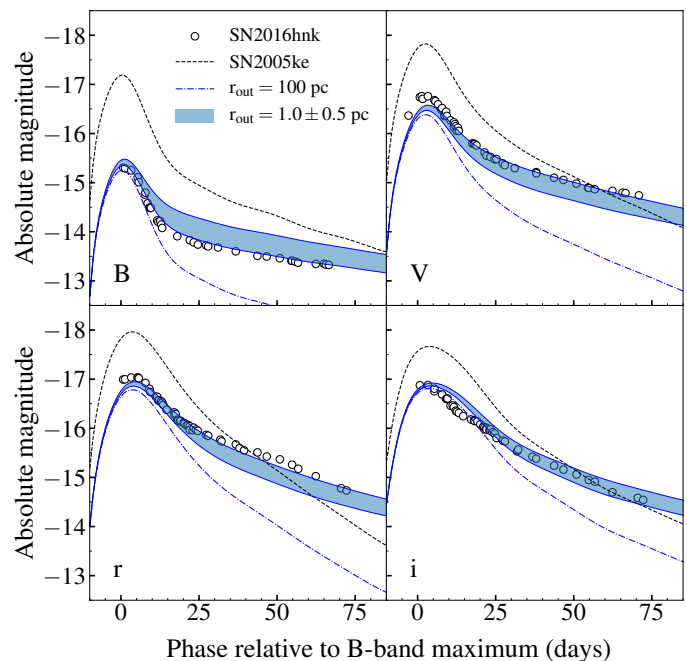
We used the Monte Carlo simulations by Bulla et al. (2018a) to estimate the distance to the intervening dust cloud and, at the same time, try to fit multiband light curves of SN 2016hnk with this model. In Table 2 we provide the dust properties of the model for each band, which are based on a LMC-type dust composition from Weingartner & Draine (2001). We use the 1991bg-like SN 2005ke (Krisciunas et al. 2017) light-curves as a reference in our simulation. In Figure 12, we show the  $BVri$  reference light curves and the resulting simulated light curves affected by nearby dust with  $E(B-V)=0.5$  mag at a distance of  $1.0 \pm 0.5$  pc from the SN (blue filled strip), and at larger distance of 100 pc (blue dashed line). While dust at larger distances would affect the light curve by dimming the light and keeping the observed light curve shape, nearby dust would broaden it at early times and provide an extra energy source at later times. SN 2016hnk observations are included in all panels to show the good agreement between data and model. Based on this simulation, SN 2016hnk would have the same  $\Delta m_{15}(B)$  than SN 2005ke, if it was not affected by the dust cloud.

## 5. Observational spectral properties

In this section, we describe our spectral data in terms of their observed properties, then present the results of a tomographic

**Table 2.** Scattering parameters (albedo, average of the cosine of the scattering angle, and absorption cross section divided by dust mass) for optical photons from Weingartner & Draine (2001) corresponding to interstellar extinction in the Large Magellanic Clouds (LMC).

Wavelength ( $\mu\text{m}$ )	Albedo $\sigma_s/(\sigma_s + \sigma_a)$	$g = \langle \cos \theta \rangle$	$\sigma_a/m_{\text{dust}}$ ( $10^3 \text{ cm}^2/\text{g}$ )	Filter
0.44	0.7159	0.6153	7.542	B
0.55	0.7631	0.6059	4.666	V
0.62	0.7765	0.5896	3.612	r
0.75	0.7749	0.5482	2.646	i

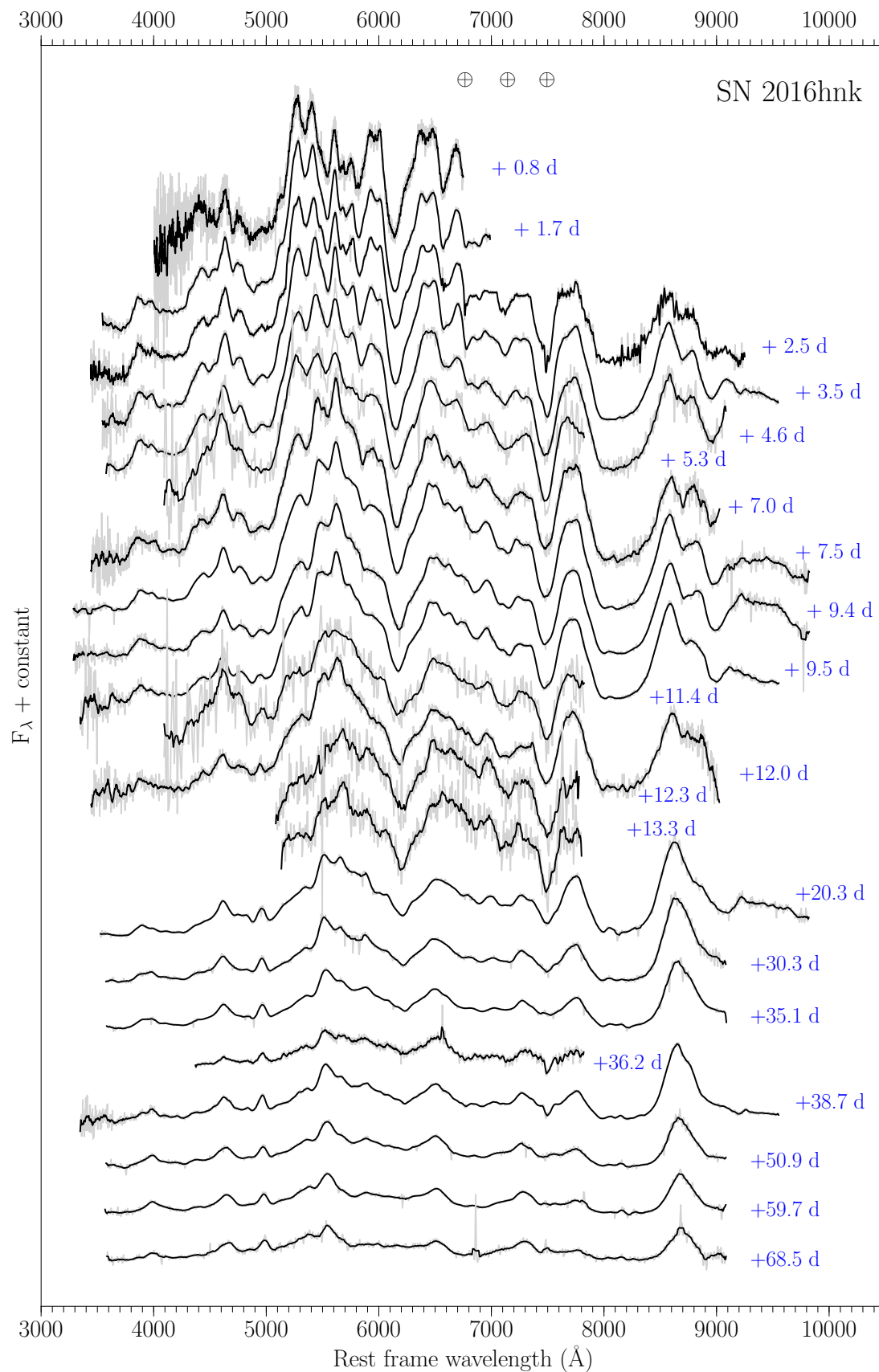


**Fig. 12.**  $BVri$  observed light curves (in circles), compared with the multi-scattering dust model of Bulla et al. (2018a) using 1991bg-like SN 2005ke as a reference (black dotted line). Simulated light curves affected by nearby dust with  $E(B-V)=0.5$  mag at a distance of  $1.0 \pm 0.5$  pc from the SN (blue filled strip) provide very reasonable fits to the observations, while a model where dust is located at larger distance (100 pc; blue dashed line) just dims the light curves by keeping the same shape.

analysis of SN 2016hnk in Section 6.1, and comparison to explosion models in §6.2.

The sequence of color corrected, Milky Way reddening corrected, restframe optical spectra is shown in Figure 13, covering epochs from +0.8 to +68.5 days. In Figure 14 we show spectra of SN 2016hnk and the subluminous SN comparison sample at three different epochs: around maximum light, at one week past maximum, and around three weeks past maximum. We also include spectra of PTF09dav (Sullivan et al. 2011) because SN 2016hnk was claimed to be similar to this Ca-rich SN, and of SN 2002es (Blondin et al. 2012) to enable comparisons to this different SN Ia subclass. In all three panels we can see the resemblance of SN 2016hnk to other 1991bg-like SNe Ia, especially with respect to the  $\text{Ti II}$  feature at  $\sim 4400 \text{ \AA}$ ,  $\text{O I } \lambda 7774$ , and other intermediate mass elements (IME) characteristic of SNe Ia. However, there are a number of features that make SN 2016hnk special:

- First, it is the object with by far the broadest NIR Ca II triplet absorption. This absorption remains wide and deep com-

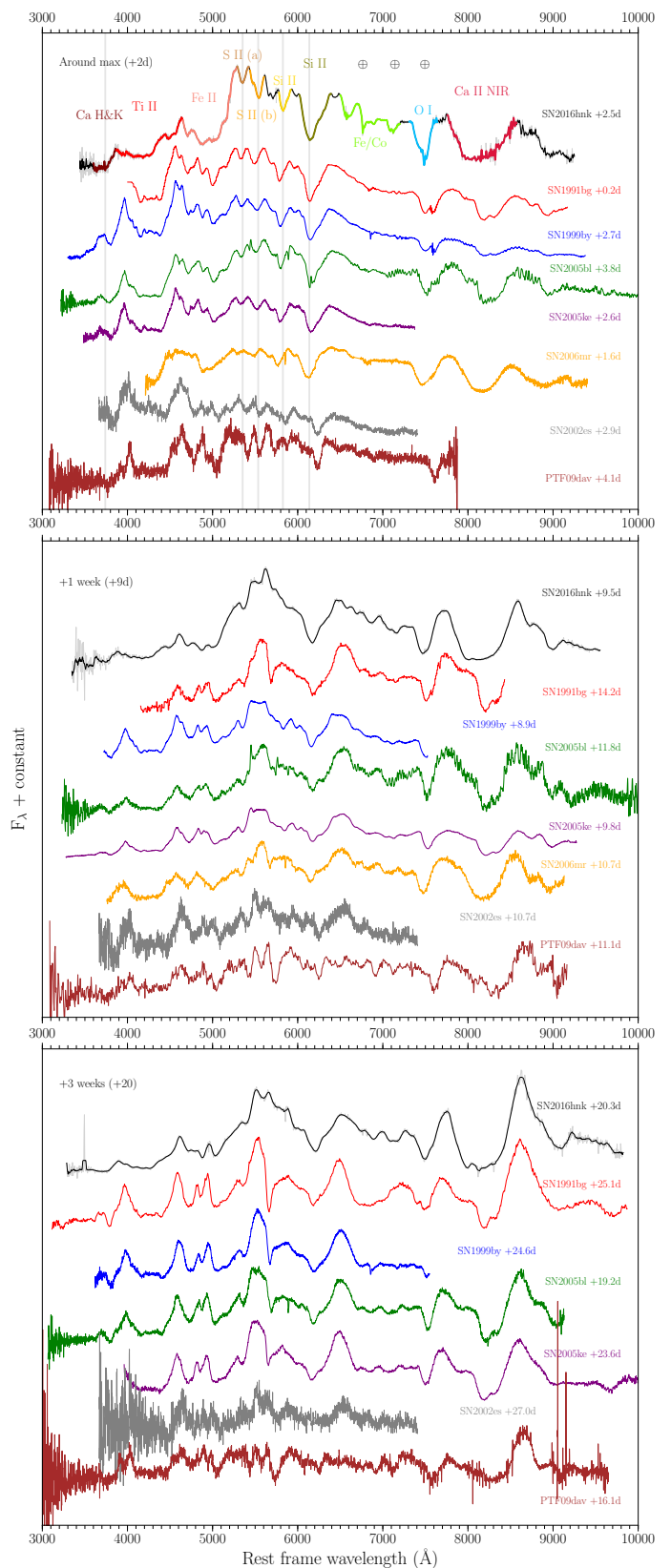


**Fig. 13.** Optical spectral sequence for SN 2016hnk including twenty-two observations listed in Table C.5 (shifted by  $1.5 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2} \text{ \AA}^{-1}$ ). In grey we present the color-corrected, Milky Way reddening corrected, and restframed spectra, and in black a smoothed version using a 10 pixel box. Rest-frame epochs for each spectrum are attached in blue.

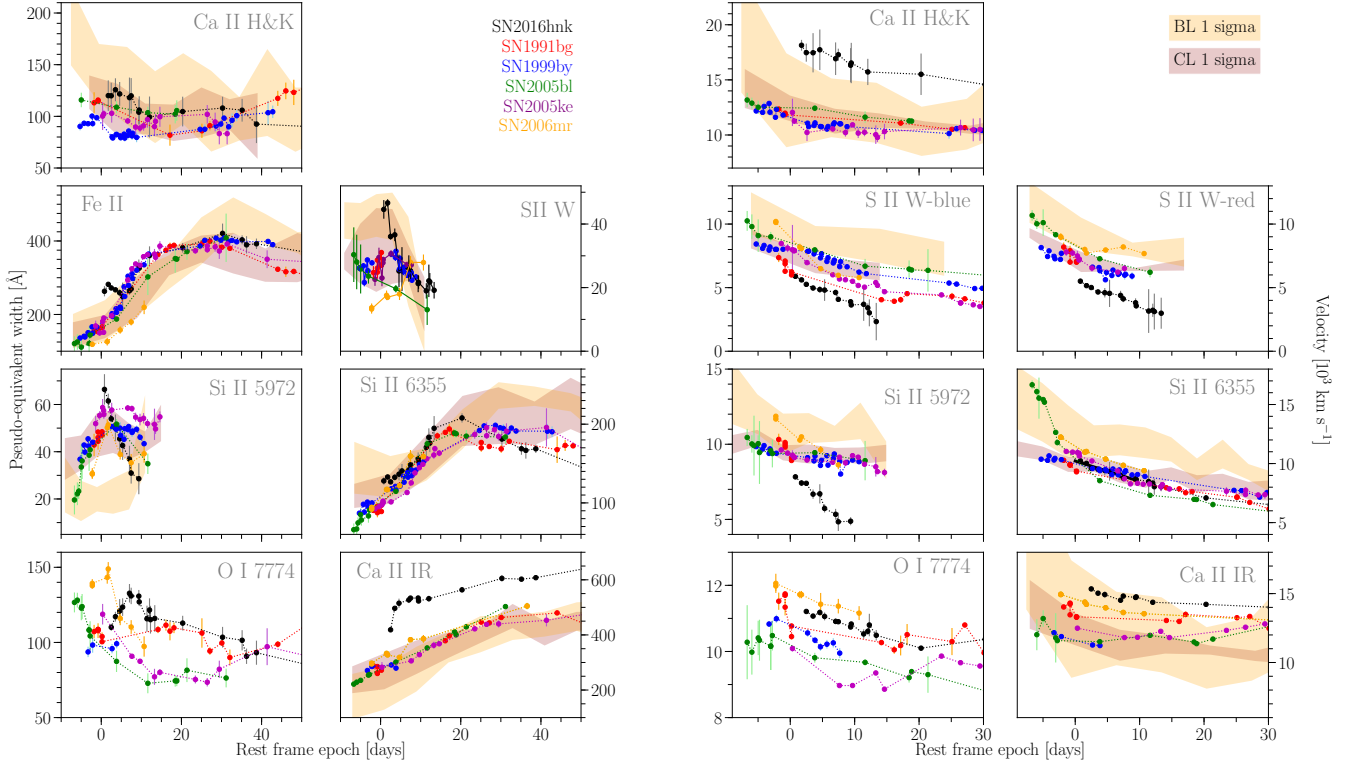
pared to other 1991bg-like objects until at least three weeks past maximum (see bottom panel of Figure 14).

- A second peculiarity are the three features between 6400 and 7300 Å, especially in the spectrum around maximum light, that fall in a wavelength region that is quite flat in other objects. We identified them to be caused by Fe II, Co I, and Co II lines (see discussion in Section 6). Some shallower features are also present in SN 2005bl, and Hachinger et al. (2009) identified them as being mostly caused by Ti II, however their Ti abundance is much larger than that which is found through explosion models and they do not reproduce all of the features.
- As Folatelli et al. (2013) previously noted, some extremely cool SNe show a double structure in the profile of the Si II  $\lambda$ 5972 absorption with a weak component on the blue side of the main absorption. This weaker absorption was attributed to the Na I D doublet that grows stronger at lower temperature, or to Fe II lines at low velocity. SN 2016hnk is the only object in the sample that has this bluer absorption detached from the main Si II feature. Related to this, in the +3 week spectra all other 1991bg-like SNe show a deep Na I D absorption, however SN 2016hnk does not show such an absorption (see also Appendix A).
- All SN 2016hnk spectra are clearly redder than those of the comparison sample. All spectra in Figure 14 have been corrected for Milky Way reddening and shifted to the rest frame, so that reddening excess has to be caused by host-galaxy reddening or intrinsic color. As shown in the previous section, we estimated a color excess of 0.4–0.5 mag, which is corroborated here when comparing SN 2016hnk to the spectra of other 1991bg-like SNe Ia.
- Another interesting thing about SN 2016hnk is the presence of two narrow lines between 5200 and 5600 Å, which we identify as the typical S II “W” feature seen in all SNe Ia, and that were used to claim similarities to the peculiar Carich SN PTF09dav (Pan et al. 2016). Sullivan et al. (2011) were able to reproduce the observed shape of these features in PTF09dav spectra by fitting synthetic spectra simulated with SYNAPPS (Thomas et al. 2011) and inputting Sc II instead of S II. We show in Section 6 that we are able to reliably reproduce these and other features with typical elements found in other 1991bg-like SNe but assuming lower ionization/excitation.

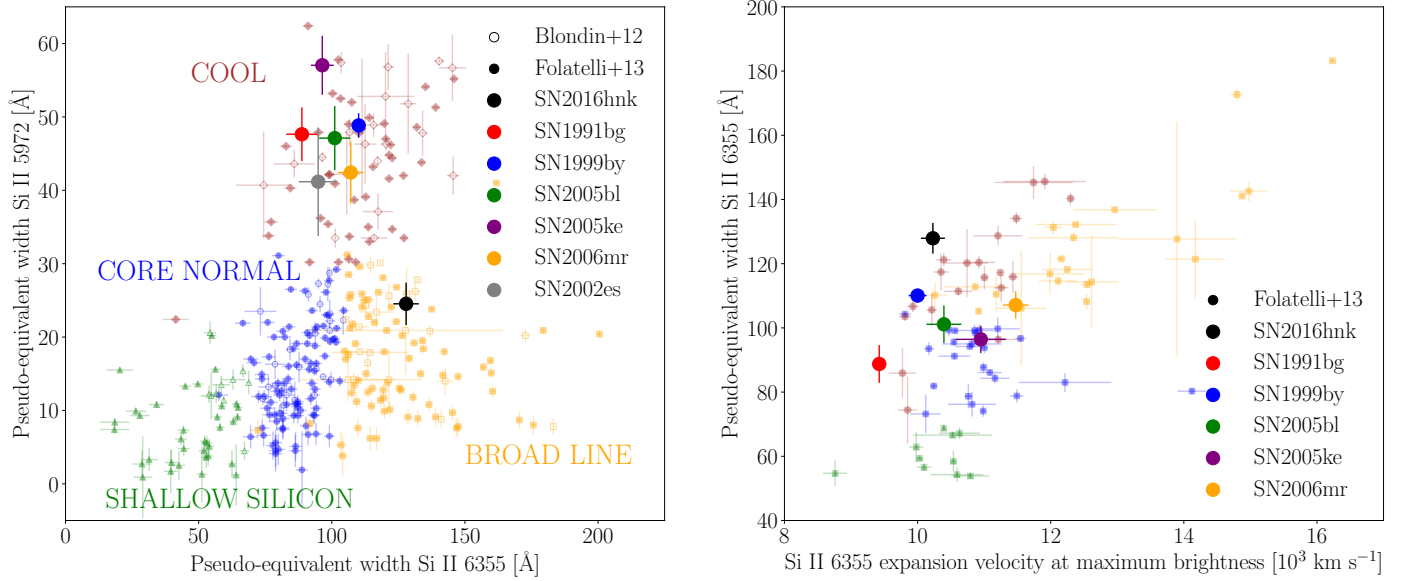
In the upper panel of Figure 14 we highlight a few features in SN 2016hnk that we characterized in the spectral sequence in terms of the evolution of their pseudo-equivalent width ( $pEW$ ) and velocities. We show our measurements in Figure 15 and list them in Tables C.6 and C.7. The most remarkable characteristics of SN 2016hnk compared to all other subluminous Ia are: (i) it shows stronger Ca II H&K and NIR features, as well as larger expansion velocities. However, we note that there are indications of a shallow feature on the blue side of the Ca NIR in the spectra of SNe 1991bg and 2005bl, which may simply be deeper in SN 2016hnk; (ii) while SN 2016hnk has larger S II  $pEW$  than other SNe, the expansion velocities are clearly lower. This can be seen in the upper panel of Figure 14 where we plotted vertical lines corresponding to the velocities of SN 2016hnk to better illustrate the difference to other objects; (iii) while the Si II velocity and  $pEW$  measured from the 6355 Å line are consistent with other subluminous SNe, the Si II  $\lambda$ 5972 velocity is significantly lower and the  $pEW$  decreases very rapidly. This may be caused by the detachment of the bluer feature; (iv) we see indications of two different behaviours regarding the O I  $\lambda$ 7774



**Fig. 14.** Spectra of SN 2016hnk at three different epochs (around maximum, after a week, and around three weeks past maximum light) compared to other subluminous 1991bg-like SNe Ia, PTF09dav (Carich), and SN 2002es (2002es-like SN Ia). The main typical features of 1991bg-like SNe Ia are identified in the upper panel. Vertical lines in the upper panel are located at the minima of a few features, and serve as a comparison of the velocities with respect other objects. All spectra have been Milky Way extinction corrected and shifted to the rest frame.



**Fig. 15.** Evolution of pseudo-equivalent widths  $pEW$  and velocities as a function of time for the most prominent features highlighted in the upper panel of Figure 14. Shaded regions represent the  $1\sigma$  dispersion of COOL (brown) and BROAD LINE (orange) evolution in the CSP-I sample from Folatelli et al. (2013).



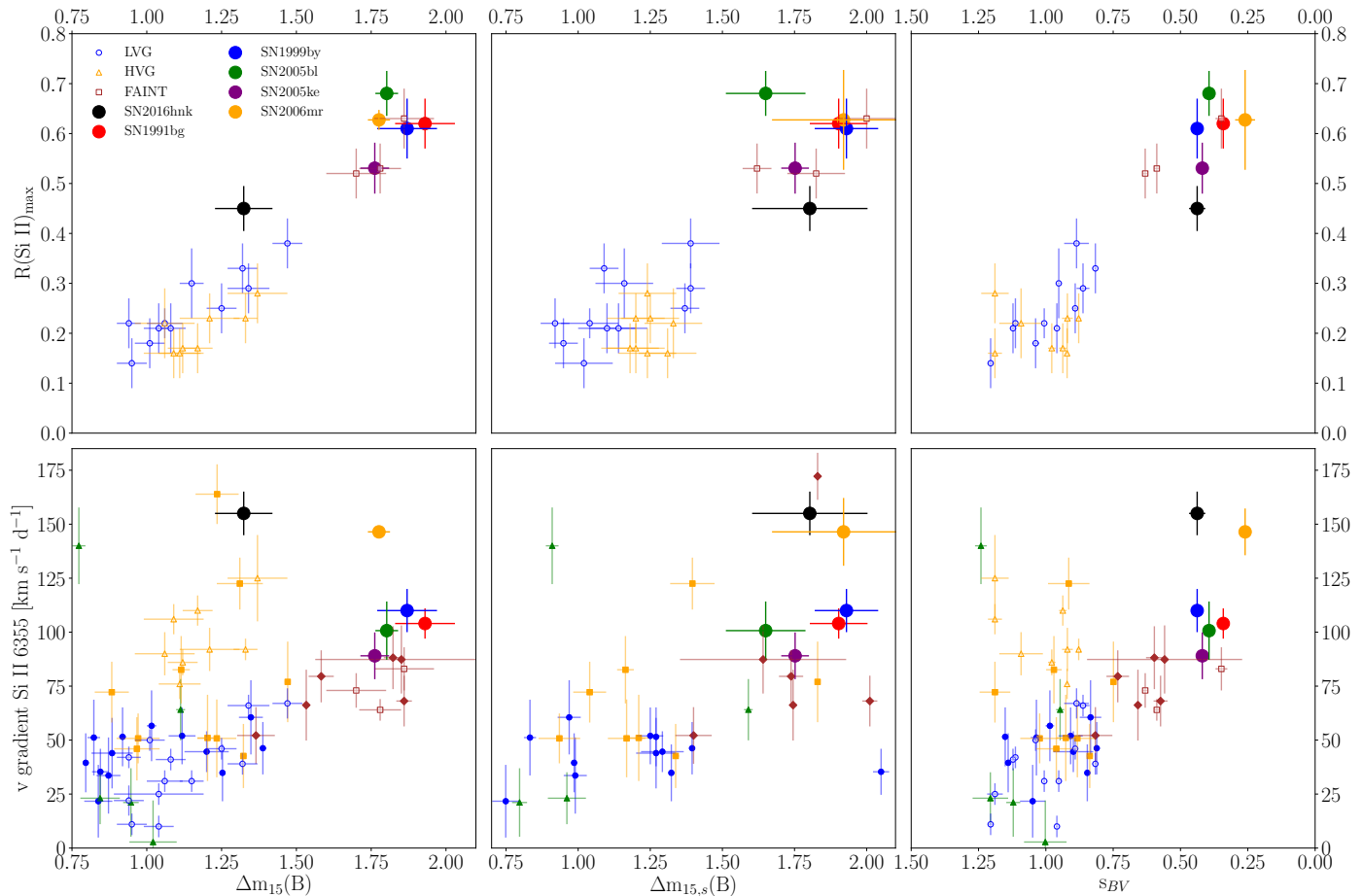
**Fig. 16.** Branch et al. (2006) and Wang et al. (2009) diagrams. The left panel has been populated with measurements from the CfA sample by Blondin et al. (2012) (open symbols) and from the CSP-I sample by Folatelli et al. (2013) (filled symbols). In the right panel only the measurements from Folatelli et al. (2013) are included. In both panels we added our own measurements of SN 2016hnk and the sample of subluminal 1991bg-like SNe Ia. SN 2002es has been added only in the left panel because with a velocity of  $\sim 5800 \text{ km s}^{-1}$  it falls outside the range shown in the right panel.

line within the group of subluminal SNe Ia: SNe 2016hnk, 1991bg, and 2006mr show larger  $pEW$  and velocities than SNe 1999by, 2005bl, and 2005ke. However, as we show in Section 6, this line is caused by a combination of O I, Mg II, and Si II, and this excess may be caused by the large concentration of Mg II in SN 2016hnk.

The  $pEW$  and velocities of the two silicon features have been widely used to identify subgroups within SNe Ia falling in different regions on a number of diagrams. For instance, Branch et al. (2006) proposed mapping SN properties using the  $pEW$ s of the two Si II features at 5972 and 6355 Å at the epoch of maximum brightness, and Wang et al. (2009) proposed a diagram of

**Table 3.** Parameters measured in this work and used in the diagrams presented in Figure 17.

Parameter	2016hnk	1991bg	1999by	2005bl	2005ke	2006mr
$v_{\text{grad}}$ [km s <sup>-1</sup> day <sup>-1</sup> ]	155.4 ± 10.1	104.0 ± 7.0	110.0 ± 10.0	100.70 ± 13.49	89.05 ± 10.83	146.45 ± 15.71
$R(\text{Si II})$	0.45 ± 0.05	0.62 ± 0.05	0.61 ± 0.06	0.68 ± 0.05	0.53 ± 0.05	0.63 ± 0.1
$\Delta m_{15}$ [mag]	1.324 ± 0.096	1.93 ± 0.10	1.87 ± 0.10	1.802 ± 0.038	1.762 ± 0.048	1.776 ± 0.037
$\Delta m_{15,s}$ [mag]	1.803 ± 0.2	1.903 ± 0.1	1.93 ± 0.11	1.650 ± 0.138	1.750 ± 0.048	1.92 ± 0.248
$s_{BV}$	0.438 ± 0.030	0.341 ± 0.014	0.438 ± 0.007	0.394 ± 0.013	0.419 ± 0.003	0.260 ± 0.004

**Fig. 17.** Benetti et al. (2005) diagrams of  $R(\text{Si II})$  and Si II velocity gradient as a function of  $\Delta m_{15}$  (left column). We replicated these diagrams using  $\Delta m_{15,s}$  (center) and  $s_{BV}$  (right column) instead as a SN light-curve width parameter. In all panels, open symbols represent measurements from the original SN sample in Benetti et al. (2005), while filled symbols are measurements from the CSP-I (Folatelli et al. 2013).

the velocity and the  $pEW$  of the same Si II  $\lambda 6355$  feature. While the latter uses the same line, the former uses two different lines, taking into account different levels of excitation. Less luminous SNe Ia have stronger Si II 5972 features at maximum light. This is caused by the recombination from Si III to Si II happening earlier in less luminous objects, and the 5972Å feature having a higher excitation level compared to the 6355Å feature, which is saturated at maximum for all SNe Ia (Hachinger et al. 2008). Furthermore, if Si lines are at low velocities the 6355Å feature can also be contaminated by Fe II  $\lambda 6516$  (see Section 6.1). Another difference between the 5972 and 6355Å Si II features is that while the range of  $pEW$  for Si 5972Å is due to differences of depth, the range of  $pEW$  in the 6355Å feature is because of width (e.g. a larger/small Si region in velocity space), as can be seen in the relation between velocity and  $pEW$  of the Wang et al. (2009) diagram.

We present both diagrams in Figure 16 with the background filled in with measurements from the CfA SN Ia sample in Blondin et al. (2012) and the CSP-I sample in Folatelli et al. (2013), and with measurements for SN 2016hnk and the comparison sample shown on top. As expected, in the Branch et al. (2006) diagram (left panel), all 1991bg-like SNe Ia fall in the COOL region, but SN 2016hnk falls in the BROAD LINE region because, although it has consistent Si II  $\lambda 6355$   $pEW$ , has a shallower Si II  $\lambda 5972$  feature compared to all other 1991bg-like SNe Ia.

Groups are a bit mixed in the Wang et al. (2009) diagram (Figure 16, right panel), with overlapping regions between COOL and NORMAL and COOL and BROAD LINE SNe Ia. Accordingly, some 1991bg-like SNe Ia fall on top of the other two groups of CSP-I SNe. In contrast to the left panel of Figure 16 (Branch et al. diagram) the location of SN 2016hnk in this diagram is more consistent with the COOL objects, having the largest Si II  $pEW$  of all objects with similar expansion veloci-

ties. While this diagram is efficient in separating high-velocity SNe (they all fall on the right side of the plot) and overluminous 1991T-like shallow silicon SNe Ia (they fall on the bottom side of the plot), it is not able to clearly separate the core normal and the cool groups, as the Branch et al. diagram does, mostly because it does not capture excitation differences.

In Figure 17 we present two additional historical diagrams from Benetti et al. (2005) (in the left column), which combine spectroscopic parameters with the light curve width parameter  $\Delta m_{15}(B)$ . In the top panels we show the  $R(\text{Si II})$  parameter, the ratio of depth of the  $\text{Si II } 5972$  over  $6355 \text{ \AA}$  features, and in the bottom panels we show the  $\text{Si II } 6355 \text{ \AA}$  velocity gradient as measured in the range from peak up to the epoch when the feature disappears. We also populated these diagrams with literature measurements from Benetti et al. (open symbols) and the CSP-I (filled symbols). Corresponding parameters for our sample of 1991bg-like SNe Ia are listed in Table 3. As before, all 1991bg-like SNe Ia fall in the region of the diagram where all other subluminal objects from the literature are located, however SN 2016hnk falls in an unpopulated region closer to NORMAL SNe Ia. Although its  $R(\text{Si II})$  would be lower but consistent with COOL SNe, the parameter that is shifting SN 2016hnk far from the COOL group is its low  $\Delta m_{15}$ .

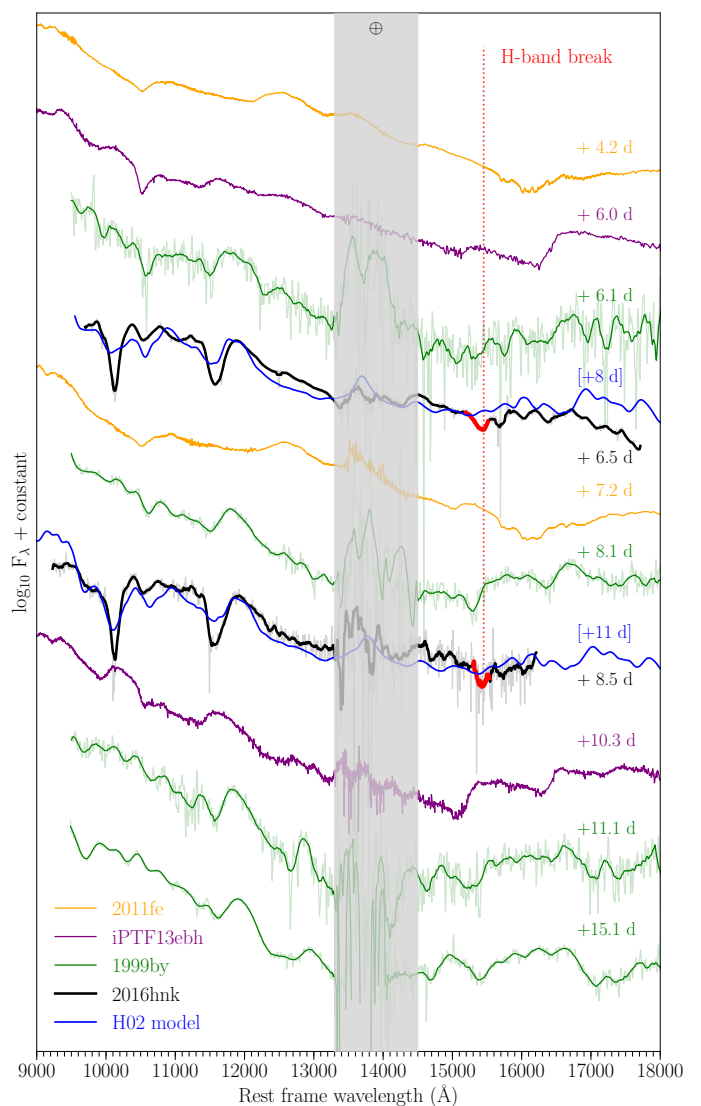
As a comparison we repeat those same two diagrams but using the  $\Delta m_{15,s}$  and  $s_{BV}$  parameters, under the hypothesis that they provide a more physically significant interpretation of the light curve width, as they measure the time at which iron recombines from  $\text{Fe III}$  to  $\text{Fe II}$  and breaks the degeneracy in the absolute magnitude vs.  $\Delta m_{15}$  (Phillips) relation. We filled the diagrams with measurements of  $\Delta m_{15,s}$  for the CSP-I from Höflich et al. (2017) and the original Benetti et al. (2005) sample measured here. Similarly, we used  $s_{BV}$  for the CSP-I sample from Burns et al. (2014) and the Benetti et al. (2005) sample measured here. In both cases, while 1991bg-like SNe Ia are kept in the region defined by the COOL objects, SN 2016hnk is also shifted towards the COOL group. It shows how the  $\Delta m_{15}$  parameter is ambiguous when the SN evolves too fast (but  $\Delta m_{15,s}$  and  $s_{BV}$  do not have this issue).

### 5.1. NIR spectra

The two NIR spectra available are presented in Figure 18 together with comparison spectra of a normal SN Ia (SN 2011fe; Mazzali et al. 2014), a transitional object (iPTF13ebh; Hsiao et al. 2015), and a SN 1991bg-like (SN 1999by; Höflich et al. 2002).

SN 2016hnk spectra show a few features between  $9,500$  and  $13,000 \text{ \AA}$ , up to the point where a telluric band starts to affect the data. According to Marion et al. (2003, 2009), the two deeper features at  $\sim 10,000$  and  $11,500 \text{ \AA}$  could be caused by  $\text{Fe II}$  but they are deeper in SN 2016hnk than in other SNe Ia. Only SN 1999by show similar shallower features at those wavelengths. However, at a week past maximum SN 1999by did not transition from the  $\text{Si/S/Ca}$  to the  $\text{Fe/Ni}$ -dominated regime (see Figure 3 in Höflich et al. 2002). SN 2016hnk shows lower velocities and the transition is faster (around 1 day), so that the  $\text{Fe}$  features in the NIR appear earlier. More detailed line identification is left for section 6 where the NIR spectra are compared in more detail to SN 1999by models from Höflich et al. (2002).

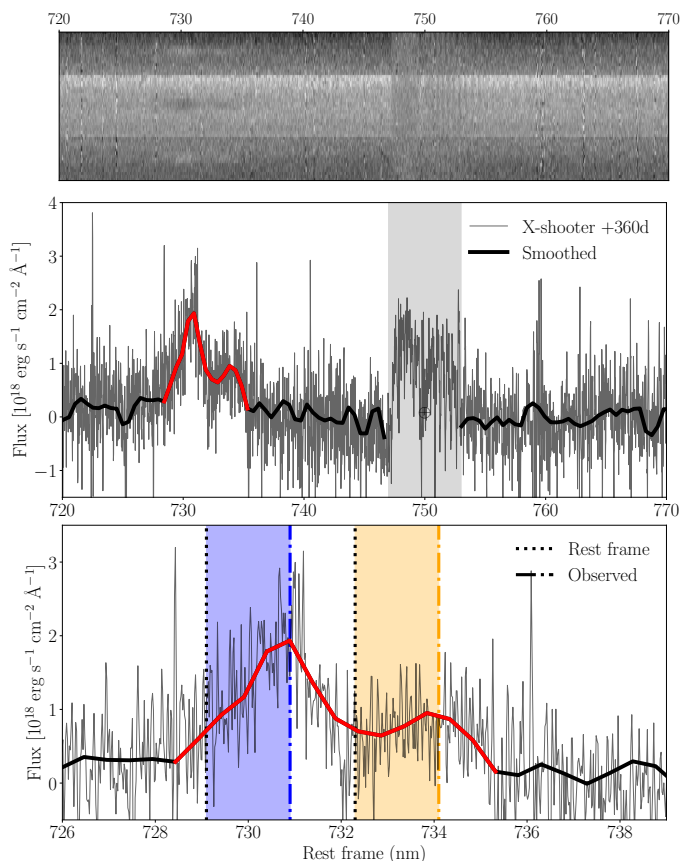
Although in the middle of the telluric feature, the  $\text{Si II } \lambda 13650$  seems to also be present at  $\sim 13,500 \text{ \AA}$ . Finally, at  $\sim 15,500 \text{ \AA}$ , we detect the  $H$ -band break, a complex formed by  $\text{Fe II/Co II/Ni II}$ . The strength of the  $H$ -band break correlates



**Fig. 18.** Near-infrared spectra of SN 2016hnk at around 6.1 and 8.5 ( $\pm 3.1$ ) days and, as a comparison, spectra of SN 2011fe (normal SN Ia), iPTF13ebh (transitional SN Ia), and SN 1999by (SN Ia 1991bg-like). We also include delayed-detonation models for SN 1999by some 24-27 days after the explosion, which corresponds to about 8-11 days past-maximum light in the  $B$  band (see section 6.2 for more details). The SN 2016hnk spectra are dominated in the  $J$  band by blends of singly-ionized iron group elements,  $\text{Ca II}$ , and  $\text{Si II}$ , while the  $H$  band appears quite flat with the  $H$ -band break marked in red. The two deep features at around  $10,000$  and  $11,500 \text{ \AA}$  are only seen shallower in SN 1999by spectra.

with  $\Delta m_{15}(B)$ , with brighter SNe Ia having the strongest  $H$ -band break (Hsiao et al. 2013), and the outer edge of the velocity of this region is correlated with light curve shape (Ashall et al. 2019b). The outer-edge of this feature has a velocity of  $v_{\text{edge}} = 5,380 \pm 490 \text{ km s}^{-1}$  as measured in the SN 2016hnk spectrum at +6.5 days. Sub-Chandrasekhar mass explosions are expected to have  $v_{\text{edge}}$  values of around  $7,000 \text{ km s}^{-1}$ , while  $M_{\text{Ch}}$  models of sub-luminous SN Ia have lower values, around  $\sim 5,000 \text{ km s}^{-1}$  (Ashall et al. 2019a). So, in this picture, SN 2016hnk is compatible with high mass, near  $M_{\text{Ch}}$  models.





**Fig. 19.** Top: 2D frame of the combined X-shooter nebular phase (+360 d) spectrum of SN 2016hnk after sky subtraction, around the region where we detected the only feature. The trace can be clearly seen around 730 nm. The other feature around 750 nm is due to a bad telluric subtraction, since it is present along the spatial direction (vertical). Middle: The combined 1D X-shooter optical spectrum showing the only two features detected in the whole UV+Opt+NIR wavelength range. Bottom: Zoom-in to the feature we identified as being due to calcium emission. The blue and orange shaded regions represent the shift of  $\sim 700 \text{ km s}^{-1}$  between the peak of the emission and the rest frame wavelength.

## 5.2. Nebular spectrum

The resulting stack of the six individual exposures results in a low signal-to-noise spectrum, at an average epoch of +360 days post-explosion. Carefully scanning all regions from 3000 to 25,000 Å, we find that the only visible feature is a double emission at  $\sim 7300 \text{ Å}$ , as reproduced in Figure 19, which we associate to the  $3 \rightarrow 1$  and  $2 \rightarrow 1$  transitions of the Ca II doublet at 7,291.5 and 7,323.9 Å, respectively. The peaks of the emission are redshifted by  $\sim 740 \text{ km s}^{-1}$ . In the upper panel we show a zoom into the 2D spectra where the trace of the SN light is clearly seen at those wavelengths. A poor telluric correction is responsible for the feature around 7,500 Å, as the 2D spectrum unequivocally demonstrates. Although similar narrow [Ca II] features have been previously seen in 2002cx-like SN Ia spectra (e.g. SN 2005hk, SN 2008A and SN 2010ae; McCully et al. 2014b; Foley et al. 2016; Stritzinger et al. 2014), and are possibly also present in the 2002es-like SN Ia spectrum of SN 2010lp (Taubenberger et al. 2013; see Figure 20), the SN 2016hnk spectrum does not show either the narrow [Fe II]  $\lambda 7155$  and [Ni II]  $\lambda 7378$ , nor the broad [Fe II]  $\lambda 7155$  accompanying features, respectively. In addition, the lack of other features at bluer wavelengths can be attributed

to the low S/N and high reddening that is affecting the SN. We provide more details in Section 6.3.

## 6. Spectral modelling

To get insights on the possible progenitor system and explosion mechanism that explain our observations we turn into modelling of the observed spectra. In section 6.1, we want to evaluate spectral line identifications and spectral fits using the method of abundance stratification. This approach is rather independent from the explosion scenario because the overall envelope structures of most explosions scenarios are similar, and nuclear physics dominates the specific explosion energy.

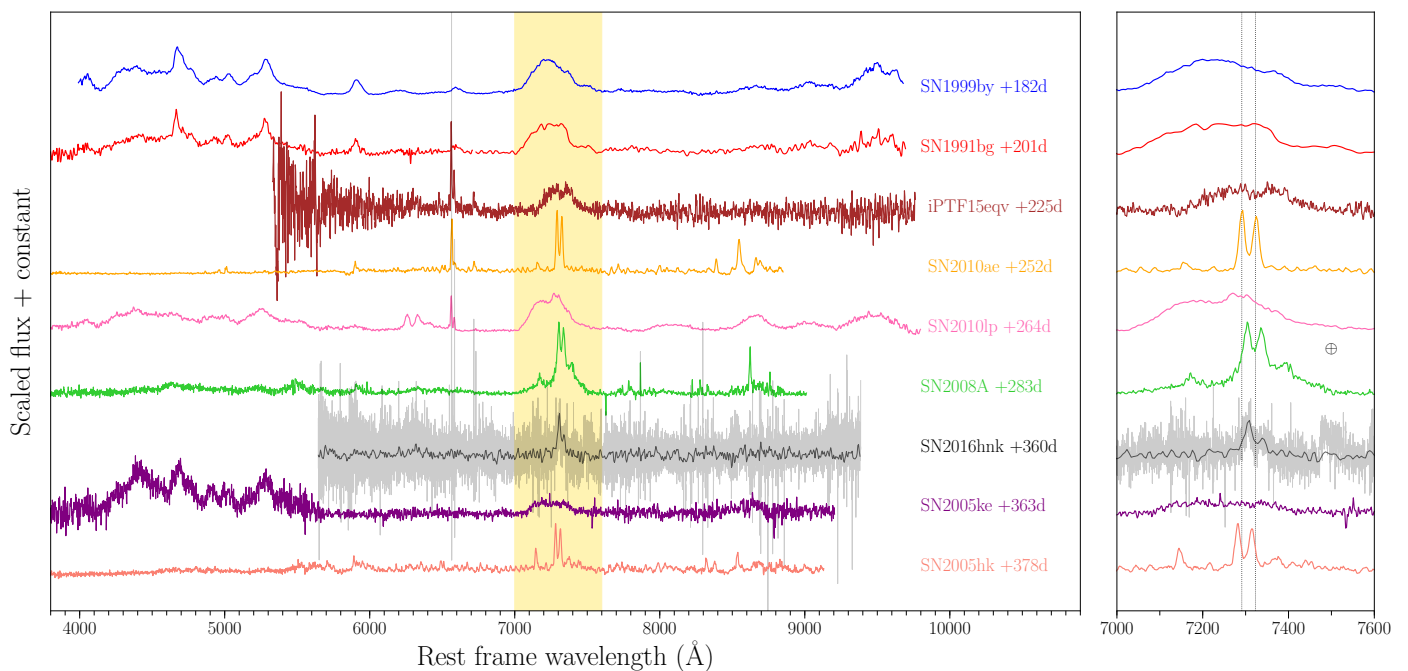
In the subsequent sections, we use detailed time-dependent non-LTE models. Based on results of the abundance stratification method and guided by the observed LCs and spectra during the photospheric phase, we identify an explosion model from literature which matches the observations of SN 2016hnk. This turns out to be the non-LTE model for SN 1999by from Höflich et al. (2002). Abundance stratification does not allow us to fit NIR spectra, due to the assumption that the photosphere can be approximated as black body. Therefore, for verification, we check that the non-LTE model mentioned above can reproduce the NIR spectrum of SN 2016hnk, without further tuning of the parameters, we also identify the features produced in the NIR (Ca I, Mg II and other IME lines). The non-LTE models are then used to revisit some of the line identifications of the previous section, and compared to the abundance stratification results.

In section 6.3, we modify our explosion model with the goal to reproduce the narrow [Ca II] feature during the nebular phase. We change the initial conditions of the WD and calculate the initial phase of the expansion up to about 10 days (Höflich et al. 2017). We also construct detailed non-LTE nebular spectra which include gamma- and positron transport but we neglect radiative and recombination time-scales. We discuss why energy input and transport by radioactive  $^{56}\text{Ni}$  decay is important in the central region.

### 6.1. Abundance stratification

The abundance stratification technique utilizes the fact that a SN Ia is in homologous expansion ( $v_{\text{ph}}(r) \propto r$ ) from  $\sim 10$  s after explosion (Röpke 2005). With a given density profile, time from explosion ( $t_{\text{exp}}$ ), photospheric velocity ( $v_{\text{ph}}$ ), bolometric luminosity ( $L_{\text{bol}}$ ), and abundance structure, optimally fitting synthetic spectra are produced. This approach has been used for many SNe including subluminous and transitional SNe Ia such as SNe 2005bl, 1986G, 2007on, and 2011iv (e.g. Hachinger et al. 2009; Ashall et al. 2016b, 2018), normal SNe Ia (e.g. Tanaka et al. 2011; Ashall et al. 2014; Mazzali et al. 2014), as well as stripped envelope SNe (e.g. Ashall et al. 2017; Prentice et al. 2018). A brief description of the code can be found in Appendix B.

The code requires an input density profile, in the case of SN 2016hnk we used the W7 density profile (Nomoto et al. 1984). The W7 density profile looks similar to delayed detonation models of sub-luminous SNe Ia from Höflich et al. (2017). This density profile comes from a fast deflagration explosion of a Ch-mass WD. Spectral models were produced between +2.5 to +7.5 days relative to *B*-band maximum, to obtain the broad structure and chemical composition of the outer layers of the ejecta. We used a rise time of 16.5 days, which gives the best fit and it is consistent with the value estimated for SN 2016hnk



**Fig. 20.** Comparison of SN 2016hnk nebular spectra (in black) to other 1991bg-like, 2002es-like, 2002cx-like, and Ca-rich objects. While all these objects show Ca emission at late times, only SN 2002cx-like (and possibly 2002es-like; e.g. SN 2010lp) objects show the narrow emission lines similar to SN 2016hnk. The yellow shaded region is zoomed in on the right panel, where the vertical dotted lines represent the restframe wavelength of the two [Ca II] emission lines. We note that while SN 2005hk lines have negative velocities, SN 2008A lines have positive velocities similarly to SN 2016hnk.

**Table 4.** Input parameters for the early-time spectral models.

$t_{\text{exp}}$ days	$t_{\text{Bmax}}$ days	$v_{\text{ph}}$ $\text{km s}^{-1}$	$L_{\text{bol}}$	$T_{\text{BB}}$ K	C %	O %	Mg %	Al %	Si %	S %	Ca %	Ti+Cr %	Fe %
-	-	15000	-	-	45.00	54.00	0.40	0.00	0.40	0.14	$6 \times 10^{-2}$	$2 \times 10^{-3}$	$2 \times 10^{-3}$
-	-	9400	-	-	9.00	83.75	4.00	0.25	2.00	0.92	$6 \times 10^{-2}$	$2 \times 10^{-3}$	$2 \times 10^{-2}$
-	-	8900	-	-	0.00	21.00	3.00	0.25	60.00	15.19	$6 \times 10^{-2}$	$2 \times 10^{-3}$	0.50
19.0	2.5	6500	8.49	7900	0.00	6.64	0.00	0.16	68.00	24.00	$5 \times 10^{-2}$	$1 \times 10^{-2}$	1.14?
21.8	5.3	5000	8.48	8400	0.00	0.00	0.00	0.00	70.00	28.74	$5 \times 10^{-2}$	$1 \times 10^{-2}$	1.20
24.0	7.5	4700	8.44	7900	0.00	0.00	0.00	0.00	72.00	26.74	$5 \times 10^{-2}$	$1 \times 10^{-2}$	1.20

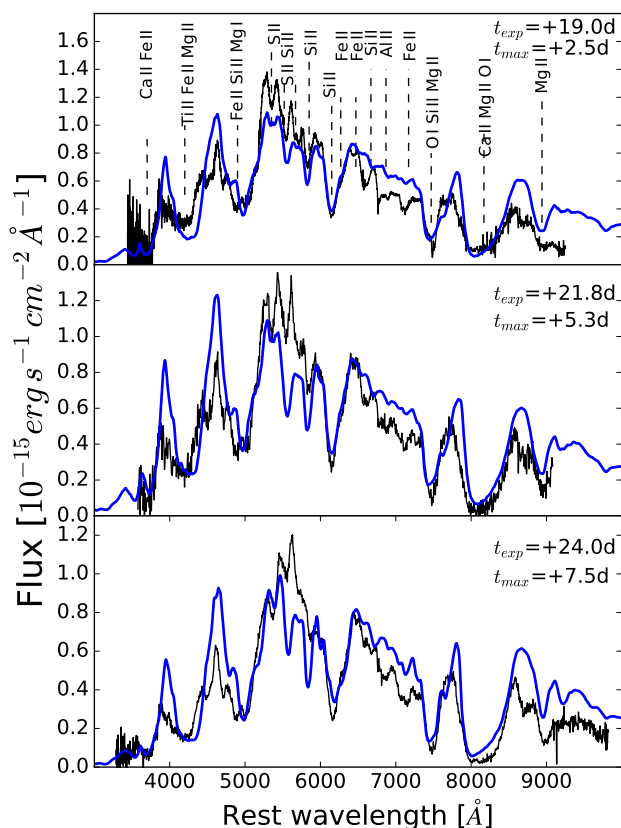
from non-detections in Section 2. Figure 21 shows the synthetic spectra and models at three epochs (corresponding to +19.0, +21.8, +24.0 days from explosion). All observed spectra have been first corrected for host galaxy and foreground galactic extinction. The strongest lines contributing to each feature are labeled in the top panel of the plot. The bluest part of the spectra ( $<3900 \text{ \AA}$ ) is dominated by blends of many Fe group and metal lines, which cause a large amount of line blanketing and reprocesses the flux into the optical. However, the overall red shape of the spectrum is intrinsic and due to SN 2016hnk being less luminous and cooler than a standard object.

The strongest lines in the  $\sim 3900 \text{ \AA}$  feature are the Ca II H&K lines  $\lambda\lambda 3934, 3969$  with some contribution from Ti II  $\lambda\lambda 3901, 3913$ . The feature at  $\sim 4200 \text{ \AA}$  is dominated by Ti II  $\lambda\lambda 4395, 4418, 4444, 4501$ , Fe II  $\lambda\lambda 4178, 4233, 4297, 4303, 4352, 4417, 4491$ , and Mg II  $\lambda 4481$ . Moving redward the feature at  $\sim 5000 \text{ \AA}$  consists of Fe II  $\lambda\lambda 4924, 5018, 5169, 5198, 5235, 5276, 5317$ , Si II  $\lambda\lambda 5041, 5056$ , and Mg I  $\lambda 5184$ . The notch at  $\sim 5300 \text{ \AA}$  is produced by S II  $\lambda\lambda 5432, 5454$ , and Si II  $\lambda\lambda 5466, 5467$ . The absorption at  $\sim 5450 \text{ \AA}$  consists of S II  $\lambda\lambda 5606, 5639$  and Si II  $\lambda\lambda 5670, 5707$ , although note that in the data there could be an emission line in this

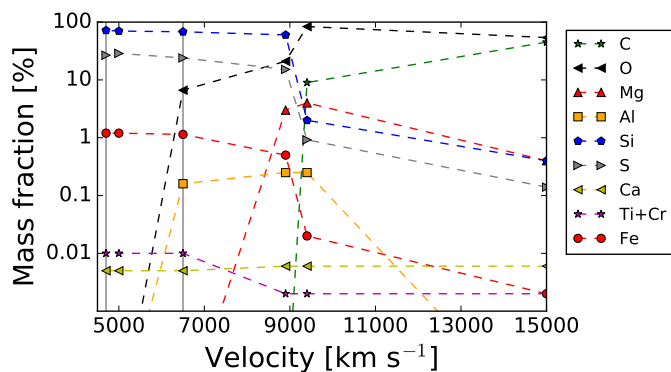
area of the spectra. The feature at  $\sim 5800 \text{ \AA}$  is produced by Si II  $\lambda\lambda 5958, 5979$ , and the small absorption between the two Si II features is made by Fe II  $\lambda\lambda 6147, 6150, 6238$ . The absorption at  $\sim 6200 \text{ \AA}$  is produced by Si II  $\lambda\lambda 6347, 6371$ , with the unusual red side of the feature being dominated by Fe II  $\lambda\lambda 6456, 6516$ , which gives it a broader asymmetrical shape. This is seen in most sub-luminous SNe Ia (see Figure 14). However, for brighter objects, where the Si region is further out in velocity space, this Fe II  $\lambda 6516$  line is on the emission component of the Si II  $\lambda 6355$  feature, as the velocity of the Fe and Si-rich regions may not scale in by the same factor. In the latter case it may be easy to misidentify this feature as C II.

Between 6200 and 7700  $\text{\AA}$  there are a number of small absorption features not typically seen in brighter SNe Ia. Many of these are low excitation Fe II lines that are produced at lower temperatures, which is consistent with the low luminosity of SN 2016hnk. The strongest lines in the models between the Si II 6355  $\text{\AA}$  feature and the O I 7774 features are Fe II  $\lambda\lambda 6456, 6516, 7308, 7462, 7711$ , Si II  $\lambda\lambda 6818, 6830$ , and Al II  $\lambda\lambda 7042, 7057$ <sup>5</sup>. The feature at  $\sim 7770 \text{ \AA}$  is dominated by

<sup>5</sup> Note that we doubt our the Al II line identification here as it requires an abundance that is larger than what is predicted by explosion models.



**Fig. 21.** A time series of data (black) and spectral models (blue) for SN 2016hnk during the photospheric phase. The data have been corrected for foreground and host galaxy extinction. The observer-frame time from  $B$  band maximum and the rest-frame time from explosion is provided for each model. Line identifications from the model are shown in the top plot.



**Fig. 22.** The abundance distribution as a function of velocity for SN 2016hnk. The vertical dashed lines denote the photosphere velocity of the +2.5 and +7.5d models. We note that due to the small time step between our first and last model our results are only sensitive between 4500-10000km s<sup>-1</sup>.

O I  $\lambda\lambda 7772, 7774, 7775$  with strong contributions from Si II  $\lambda\lambda 7849, 7850$ , and Mg II  $\lambda\lambda 7877, 7896, 7896$ . The deep feature at 8200 Å, often attributed to only Ca II, is in fact produced by Ca II  $\lambda\lambda 8498, 8542, 8662$ , O I  $\lambda\lambda 8446, 8446, 8447$ , and Mg II  $\lambda\lambda 8213, 8235$ . It is important to note that the blue side of this feature consists of Mg II, Ca II, and O I, however Mg II is not one of the main sources of opacity in the feature and its contribution is minor. Finally the absorption at 8900 Å is produced by

Mg II  $\lambda\lambda 9218, 9244$ . Throughout the spectra modeled in Figure 21 the line identification does not change significantly, it is only the strength of each line that varies.

Table 4 shows the abundances as a function of velocity that were used to compute the synthetic spectra. We have also provided the abundances in the shells outside the formation region of the first spectrum. These shells are required to produce the stratified abundances and to match the observations. For example, if we assume the abundances in the 6500 km s<sup>-1</sup> shell are the same as the very outer layer, the models would contain high velocity Si absorption. This would be caused by the ejecta at high velocities containing 68% Si.

In the first epoch the modeled  $v_{\text{ph}}$  is 6500 km s<sup>-1</sup>, the abundances near the photosphere are dominated by Si (~68%), S (~24%), and C+O (7%), with Ca, Ti, Cr, V, and Fe making up the remaining 1%. The next epoch has a  $v_{\text{ph}}$  of 5000 km s<sup>-1</sup>, the Si and S abundances have increased slightly (~70% and ~29%, respectively). This model has no C or O at the photosphere, with Ca, Ti, Cr, V, and Fe making up the rest of the mass. By the third epoch the  $v_{\text{ph}}$  is 4700 km s<sup>-1</sup> and the Si abundance is ~72%, S is ~27% and the remaining mass consists of Ca, Ti, Cr, V, and Fe. For a summary of errors using this method see Mazzali et al. (2008).

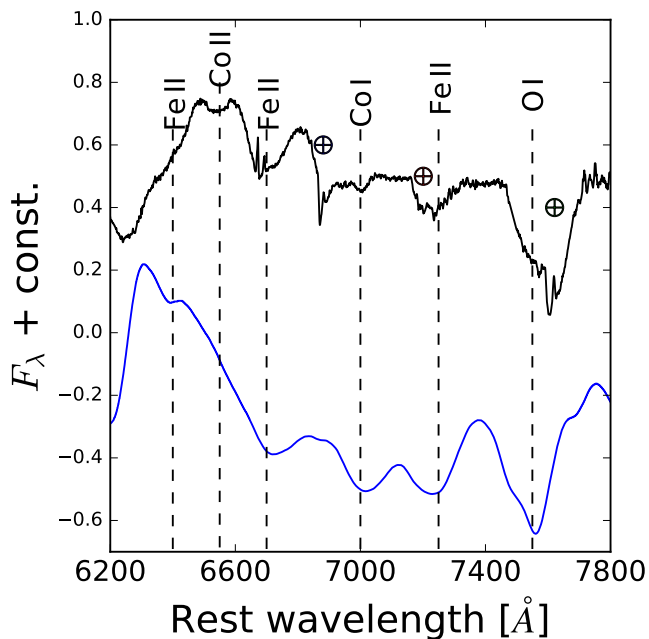
The derived abundance distribution obtained through spectral modeling is consistent with a low luminosity delayed-detonation SN Ia. Our models are sensitive out to velocities of ~15000 km s<sup>-1</sup>. Intermediate mass elements dominate the spectra at the epochs observed, however the presence of significant O is required above 9000 km s<sup>-1</sup>. It can be assumed that the bulk of the <sup>56</sup>Ni is located below our innermost shell. There could be a small abundance of <sup>56</sup>Ni above the innermost photosphere, but it is not large enough to be optically thick. There is less effective burning in this SN explosion compared to a more luminous event.

In these models at +24 d past explosion there are 0.8 $M_{\odot}$  above the photosphere. As there is a large amount of material below the photosphere, it is not possible for us to constrain the total mass of the explosion. However the velocity of the  $H$ -band break (see above) and the nebular phase spectrum will help, see Section 6.2.

A comparison of Figure 22 to the abundance distributions of other transitional and sub-luminous SNe Ia show similar results. However, for SN 2016hnk we find no evidence of <sup>56</sup>Ni in our photospheric models, whereas in the sub-luminous SN 2005bl Hachinger et al. (2009) found <sup>56</sup>Ni out to ~7000km s<sup>-1</sup>. A comparison to the abundance distribution to the transitional SN 1986G (Ashall et al. 2016b) shows similar results to SN 2016hnk, with the only discrepancy being that SN 1986G has <sup>56</sup>Ni located out to 5500km s<sup>-1</sup>. As discussed below, we note that our abundance distribution is similar to the delayed detonation model of SN 1999by of Höflich et al. (2002).

## 6.2. Explosion models photospheric phase

We calculate the final phase of the thermonuclear runaway, the explosion, LCs and spectra with our code HYDRA Höflich et al. (2017), and references therein. Here, we use delayed detonation models Khokhlov (1989) in which the explosion starts as a slow laminar deflagration. The free parameters are the main sequence mass  $M_{\text{MS}}$  and metallicity  $Z$  of the progenitor, the central density of the explosion  $\rho_c$ . The flame starts as laminar deflagration and transitions to a detonation at a density  $\rho_{\text{tr}}$  which, in spherical geometry is equivalent to the amount of burning during the deflagration phase. To setup the explosion, we calculate the final



**Fig. 23.** Comparison of SN 2016hnk (black) +2.5 d past maximum light and the non-LTE model of SN 1999by at +7 days past maximum light (blue) from Höflich et al. (2002). The telluric features are marked in the observed spectra.

phase of the accretion towards the thermonuclear runaway with our code but neglect the second derivatives to allow for large time steps similar to Nomoto (1982). Here, the constant accretion rate is adjusted so that the thermonuclear runaway occurs at a given central density  $\rho_c$  (Höflich et al. 2017). To first order, the explosion is triggered when compressional heating due to increasing mass by accretion from the companion star cannot be balanced by heat conduction and neutrino cooling (including the convective URCA process). Close to the AIC, the detailed parameters such as the accretion rate depends sensitively on the equation of state and neutrino cooling, and the history of the binary evolution (see final discussion and conclusions).

As mentioned in Sect. 6.1, SN 2016hnk has many observational characteristics similar to the subluminous SN 1999by. To produce sub-luminous SNe Ia in the framework of DDT models, the pre-expansion of the WD needs to be sufficiently large during the deflagration phase, that the density has dropped enough so that little  $^{56}\text{Ni}$  is produced during the subsequent detonation phase, and we need a partial suppression of Rayleigh-Taylor (RT) mixing (Höflich et al. 2002, 2017; Hristov et al. 2018). The model for SN1999by had  $\rho_c = 2 \times 10^9 \text{ g/cm}^3$ ,  $\rho_{tr} = 8 \times 10^6 \text{ g/cm}^3$ , and  $MS = 5M_\odot$ ,  $Z = Z_\odot$ . Our tomography approach for the early spectra resulted in an abundance structure that closely resembled the one of the delayed-detonation model for SN1999by which shows that the NSE elements are within  $\approx 4500 \text{ km/sec}$ , followed by a Si/S-rich layer and a transition to layers of explosive carbon burning (Mg/O) at  $\approx 9000 \text{ km/sec}$ .

NIR observations provide a useful constraint on the explosion because its sensitivity to whether the photosphere is formed in a mostly C/O rich mixture, IME elements, or NSE elements. As such, it provides a useful probe for the time of the explosion, and the explosion model. In Figure 18 a NIR spectrum of SN 2016hnk is compared to the non-LTE delayed-detonation model of SN 1999by at 26 days past explosion, corresponding to about 11 days past maximum. Overall, the observed and the-

oretical spectra agree. The theoretical spectrum is dominated by blends of singly ionized Si, Ca, Fe, and Co. The strongest features are Fe II ( $\approx 1.02\mu\text{m}$ ), Si II ( $\approx 1.20\mu\text{m}$ ), Fe II ( $\approx 1.06\mu\text{m}$ ), Ca II, Fe II and Co II ( $1.16\mu\text{m}$ ), Co II and Fe II ( $1.15$  and  $1.38\mu\text{m}$ ), as well as a multiplet of Fe II/Co II/Ni II which is weak but appears in the *H*-band region. Note that the feature at  $0.98\mu\text{m}$  is a Fe II dominated blend of Ca II/Fe II/Co II. In the model it is more defined or in the observation the suppression of flux is stronger. Because it is formed in a steeply declining wing, it is ambiguous whether the metallicity,  $Z$ , of SN 2016hnk is significantly larger than  $Z_\odot$  in the model leading to a suppression of flux and a widening of the adjoining feature to form a shoulder, or if  $Z$  is significantly smaller than  $Z_\odot$  resulting in a narrower feature. This strongly places SN 2016hnk into the realm of the sub-luminous 1991bg-like SNe Ia and emphasizes the NIR as a chronometer for the time since maximum because of the rapid spectral changes. The time of this model is consistent with our lightcurve-based estimates for maximum light, as well as the abundance stratification models.

The features between  $6000\text{--}7500 \text{ \AA}$  in the photospheric phase spectra are not normally seen in bright SNe Ia. With the abundance stratification models we produce many of these lines but at different flux levels. Therefore, to see if our line identifications are correct, we compare SN 2016hnk to the full non-LTE models of SN 1999by from Höflich et al. (2002).

To conclude the photospheric section, the element identification should be revisited in light of the explosion based non-LTE models. Figure 23 shows the  $+2.5 \pm 3.3 \text{ d}$  spectra of SN 2016hnk and the non-LTE model of SN 1999by at +7 d in the wavelength region  $6200\text{--}7800 \text{ \AA}$ . Note that we have limited our analysis to the published spectra from Höflich et al. (2002). We note that in the model at +2 d these lines have not yet formed, but given the error on the time of maximum from the observations and the quick evolution of SN 2016hnk, we speculate that these lines begin to emerge in the models of SN 1999by between +2 and +7 d, which is consistent with the observations and NIR models. The model and data have similar spectral features, shown with the vertical dashed lines. The line identification is mainly consistent with the abundance stratification models, where Fe II lines dominate the spectra in the  $6000\text{--}7500 \text{ \AA}$  region.

Some of the other features in the  $6200\text{--}7800 \text{ \AA}$  region are Fe II  $\lambda 6664$  and Co I  $\lambda 7084$ . The one feature which cannot be explained by the models is the small absorption at  $6550 \text{ \AA}$ . However, the Höflich et al. (2002) models assume no mixing. With a small amount of mixing it could be expected that this feature is caused by Co II  $\lambda \lambda 6571, 6576$ , as mixing will enhance the Co abundance in these layers. In the wavelength region  $6000\text{--}7500 \text{ \AA}$  we favour the line identifications from the non-LTE models, as these solve for atomic level populations explicitly and are produced through full radiation-hydrodynamical simulations and not a parameterized to fit the data.

### 6.3. Explosion models for the nebular phase

Due to the similarity of SN 2016hnk to SN 1999by, we have adapted the previous  $M_{\text{Ch}}$  mass spectral simulations of SN 1999by (Höflich et al. 2002, 2017) to produce models for SN 2016hnk. In SN 2016hnk an interesting observation in the nebular spectrum is the appearance of a single, narrow dominating spectral feature at  $\approx 7200 \text{ \AA}$  that has a Doppler width of  $\approx 1000 \text{ km s}^{-1}$ . This is very different from late-time spectra of normal bright SNe Ia that are dominated by iron-group elements. Because the delayed-detonation model for SN 1999by does not

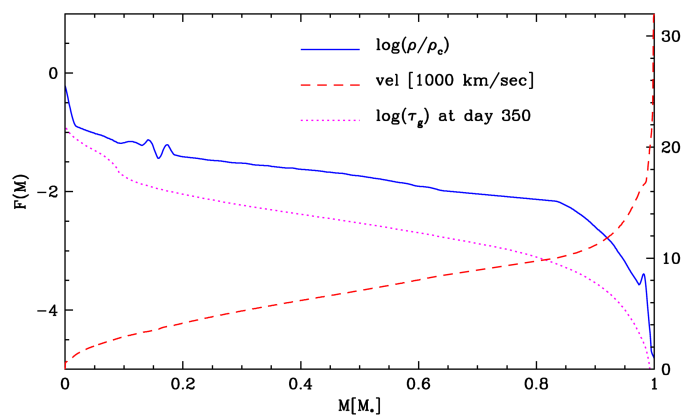
produce narrow features on less than a  $2000\text{--}3000\text{ km s}^{-1}$  scale, the inner region of the original model needs to be modified.

We note that, in principle, multidimensional effects such as Rayleigh-Taylor (RT) instabilities produce plumes with an individual velocity spread on the scale of  $\approx 1000\text{ km s}^{-1}$ . However they produce multiple plumes which would make the total velocity spread too large to explain the observation of SN 2016hbk. Moreover, various observations and analyses strongly suggest that the RT instabilities are present but need to be partially suppressed (Höflich et al. 2006; Fesen et al. 2015; Höflich et al. 2017; Diamond et al. 2018; Hristov et al. 2018). Therefore, in this work spherical delayed-detonation models are used, and we focus on nuclear physics based effects in combination with progenitor parameters to explain the observations of SN 2016hbk. The hydrodynamical simulations are in the comoving (mass) frame to avoid numerical advection of abundances, which may affect both nuclear burning and small scale structures. The simulations utilize a nuclear network of 315 isotopes during the early phase of the explosion based on the nuclear reaction network using the library REACLIB (Cyburt et al. 2010) and weak rates from Martínez-Pinedo et al. (2014), Langanke & Martínez-Pinedo (2003), and Fuller et al. (1985). We use detailed, time-dependent non-LTE models for atomic level populations, including  $\gamma$ -ray and positron transport and radiation-hydrodynamics (RH; Höflich 1995, 2003; Penney & Höflich 2014; Höflich et al. 2017). After an initial deflagration phase, a detonation is triggered when the density at the flame front drops to a transition density,  $\rho_{\text{tr}}$ , which is a parameterization of the amount of deflagration burning.

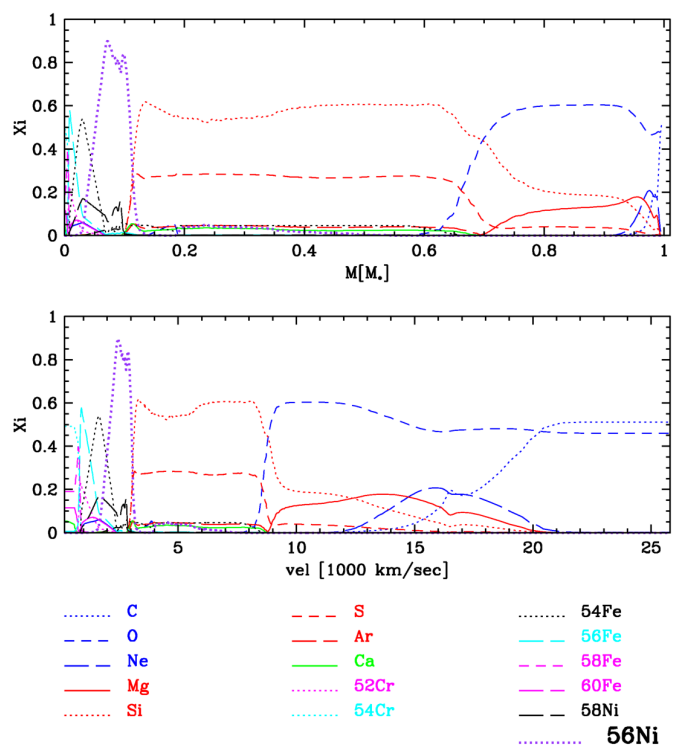
The model presented originates from a progenitor with a main sequence mass ( $M_{\text{MS}}$ ) of  $7 M_{\odot}$  with solar metallicity. The models of SN 2016hbk require a higher central density ( $\rho_c = 6 \times 10^9\text{ g cm}^{-3}$ ) than that of SN 1999by, which implies a larger binding energy, and to compensate for the longer phase to lift the WD, we used a slightly higher  $\rho_{\text{tr}} = 10^7\text{ g cm}^{-3}$ . The rate of subluminous SNe Ia diminish at high redshifts indicating a long evolutionary time scale (González-Gaitán et al. 2011). Thus, we assume an increased  $^{22}\text{Ne}/^{20}\text{Ne}$  ratio ( $\approx$  factor of 2) in the center of the WD due to gravitational settling over billions of years (Deloye & Bildsten 2002). Note that most of the Ne is inherited from the formation of the progenitor star in the molecular cloud. In effect, our model starts at a slightly lower  $Y_e$  compared to models without gravitational settling. The front starts with a laminar speed. To increase the electron capture, we assume a slow deflagration front ( $C1 = 0.1$  of Eq. 1 in Domínguez & Höflich 2000).

The density and velocity structure of the model during the homologous expansion, after  $\approx 10$  days, is shown in Figures 24 and 25. The density structure is very similar to the model used for SN 1999by in most respects. As in Höflich et al. (2002) models and also in W7 (Nomoto et al. 1984), a local maximum in the outer density structure,  $M \approx 0.98M_{\text{WD}}$ , marks the transition from burned to unburned material. The feature at  $M \approx 0.15M_{\text{WD}}$  marks the imprint of the refraction wave during the deflagration to detonation transition, which is inherent to spherical models. The increased density in the central region of the explosion<sup>6</sup> is not present in the Höflich et al. (2002) models, as it is caused by radiation hydrodynamical effects. After the explosion, the envelope enters the phase of ‘almost’ free expansion on time scales of seconds to minutes. Adiabatic cooling results in a rapid drop of temperature in the envelope and an overall free expansion. This explains the similarity in the outer layers between the new model

<sup>6</sup> Note this is different from the central density of the exploding WD.



**Fig. 24.** Structure during the phase of homologous expansion as a function of the mass coordinate in WD-masses for our subluminous model with high central density. We give the normalized density as  $\log(\rho/\rho_c)$  (left scale), velocity (right scale) and the mean optical depth  $\log(\tau_\gamma)$  (left scale) in  $\gamma$ -rays at day 350.



**Fig. 25.** Same as Figure 24 but showing the overall abundance structure at large times as a function of mass coordinate (top) and expansion velocity (bottom). It is very similar to the subluminous model for SN 1999by and consistent with the chemical structure based on our abundance tomography given in Table 3. In addition, we show the initial  $^{56}\text{Ni} \rightarrow ^{56}\text{Co} \rightarrow ^{56}\text{Fe}$  at the time of the explosion (purple dotted line).

and the old SN 1999by model. However, the energy production by radioactive  $^{56}\text{Ni}$  decay corresponds to a change of the specific energy equivalent to a velocity of  $\approx 3000\text{ km s}^{-1}$ . This effect changes the density structure of low-velocity layers because the ‘dynamics’ is dominated by pressure equilibrium. In models with lower  $\rho_c$ , the central region becomes partially optically thin for  $\gamma$ -rays and low energy photons. Thus, the entire inner region is affected by the heating. However, in our high-density model with  $\rho_c = 6 \times 10^9\text{ g cm}^{-3}$ , the average  $\gamma$ -ray optical depth  $\tau_\gamma$  stays larger than 2–3, even at twice the decay time of  $^{56}\text{Ni}$ ,

**Table 5.** Asymptotic isotopic yields at large times for the delay-detonation model for SN 2016hnk discussed in Section 6.2. The mass is given in units of  $M_{\text{Ch}}$ .

$^{12}\text{C}$	$8.109 \times 10^{-3}$	$^{16}\text{O}$	$1.891 \times 10^{-1}$	$^{20}\text{Ne}$	$7.443 \times 10^{-3}$
$^{22}\text{Ne}$	$1.351 \times 10^{-4}$	$^{23}\text{Na}$	$7.526 \times 10^{-5}$	$^{24}\text{Mg}$	$3.371 \times 10^{-2}$
$^{25}\text{Mg}$	$9.375 \times 10^{-5}$	$^{26}\text{Mg}$	$1.273 \times 10^{-4}$	$^{27}\text{Al}$	$1.503 \times 10^{-3}$
$^{28}\text{Si}$	$3.943 \times 10^{-1}$	$^{29}\text{Si}$	$1.492 \times 10^{-3}$	$^{30}\text{Si}$	$3.393 \times 10^{-3}$
$^{31}\text{P}$	$7.924 \times 10^{-4}$	$^{32}\text{S}$	$1.669 \times 10^{-1}$	$^{33}\text{S}$	$5.248 \times 10^{-4}$
$^{34}\text{S}$	$3.415 \times 10^{-3}$	$^{35}\text{Cl}$	$2.136 \times 10^{-4}$	$^{37}\text{Cl}$	$3.536 \times 10^{-5}$
$^{36}\text{Ar}$	$2.413 \times 10^{-2}$	$^{38}\text{Ar}$	$1.477 \times 10^{-3}$	$^{39}\text{K}$	$1.217 \times 10^{-4}$
$^{40}\text{Ca}$	$1.597 \times 10^{-2}$	$^{42}\text{Ca}$	$3.616 \times 10^{-5}$	$^{48}\text{Ca}$	$8.148 \times 10^{-4}$
$^{46}\text{Ti}$	$2.200 \times 10^{-4}$	$^{48}\text{Ti}$	$2.358 \times 10^{-4}$	$^{49}\text{Ti}$	$1.098 \times 10^{-4}$
$^{47}\text{V}$	$2.826 \times 10^{-3}$	$^{48}\text{V}$	$2.076 \times 10^{-4}$	$^{51}\text{V}$	$1.183 \times 10^{-2}$
$^{50}\text{Cr}$	$3.032 \times 10^{-3}$	$^{52}\text{Cr}$	$3.272 \times 10^{-3}$	$^{53}\text{Cr}$	$2.318 \times 10^{-3}$
$^{55}\text{Mn}$	$7.401 \times 10^{-4}$	$^{54}\text{Fe}$	$1.572 \times 10^{-3}$	$^{56}\text{Fe}$	$1.045 \times 10^{-1}$
$^{57}\text{Fe}$	$1.710 \times 10^{-3}$	$^{58-64}\text{Ni}$	$1.170 \times 10^{-2}$		

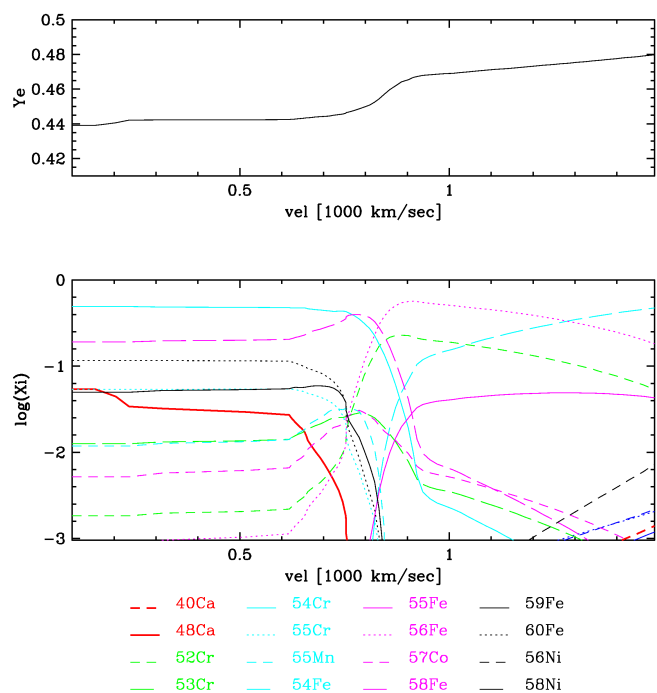
leading to a compression of layers with  $\lesssim 1,500 \text{ km s}^{-1}$  (Figure 24). Note that these layers are well inside the  $^{56}\text{Ni}$  layers as discussed below. However, this increase will have severe effects on the  $\gamma$ -ray heating, because the central  $\tau_g$  is an order of magnitude larger than the rest of the envelope even at +350 d. In part, this is responsible for the late-time spectral formation discussed below.

### 6.3.1. Ca line in the late-time spectrum

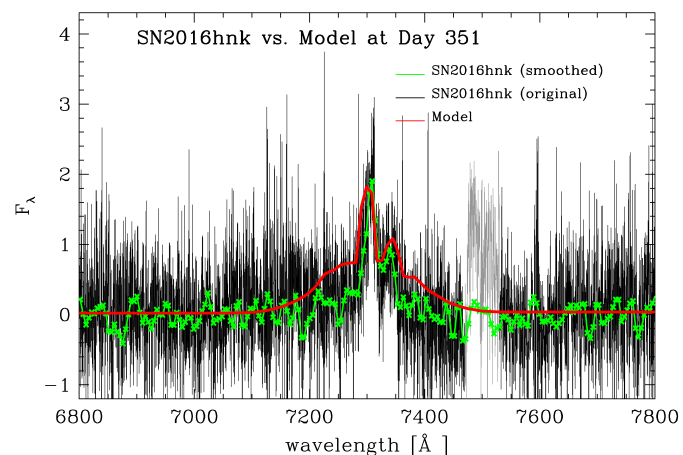
The main effect on  $Y_e$  is the density of burning. A central density of  $6 \times 10^9 \text{ g cm}^{-3}$  is required to reduce  $Y_e$  to values that produce centrally located Ca as observed. With regard to the electron capture, the settling of Ne and slow deflagration phase has the same effect of an increased  $\rho_c$  by  $\approx 10\%$ . The transition to a detonation is triggered at  $\rho_{\text{tr}} = 10^7 \text{ g cm}^{-3}$  to compensate for the larger binding energy of the initial WD. The WD has a final explosion energy  $E_{\text{kin}} = 8.83 \times 10^{50} \text{ erg}$  and produces  $M(^{56}\text{Ni}) = 0.108 M_{\odot}$ . The overall density and chemical structure is given in Figure 25, and the element production is shown in Table 5. The overall chemical structure resembles the Höflich et al. (2002) model for SN 1999by and, more importantly, the results of the abundance tomography analysis (Table 4): a) the unburned C-O layers reach down to about  $15,000 \text{ km s}^{-1}$ ; b) products of explosive carbon burning (O-Ne-Mg) are seen down to  $\approx 9500 \text{ km s}^{-1}$ ; c) followed by incomplete burning (Si-S) down to about  $3500 \text{ km s}^{-1}$ .

In the inner region, electron capture shifts the nuclear statistical equilibrium from  $^{56}\text{Ni}$  to neutron rich isotopes, as the  $Y_e$  decreases towards the central region, see Figure 26. The decrease is from a  $Y_e$  of 0.497, which produces  $^{56-58}\text{Ni}$ ,  $^{55}\text{Co}$ ,  $^{54}\text{Fe}$ , to a  $Y_e$  of 0.43 – 0.44, which produces  $^{58-60}\text{Fe}$ ,  $^{54-56}\text{Mn}$ ,  $^{52-56}\text{Cr}$  and most importantly  $^{47-48}\text{Ca}$ .

In Figure 27, the comparison between the theoretical and observed late-time spectrum is given. The observation shows narrow emission from the forbidden [Ca II] doublet at  $\lambda\lambda 7291, 7324$  whose lower level is the ground-state of Ca II. This narrow [Ca II] doublet dominates with respect to the other features because most of the energy emitted in Ca is confined to a very narrow region of central Ca powered by radioactive decay. The strength of these other features is only up to the 5% of the Ca doublet and are lost below the noise level. The central energy deposition is large because that region is still optically thick to  $\gamma$ -rays, whereas  $\gamma$ -rays from the outer layers mostly escape (Figure 24). The underlying continuum is formed by many overlapping allowed and forbidden transitions of mostly Cr, Mn, and Fe group-elements, which distribute the energy over a wide wavelength range from



**Fig. 26.** Distribution of most abundant isotopes in  $\log(X_i)$  and  $Y_e$  as a function of velocity for the central region are dominated by neutron-rich isotopes at several minutes after the explosion. Subsequently, radioactive isotopes decay  $^{55}\text{Cr}$  (3.5 m)  $\rightarrow$   $^{55}\text{Mn}$ ,  $^{55}\text{Fe}$  (2.7 y)  $\rightarrow$   $^{55}\text{Mn}$ ,  $^{59}\text{Fe}$  (44.5 d)  $\rightarrow$   $^{59}\text{Co}$ ,  $^{56}\text{Ni}$  (6 d)  $\rightarrow$   $^{56}\text{Co}$  (77 d)  $\rightarrow$   $^{56}\text{Fe}$ ,  $^{57}\text{Co}$  (271 d)  $\rightarrow$   $^{57}\text{Fe}$ .



**Fig. 27.** Comparison of the nebular spectrum of SN 2016hnk with our high-density delayed-detonation model in the rest frame assuming a redshift of  $z = 0.01610$ . The flux is in  $10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$  on the same scale as Figure 19. The telluric region in the original spectrum is plotted in gray, around  $7,500 \text{ \AA}$ . The theoretical spectrum is dominated by a narrow, forbidden [Ca II] doublet at  $7291$  and  $7324 \text{ \AA}$  above a quasi-continuum at about the 3% level formed by many transitions of mostly heavier elements including Fe, Co, Ni, Cr, etc.

the UV to the mid-IR and produce many lines commonly seen in SNe Ia, but on much lower levels of about  $\sim 3\%$  relative to the strength of the dominant narrow Ca feature. Note that the density at the inner region is still high at  $\approx 3 \times 10^7 \text{ particles cm}^{-3}$ . Thus, forbidden lines dominate for Ca for transitions to the ground state but not for the more complex ions. In the model, the central feature is formed by  $^{48}\text{Ca}$  with expansion velocities  $v_{\text{exp}}$  up to about  $700 \text{ km s}^{-1}$  whereas a shoulder is the result of

$^{40}\text{Ca}$  of the region expanding at  $v \approx 3000\text{-}9000 \text{ km s}^{-1}$ . Within the scatter in the observed spectrum (black, Figure 27), the synthetic spectra agree with the observations.

The comparison with the smoothed spectrum (green, Figure 27) shows some discrepancies, namely the hint of the  $^{40}\text{Ca}$  shoulder in the models as just discussed and, in particular, a Doppler shift of the central Ca feature by about  $700 \text{ km s}^{-1}$ . Possible factors that could contribute to this offset and that can be discriminated by future observations of high-resolution spectra include a peculiar motion of the SN progenitor within the galaxy or a passive drag of pre-existing turbulent fields during the smoldering phase (Höflich & Stein 2002; Yoon & Langer 2004; Zingale & Dursi 2007; Zingale et al. 2011) during the 2-3 seconds of a typical deflagration phase. The former would show variations but would be correlated with the rotation of the host galaxy, while the latter would show variations between objects. We note that the smoothed spectrum shows structure on the 10% level that is not predicted, but is also not at a significant level. The strong jump at  $7500 \text{ \AA}$  is telluric.

## 7. Discussion and conclusions

We presented photometric and spectroscopic observations of SN 2016hnk and IFS of its type SB host galaxy MCG -01-06-070. For the first time, these data are combined with a unique late-time spectrum at  $\approx 360 \text{ d}$  and with IFS data providing a unique link to the progenitor and environment of sub-luminous SNe Ia and, possibly, their nature. The IFS places SN 2016hnk close to a region with significant  $\text{H}\alpha$  emission and with high dust content near the nucleus of MCG -01-06-070. For the ISM, we consistently found  $E(B - V) \sim 0.4 \text{ mag}$  with three different methods and a low  $R_V = 2.1 \pm 0.6$  typical of highly extinguished SNe Ia (Burns et al. 2014).

The SN 2016hnk light curve was most similar to 2002es-like SNe Ia up to +15 days past maximum, however, optical and NIR spectra of SN 2016hnk during the photospheric phase show that the SN closely resembles the prototypes for sub-luminous SNe Ia, namely SNe 1991bg and 1999by. The  $\Delta m_{15,s}(B)$  of  $1.80 \pm 0.20 \text{ mag}$  and  $s_{BV}$  of  $0.438 \pm 0.030$  confirms the similarity between SN 2016hnk and sub-luminous SNe Ia. Based on detailed analyses of the optical and NIR light curves, we detect a flux excess after  $\approx 15 \text{ d}$  caused by scattering processes with an interstellar medium at around  $1.0 \text{ pc}$  from the SN, suggesting that the sub-luminous SN 2016hnk is embedded in a dense environment. We note that the extinction by dust is significantly larger in our object than in SNe 1991bg, 1999by and 2005ke (Lucy 1999; Höflich et al. 2002; Patat et al. 2012a; Penney & Höflich 2014). The under-representation of sub-luminous SNe Ia in high-redshift samples is usually attributed to long progenitor evolution, but may also be magnified by sub-luminous SNe Ia occurring in dusty regions or at low galactocentric distances (Galbany et al. 2012), or simply be attributed to selection effects (González-Gaitán et al. 2011).

The early-time spectra of SN 2016hnk have unique absorption features between  $6000\text{-}7000 \text{ \AA}$ , which we identified as low excitation Fe II, Co II and Co I lines. Furthermore, it was shown that the line on the redward side of the Si II 6355Å feature is not produced by C II but rather by Fe II  $\lambda 6516$ . In the instance of sub-luminous SNe Ia the ionization state of carbon at the epochs studied would be neutral and carbon would not be present in the optical spectrum, furthermore it could be expected to be further out in velocity space and not visible at these epochs. This Fe II line has also been seen by a few groups for normal bright SNe

Ia (see Höflich 1995) and with the PHOENIX code (E. Baron private communication).

We performed a detailed comparison of the optical spectra to spectral abundance stratification/tomography. The results are consistent with a  $M_{\text{Ch}}$  mass WD with characteristics found in a previous analysis of SN 1999by based on delayed-detonation models (Höflich et al. 2002). These characteristics include: a) low maximum-light luminosity and ionization state; b) low photospheric velocities; c) low-velocity iron-group elements; and d) an extended region (in velocity space) of products of explosive C-burning and incomplete O- and Si-burning. The model spectrum places our early-time NIR spectrum at +11 days after maximum light, which is  $\approx 2$  days later than inferred from maximum-light curve fits. Given the fast evolution in the NIR and the large uncertainty in  $t_{\text{max}}$ , we show that NIR spectra can be used efficiently to refine the time of maximum light and place the first optical spectrum at about +2-3 days after maximum. Both the tomography and the previous analysis of SN 1999by provide consistent line identifications for spectra dominated by neutral and single-ionized Fe, S/Si II, and O I.

Moreover, and unlike other SN Ia subtypes, there are indications of asphericity for sub-luminous SNe Ia. The early continuum polarization of SN 1999by shows a good agreement with the same models from Höflich et al. (2002) used in this work when an asphericity of about  $\sim 20\%$  is included (Howell et al. 2001). However, the resulting anisotropy in the luminosity is about  $\sim 5\%$  and hardly effects the spectra (Patat et al. 2012a).

The late-time nebular spectrum obtained by X-shooter shows only one prominent, narrow feature with a width of  $\approx 700 \text{ km s}^{-1}$ . This feature has not been predicted by the delayed-detonation model for SN 1999by despite the fact that most of the optical and NIR features at early phases are rather similar between SN 2016hnk and SN 1999by, as discussed above. This narrow late-time feature can be attributed to the [Ca II] doublet at  $\lambda 7291, 7324$ . It can be understood as a result of a nucleosynthesis effect, namely electron capture. For high central densities of a WD close to the limit for an accretion induced collapse, with  $Y_e$  less than 0.44, results in a shift of the nuclear statistical equilibrium towards  $^{48}\text{Ca}$ . In our model, a separate Ca-layer is produced at expansion velocities less than  $750 \text{ km s}^{-1}$ . As discussed in Section 6.2, the exact density depends on the initial conditions. For example, the gravitational settling of heavy elements (i.e.  $^{22}\text{Ne}$  and  $^{20}\text{Ne}$ ) over long time-scales may increase the central density by about 10 % to achieve lower  $Y_e$ . There is a high ratio of  $^{48}\text{Ca}/^{40}\text{Ca}$  in solar composition (Thielemann et al. 2002), and the result here may be suggestive of some sub-luminous SNe Ia as a main production sites for  $^{48}\text{Ca}$ . However, we have late-time, high-resolution observations only for SN 2016hnk and for none of the other sub-luminous SNe Ia, and we lack high S/N observations of the underlying iron-group lines that are predicted on the few percent level compared to the [Ca II] feature. We note that sub- $M_{\text{Ch}}$  models seem to fall well short of producing  $^{48}\text{Ca}$  (Brachwitz et al. 2000), despite having a compressional detonation wave, which increases the peak density in a WD (Höflich et al. 2018).

We note our interpretation of the narrow nebular [Ca II] feature may not be unique. However, RT instabilities are unlikely to cause this feature. As discussed in Section 6.2, RT instabilities would produce many plumes rather than a single narrow feature and, even if one plume were to dominate from the region of incomplete burning, we would expect a mixture of Ca/Fe and, potentially, S/Si. Conversely, a narrow feature similar to the observations could be produced within the ejecta in a narrow jet-like Ca-rich structure. In this scenario, the single strong feature

does not consist of iron-group elements and must be viewed almost equatorially. If viewed pole on we would not see a narrow feature, but rather a high-velocity one.

Another explanation could be in the framework of interacting binaries with material stripped from the donor star. This can occur within both the  $M_{\text{Ch}}$  and sub- $M_{\text{Ch}}$  scenarios. In principle, the low-velocity Ca may be attributed to material stripped from a companion star. However, in this case, we may expect very strong forbidden [O I] lines at 6300, 6364 (and 6393) Å regardless of whether the accreted material is H, He, or C-O rich. To rule out or confirm this possibility we would require a time series of spectra of more objects at those late epochs.

Finally, we want to address the possible link to progenitor systems. Potential progenitor systems may either consist of systems of two WDs (called double degenerate), a single WD with a donor which may be main sequence, red giant, or He star (called a single degenerate), or explosions within an asymptotic red giant branch star (Branch et al. 1995; Han et al. 2010; Piersanti et al. 2014; Denissenkov et al. 2015; Nomoto & Leung 2018; Yoon et al. 2007).

However, the link of SN2016hnk is highly uncertain. As discussed, SN2016hnk is in a galaxy of low star formation rate. This may point towards a progenitor scenario with a long evolutionary time scale. Though, there is still some ongoing star formation. This problem cannot be solved by direct analysis of the ejecta due to the lack of 'classical' observations such as high-precision light curves, as well as flux and polarization spectra well before maximum light (Patat et al. 2012b; Hoefflich et al. 2013; Gall et al. 2018a). The lack of early time LCs and optical and IR spectra does not allow us to probe the outermost layers to measure  $Z$  in the unburned layers, or the amount of stable iron-group elements.

This leaves the the high central density of the WD and the nearby material in the environment as a hint towards the progenitor. Within our models, we need a low accretion rate during the final phase of accretion to achieve a high central density, which results in a low  $Y_e$ . The nearby material may indicate a high mass loss rate prior to the explosion, e.g. during a common envelope phase. For many of the common progenitor systems, a drop in the final accretion rate may be expected because binary systems tend to become wider due to mass-loss. However, the low accretion rate depends on the treatment of the convective URCA process (Paczynski 1972; Barkat & Wheeler 1990; Stein & Wheeler 2006; Denissenkov et al. 2015; Piersanti et al. 2017), and a faster accretion rate may automatically lead to an explosion close to the AIC.

*Acknowledgements.* Endavant les atxes. LG would like to thank the organizers of the *Transients in New Surveys: the Undiscovered Country* workshop held in July 2018 in Leiden where this work started as a result of a huge collaborative effort, and the Lorentz Center staff for providing such a great work environment. CA and EYH acknowledge the support provided by the National Science Foundation under grant No. AST-1613472. PH acknowledges support by the National Science Foundation (NSF) grant 1715133, and thank thank F.K. Thielemann and G. Martinez-Pinedo for helpful discussions and advice on the current state of electron-capture rates. MS is supported in part by a generous grant (13261) from VILLUM FONDEN and a project grant from the Independent Research Fund Denmark. EB acknowledges partial support from NASA grant NNX16AB25G, and thanks the Aarhus University Research Fund (AUFF) for a guest researcher grant. SB is partially supported by China postdoctoral science foundation grant No.2018T110006. MB acknowledges support from the Swedish Research Council (Vetenskapsrådet), the Swedish National Space Board and the research environment grant "Gravitational Radiation and Electromagnetic Astrophysical Transients (GREAT)". CPG and MS acknowledge support from EU/FP7-ERC grant number 615929. JH acknowledges financial support from the Finnish Cultural Foundation. DAH, CM, and GH were supported by NSF grant AST-1313484. Support for TdJ have been provided by US NSF grant AST-1211916, the TABASGO Foundation, Gary and Cynthia

Bengier. SWJ acknowledges support from NSF award AST-1615455. KM acknowledges support from STFC (ST/M005348/1) and from H2020 through an ERC Starting Grant (758638). Research by DJS is supported by NSF grants AST-1821987 and 1821967. XW is supported by the National Natural Science Foundation of China (NSFC grants 11325313 and 11633002) JZ is supported by the National Natural Science Foundation of China (NSFC, grants 11773067, 11403096), the Youth Innovation Promotion Association of the CAS (grants 2018081), and the Western Light Youth Project of CAS. Based in part on data obtained with the LCO network under programmes CLN2016A-005 (PI: Prieto), NAOC2017AB-001 (PI: Dong), CON2016A-008 (PI: Stanek), and ARI2017AB-001 (PI: Bersier). This work has made use of data from the Asteroid Terrestrial-impact Last Alert System (ATLAS) project. ATLAS is primarily funded to search for near earth asteroids through NASA grants NN12AR55G, 80NSSC18K0284, and 80NSSC18K1575; byproducts of the NEO search include images and catalogs from the survey area. The ATLAS science products have been made possible through the contributions of the University of Hawaii Institute for Astronomy, the Queen's University Belfast, and the Space Telescope Science Institute. ASAS-SN is supported by the Gordon and Betty Moore Foundation through grant GBMF5490 to the Ohio State University and NSF grant AST-1515927. Development of ASAS-SN has been supported by NSF grant AST-0908816, the Mt. Cuba Astronomical Foundation, the Center for Cosmology and AstroParticle Physics at the Ohio State University, the Chinese Academy of Sciences South America Center for Astronomy (CASSACA), the Villum Foundation, and George Skestos. We thank the Las Cumbres Observatory (LCO) and its staff for its continuing support of the ASAS-SN project. NUTS is supported in part by the The Instrument Center for Danish Astrophysics (IDA). This research uses data obtained through the Telescope Access Program (TAP), which has been funded by "the Strategic Priority Research Program- The Emergence of Cosmological Structures" of the Chinese Academy of Sciences (Grant No.11 XDB09000000), the Special Fund for Astronomy from the Ministry of Finance, and the National Astronomical Observatories of China. Funding for the LJT has been provided by the Chinese Academy of Sciences (CAS) and the People's Government of Yunnan Province. Based on observations obtained at the Gemini Observatory under program GS-2016B-Q-22 (PI: Sand). Gemini is operated by the Association of Universities for Research in Astronomy, Inc., under a cooperative agreement with the NSF on behalf of the Gemini partnership: the NSF (United States), the National Research Council (Canada), CONICYT (Chile), Ministerio de Ciencia, Tecnología e Innovación Productiva (Argentina), and Ministério da Ciência, Tecnologia e Inovação (Brazil). The data were processed using the Gemini IRAF package. We thank the queue service observers and technical support staff at Gemini Observatory for their assistance.

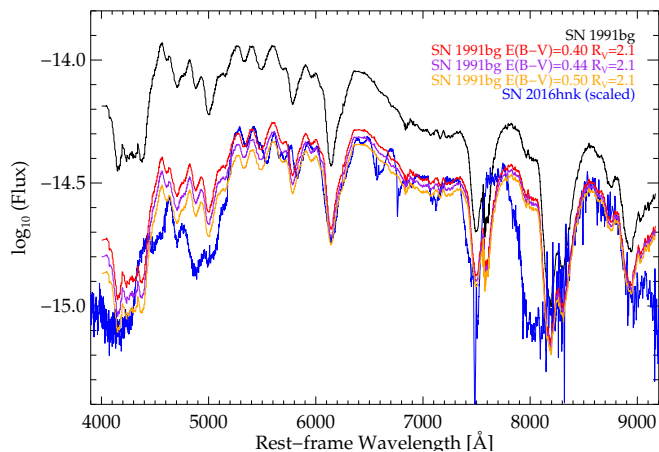
## References

- Abolfathi, B., Aguado, D. S., Aguilar, G., et al. 2018, *ApJS*, 235, 42  
Aguado, D. S., Ahumada, R., Almeida, A., et al. 2019, *ApJS*, 240, 23  
Amanullah, R. & Goobar, A. 2011, *ApJ*, 735, 20  
Arnett, W. D. 1982, *ApJ*, 253, 785  
Ashall, C., Hoefflich, P., Hsiao, E. Y., et al. 2019a, arXiv e-prints [arXiv:1904.01633]  
Ashall, C., Hsiao, E. Y., Hoefflich, P., et al. 2019b, *ApJ*, 875, L14  
Ashall, C., Mazzali, P., Bersier, D., et al. 2014, *MNRAS*, 445, 4427  
Ashall, C., Mazzali, P., Sasdelli, M., & Prentice, S. J. 2016a, *MNRAS*, 460, 3529  
Ashall, C., Mazzali, P. A., Pian, E., & James, P. A. 2016b, *MNRAS*, 463, 1891  
Ashall, C., Mazzali, P. A., Stritzinger, M. D., et al. 2018, *MNRAS*, 477, 153  
Ashall, C., Pian, E., Mazzali, P. A., et al. 2017, ArXiv e-prints [arXiv:1702.04339]  
Barbon, R., Benetti, S., Cappellaro, E., Rosino, L., & Turatto, M. 1990, *A&A*, 237, 79  
Barkat, Z. & Wheeler, J. C. 1990, *ApJ*, 355, 602  
Becker, A. 2015, HOTPANTS: High Order Transform of PSF ANd Template Subtraction, Astrophysics Source Code Library  
Benetti, S., Cappellaro, E., Mazzali, P. A., et al. 2005, *ApJ*, 623, 1011  
Betoule, M., Kessler, R., Guy, J., et al. 2014, *A&A*, 568, A22  
Blondin, S., Dessart, L., Hillier, D. J., & Khokhlov, A. M. 2017, *MNRAS*, 470, 157  
Blondin, S., Matheson, T., Kirshner, R. P., et al. 2012, *AJ*, 143, 126  
Brachwitz, F., Dean, D. J., Hix, W. R., et al. 2000, *ApJ*, 536, 934  
Branch, D., Dang, L. C., Hall, N., et al. 2006, *PASP*, 118, 560  
Branch, D., Livio, M., Yungelson, L. R., Boffi, F. R., & Baron, E. 1995, *PASP*, 107, 1019  
Brown, T. M., Baliber, N., Bianco, F. B., et al. 2013, *PASP*, 125, 1031  
Bulla, M., Goobar, A., Amanullah, R., Feindt, U., & Ferretti, R. 2018a, *MNRAS*, 473, 1918  
Bulla, M., Goobar, A., & Dhawan, S. 2018b, *MNRAS*, 479, 3663  
Burns, C. R., Parent, E., Phillips, M. M., et al. 2018, *ApJ*, 869, 56  
Burns, C. R., Stritzinger, M., Phillips, M. M., et al. 2014, *ApJ*, 789, 32



- Burns, C. R., Stritzinger, M., Phillips, M. M., et al. 2011, *AJ*, 141, 19
- Buzzoni, B., Delabre, B., Dekker, H., et al. 1984, *The Messenger*, 38, 9
- Cannizzaro, G., Fraser, M., Benetti, S., et al. 2016a, *Transient Name Server Classification Report*, 872
- Cannizzaro, G., Fraser, M., Benetti, S., et al. 2016b, *The Astronomer's Telegram*, 9703
- Cao, Y., Kulkarni, S. R., Howell, D. A., et al. 2015, *Nature*, 521, 328
- Cid Fernandes, R., Mateus, A., Sodré, L., Stasińska, G., & Gomes, J. M. 2005, *MNRAS*, 358, 363
- Contreras, C., Hamuy, M., Phillips, M. M., et al. 2010, *AJ*, 139, 519
- Cybur, R. H., Amthor, A. M., Ferguson, R., et al. 2010, *ApJS*, 189, 240
- Deloye, C. J. & Bildsten, L. 2002, *ApJ*, 580, 1077
- Denissenkov, P. A., Truran, J. W., Herwig, F., et al. 2015, *MNRAS*, 447, 2696
- DES Collaboration, Abbott, T. M. C., Allam, S., et al. 2018, *arXiv e-prints*, arXiv:1811.02374
- Diamond, T. R., Höflich, P., Hsiao, E. Y., et al. 2018, *ApJ*, 861, 119
- Dimitriadis, G., Pursiainen, M., Smith, M., et al. 2016, *The Astronomer's Telegram*, 9704
- Domínguez, I. & Höflich, P. 2000, *ApJ*, 528, 854
- Dopita, M. A., Kewley, L. J., Sutherland, R. S., & Nicholls, D. C. 2016, *Ap&SS*, 361, 61
- Fesen, R. A., Höflich, P. A., & Hamilton, A. J. S. 2015, *ApJ*, 804, 140
- Filippenko, A. V., Richmond, M. W., Branch, D., et al. 1992, *AJ*, 104, 1543
- Fitzpatrick, E. L. 1999, *PASP*, 111, 63
- Folatelli, G., Morrell, N., Phillips, M. M., et al. 2013, *ApJ*, 773, 53
- Foley, R. J., Jha, S. W., Pan, Y.-C., et al. 2016, *MNRAS*, 461, 433
- Förster, F., González-Gaitán, S., Folatelli, G., & Morrell, N. 2013, *ApJ*, 772, 19
- Friedman, A. S., Wood-Vasey, W. M., Marion, G. H., et al. 2015, *ApJS*, 220, 9
- Fuller, G. M., Fowler, W. A., & Newman, M. J. 1985, *ApJ*, 293, 1
- Galbany, L., Anderson, J. P., Rosales-Ortega, F. F., et al. 2016a, *MNRAS*, 455, 4087
- Galbany, L., Anderson, J. P., Sánchez, S. F., et al. 2018, *ApJ*, 855, 107
- Galbany, L., Miquel, R., Östman, L., et al. 2012, *ApJ*, 755, 125
- Galbany, L., Stanishchev, V., Mourão, A. M., et al. 2014, *A&A*, 572, A38
- Galbany, L., Stanishchev, V., Mourão, A. M., et al. 2016b, *A&A*, 591, A48
- Gall, C., Stritzinger, M. D., Ashall, C., et al. 2018a, *A&A*, 611, A58
- Gall, C., Stritzinger, M. D., Ashall, C., et al. 2018b, *A&A*, 611, A58
- Ganeshalingam, M., Li, W., Filippenko, A. V., et al. 2010, *ApJS*, 190, 418
- Ganeshalingam, M., Li, W., Filippenko, A. V., et al. 2012, *ApJ*, 751, 142
- García-Berro, E. & Lorén-Aguilar, P. 2017, *Dynamical Mergers (Alsabti, A. W. and Murdin, P.)*, 1237
- Garnavich, P. M., Bonanos, A. Z., Krisciunas, K., et al. 2004, *ApJ*, 613, 1120
- Goldhaber, G., Groom, D. E., Kim, A., et al. 2001, *ApJ*, 558, 359
- González Delgado, R. M., García-Benito, R., Pérez, E., et al. 2015, *A&A*, 581, A103
- González-Gaitán, S., Perrett, K., Sullivan, M., et al. 2011, *ApJ*, 727, 107
- Goobar, A. 2008, *ApJ*, 686, L103
- Hachinger, S., Mazzali, P. A., Tanaka, M., Hillebrandt, W., & Benetti, S. 2008, *MNRAS*, 389, 1087
- Hachinger, S., Mazzali, P. A., Taubenberger, S., Pakmor, R., & Hillebrandt, W. 2009, *MNRAS*, 399, 1238
- Hamuy, M., Phillips, M. M., Suntzeff, N. B., et al. 2003, *Nature*, 424, 651
- Hamuy, M., Phillips, M. M., Suntzeff, N. B., et al. 1996, *AJ*, 112, 2391
- Han, D.-H., Park, C., Choi, Y.-Y., & Park, M.-G. 2010, *ApJ*, 724, 502
- Henden, A. A., Levine, S., Terrell, D., & Welch, D. L. 2015, in *American Astronomical Society Meeting Abstracts*, Vol. 225, *American Astronomical Society Meeting Abstracts #225*, 336.16
- Hicken, M., Challis, P., Jha, S., et al. 2009, *ApJ*, 700, 331
- Hinshaw, G., Larson, D., Komatsu, E., et al. 2013, *The Astrophysical Journal Supplement Series*, 208, 19
- Höflich, P., Dragulin, P., Mitchell, J., et al. 2013, *Frontiers of Physics*, 8, 144
- Höflich, P. 1995, *ApJ*, 443, 89
- Höflich, P. 2003, in *Astronomical Society of the Pacific Conference Series*, Vol. 288, *Stellar Atmosphere Modeling*, ed. I. Hubeny, D. Mihalas, & K. Werner, 185
- Höflich, P., Ashall, C., Fisher, A., et al. 2018, *ArXiv e-prints [arXiv:1810.05910]*
- Höflich, P., Gerardy, C. L., Fesen, R. A., & Sakai, S. 2002, *ApJ*, 568, 791
- Höflich, P., Gerardy, C. L., Marion, H., & Quimby, R. 2006, *New Astronomy Review*, 50, 470
- Höflich, P., Hsiao, E. Y., Ashall, C., et al. 2017, *ApJ*, 846, 58
- Höflich, P. & Stein, J. 2002, *ApJ*, 568, 779
- Howell, D. A. 2011, *Nature Communications*, 2, 350
- Howell, D. A., Höflich, P., Wang, L., & Wheeler, J. C. 2001, *ApJ*, 556, 302
- Hoyle, F. & Fowler, W. A. 1960, *ApJ*, 132, 565
- Hristov, B., Collins, D. C., Höflich, P., Weatherford, C. A., & Diamond, T. R. 2018, *ApJ*, 858, 13
- Hsiao, E. Y., Burns, C. R., Contreras, C., et al. 2015, *A&A*, 578, A9
- Hsiao, E. Y., Conley, A., Howell, D. A., et al. 2007, *ApJ*, 663, 1187
- Hsiao, E. Y., Marion, G. H., Phillips, M. M., et al. 2013, *ApJ*, 766, 72
- Hsiao, E. Y., Phillips, M. M., Marion, G. H., et al. 2019, *PASP*, 131, 014002
- Jester, S., Schneider, D. P., Richards, G. T., et al. 2005, *AJ*, 130, 873
- Kasen, D. & Woosley, S. E. 2007, *ApJ*, 656, 661
- Kasliwal, M. M., Kulkarni, S. R., Gal-Yam, A., et al. 2012, *ApJ*, 755, 161
- Kelz, A., Verheijen, M. A. W., Roth, M. M., et al. 2006, *PASP*, 118, 129
- Khokhlov, A. M. 1989, *MNRAS*, 239, 785
- Krisciunas, K., Contreras, C., Burns, C. R., et al. 2017, *AJ*, 154, 211
- Krisciunas, K., Suntzeff, N. B., Phillips, M. M., et al. 2004, *AJ*, 128, 3034
- Kromer, M., Pakmor, R., Taubenberger, S., et al. 2013, *ApJ*, 778, L18
- Kromer, M., Sim, S. A., Fink, M., et al. 2010, *ApJ*, 719, 1067
- Langanke, K. & Martínez-Pinedo, G. 2003, *Reviews of Modern Physics*, 75, 819
- Leibundgut, B., Kirshner, R. P., Phillips, M. M., et al. 1993, *AJ*, 105, 301
- Lira, P., Suntzeff, N. B., Phillips, M. M., et al. 1998, *AJ*, 115, 234
- López Fernández, R., Cid Fernandes, R., González Delgado, R. M., et al. 2016, *MNRAS*, 458, 184
- Lucy, L. B. 1999, *A&A*, 345, 211
- Lyman, J. D., Taddia, F., Stritzinger, M. D., et al. 2018, *MNRAS*, 473, 1359
- Mannucci, F., Della Valle, M., & Panagia, N. 2006, *MNRAS*, 370, 773
- Marino, S., González-Gaitán, S., Förster, F., et al. 2015, *ApJ*, 806, 134
- Marion, G. H., Höflich, P., Gerardy, C. L., et al. 2009, *AJ*, 138, 727
- Marion, G. H., Höflich, P., Vacca, W. D., & Wheeler, J. C. 2003, *ApJ*, 591, 316
- Martínez-Pinedo, G., Lam, Y. H., Langanke, K., Zegers, R. G. T., & Sullivan, C. 2014, *Phys. Rev. C*, 89, 045806
- Martini, P., Stoll, R., Derwent, M. A., et al. 2011, *PASP*, 123, 187
- Mazzali, P. A. 2000, *A&A*, 363, 705
- Mazzali, P. A., Ashall, C., Pian, E., et al. 2018, *MNRAS*, 476, 2905
- Mazzali, P. A. & Lucy, L. B. 1993, *A&A*, 279, 447
- Mazzali, P. A., Nomoto, K., Cappellaro, E., et al. 2001, *ApJ*, 547, 988
- Mazzali, P. A., Sauer, D. N., Pastorello, A., Benetti, S., & Hillebrandt, W. 2008, *MNRAS*, 386, 1897
- Mazzali, P. A., Sullivan, M., Hachinger, S., et al. 2014, *MNRAS*, 439, 1959
- McCully, C., Jha, S. W., Foley, R. J., et al. 2014a, *Nature*, 512, 54
- McCully, C., Jha, S. W., Foley, R. J., et al. 2014b, *ApJ*, 786, 134
- Monet, D. G., Levine, S. E., Canzian, B., et al. 2003, *AJ*, 125, 984
- Moorwood, A., Cuby, J.-G., & Lidman, C. 1998, *The Messenger*, 91, 9
- Nomoto, K. 1982, *ApJ*, 253, 798
- Nomoto, K. & Leung, S.-C. 2018, *Space Sci. Rev.*, 214, 67
- Nomoto, K., Thielemann, F.-K., & Yokoi, K. 1984, *ApJ*, 286, 644
- Osterbrock, D. E. & Ferland, G. J. 2006, *Astrophysics of gaseous nebulae and active galactic nuclei (University Science Books)*
- Paczynski, B. 1972, *Astrophys. Lett.*, 11, 53
- Pan, Y.-C., Duarte, A. S., Foley, R. J., et al. 2016, *The Astronomer's Telegram*, 9705
- Patat, F., Höflich, P., Baade, D., et al. 2012a, *A&A*, 545, A7
- Patat, F., Höflich, P., Baade, D., et al. 2012b, *A&A*, 545, A7
- Penney, R. & Höflich, P. 2014, *ApJ*, 795, 84
- Perlmutter, S., Aldering, G., Goldhaber, G., et al. 1999, *ApJ*, 517, 565
- Phillips, M. M. 1993, *ApJ*, 413, L105
- Phillips, M. M., Lira, P., Suntzeff, N. B., et al. 1999, *AJ*, 118, 1766
- Phillips, M. M., Phillips, A. C., Heathcote, S. R., et al. 1987, *PASP*, 99, 592
- Phillips, M. M., Simon, J. D., Morrell, N., et al. 2013, *ApJ*, 779, 38
- Piasticik, A. S., Steele, I. A., Bates, S. D., et al. 2014, in *Proc. SPIE*, Vol. 9147, *Ground-based and Airborne Instrumentation for Astronomy V*, 91478H
- Piersanti, L., Bravo, E., Cristallo, S., et al. 2017, *ApJ*, 836, L9
- Piersanti, L., Tornambé, A., & Castellani, V. 2004, *MNRAS*, 353, 243
- Piersanti, L., Tornambé, A., & Yungelson, L. R. 2014, *MNRAS*, 445, 3239
- Pinto, P. A. & Eastman, R. G. 2000, *ApJ*, 530, 744
- Poznanski, D., Ganeshalingam, M., Silverman, J. M., & Filippenko, A. V. 2011, *MNRAS*, 415, L81
- Poznanski, D., Prochaska, J. X., & Bloom, J. S. 2012, *MNRAS*, 426, 1465
- Prentice, S. J., Ashall, C., Mazzali, P. A., et al. 2018, *MNRAS*, 478, 4162
- Richmond, M. W. & Smith, H. A. 2012, *Journal of the American Association of Variable Star Observers (JAAVSO)*, 40, 872
- Riess, A. G., Filippenko, A. V., Challis, P., et al. 1998, *AJ*, 116, 1009
- Riess, A. G., Press, W. H., & Kirshner, R. P. 1996, *ApJ*, 473, 88
- Röpke, F. K. 2005, *A&A*, 432, 969
- Roth, M. M., Kelz, A., Fechner, T., et al. 2005, *PASP*, 117, 620
- Sánchez, S. F., García-Benito, R., Zibetti, S., et al. 2016, *A&A*, 594, A36
- Sánchez, S. F., Rosales-Ortega, F. F., Marino, R. A., et al. 2012, *A&A*, 546, A2
- Scalzo, R., Aldering, G., Antilogus, P., et al. 2014, *MNRAS*, 440, 1498
- Schlaflly, E. F. & Finkbeiner, D. P. 2011, *ApJ*, 737, 103
- Schmidt, G. D., Weymann, R. J., & Foltz, C. B. 1989, *PASP*, 101, 713
- Sell, P. H., Arur, K., Maccarone, T. J., et al. 2017, *ArXiv e-prints [arXiv:1712.02799]*
- Shappee, B. J., Prieto, J. L., Grupe, D., et al. 2014, *ApJ*, 788, 48
- Shen, K. J. & Moore, K. 2014, *ApJ*, 797, 46
- Skrutskie, M. F., Cutri, R. M., Stiening, R., et al. 2006, *AJ*, 131, 1163
- Smartt, S. J., Valenti, S., Fraser, M., et al. 2015, *A&A*, 579, A40
- Stalder, B., Tonry, J., Smartt, S. J., et al. 2017, *ApJ*, 850, 149

- Steele, I. A., Smith, R. J., Rees, P. C., et al. 2004, in Proc. SPIE, Vol. 5489, Ground-based Telescopes, ed. J. M. Oschmann, Jr., 679–692
- Stehle, M., Mazzali, P. A., Benetti, S., & Hillebrandt, W. 2005, MNRAS, 360, 1231
- Stein, J. & Wheeler, J. C. 2006, ApJ, 643, 1190
- Stritzinger, M., Burns, C. R., Phillips, M. M., et al. 2010, AJ, 140, 2036
- Stritzinger, M., Mazzali, P. A., Sollerman, J., & Benetti, S. 2006, A&A, 460, 793
- Stritzinger, M. D., Hsiao, E., Valenti, S., et al. 2014, A&A, 561, A146
- Stritzinger, M. D., Valenti, S., Hoeflich, P., et al. 2015, A&A, 573, A2
- Sullivan, M., Kasliwal, M. M., Nugent, P. E., et al. 2011, ApJ, 732, 118
- Tanaka, M., Mazzali, P. A., Stanishev, V., et al. 2011, MNRAS, 410, 1725
- Taubenberger, S. 2017, ArXiv e-prints [arXiv:1703.00528]
- Taubenberger, S., Hachinger, S., Pignata, G., et al. 2008, MNRAS, 385, 75
- Taubenberger, S., Kromer, M., Pakmor, R., et al. 2013, ApJ, 775, L43
- Thielemann, F.-K., Argast, D., Brachwitz, F., et al. 2002, Ap&SS, 281, 25
- Thomas, R. C., Nugent, P. E., & Meza, J. C. 2011, PASP, 123, 237
- Thompson, T. A. 2011, ApJ, 741, 82
- Tonry, J. L., Denneau, L., Heinze, A. N., et al. 2018, PASP, 130, 064505
- Tripp, R. 1998, A&A, 331, 815
- Tucker, M. A., Shappee, B. J., Vallely, P. J., et al. 2019, arXiv e-prints, arXiv:1903.05115
- Turatto, M., Benetti, S., Cappellaro, E., et al. 1996, MNRAS, 283, 1
- Verheijen, M. A. W., Bershady, M. A., Andersen, D. R., et al. 2004, Astronomische Nachrichten, 325, 151
- Vernet, J., Dekker, H., D’Odorico, S., et al. 2011, A&A, 536, A105
- Villanueva, S., Eastman, J. D., Gaudi, B. S., et al. 2016, in Proc. SPIE, Vol. 9906, Ground-based and Airborne Telescopes VI, 99062L
- Villanueva, Jr., S., Gaudi, B. S., Pogge, R. W., et al. 2018, PASP, 130, 015001
- Vorontsov-Vel’Yaminov, B. A. & Arhipova, V. P. 1962, in Morphological catalogue of galaxies., 1 (1962)
- Wang, X., Filippenko, A. V., Ganeshalingam, M., et al. 2009, ApJ, 699, L139
- Wang, X., Li, W., Filippenko, A. V., et al. 2008, ApJ, 675, 626
- Weingartner, J. C. & Draine, B. T. 2001, ApJ, 548, 296
- Wygoda, N., Katz, B., & Elbaz, Y. 2018, ArXiv e-prints [arXiv:1805.06907]
- Yoon, S.-C. & Langer, N. 2004, A&A, 419, 623
- Yoon, S.-C., Podsiadlowski, P., & Rosswog, S. 2007, MNRAS, 380, 933
- Zhang, J.-J., Wang, X.-F., Bai, J.-M., et al. 2014, AJ, 148, 1
- Zingale, M. & Dursi, L. J. 2007, ApJ, 656, 333
- Zingale, M., Nonaka, A., Almgren, A. S., et al. 2011, ApJ, 740, 8
- 
- <sup>1</sup> PITT PACC, Department of Physics and Astronomy, University of Pittsburgh, Pittsburgh, PA 15260, USA.
- <sup>2</sup> Departamento de Física Teórica y del Cosmos, Universidad de Granada, E-18071 Granada, Spain.
- <sup>3</sup> Department of Physics, Florida State University, Tallahassee, FL 32306, USA.
- <sup>4</sup> CENTRA/COSTAR, Instituto Superior Técnico, Universidade de Lisboa, Av. Rovisco Pais 1, 1049-001 Lisboa, Portugal.
- <sup>5</sup> European Southern Observatory, Karl-Schwarzschild-Straße 2, D-85748 Garching bei München, Germany.
- <sup>6</sup> Max-Planck-Institut für Astrophysik, Karl-Schwarzschild-Straße 1, D-85748 Garching bei München, Germany.
- <sup>7</sup> Department of Physics and Astronomy, Aarhus University, Ny Munkegade 120, DK-8000 Aarhus C, Denmark.
- <sup>8</sup> Astrophysics Research Institute, Liverpool John Moores University, IC2, Liverpool Science Park, 146 Brownlow Hill, Liverpool L3 5RF, UK
- <sup>9</sup> Homer L. Dodge Department of Physics and Astronomy, University of Oklahoma, Norman, OK, USA.
- <sup>10</sup> Aix Marseille Univ, CNRS, CNES, Laboratoire d’Astrophysique de Marseille (LAM), Marseille, France.
- <sup>11</sup> Unidad Mixta Internacional Franco-Chilena de Astronomía, CNRS/INSU UMI 3386, and Departamento de Astronomía, Universidad de Chile, Camino El Observatorio 1515, Las Condes, Santiago, Chile.
- <sup>12</sup> Kavli Institute for Astronomy and Astrophysics, Peking University, Yi He Yuan Road 5, Hai Dian District, Beijing 100871, People’s Republic of China
- <sup>13</sup> The Oskar Klein Centre, Department of Physics, Stockholm University, AlbaNova, SE-106 91 Stockholm, Sweden.
- <sup>14</sup> Las Cumbres Observatory, 6740 Cortona Dr. Suite 102, Goleta, CA, USA 93117.
- <sup>15</sup> University of California, Santa Barbara, Department of Physics, Broida Hall, Santa Barbara, CA, USA 93111.
- <sup>16</sup> The Observatories of the Carnegie Institution for Science, 813 Santa Barbara St., Pasadena, CA 91101, USA.
- <sup>17</sup> Cerro Tololo Inter-American Observatory, National Optical Astronomy Observatory, Casilla 603, La Serena, Chile.
- <sup>18</sup> Department of Astronomy, School of Physics, Peking University, Yi He Yuan Road 5, Hai Dian District, Beijing 100871, China.
- <sup>19</sup> Capodimonte Observatory, INAF-Naples , Salita Moiariello 16, 80131-Naples, Italy.
- <sup>20</sup> Astrophysics Science Division, NASA Goddard Space Flight Center, Greenbelt, MD 20771, USA.
- <sup>21</sup> School of Physics and Astronomy, University of Southampton, Southampton, SO17 1BJ, UK.
- <sup>22</sup> Tuorla Observatory, Department of Physics and Astronomy, FI-20014 University of Turku, Finland.
- <sup>23</sup> Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138.
- <sup>24</sup> University of California, San Diego, La Jolla, CA 92093.
- <sup>25</sup> School of Physics & Astronomy, Cardiff University, Queens Buildings, The Parade, Cardiff, CF24 3AA, UK.
- <sup>26</sup> Department of Astronomy, University of California, Berkeley, CA 94720-3411, USA.
- <sup>27</sup> Department of Physics and Astronomy, Rutgers the State University of New Jersey, 136 Frelinghuysen Road, Piscataway, NJ 08854 USA.
- <sup>28</sup> Space Telescope Science Institute, 3700 San Martin Drive, Baltimore, MD 21218, USA.
- <sup>29</sup> Zentrum für Astronomie der Universität Heidelberg, Institut für Theoretische Astrophysik, Philosophenweg 12, 69120, Heidelberg, Germany.
- <sup>30</sup> Heidelberger Institut für Theoretische Studien, Schloss-Wolfsbrunnengasse 35, 69118 Heidelberg, Germany.
- <sup>31</sup> Department of Physics, University of Warwick, Coventry, CV4 7AL, UK.
- <sup>32</sup> Astrophysics Research Centre, School of Mathematics and Physics, Queen’s University Belfast, BT7 1NN, UK.
- <sup>33</sup> School of Physics, Trinity College Dublin, Dublin 2, Ireland.
- <sup>34</sup> University of Texas at Austin, 1 University Station C1400, Austin, TX, 78712-0259, USA.
- <sup>35</sup> Department of Physics and Astronomy, Purdue University, 525 Northwestern Avenue, West Lafayette, IN, 47907, USA.
- <sup>36</sup> European Southern Observatory, Alonso de Córdova 3107, Casilla 19, Santiago, Chile.
- <sup>37</sup> Departamento de Astronomía, Universidad de Chile, Camino El Observatorio 1515, Las Condes, Santiago, Chile.
- <sup>38</sup> Steward Observatory, University of Arizona, 933 North Cherry Avenue, Tucson, AZ 85721-0065, USA.
- <sup>39</sup> Institute for Astronomy, University of Hawai’i, 2680 Woodlawn Drive, Honolulu, HI 96822, USA.
- <sup>40</sup> Department of Physics, BITS Pilani K.K. Birla Goa Campus.
- <sup>41</sup> Vanderbilt University, Department of Physics & Astronomy, Nashville, TN 37235, USA.
- <sup>42</sup> Department of Physics, University of California, Davis, CA 95616, USA.
- <sup>43</sup> Massachusetts Institute of Technology, Cambridge, MA 02139, USA.
- <sup>44</sup> Pappalardo Fellow.
- <sup>45</sup> Physics Department/Tsinghua Center for Astrophysics, Tsinghua University, Beijing, 100084, China.
- <sup>46</sup> Department of Astronomy, University of Texas at Austin, Austin, TX 78712, USA.
- <sup>47</sup> Yunnan Observatories, Chinese Academy of Sciences, Kunming 650216, China.
- <sup>48</sup> Key Laboratory for the Structure and Evolution of Celestial Objects, Chinese Academy of Sciences, Kunming 650216, China.



**Fig. A.1.** SN 1991bg (black) and SN 2016hnk (blue) spectra at maximum, together with three artificially reddened SN 1991bg spectra using an  $R_V=2.1$  and  $E(B-V)=0.40$  (red), 0.45 (purple), and 0.50 (orange) mag. All spectra have been first corrected for Milky Way extinction and rest-framed.

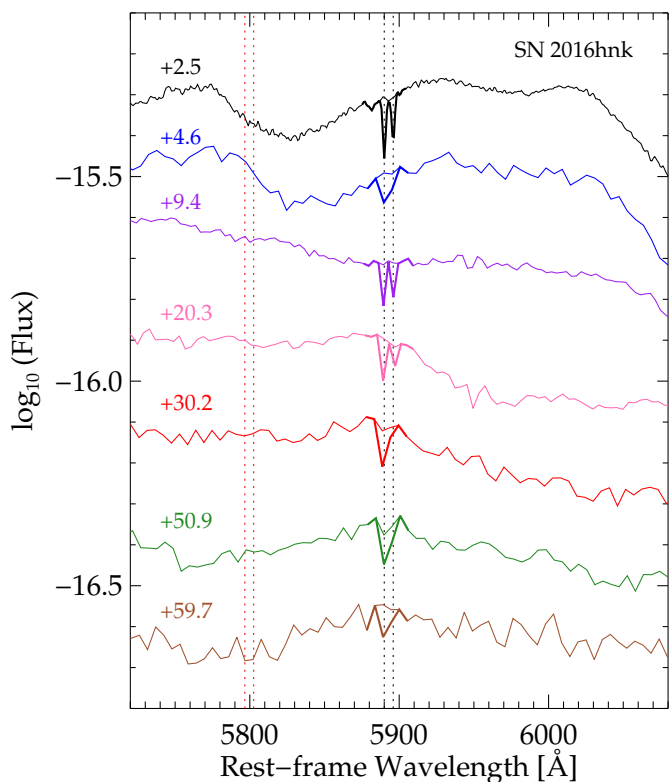
## Appendix A: Light curve excess tests

We performed a number of tests to try to explain the late light curve excess observed in SN 2016hnk. In the main text we provided the most plausible explanation, multi-scattering by dust clouds at  $1.0 \pm 0.5$  pc from the SN, while here we give more details of other tests.

### Appendix A.1: Reddening SN 1991bg spectrum and Na I D absorption

From Figures 8 and 14, it is obvious that SN 2016hnk light and color curves show some peculiarities. The first question to address is: is SN 2016hnk actually reddened, or is its red color intrinsic? For that, we took SN 1991bg spectrum at maximum light and, after correcting for Milky Way extinction and shifting to the restframe, we artificially applied the same reddening we found for SN 2016hnk ( $E(B-V)=0.4-0.5$  mag with an  $R_V=2.1$ ). Figure A.1 shows the SN 1991bg spectrum at maximum (in black), together with the corresponding SN 2016hnk spectrum also at maximum (in blue; also corrected for Milky Way reddening and shifted to the restframe), and three SN 1991bg reddened spectra with  $E(B-V)=0.40$ , 0.45, and 0.50 mag (in red, purple, and orange, respectively). Besides the differences already pointed out in Section 2.2 (broader NIR Calcium absorption, and features in the range 6400–7300 Å), we can see an almost perfect match between SN 2016hnk and all reddened SN 1991bg spectra redwards of 5200 Å. However, the blue part of the spectrum is where most of the discrepancy comes from: both the aggregate Fe II, Si II, and Mg II features at  $\sim 5000$  Å, and the Ti II at  $\sim 4200$  Å, are deeper in SN 2016hnk than in the reddened SN 1991bg spectra, which makes SN 2016hnk color curves even redder than expected by simple dust extinction. So, we conclude that there should be an intrinsic effect, in addition to reddening responsible for the large red colors of SN 2016hnk.

Another peculiarity of SN 2016hnk is the lack of Na I D absorption at any epoch, even when we obtain large  $E(B-V)$  ( $>0.4$  mag) with different methods. Although the relation between extinction and Na I D  $pEW$  has been put in question when measured in low-resolution spectra (Poznanski et al. 2011), it is usually accepted that the presence of Na I D  $pEW$  absorption is an indication of extinction by ISM (Phillips et al. 2013). The rela-



**Fig. A.2.** Zoom into SN 2016hnk spectra around the Na I D doublet (restframe 5890 and 5896 Å; in vertical black dotted lines) absorption, at seven epochs from +2.5 to +59.7 days after maximum. Thick lines represent Na I D absorption with  $pEW=1.3$  Å at the spectral resolution of each spectrum. Red vertical lines mark the position where the Na I D doublet from the Milky Way should be found. In all cases, we do not detect the Na I D lines in SN 2016hnk spectra.

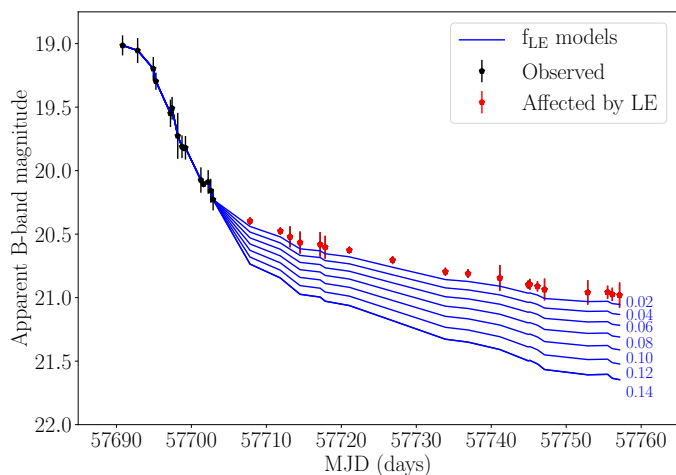
tions from Barbon et al. (1990) and Poznanski et al. (2012), predict both Na I D  $pEW$  of at least  $1.3$  Å, from the  $E(B-V) \sim 0.4$  mag measured in SN 2016hnk. We artificially input Na I D absorption with  $pEW$  of  $1.3$  Å in several SN 2016hnk spectra at different epochs, as shown in Figure A.2, taking into account the spectral resolution of each spectrum. Even for the latter spectrum at +59.7 d that has the lowest signal-to-noise ratio, we can significantly discard any indication of Na I D absorption. Therefore, in the case of SN 2016hnk, the large extinction does not come together with Na I D absorption.

### Appendix A.2: Simple light echo model

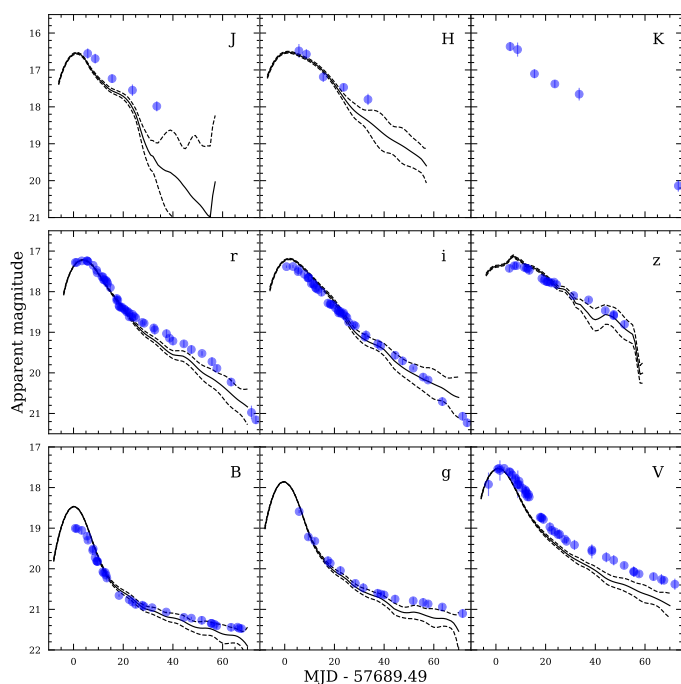
Another possible explanation for the light excess in the SN 2016hnk light curve after +15 days past maximum would be an intervening dust cloud close to the line-of-sight that is producing a light echo (LE; Marino et al. 2015). To test this possibility, we constructed a simple model accounting for different contributions of the LE to the total observed light, in the form,

$$f_{\text{SN}}^X(t) = f_{\text{obs}}^X(t) - \alpha f_{\text{LE}}^X S(t, t_{\text{delay}}, t_{\text{end}}), \quad (\text{A.1})$$

where  $f_{\text{SN}}^X$  is the actual emitted flux from the SN in the X band,  $f_{\text{obs}}^X$  is the flux we observed in the same band,  $f_{\text{LE}}^X$  is the re-emitted flux from the cloud,  $\alpha$  is a multiplicative factor that corresponds to the contribution of the LE with respect to the actual emission from the SN (and which we assumed to be band independent) in the sense that  $\alpha=0$  means no contribution at all and



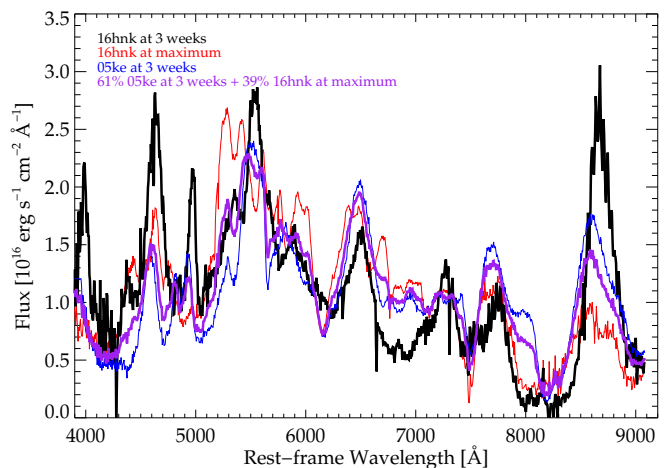
**Fig. A.3.** Effect of subtracting different fractional contribution of the light-echo to the SN 2016hbk *B* band observed light curve.



**Fig. A.4.** Multiband SN 2016hbk SNooPy fits using a 91bg template and with a 10% contribution of a light echo starting at +15 days until the end of the light curve. Fits in all band are reasonably better than those without a LE shown in Figure 7.

$\alpha=1$  corresponds to the extreme case of perfect reflection, resulting in  $f_{\text{SN}}=f_{\text{LE}}$ , and  $S$  is a step function that has a value of one for  $t_{\text{delay}} < t < t_{\text{end}}$  and zero at other epochs.

To estimate  $f_{\text{LE}}^X$  we assumed that early epochs are free of LE contribution and performed a SNooPy fit (fixing  $t_{\text{max}}$  and  $R_V$  to the same values found on Section 2.1) but only up to +15 days past maximum. Then, the peak flux of the best-fit model was taken as  $f_{\text{LE}}^X$  for each band independently, and used to generate a series of LE-free light curves with different combinations of:  $\alpha$  parameters from 0.02 to 0.14, in steps of 0.02;  $t_{\text{delay}}$  from 5 to 30 days in steps of 5 days; and  $t_{\text{end}}$  from 40 days on in steps of 5 days. See Figure A.3 as an example. Those sets of light curves were fitted with SNooPy keeping the assumptions on the  $t_{\text{max}}$  and  $R_V$ , but extending the fit to the whole time extent. The



**Fig. A.5.** Simulation of a light echo contaminated spectrum at three weeks past maximum, based on a fiducial subluminous spectrum at +3 weeks past maximum from SN 2005ke (contributing 61% to the total), and adding 39% of the SN 2016hbk spectrum at maximum light, which represents the contamination from the light echo. Comparing this LE-affected spectrum (in purple) to the observed SN 2016hbk (in black) we do not see any feature that definitely confirms that the latter is in fact affected by any feature reflected from earlier epochs.

best fit, shown in Figure A.4, was for a combination of parameters  $(\alpha, t_{\text{delay}}, t_{\text{end}}) = (0.10, 15, 80)$ . Note that the best  $t_{\text{end}}$  value is right after the last photometric point, meaning that the LE is contributing flux at least until the epoch when we stopped our follow up campaign. The model also provided the best estimates of  $s_{\text{BV}} = 0.380 \pm 0.032$  and  $E(B - V) = 0.458 \pm 0.064$  mag, the former being a bit lower than the reported in the main text, and the latter being consistent with the value found using CDMagic.

Although light curve fits improve significantly when this simple model of the LE is removed from observations, if this excess is due to a light echo, we should further see a mixture of spectral features in the spectra. At later epochs when the LE is affecting the observations ( $>+15$  d in our model), we should see spectral features from earlier epochs that are being reflected and that are contaminating the light directly received by the observer from the SN. This contamination is strongly dependent on the configuration of the ISM around the SN but, again, in the simplest scenario (single cloud at a certain distance from the SN) this would correspond to a mixture of the spectrum at the actual epoch and a spectrum at an early epoch equal to the delayed time.

To test if these features are present at late epochs we took the spectrum of a normal 91bg-like SNIa, SN 2005ke at +23.6 days, and add up incremental contributions of SN 2016hbk spectrum at maximum light, which represents the contamination from the light echo. The resulting mixed spectrum is compared to the observed spectrum of SN 2016hbk at +20.3 days to check if the mixed features from those two spectra reproduce those seen at late epochs in SN 2016hbk. Starting from no contribution from the LE, and increasing in steps of 1% up to 50% contribution to the total light, we find the most similar spectrum for a 39% LE contribution (based on  $\chi^2$  minimization). In this best case, shown in Figure A.5, we do not find a convincing resemblance between the two spectra (observed, in black, and simulated, in purple) that confirms the hypothesis of a LE causing those features. For this reason, we do not consider the light echo explanation for the late-time light excess, and favor the multiscattering approach detailed in the main text.

**Table C.1.** ASASSN *V*-band photometry and upper limits.

MJD	Epoch	<i>V</i>
57652.435	-36.46	>17.72
57653.253	-35.65	>18.00
57655.145	-33.79	>18.31
57660.479	-28.54	>18.36
57663.160	-25.90	>17.87
57664.204	-24.88	>18.33
57668.340	-20.81	>17.57
57669.340	-19.82	>17.94
57670.382	-18.80	>18.01
57671.476	-17.72	>17.97
57672.324	-16.88	>17.91
57673.521	-15.71	>17.83
57685.513	-3.91	>16.74
57686.507	-2.93	17.92 (0.290)
57689.498	+0.02	17.25 (0.210)
57691.130	+1.62	17.58 (0.250)
57698.282	+8.66	17.92 (0.340)

## Appendix B: Spectral modeling technique

The code is a 1D Monte Carlo radiation transport code that follows the propagation of photon packets through a SN atmosphere. The photon packets can undergo Thomson scattering and line absorption. Photon packets have two fates: they either re-enter the photosphere in a process known as back scattering or escape the SN atmosphere after possibly undergoing scattering or line absorption. A modified nebular approximation is used to account for non-local thermodynamical equilibrium effects. The code makes use of the Schuster-Schwarzschild approximation, which assumes that radiative energy is emitted from an inner blackbody. This has the advantage that the radiation transport below the photosphere is not required to be known, but can cause excess flux in the red ( $>6500 \text{ \AA}$ ) past maximum light. We note that the values we obtain for  $v_{\text{ph}}$  will in most cases be lower than those measured from the individual ions in the spectra in Section 5, as the lines are produced in the ejecta above the photosphere. For further details on the abundance stratification code and method used in this work see [Mazzali & Lucy \(1993\)](#); [Lucy \(1999\)](#); [Mazzali \(2000\)](#); [Stehle et al. \(2005\)](#).

## Appendix C: Photometry, local sequence, and spectral log

**Table C.2.** ATLAS *c*- and *o*-band photometry and upper limits.

MJD	Epoch	<i>c</i>	MJD	Epoch	<i>o</i>
57688.474	-0.99	17.69 (0.07)	57643.603	-45.15	>19.20
57688.492	-0.97	17.91 (0.07)	57643.617	-45.14	>19.53
57688.506	-0.96	17.77 (0.07)	57643.633	-45.12	>19.71
57688.523	-0.94	17.68 (0.07)	57671.530	-17.66	>20.18
57688.542	-0.92	17.81 (0.08)	57671.545	-17.65	>19.09
57696.448	+6.86	17.83 (0.07)	57671.560	-17.64	>20.50
57696.463	+6.87	17.70 (0.06)	57680.492	-8.85	18.74 (0.29)
57696.482	+6.89	17.65 (0.06)	57680.507	-8.83	19.06 (0.41)
57696.489	+6.90	17.82 (0.07)	57680.526	-8.81	18.67 (0.31)
57696.509	+6.92	17.83 (0.07)	57680.533	-8.81	18.37 (0.24)
57699.356	+9.72	17.91 (0.13)	57680.551	-8.79	19.32 (0.49)
57699.374	+9.74	18.08 (0.10)	57700.433	+10.78	17.80 (0.14)
57699.394	+9.76	17.92 (0.09)	57700.462	+10.81	17.53 (0.09)
57699.408	+9.77	17.79 (0.08)	57700.477	+10.82	17.45 (0.08)
57699.422	+9.78	17.90 (0.12)	57700.497	+10.84	17.45 (0.10)
57716.402	+26.50	18.48 (0.13)	57704.387	+14.67	17.49 (0.16)
57716.417	+26.51	18.95 (0.21)	57704.406	+14.69	17.59 (0.18)
57716.433	+26.53	18.89 (0.23)	57704.415	+14.70	18.06 (0.27)
57716.449	+26.54	18.42 (0.40)	57704.436	+14.72	17.76 (0.21)
57716.465	+26.56	19.03 (0.71)	57704.443	+14.72	18.04 (0.28)
57744.335	+53.99	19.81 (0.38)	57712.443	+22.60	17.73 (0.15)
57744.348	+54.00	20.80 (0.70)	57736.351	+46.13	18.17 (0.25)
57744.388	+54.04	19.46 (0.29)	57736.364	+46.14	18.88 (0.42)
57756.336	+65.80	20.39 (0.50)	57736.376	+46.15	19.27 (0.57)
57756.346	+65.81	20.57 (0.59)	57736.389	+46.17	18.39 (0.34)
57780.272	+89.35	20.23 (0.47)	57743.343	+53.01	19.61 (0.40)
57780.285	+89.37	20.71 (0.71)	57743.356	+53.02	18.96 (0.23)
57780.318	+89.40	20.52 (0.63)	57743.369	+53.03	19.14 (0.27)
57784.309	+93.33	20.16 (0.52)	57743.382	+53.05	18.73 (0.20)

Table C.3. Photometry for SN 2016hnk. Epochs are given in the rest frame.

MJD	Epoch	<i>B</i>	<i>V</i>	<i>u</i>	<i>g</i>	<i>r</i>	<i>i</i>	<i>z</i>	Source
57690.200	+0.71	19.003 (0.088)	17.544 (0.054)	...	...	17.282 (0.046)	17.384 (0.030)	...	fl03
57690.815	+1.31	19.014 (0.079)	17.526 (0.061)	...	...	17.278 (0.037)	...	...	fl14
57692.830	+3.29	19.054 (0.098)	17.526 (0.041)	...	...	17.241 (0.037)	17.380 (0.039)	...	fl06
57694.920	+5.36	19.197 (0.092)	17.607 (0.088)	...	...	17.243 (0.038)	17.455 (0.082)	17.426 (0.088)	IO:O
57694.930	+5.36	...	...	...	...	17.240 (0.063)	17.508 (0.059)	...	fl06
57695.270	+5.70	19.296 (0.068)	17.628 (0.034)	...	18.591 (0.038)	17.257 (0.033)	...	...	fl15
57696.130	+6.54	...	17.690 (0.082)	...	...	...	...	...	FLI
57696.970	+7.37	...	...	>18.832	...	...	...	17.361 (0.089)	fl16
57697.205	+7.60	19.549 (0.108)	17.802 (0.082)	...	...	...	...	...	FLI
57697.435	+7.83	19.509 (0.089)	17.768 (0.047)	...	...	17.351 (0.035)	17.571 (0.039)	...	fl12
57697.835	+8.22	...	...	>18.957	...	...	...	17.358 (0.063)	fl06
57698.175	+8.56	19.726 (0.181)	17.929 (0.093)	...	...	...	...	...	FLI
57698.755	+9.13	19.812 (0.090)	17.842 (0.065)	...	...	17.471 (0.047)	17.663 (0.053)	...	fl06
57699.000	+9.37	19.821 (0.026)	17.927 (0.088)	...	19.215 (0.081)	17.534 (0.046)	17.659 (0.045)	...	fl14
57699.210	+9.58	19.820 (0.093)	...	...	...	...	17.653 (0.073)	...	FLI
57700.195	+10.55	...	18.017 (0.091)	...	...	...	17.801 (0.111)	...	FLI
57700.890	+11.23	...	...	...	...	17.634 (0.100)	17.796 (0.101)	17.402 (0.089)	IO:O
57701.270	+11.60	20.074 (0.100)	18.110 (0.098)	...	...	17.632 (0.050)	17.830 (0.051)	...	fl04
57701.630	+11.96	20.106 (0.026)	18.094 (0.109)	...	19.318 (0.075)	17.709 (0.069)	...	...	fl11
57701.785	+12.11	...	...	>19.174	...	...	...	17.425 (0.058)	fl16
57701.870	+12.19	...	18.138 (0.099)	...	...	17.706 (0.024)	17.902 (0.043)	...	IO:O
57702.235	+12.55	20.089 (0.093)	18.231 (0.087)	...	...	...	17.923 (0.097)	...	FLI
57702.405	+12.72	...	...	>19.542	...	...	...	17.430 (0.080)	fl12
57702.580	+12.89	20.158 (0.094)	18.168 (0.089)	...	...	17.724 (0.048)	...	...	fl11
57702.900	+13.21	20.228 (0.085)	18.232 (0.041)	...	...	17.751 (0.033)	17.948 (0.035)	17.475 (0.095)	IO:O
57703.225	+13.53	...	...	...	...	17.796 (0.134)	...	...	fl04
57704.320	+14.60	...	...	...	...	17.899 (0.086)	18.014 (0.092)	...	fl04
57706.860	+17.10	...	...	...	...	17.952 (0.090)	...	...	fl06
57706.890	+17.13	...	...	...	19.570 (0.102)	17.993 (0.101)	18.083 (0.101)	...	fl16
57707.290	+17.53	...	18.487 (0.115)	...	...	17.983 (0.084)	...	...	fl15
57707.815	+18.05	20.397 (0.030)	18.491 (0.117)	...	19.620 (0.104)	18.101 (0.102)	18.101 (0.102)	...	fl06
57708.030	+18.26	...	18.501 (0.084)	...	...	18.117 (0.048)	18.113 (0.062)	17.540 (0.088)	IO:O
57708.885	+19.10	...	...	...	...	18.122 (0.024)	18.121 (0.025)	17.584 (0.079)	IO:O
57709.870	+20.07	...	...	...	...	18.162 (0.030)	18.181 (0.046)	17.611 (0.088)	IO:O
57710.180	+20.37	...	...	...	...	18.176 (0.097)	...	...	fl04
57710.265	+20.46	...	...	>19.462	...	...	...	17.613 (0.062)	fl03
57710.850	+21.03	...	...	...	...	18.218 (0.022)	18.227 (0.032)	17.617 (0.089)	IO:O
57711.205	+21.38	...	...	...	...	...	18.281 (0.095)	...	fl04
57711.250	+21.42	...	18.671 (0.112)	...	...	18.205 (0.086)	18.240 (0.092)	...	fl03
57711.790	+21.96	...	...	...	...	...	...	17.639 (0.090)	fl06
57711.870	+22.04	20.476 (0.030)	18.734 (0.105)	...	19.762 (0.102)	18.308 (0.057)	18.285 (0.067)	...	fl14
57711.875	+22.04	...	...	...	...	18.234 (0.028)	18.254 (0.041)	...	IO:O
57712.845	+22.99	...	...	...	...	18.315 (0.029)	18.303 (0.033)	17.620 (0.066)	IO:O
57713.160	+23.30	20.521 (0.083)	18.760 (0.072)	...	...	18.262 (0.059)	18.281 (0.062)	...	fl04
57713.865	+24.00	...	...	...	...	18.301 (0.029)	18.341 (0.033)	17.672 (0.084)	IO:O
57714.505	+24.63	20.566 (0.087)	18.810 (0.095)	...	...	18.326 (0.061)	18.369 (0.066)	...	fl11
57715.155	+25.27	...	18.805 (0.123)	...	...	...	18.459 (0.084)	...	FLI
57717.175	+27.26	20.581 (0.097)	18.892 (0.083)	...	...	18.410 (0.053)	18.511 (0.066)	...	fl04
57717.850	+27.92	20.603 (0.090)	18.929 (0.059)	...	19.995 (0.080)	18.425 (0.042)	18.529 (0.053)	...	fl14
57720.820	+30.84	...	...	...	...	...	...	17.905 (0.053)	fl16
57721.065	+31.09	20.626 (0.028)	18.990 (0.126)	...	20.074 (0.115)	18.507 (0.095)	18.729 (0.101)	...	fl14
57722.210	+32.21	...	...	...	...	18.542 (0.113)	18.694 (0.115)	...	fl04
57726.820	+36.75	...	...	...	...	...	...	17.993 (0.057)	fl06
57726.840	+36.77	20.704 (0.032)	...	...	20.164 (0.102)	18.604 (0.101)	18.841 (0.102)	...	fl16
57728.050	+37.96	...	19.070 (0.099)	>19.978	20.178 (0.053)	18.677 (0.074)	18.879 (0.047)	...	ALFOSC_FASU
57728.130	+38.04	...	19.095 (0.174)	...	...	...	...	...	FLI
57729.545	+39.43	...	...	...	20.192 (0.104)	18.723 (0.102)	...	...	fl11
57733.885	+43.70	20.796 (0.032)	19.185 (0.125)	...	20.258 (0.102)	18.766 (0.086)	19.021 (0.094)	18.205 (0.061)	fl06
57736.910	+46.68	20.810 (0.034)	19.230 (0.127)	...	...	18.850 (0.101)	19.102 (0.102)	18.291 (0.092)	fl16
57741.155	+50.86	20.845 (0.101)	19.305 (0.077)	...	20.284 (0.084)	18.907 (0.056)	19.201 (0.039)	18.463 (0.086)	fl04
57744.930	+54.57	20.897 (0.033)	19.388 (0.128)	...	...	...	...	...	fl16
57745.155	+54.79	20.895 (0.043)	19.393 (0.122)	...	20.310 (0.103)	19.017 (0.104)	19.313 (0.097)	...	fl03
57746.170	+55.79	20.911 (0.040)	...	...	...	...	...	...	fl14
57747.105	+56.71	20.936 (0.086)	19.419 (0.076)	...	20.332 (0.082)	19.099 (0.079)	19.347 (0.072)	...	fl04
57752.870	+62.39	20.959 (0.096)	19.453 (0.093)	...	20.378 (0.082)	19.250 (0.101)	19.563 (0.102)	...	fl14
57755.520	+64.99	20.957 (0.051)	...	...	...	...	...	...	fl12
57756.155	+65.62	20.974 (0.052)	19.488 (0.134)	...	...	...	...	...	fl04
57757.120	+66.57	20.980 (0.099)	19.498 (0.077)	...	...	...	...	...	fl03
57761.100	+70.49	...	...	...	20.470 (0.105)	19.496 (0.168)	19.676 (0.103)	...	fl04
57761.485	+70.86	...	19.544 (0.138)	...	...	...	...	...	fl12
57762.877	+72.24	...	...	...	...	19.540 (0.101)	19.718 (0.101)	...	fl06
57799.780	+108.55	...	...	...	...	20.458 (0.332)	...	...	fl16

Notes: fl03, LCOGT node at Cerro Tololo Inter-American Observatory; fl04, LCOGT node at Cerro Tololo Inter-American Observatory; fl06, LCOGT node at SAAO; fl11, LCOGT node at Siding Spring Observatory; fl12, LCOGT node at Siding Spring Observatory; fl14, LCOGT node at SAAO; fl15, LCOGT node at Cerro Tololo Inter-American Observatory; fl16, LCOGT node at SAAO; FLI, DEMONEXT; IO:O, Liverpool Telescope; ALFOSC\_FASU, Nordic Optical Telescope.

**Table C.4.** Near-infrared Photometry for SN 2016hnk. Epochs from peak are given in the rest frame.

<b>ID</b>	<b>Epoch</b>	<i>J</i>	<i>H</i>	<i>K</i>	<b>Source</b>
57695.193	+5.62	16.555 (0.139)	16.483 (0.186)	16.366 (0.124)	SOFI
57698.163	+8.55	16.694 (0.132)	16.573 (0.136)	16.448 (0.191)	SOFI
57705.030	+15.30	17.059 (0.112)	17.005 (0.137)	16.915 (0.121)	NOTCAM
57713.220	+23.36	17.320 (0.156)	17.239 (0.100)	17.141 (0.104)	SOFI
57722.957	+32.95	17.658 (0.137)	17.495 (0.154)	17.358 (0.172)	NOTCAM
57762.930	+72.29	19.995 (0.176)	19.193 (0.161)	18.604 (0.158)	NOTCAM

Notes: SOFI, New Technology Telescope; NOTCAM, Nordic Optical Telescope.

Table C.5. Journal of spectroscopic observations for SN 2016hmk.

UT date	MJD	Epoch	Exposure	Airmass	Telescope	Spectrograph	Grism	Range	Resolution
2016-10-29 06:51:58.975	57690.29	+0.80	900	1.30	HILTNER	OSMOS	VPH Blue	3960-6880	1600
2016-10-30 06:03:08.180	57691.25	+1.74	1200	1.19	SOAR	Goodman	G400	3600-7040	1850
2016-10-31 00:20:51.571	57692.01	+2.49	2600	1.26	SALT	RSS	PG0900	3500-9400	900
2016-11-01 01:52:29.692	57693.08	+3.54	1800	1.30	NOT	ALFOSC	Gr#4	3400-9400	700
2016-11-02 03:21:06.248	57694.14	+4.59	600	1.09	NTT	EFOOSC2	Gr#13	3640-8930	2400
2016-11-02 21:32:08.720	57694.90	+5.33	1500	1.74	LT	SPRAT	Wasatch600	4020-6950	1300
2016-11-04 01:57:23.700	57696.08	+6.49	12x150	1.245	GEMINI	FLAMINGOS-2	G5801	9900-18000	900
2016-11-04 15:10:05.940	57696.63	+7.04	2100	1.30	LJT	YFOSC	G3	3500-9170	2000
2016-11-05 01:59:35.148	57697.08	+7.48	1800	1.35	NTT	EFOOSC2	Gr#11+Gr#16	3340-9650	3600
2016-11-06 02:05:03.429	57698.09	+8.47	3240	1.19	NTT	SOFI	Gr#BG	9340-16480	1000
2016-11-07 00:05:48.541	57699.00	+9.37	1800	1.30	NOT	ALFOSC	Gr#4	3400-9700	700
2016-11-07 02:06:58.316	57699.09	+9.46	1800	1.15	NTT	EFOOSC2	Gr#11+Gr#16	3340-9980	3600
2016-11-09 01:34:17.864	57701.07	+11.41	1900	1.38	LT	SPRAT	Wasatch600	4020-8000	1300
2016-11-09 17:21:16.187	57701.72	+12.04	2100	1.30	LJT	YFOSC	G3	3500-9160	2000
2016-11-09 23:05:53.854	57701.96	+12.28	2850	1.26	LT	SPRAT	Wasatch600	5170-7900	1300
2016-11-11 00:59:54.332	57703.04	+13.34	2850	1.33	LT	SPRAT	Wasatch600	5220-7920	1300
2016-11-18 02:00:24.765	57710.08	+20.27	1800x2	1.09	NTT	EFOOSC2	Gr#11+Gr#16	3340-9980	3600
2016-11-28 04:39:23.638	57720.19	+30.22	3600+2x1800	1.74	NTT	EFOOSC2	Gr#13	3640-9220	2400
2016-12-03 03:46:45.336	57725.16	+35.11	2700x3	1.41	NTT	EFOOSC2	Gr#13	3640-9230	2400
2016-12-04 07:22:48.000	57726.31	+36.24	1200	1.73	MMT	Blue-Channel	G300	4440-8490	740
2016-12-06 19:49:23.524	57728.83	+38.72	3600	1.60	NOT	ALFOSC	Gr#4	3400-9700	700
2016-12-19 04:05:09.167	57741.17	+50.87	2700x2	1.84	NTT	EFOOSC2	Gr#13	3640-9220	2400
2016-12-28 03:49:27.464	57750.16	+59.72	2700+1900	2.12	NTT	EFOOSC2	Gr#13	3630-9220	2400
2017-01-06 01:14:34.082	57759.05	+68.47	2700x2	1.30	NTT	EFOOSC2	Gr#13	3640-9230	2400
2017-10-16 03:21:08.439	58042.14	+347.07	3060(UV)+3000(Opt)+3000(NIR)	1.16	VLT	X-Shooter	...	3000-25000	4100/6500/5600
2017-10-17 03:25:50.270	58043.14	+348.06	3060(UV)+3000(Opt)+3000(NIR)	1.13	VLT	X-Shooter	...	3000-25000	4100/6500/5600
2017-10-17 04:35:12.328	58043.19	+348.11	3060(UV)+3000(Opt)+3000(NIR)	1.05	VLT	X-Shooter	...	3000-25000	4100/6500/5600
2017-11-10 02:35:30.919	58067.11	+371.65	3060(UV)+3000(Opt)+3000(NIR)	1.07	VLT	X-Shooter	...	3000-25000	4100/6500/5600
2017-11-10 03:46:08.985	58067.16	+371.70	3060(UV)+3000(Opt)+3000(NIR)	1.05	VLT	X-Shooter	...	3000-25000	4100/6500/5600
2017-11-15 04:31:21.627	58072.19	+376.65	3060(UV)+3000(Opt)+3000(NIR)	1.12	VLT	X-Shooter	...	3000-25000	4100/6500/5600

HILTNER, Hiltner Telescope 3.58m, LT, Liverpool Telescope 2.0m, GEMINI, Gemini South Telescope 8.1m, LJT, Lijiang Telescope 6.5m, VLT, Very Large Telescope 8.2m, SOAR, Southern Astrophysical Research Telescope 4.1m, SALT, Southern African Large Telescope 11m, NOT, Nordic Optical Telescope 2.6m, NTT, New Technology Telescope 3.58m, LT, Liverpool Telescope 2.0m, GEMINI, Gemini South Telescope 8.1m, LJT, Lijiang Telescope 6.5m, VLT, Very Large Telescope 8.2m.



**Table C.6.** Measured velocities (in units of  $10^3 \text{ km s}^{-1}$ ). Epochs are measured from the *B*-band maximum brightness.

Epoch	Ca II H&K	Fe II	S II W1	S II W2	Si II 5972Å	Si II 6355Å	O I	Ca II NIR
+0.80	–	12.74(0.51)	5.89(0.16)	5.51(0.19)	7.83(0.16)	10.23(0.18)	–	–
+1.74	18.14(0.49)	12.46(0.14)	5.61(0.05)	5.16(0.06)	7.42(0.20)	10.02(0.12)	–	–
+2.49	17.47(0.82)	12.33(0.11)	5.26(0.07)	5.01(0.06)	7.40(0.10)	9.59(0.07)	11.21(0.04)	15.34(0.07)
+3.54	17.46(1.78)	12.67(0.14)	4.97(0.13)	4.66(0.12)	6.68(0.31)	9.51(0.20)	11.07(0.04)	15.03(0.07)
+4.59	17.73(1.84)	12.73(0.36)	4.85(0.25)	4.60(0.37)	6.70(0.64)	9.42(0.38)	11.18(0.09)	14.93(0.22)
+5.33	–	13.04(0.87)	4.81(0.85)	4.55(1.07)	5.73(0.27)	9.34(0.79)	11.06(0.12)	–
+7.04	16.92(1.48)	13.34(0.31)	4.62(0.21)	4.39(0.36)	5.34(0.36)	9.02(0.22)	10.91(0.09)	14.52(0.15)
+7.48	17.28(0.55)	13.47(0.14)	4.08(0.14)	4.12(0.26)	4.84(0.61)	8.76(0.29)	10.90(0.06)	14.82(0.10)
+9.37	16.33(1.55)	17.19(0.25)	3.89(0.15)	3.81(0.35)	4.88(0.27)	8.71(0.21)	10.74(0.05)	14.72(0.07)
+9.46	16.53(1.83)	17.52(0.17)	3.64(0.12)	3.57(0.21)	–	8.64(0.23)	10.83(0.04)	14.76(0.06)
+11.41	–	13.73(1.73)	3.69(1.30)	3.16(1.72)	–	8.48(1.01)	10.53(0.14)	–
+12.04	15.72(1.19)	13.76(0.62)	3.41(0.35)	3.22(0.52)	–	8.34(0.26)	10.61(0.11)	14.37(0.14)
+12.28	–	–	3.02(1.06)	3.12(1.41)	–	8.04(0.68)	10.80(0.34)	–
+13.34	–	–	2.33(1.47)	2.99(1.22)	–	8.04(0.38)	10.49(0.16)	–
+20.27	15.51(1.88)	13.27(0.50)	–	–	–	7.09(0.23)	10.10(0.07)	14.22(0.12)
+30.22	14.57(0.84)	10.04(1.28)	–	–	–	6.52(0.25)	10.37(0.20)	14.01(0.17)
+35.11	14.52(0.36)	9.43(1.54)	–	–	–	6.08(0.20)	9.76(0.19)	13.55(0.08)
+36.24	–	8.61(1.24)	–	–	–	5.73(0.57)	10.26(0.21)	–
+38.72	15.42(1.80)	8.72(1.42)	–	–	–	5.76(0.21)	10.27(0.07)	13.66(0.07)
+50.87	14.50(1.56)	7.25(1.20)	–	–	–	5.30(0.34)	11.59(0.67)	13.15(0.24)
+59.72	13.88(1.52)	5.54(1.77)	–	–	–	5.05(1.41)	10.03(0.82)	12.79(0.19)
+68.47	12.09(1.25)	4.77(1.97)	–	–	–	–	9.64(2.85)	12.59(0.42)

**Table C.7.** Measured pseudo-equivalent widths (in Å). Epochs are measured from the *B*-band maximum brightness.

Epoch	Ca II H&K	Fe II	S II W1	S II W2	Si II 5972Å	Si II 6355Å	O I	Ca II NIR
+0.80	–	263.17(13.86)	10.34(2.01)	34.26(3.85)	66.39(6.45)	127.92(4.78)	–	–
+1.74	120.05(14.80)	281.95(4.54)	12.27(0.78)	34.37(1.51)	61.48(3.10)	132.88(3.16)	–	–
+2.49	119.88(12.61)	272.21(2.86)	10.92(0.75)	25.15(1.15)	53.96(1.47)	126.71(1.12)	110.08(2.37)	418.58(5.14)
+3.54	125.70(11.11)	267.05(4.59)	12.38(1.58)	24.22(2.40)	50.08(5.67)	132.54(4.21)	117.13(3.26)	496.48(4.12)
+4.59	121.70(11.54)	259.03(9.63)	10.95(3.64)	14.29(6.56)	45.46(6.68)	136.51(7.87)	120.91(7.96)	515.12(16.60)
+5.33	–	248.69(37.93)	12.69(6.52)	13.13(7.30)	42.64(6.10)	140.62(8.82)	123.50(5.85)	–
+7.04	118.09(19.43)	264.50(10.92)	16.57(3.44)	9.44(4.85)	36.85(5.67)	141.86(8.92)	132.73(6.14)	526.95(11.37)
+7.48	119.84(17.19)	269.35(5.41)	16.31(2.10)	8.97(2.98)	30.96(4.45)	148.26(5.46)	130.95(5.47)	533.60(7.44)
+9.37	104.04(12.02)	332.98(16.58)	14.97(2.89)	6.31(2.54)	28.58(6.58)	156.34(7.80)	130.80(4.11)	533.86(5.69)
+9.46	106.07(13.34)	326.13(13.35)	15.86(2.33)	5.79(1.74)	–	152.67(5.51)	126.98(2.79)	523.19(4.20)
+11.41	–	–	13.18(3.77)	5.80(1.07)	–	170.60(11.55)	115.76(8.71)	–
+12.04	99.36(19.81)	360.20(25.19)	16.16(5.83)	5.97(4.38)	–	183.31(11.39)	121.58(7.75)	531.55(10.94)
+12.28	–	–	16.04(3.70)	5.75(1.89)	–	175.90(11.35)	115.14(9.23)	–
+13.34	–	–	13.14(3.69)	6.03(1.52)	–	194.76(15.53)	115.96(8.84)	–
+20.27	104.64(14.39)	381.58(13.18)	–	–	–	207.95(5.89)	112.86(4.87)	563.62(8.96)
+30.22	108.01(10.98)	420.72(26.65)	–	–	–	181.10(13.57)	103.47(6.88)	605.20(12.11)
+35.11	105.92(11.02)	403.08(12.29)	–	–	–	168.78(7.64)	101.44(6.73)	602.40(6.01)
+36.24	–	389.52(23.55)	–	–	–	166.72(10.50)	90.80(5.09)	–
+38.72	92.52(18.40)	392.41(14.21)	–	–	–	168.73(5.28)	93.18(8.02)	607.72(3.96)
+50.87	90.27(10.21)	369.78(19.98)	–	–	–	143.79(13.15)	85.50(7.27)	640.54(17.37)
+59.72	129.27(11.52)	361.46(11.38)	–	–	–	125.17(14.53)	79.72(6.44)	657.21(14.31)
+68.47	122.63(18.45)	306.01(20.03)	–	–	–	–	74.19(7.12)	650.72(33.03)