

OzDES multi-object fibre spectroscopy for the Dark Energy Survey: Results and second data release

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ABSTRACT

We present a description of the Australian Dark Energy Survey (OzDES) and summarise the results from its six years of operations. Using the 2dF fibre positioner and AAOmega spectrograph on the 3.9-m Anglo-Australian Telescope, OzDES has monitored 771 AGN, classified hundreds of supernovae, and obtained redshifts for thousands of galaxies that hosted a transient within the 10 deep fields of the Dark Energy Survey. We also present the second OzDES data release, containing the redshifts of almost 30,000 sources, some as faint as $r_{AB} = 24$ mag, and 375,000 individual spectra. These data, in combination with the time-series photometry from the Dark Energy Survey, will be used to measure the expansion history of the Universe out to $z \sim 1.2$ and the masses of hundreds of black holes out to $z \sim 4$. OzDES is a template for future surveys that combine simultaneous monitoring of targets with wide-field imaging cameras and wide-field multi-object spectrographs.

Key words: transients: supernovae; quasars: supermassive black holes; cosmology: dark energy; surveys; catalogues; techniques: spectroscopic

1 INTRODUCTION

Over the next ten years, surveys using the next generation of multi-object spectroscopic facilities fed by rapidly con-

figurably fibre positioning systems will obtain the spectra of tens of millions of sources. Facilities such as the Subaru Prime Focus Spectrograph (PFS, Tamura et al. 2018)

on the 8.2-m Subaru telescope, the Dark Energy Spectroscopic Instrument (DESI, [DESI Collaboration 2016](#); [Vargas-Magana et al. 2019](#)) on the 4-m Mayall telescope, and the 4-metre Multi-Object Spectroscopic Telescope (4MOST, [de Jong et al. 2019](#)) on the 4-m VISTA telescope are either being built or on the cusp of entering operation. Other facilities, such as Mauna Kea Spectroscopic Explorer (MSE, [The MSE Science Team et al. 2019](#)) and MANIFEST ([Lawrence et al. 2018](#)) for the Giant Magellan Telescope, are being planned. Over the same time frame, the Legacy Survey of Space and Time (LSST; [LSST Science Collaboration 2017](#)) at the Rubin Observatory will start imaging the entire southern sky multiple times in multiple pass-bands over 10 years. The advent of the LSST in combination with these spectroscopic facilities will allow contemporaneous imaging and spectroscopic studies of tens of thousands of targets.

As a herald to these future surveys, OzDES¹ has used the 2dF fibre positioner on the Anglo-Australian Telescope (AAT) to obtain spectra of thousands of sources in the 10 deep fields of the Dark Energy Survey (DES) over the 6 years than DES ran. OzDES has two main scientific goals: i) constraining the dark energy equation of state parameter using type Ia supernovae (SNe Ia), and ii) measuring the mass of super-massive black holes over a broad range of redshifts using AGN reverberation mapping. Additionally, OzDES has obtained spectra for a number of ancillary projects, such as improving and quantifying the accuracy of redshifts obtained from broad-band photometry. In total, about 40% of the fibre hours not used for active supernovae, host galaxies and AGN were used for these projects.

This paper presents results from the full six years of OzDES, and describes the second OzDES data release². After giving a brief summary of the principal scientific aims of OzDES in Sec. 2, we provide an overview of OzDES operations for the last three years of the survey. Throughout the paper we will use Y4, Y5, and Y6 to denote the fourth, fifth and sixth years of OzDES. The observing strategy for the first three years of OzDES (Y1, Y2 and Y3) and the precursor survey are discussed in two papers: [Yuan et al. \(2015\)](#) gives an overview up to the end of Y1 (2013), including the OzDES precursor survey, and [Childress et al. \(2017\)](#) gives an overview up until the end of Y3 (2015). We will refer to these papers as Yu15 and Ch17 henceforth. The reliability and completeness of OzDES redshifts are analysed in Sec. 4, and the data release is described in Sec. 5. In the sections that follow, we examine various aspects of the survey, such as the frequency at which we were able to target AGN, and how the signal-to-noise ratio behaves with time for objects that have the longest integration times: the host galaxies of transients. Before giving the summary of the paper in the final section, we compare OzDES to TiDES, a OzDES-like survey using the 4MOST facility ([Swann et al. 2019](#)). All magnitudes listed in this paper are measured using 2'' diameter apertures and are on the AB magnitude system.

2 DES AND OZDES

The Dark Energy Survey (DES) was a six-year programme, ending in 2019, using the DECam imager ([Flaugher et al. 2015](#)) on the CTIO Blanco 4m telescope in Chile. DES consisted of two surveys: a wide-field survey covering 5,000 square degrees, and a deep survey consisting of 10 DECam fields located in several well-studied extra-galactic fields. These 10 fields, the coordinates of which can be found in Table 2 of [D'Andrea et al. \(2018\)](#), were targeted repeatedly in *griz* over the course of the first five years of DES with a cadence of approximately 6 days ([Diehl et al. 2018](#)), and consist of 8 relatively shallow fields and 2 extra deep fields (known as X3 and C3). The regular cadence facilitated the discovery and follow-up of transient sources, such as supernovae (SNe), and the monitoring of variable objects, such as AGN.

DES aims to place constraints on the dark energy equation of state parameter and deviations from General Relativity using four astronomical probes: weak gravitational lensing, galaxy clusters, the large scale distribution of galaxies, and type Ia supernovae. Results from small subset of the DES data set have been published in a number of papers, the most recent of which are constraints that combine supernovae, gravitational lensing, and galaxy clustering ([Dark Energy Survey 2019a](#)).

DES data have been used in many studies covering a diverse range of fields ([Dark Energy Survey 2016](#)). Notable examples include studies of bodies in the outer solar system ([Khain et al. 2018](#), for example), AGN accretion disk sizes ([Mudd et al. 2018](#); [Yu et al. 2020](#)), dwarf galaxies and stellar streams in the halo of the Milky Way ([Marshall et al. 2019](#); [Li et al. 2019](#), for example) and rare transients such as merging neutron stars ([Palmese et al. 2017](#)), super-luminous supernovae ([Smith et al. 2018](#); [Angus et al. 2019](#)), and rapidly evolving transients ([Pursiainen et al. 2018](#)).

Since 2012, OzDES has been targeting sources in the 10 DES deep fields with the 2dF fibre positioner on the AAT. As shown in Fig. 1 of [Yuan et al. \(2015\)](#), the fields of view of 2dF and DECam are very similar. In addition to the 10 DES deep fields, OzDES targeted the DES MaxVis field during one of the OzDES observing runs. Centred at RA=97.5° and Dec=-58.75°, it is observable during the entire DES observing season. It is one of the DES standard star fields that was observed during twilight. [Yu et al. \(2020\)](#) used quasars in this field to study continuum reverberation mapping.

The scientific aims of OzDES focus on two areas: i) using redshifts and distances of type Ia supernovae to measure the expansion history of the Universe and thereby constrain the dark energy equation of state parameter; and ii) measuring the time lags between the continuum and the broad lines in a sample of AGN to measure their central black hole masses and to investigate the possibility of using AGN as a luminosity distance indicator that would explore a redshift range that is currently beyond the reach of supernovae.

Over the six years that OzDES ran, OzDES obtained redshifts for 7,000 candidate supernova hosts ([D'Andrea et al. 2018](#)) and spectroscopically confirmed several hundred supernovae (SNe). These redshifts will be used with distances to the supernovae inferred from broad-band lightcurves obtained by DES to measure the expansion history of the Universe and to constrain the dark energy equation

¹ Australian (aka "Oz") Dark Energy Survey.

² Available from Data Central: <https://datacentral.org.au>. Future releases will be announced on Data Central

tion of state parameter. Using an initial sample of 207 spectroscopically confirmed type Ia supernovae from the first three years of data in combination with 122 low redshift supernova from the literature and constraints from the Cosmic Microwave Background (CMB), the [Dark Energy Survey \(2019b\)](#) obtained a 7% constraint on a constant dark energy equation-of-state parameter. Analysis of the full DES sample is ongoing, and the final cosmological results will include approximately an order of magnitude more supernovae, the majority of which will be classified photometrically.

At the same time, OzDES has monitored a sample of 771 AGN up to $z \sim 4$ with the goal of measuring black hole masses using the time lag between variations in the continuum, measured from the broad-band photometry obtained by DES, and the broad lines, measured from the spectra obtained by OzDES. The first time lags, based on the first four years of data, were published in [Hoormann et al. \(2019\)](#).

In addition to these two main projects, OzDES is facilitating the studies of a number of areas, both within DES and outside of DES, by obtaining redshifts of targets of interest in these 10 fields using the “spare” fibres that are not allocated to an AGN, an active transient, or a host galaxy. The most notable examples are: radio galaxies from the ATLAS survey ([Franzen et al. 2015](#)); brightest cluster galaxies ([Webb et al. 2015](#)); redMaGiC galaxies ([Rozo et al. 2016](#); [Prat et al. 2018](#)); and luminous red galaxies (LRGs), which are used in combination with redshifts from other sources to train photometric redshift algorithms ([Sánchez et al. 2014](#); [Bonnett et al. 2016](#); [Gschwend et al. 2018](#)).

3 OZDES OPERATIONS FOR Y4, Y5 AND Y6

Starting in 2013, OzDES ran for six observing seasons, with an initial allocation of 12 nights in Y1 (AAT semester 2013B), and allocations of 16, 20, 20, 20, and 12 nights in subsequent years. In 2012, when aspects of the DES survey strategy were being tested during a period of science verification (SV), 5 nights were allocated to an OzDES precursor survey. As for DES, these 5 nights were used to develop and test the observing strategy that would be used for the next six years.

Observations typically started in late August or early September and terminated four to five months later in December or January the following year. The observing log for Y4, Y5, and Y6 can be found in Appendix A and follows the format presented for earlier years in Yu15 and Ch17.

In total, there were 41 numbered runs, with the last night of run 41 occurring on 09 January 2019. Some of the nights during these runs were shared with other programmes, such as 2dFLens ([Johnson et al. 2017](#)) during Y2 and Y3, XXL ([Pierre et al. 2016](#)) in Y4, DEVILS ([Davies et al. 2018](#)) in Y5, and the HSC Transient Survey ([Yasuda et al. 2019](#)) in Y6. During Y2 and Y3, OzDES and 2dFLens shared a common run numbering scheme, and there are a number of runs during those years in which only 2dFLens fields were observed.

The end of Y5 marked the end of the DES programme to discover transients in the 10 DES deep fields. DES continued to target fields in the wide survey during Y6. OzDES also continued into a sixth year, but was no longer targeting transients. Instead, OzDES continued with its programme

to monitor AGN ([King et al. 2015](#); [Hoormann et al. 2019](#)) and to obtain redshifts of galaxies that hosted transients in earlier years. DES continued to target the 10 DES supernova fields during Y6, but with a reduced frequency of about once a month.

These data are used to spectrophotometrically calibrate the spectra of the AGN in the deep fields that were being monitored by OzDES.

3.1 Instrumental Setup

Apart from the final OzDES run in Y6 (run 041), the instrumental setup used in Y4, Y5 and Y6 was identical to the set up used in Y3. The set up in Y1 and Y2 and the precursor survey were slightly different. During these years, the central wavelength of the blue grating was set 20\AA to the blue. Details of the setup and the reasons for the change between Y2 and Y3 can be found in Ch17.

The final OzDES run was not officially scheduled as an OzDES run. At the beginning of 2019, an opportunity arose to take additional data in the C3 and X3 fields during time that was allocated to the HSC Transient Survey, which has similar aims to the SN programme in OzDES, but is targeting fainter, more distant host galaxies. For this run, and this run only, we used the setup used by the HSC Transient Survey, which uses the x6700 dichroic (OzDES uses the x5700 dichroic) and a wavelength setting that results in a wavelength coverage that is shifted about 1000\AA to the red, from 3700\AA – 8800\AA to 4700\AA – 9800\AA .

3.2 OzDES Target Allocation

OzDES targeted a wide range of sources over the six years it ran, with active transients, AGN, and host galaxies with $m_r < 24$ having the highest priority and occupying most of the fibres. A full listing of the source types that were observed in the last three years, together with their priorities, is provided in Table 1.

In addition to a priority, each type had a quota. Starting with the highest priority (priority 8), targets were randomly selected until their was no more targets to select or the quota for that type of source had been reached. Further details on the target selection algorithm are described in Y15.

If conditions were poor because of low transparency (generally more than 2 mag of extinction) and/or poor seeing (generally poorer than 3 arc-seconds) we switched to a backup programme that mostly consisted of bright galaxies. However, transients and AGN in the RM programme were always targeted.

Over the seasons, the composition of sources and their priorities changed. The biggest changes occurred at the end of Y5 and the beginning of Y6. In Y6, no new transients were being discovered, as DES stopped monitoring the 10 deep fields for transients; however, during Y6, we did continue to observe a small number of transients that were discovered in earlier years.

There were other changes too. At the end of Y5, the quota for SN hosts increased from 100 (200 for X3 and C3, the two extra deep fields) in the first five years to 200 (250 for X3 and C3), and two new classes of sources were added. Given the success at obtaining redshifts for sources as faint

Table 1. A listing of the source types observed during Y4, Y5, and Y6, together with the nominal priority. The list is ordered in priority from highest (8) to lowest (1). Also listed are source types (noted with asterisks) from the first three years of OzDES that were discontinued before the start of Y3. Source types for the MaxVis field are listed in Appendix D.

Type	Priority	Comment
Transient	8	Active transients
AGN	7	Reverberation mapping programme (Hoonmann et al. 2019)
SN_host ^a	6	Hosts with $m_r < 24$
Cooke_host	6	Host galaxy candidates of transients found in deep SNLS ^b stacks
StrongLens*	6	Candidate strong lenses (Nord et al. 2016; Jacobs et al. 2019)
White Dwarfs*	6	White dwarfs to aid calibration
DEVILS ^c	5	Objects in the DEVILS survey (Davies et al. 2018)
SNLS	5	SN hosts from the SNLS survey data (Astier et al. 2006; Betoule et al. 2014)
Tertiary*	5	Stars with a broad range of colours to aid calibration
Cluster Galaxies	4	Cluster Galaxies (some targets were observed at higher priority)
Radio Galaxies	4	Radio Galaxies (some targets were observed at higher priority)
QSO	4	Faint QSOs in the S1 and S2 fields
XXL_QSO	4	QSOs in the XXL fields (Pierre et al. 2016)
SN_host_faint	4	Hosts with $24 < m_r < 25$
RedMaGiC	4	For calibrating photometric redshifts of redMaGiC galaxies (Roza et al. 2016)
SpARCS ^d	4	Brightest Cluster Galaxies from the SpARCS survey (Wilson et al. 2009)
ELG*	4	Emission Line Galaxies
SN_free_host	3	Host galaxies of previously targeted transients
Photo-z	3	For training and testing photometric redshift algorithms
LRG	2	Luminous Red Galaxies
Bright galaxies	1	Back up targets during poor conditions
Bright stars*	1	Back up targets during poor conditions

^a The type "SN host" not only includes galaxies that hosted a SN, but also galaxies that hosted transients that were not SNe.

^b SNLS - Supernova Legacy Survey

^c DEVILS - Deep Extragalactic Visible Legacy Survey

^d SpARCS - Spitzer Adaptation of the Red-Sequence Cluster Survey

as $r_{AB} \sim 24$ and no strong evidence to a floor in the noise (Ch17), we added galaxies with $24 < r_{AB} < 25$. We labelled these as `SN_host_faint`. We had targeted hosts as faint as this in past seasons, but the aim then was different. Then, if we did not see an emission line that enabled us to get a redshift after one run, we discarded the source. During Y6, we continued observing these faint hosts until a secure redshift (see Sec. 4) was obtained.

We also added host galaxies of transients that were previously observed when the transient was still bright enough to contribute to the flux density. We labelled these as `SN_free_host`. We will use these spectra to examine the properties of host galaxies without light from the supernova, and to subtract the galaxy spectrum from earlier spectra that contained light from the both the galaxy and the transient. The aim here is to see if one can identify the type of transient in spectra where this was not possible. This has been proven to work for spectra taken with slits (Dawson et al. 2009), but has never been tried for spectra obtained with fibres.

Fibre allocations for the last three years of OzDES are listed in Table 2 and shown as plots in Fig 1. A similar table in for the first three years is provided in Ch17. After Y1, a quarter of all fibres were allocated to AGN that were being followed in the OzDES reverberation mapping. The number of fibres allocated to SN hosts steadily increased from 12% in Y1 to 39% in Y5. The large jump from Y5 to Y6 in the number of SN hosts reflects the increased quota for SN hosts in Y6 and the addition of two new SN host categories.

Almost all fibres in Y6 were allocated to an AGN or a SN host.

Fibre allocations over the full OzDES survey are shown in Fig 2. The three most commonly observed sources were SN hosts (35%), AGN (26%) and galaxies used to train photometric redshift algorithms: LRGs, ELGs, and redMaGiC galaxies (21%).

There can be considerable overlap in the individual target catalogues that make up the OzDES target input catalogue. For example, a SN host can also be listed as a radio galaxy, LRG, or redMaGiC galaxy. In Table 2, Fig 1, and Fig 2 objects retain the type that has the highest priority. For example, a source that is listed in the OzDES target input catalogue as a SN host, radio galaxy, an LRG, and a redMaGiC galaxy will be classified and counted as a SN host.

3.3 Observing Strategy, Calibration and Data Reduction

During each run, we aimed to get at least two 2400 second exposures per field. This was not always possible due to the length of some runs and weather conditions. Of the 31 runs that were scheduled for OzDES, we were able to observe all 10 fields on 16 occasions. There were only two, relatively short runs where we did not observe any fields because of poor weather.

Overall, $\sim 25\%$ of time was lost to poor observing conditions. Another $\sim 5\%$ of the time was sub-optimal (thick

Table 2. Fibre allocations for the last three years of OzDES (Y4, Y5, and Y6)

Object	Y4		Y5		Y6	
	Fibre Hours	Fraction of Fibre Hours	Fibre Hours	Fraction of Fibre Hours	Fibre Hours	Fraction of Fibre Hours
Transients	939	2.7%	1195	3.2%	29	0.1%
AGN	8191	23.2%	8578	23.1%	6115	22.8%
SN hosts	13333	37.8%	14413	38.9%	18513	69.1%
Cluster Galaxies	1219	3.5%	1048	2.8%	449	1.7%
Radio Galaxies	1989	5.6%	1455	3.9%	0	0.0%
DEVILS	0	0.0%	720	1.9%	0	0.0%
RedMaGiC	2220	6.3%	3481	9.4%	0	0.0%
LRGs	4235	12.0%	3137	8.5%	0	0.0%
Other	3107	8.8%	3071	8.3%	1688	6.3%
Total	35232		37097		26793	

Table 3. Fibre allocations for the entire OzDES survey

Object	All years	
	Fibre Hours	Fraction of Fibre Hours
Transients	4438	2.3%
AGN	49548	25.5%
SN hosts	68028	35.0%
Cluster Galaxies	5523	2.8%
DEVILS	720	0.4%
Radio Galaxies	9271	4.8%
RedMaGiC	8394	4.3%
F stars	6247	3.2%
LRGs	25521	13.1%
ELGs	6175	3.2%
Other	3043	1.6%
Bright Galaxies	7298	3.8%
Total	194207	

cirrus and/or poor seeing), and was used for the back up programme. Observing logs for the final three years of OzDES, including the amount of time lost for poor weather, are included in the appendixes.

The two extra deep fields, C3 and X3, had the highest priority, so these fields were the most frequently observed, as can be seen in Fig. 3. The two E fields, E1 and E2 were also observed frequently, since these fields were often the only ones visible at the beginning of the night during the start of the observing season and are the most southern of the ten deep fields, which results in a longer window of observability from the AAT. Next in priority were the six remaining fields: X1, X2, C1, C2, S1, and S2.

In addition to the sources listed in Table 1, about a dozen fibres per field were allocated to F stars, and another 25 “sky” fibres were placed on positions that were free of sources. The targeting priority for both the F stars and the “sky” fibres was set to 5, which is below that of Transients, AGN and SN hosts, but higher than other types of objects. See Table 1 for details.

The F stars were primarily used to estimate throughput. The throughput was measured by integrating the extracted and wavelength calibrated spectra through the DECam filter band passes (Flaugher et al. 2015). For the red arm, we used the DECam *r* filter; for the blue arm, we used the DE-

Cam *g* filter. The relative throughput of the red arm over the course of OzDES is shown in Fig. 4. The throughput is reported in magnitudes, so each tick on the vertical access corresponds to a difference of about 2.5 in throughput. The vertical scatter is largely driven by transparency and atmospheric seeing.

Spectra that were taken when the throughput was excessively low (a couple of magnitudes lower than the medians shown in Fig. 4) are not used in the coadded spectra; however, they are included in the data release. For such spectra, the QC keyword in the FITS header is given the value `poor-Conditions`.

In the early years, we used the F-stars to derive the transfer functions for the red and blue arms that were then used throughout the survey, and to develop techniques that used the broadband DES photometry to spectrophotometrically calibrate the spectra of AGN used for the OzDES RM project (Hoormann et al. 2019).

All data were processed with the OzDES pipeline, which uses a modified copy of v6.46 of 2dfdr³ (Croom et al. 2004) and our own bespoke python scripts. The steps are described in Yu15 and Ch17, so we do not repeat them here. Since Ch17, a small number of improvements to the pipeline have been made, the most significant of which were a more robust algorithm to determine the factor that scales spectra in the blue and red arms, modifications to the PCA sky subtraction algorithm to make it less sensitive to cosmic rays that were not fully masked, and more complete masking of artifacts. These artifacts are described in section C3.

All data were processed at the telescope, and combined with data taken during earlier observing runs. Once processed, all raw and reduced files were then archived on a server in Canberra. Typically, this was done daily, after the night had ended, but before the start of the next night. After each field was processed, the data was made available for OzDES team members⁴ dedicated to the task of measuring redshifts within 24 to 48 hours of the data being taken. With the exception of transients and AGN, sources with secure redshifts (see the next section) were removed from the target catalogues and not re-observed during subsequent nights. During some of the longer OzDES runs, we observed

³ <https://aat.anu.edu.au/science/software/2dfdr>

⁴ Colloquially referred to as “redshifters.”

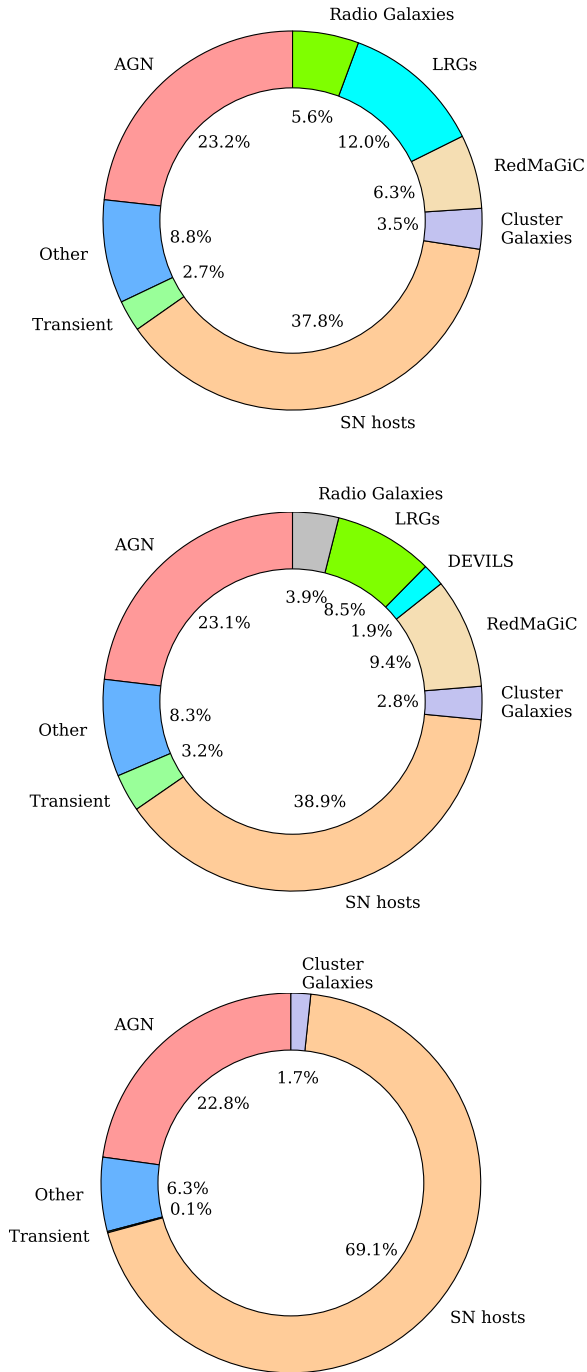


Figure 1. Plots showing how the allocation changed in the last three years of OzDES: Y4 at the top, Y5 in the middle, and Y6 at the bottom. The percentages represent the fraction of fibre hours. In the final year almost all targets were either AGN or SN hosts.

a number of the DES fields on multiple occasions. If this occurred, then all objects that were observed on the first occasion would be given higher priority in subsequent observations, unless they had a secure redshift and had been removed from the target catalogue. We have reprocessed all the data with the same version of the pipeline on a couple

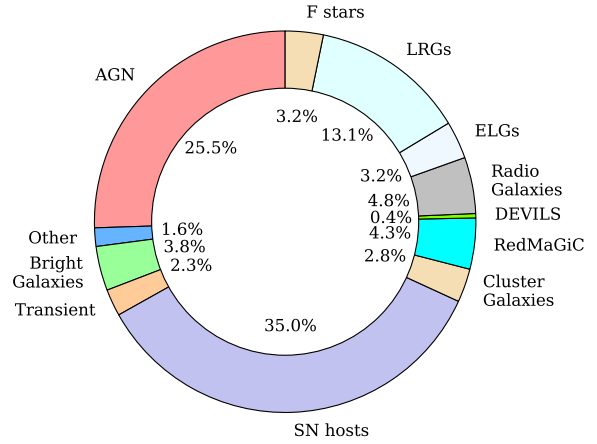


Figure 2. Plot showing the relative allocation for the full six years of OzDES. The percentages represent the fraction of fibre hours.

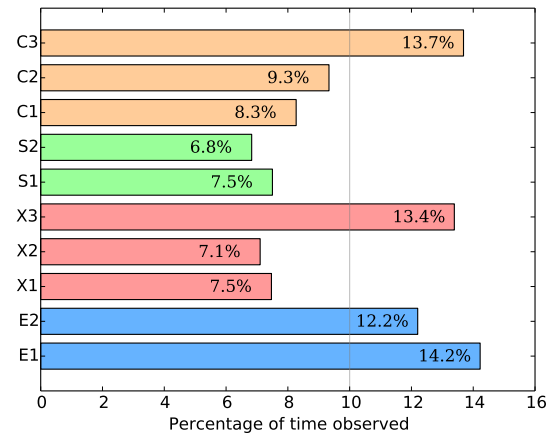


Figure 3. Bar plot showing the percentage of time each field was observed. The vertical line sits at 10%, which is where all fields would land if they had been observed with the same frequency. The X3, C3, E1, and E2 fields were the ones most frequently observed, reflecting the priorities these fields had and their locations in the sky.

of occasions. This ensures that all data are processed in a consistent manner. The latest version of the OzDES pipeline is version 18.21, and all data presented in this paper have been processed with that version.

4 REDSHIFTS

As noted in Ch17, redshifts have been measured using MARZ (Hinton et al. 2016) since the start of Y3. Earlier than this RUNZ, which was developed by Will Sutherland, had been used. Measuring redshifts is a task that is performed by two redshifters, whose results are collated, scrutinised and then merged by the chief redshift whip. Merging is also done using MARZ. The merged results then enter the OzDES database, from which all other products are derived.

OzDES uses a redshift flag to quantify the confidence

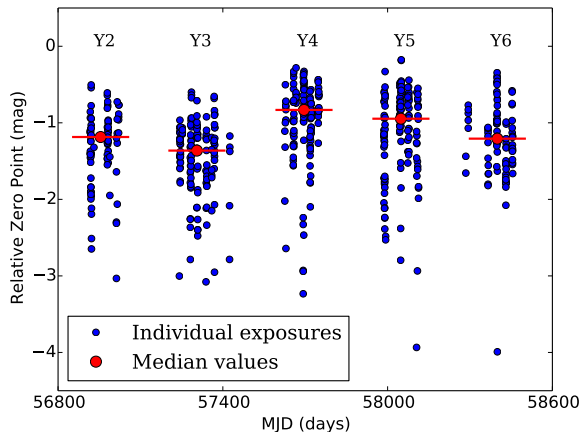


Figure 4. Plot of the relative throughput (zero point) of data taken from Y2 to Y6. The median value for each year is shown as the red dot and the red line. The scale is in magnitudes and more positive numbers indicate higher throughput. In Y1, very few F-stars were observed. Consequently there are no reliable throughput measurements for the first year of OzDES. The relative zero point plotted here differs from the zero point recorded in the FITS headers by an additive constant of 33. See footnote d to Table C3 for details on how the zero point is computed.

that the reported redshift is correct. As reported in Yu15 and Ch17, the flag can take on five main values.

- $Q = 4$, redshift based on multiple strong spectroscopic features matched, $> 99\%$ confidence.
- $Q = 3$, redshift based typically on a single strong spectroscopic feature or multiple weak features, $> 95\%$ confidence.
- $Q = 2$, potential redshift associated with typically a single weak feature, low confidence, not to be used for science.
- $Q = 1$, no matching features, thus no constraints on redshift.
- $Q = 6$, securely classified star.

As the survey progressed, we monitored the growing redshift catalogue for sources that had conflicting redshifts. These are sources in which the redshift changes between observing runs but the redshift quality flag, Q , remains unchanged and $Q > 2$. An example is a source that is assigned a $Q = 3$ redshift during one run, but with the addition of more data in a later run, is assigned a different redshift ($\Delta z > 0.05$) with the same quality flag.

On two occasions, once at the end of Y4 once at the end of Y6, we collated all cases where there was conflict, defined on these occasions having a redshift difference larger than 0.05, and reprocessed and re-redshifted these sources in the same way we redshifted sources during each observing run. The revised redshifts were then added to our redshift catalog. At the end of Y4, there were a total of 191 sources that had conflicting redshifts. At the end Y6, there were a total of 49, perhaps illustrating a degree of improvement by the human redshifters.

In preparation for the second OzDES data release, we revisited cases with conflicting redshifts, this time decreasing the threshold from 0.05 to 0.01 in redshift. We also included sources that were classified with $Q = 3$ quality redshift at

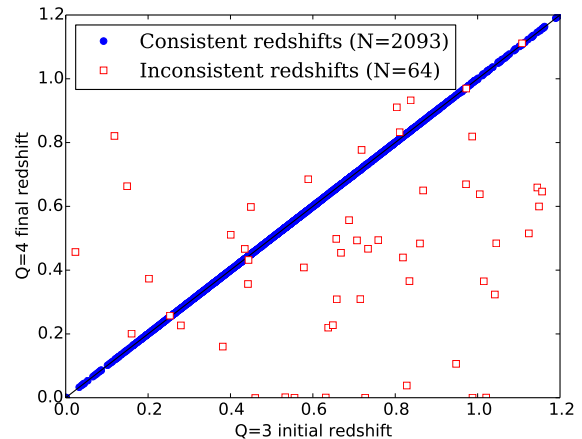


Figure 5. Redshift comparison for OzDES SN hosts that attained a $Q = 3$ quality redshift before attaining a $Q = 4$ quality redshift after further observations. Outliers (defined as $\Delta z > 0.003$) are shown as red squares. Sources with concordant redshifts are shown as blue points. The black line represents the one-to-one relationship. Only 3.1% of sources changed redshifts when moving from $Q = 3$ to $Q = 4$.

some stage during the survey but ended up with a lower redshift quality flag once all data were combined. Combined, about 820 sources were re-examined.

4.1 Redshift Reliability

Using data taken up until the end of Y1, Yu15 estimated the reliability of redshifts that are assigned $Q = 3$ and $Q = 4$ using external catalogues and repeat observations of the same target. They find that $Q = 4$ quality redshifts are 100% reliable and $Q = 3$ quality redshifts are $\sim 90\%$ reliable. After Y1, a number of measures were put in place to lift the reliability of the $Q = 3$ quality redshifts. Our target for $Q = 3$ quality redshifts at the start of the OzDES survey was $\sim 95\%$. As discussed below, we have reached this goal.

Ch17 used a different approach to assessing redshift reliability, taking advantage of the observing strategy applied to SN hosts. Unlike other target types, SN hosts were not removed from the input target catalogues until they attained a $Q = 4$ quality redshift. For those SN hosts with $Q = 4$ quality redshifts that initially attained a $Q = 3$ quality redshift, one can compare the initial redshift with the final redshift and get an estimate of reliability of the $Q = 3$ quality redshifts, assuming that the $Q = 4$ quality redshifts have a known reliability, which we take to be 100%, following that found in Yu15.

Using data up until the end of Y3, Ch17 found that the redshifts of 17 out of 443 sources (3.8%) changed between the initial $Q = 3$ quality redshift and the final $Q = 4$ quality redshift. We repeated this analysis, using a sample that is almost five times larger. We find that 64 of 2093 sources (3.1%) changed redshifts, a slightly smaller fraction than found in Ch17. The results are shown in Fig. 5.

The SN hosts allow one to examine the reliability of the $Q = 3$ quality redshifts in a second, complimentary way. A number of the fainter SN hosts have been observed over several years and 20 or more observing runs. As more data

were obtained and added to earlier data, they would have been re-redshifted, sometimes twice or more per run, which can occur if they were observed over multiple times during a run. Some sources have dozens of redshift measurements. Every time a source was redshifted, there is a chance of an erroneous redshift being assigned to that source. We can get a rough estimate of how often this happens by searching for objects that obtained a $Q = 3$ quality redshift during one run, but did not subsequently obtain a $Q = 3$ or $Q = 4$ quality redshift with more data.

We search the database for examples where a $Q = 3$ quality redshift was assigned once and not assigned again, even if the exposure time had increased by a factor of two or three by the end of Y6. For these two cases, we identified 31 and 11 SN hosts, respectively. For these sources, we reduce the redshift quality flag from $Q = 3$ to $Q = 2$. However, this can only be applied to SN hosts, as all other objects are deselected from the input target catalogue once a $Q = 3$ quality redshift is obtained. There are 1100 SN hosts with a $Q = 3$ quality redshift, potentially indicating an error rate in addition to that noted above of 1-3%.

The SN hosts also allow one to get a handle on the repeatability of the redshift measurement. We first selected SN hosts that have at least four redshift measurements with a quality flag of 3 or 4. We then compute the RMS of the redshifts for each host. The median RMS is 0.00015. While this is indicative of the redshift uncertainty, it is likely that it is a lower limit, as the measurements are not fully independent: redshifts derived from data taken later include earlier data. Indeed, the median RMS is a factor of two smaller than the uncertainty reported in Yu15, who report 0.0003 for SN hosts using redshifts that were obtained from data that were fully independent.

The impact of a redshift uncertainty of 0.0003 on the dark energy equation of state parameter is negligible if the uncertainty is not systematic. If the uncertainty is systematic, then the dark energy equation of state parameter could be biased (Calcino & Davis 2017).

4.2 Redshift Outcomes

The diversity of objects that were observed over the six years by OzDES is large; from radio sources covering a broad range in magnitude and redshift, to more restricted classes of objects, such as LRGs, covering smaller magnitude and redshift ranges.

In Figs 6, 7, and 8, we provide magnitude and redshift distributions for RedMaGiC galaxies, LRGs and radio galaxies. As noted at the end of Sec. 3.2, some targets had multiple types. An LRG may have been classified as a redMaGiC galaxy and visa-versa. About 40% of all sources observed by OzDES had multiple types. In these plots, we include all objects of the indicated type, even if it also had another type.

Also shown as an insert to these figures and Fig. 10 is the redshift completeness, defined as the number of objects with a redshift divided by the number of objects that were targeted. A number of targets were spectroscopically classified as stars. These were removed before computing the redshift completeness.

4.3 Redshift Completeness

4.3.1 Completeness as a function of magnitude and exposure time

A little more than one third of the 192,000 fibre hours that were spent observing all targets were spent on SN hosts (see Table 3). SN hosts were targeted during the entire duration of OzDES and were not deselected from the input target catalogue until a $Q = 4$ quality redshift was obtained. They are therefore useful in examining how redshift completeness depends on magnitude and exposure time. However, caution is needed before applying the results from SN hosts to the general population of field galaxies, because SN hosts are a biased subset. SN hosts are more likely to be $z < 1.35$, and therefore within the spectral range covered by OzDES.

All SN hosts were selected on the basis of the detection of a transient in the DECam images. With the exception of super-luminous supernovae, which are rare, there would be few SN hosts that are at redshifts that places the [OII] $\lambda\lambda 3727\text{\AA}$ doublet beyond the red end of the spectral range covered by the setting we used in AAOmega, which is at 8800\AA . The red end corresponds to a redshift of $z \sim 1.35$ for [OII]. Nearly all of the SN Ia detected by DES will be below this redshift (Bernstein et al. 2012). Hence, when interpreting the completeness of OzDES as a function of magnitude, the completeness will be different than for a survey that targeted field galaxies.

The redshift completeness of SN hosts as a function of exposure time is shown in Fig. 9. The completeness is shown for different magnitude bins and two redshift quality flags. As expected, completeness increases with exposure time and is highest for the brightest sources. It is also higher for lower quality redshift flags.

For the faintest bin ($23.5 < r_{AB} < 24$), the redshift completeness reaches $\sim 95\%$ for quality 3 redshifts after 40 hours of integration. This does not mean that we reached this level of completeness in the survey, as not all SN hosts have been observed for that long. We anticipate that if we were to observe this long for all sources, then we would reach this level of completeness. As shown in Fig 10 the completeness in the faintest bin is 65%.

4.3.2 Completeness as a function of plate position

The positioning of fibres relies on a series of transformations between sky coordinates and plate coordinates. The transformations include the `poscheck`, which is updated by observatory staff, typically during the first night after 2dF is remounted on the telescope, and the `plate survey`, which is run every time a field is configured. The accuracy of these transformations has a direct impact on the amount of flux entering fibres and on the redshift completeness.

We examine how the redshift completeness varies across the field for SN hosts and LRGs. For LRGs, we include redMaGiC galaxies. Any variation in the redshift completeness across the field would be indicative of a problem in the accuracy of the transformations.

In the two panels of Fig. 11, we plot the position of SN hosts with and without redshifts, respectively. One can clearly see the boundary of the DECam chips traced by the SN hosts. The 10 fields in the DES deep survey were observed with minimal dithering. Consequently, chips gaps be-

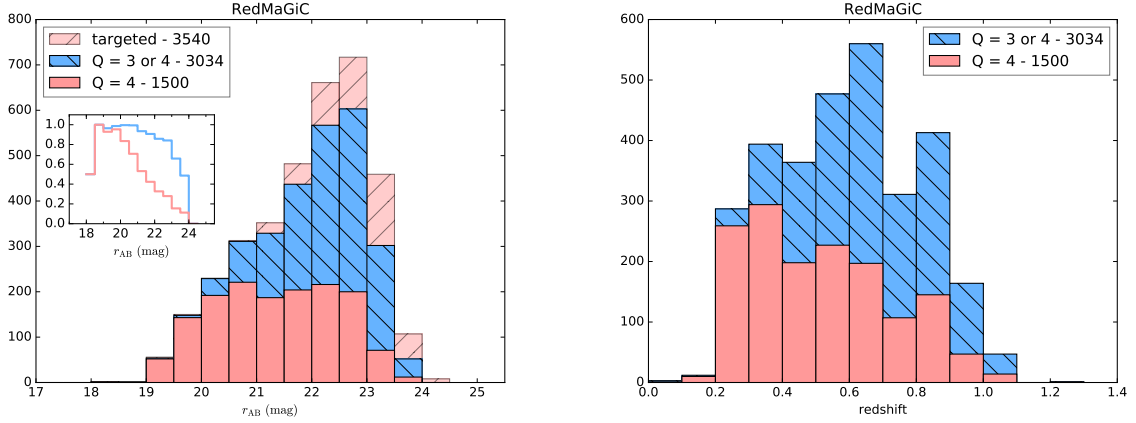


Figure 6. Left: The number of redshifts as a function of r_{AB} for redMaGiC galaxies. The inset shows the fractional completeness as a function of magnitude. The redshift completeness is the number of objects with $Q=3$ (or $Q=4$) divided by the number of objects that were targeted. For all object types other than host galaxies, sources were removed from the target input catalogues once they reached a $Q=3$ redshift quality. Right: Redshift histogram. RedMaGiC galaxies tend to be massive galaxies with little ongoing star formation, so their spectra are usually devoid of strong emission lines. Most of the redshifts were obtained from the calcium H and K lines in absorption, which are usually quite distinct in redMaGiC galaxies.

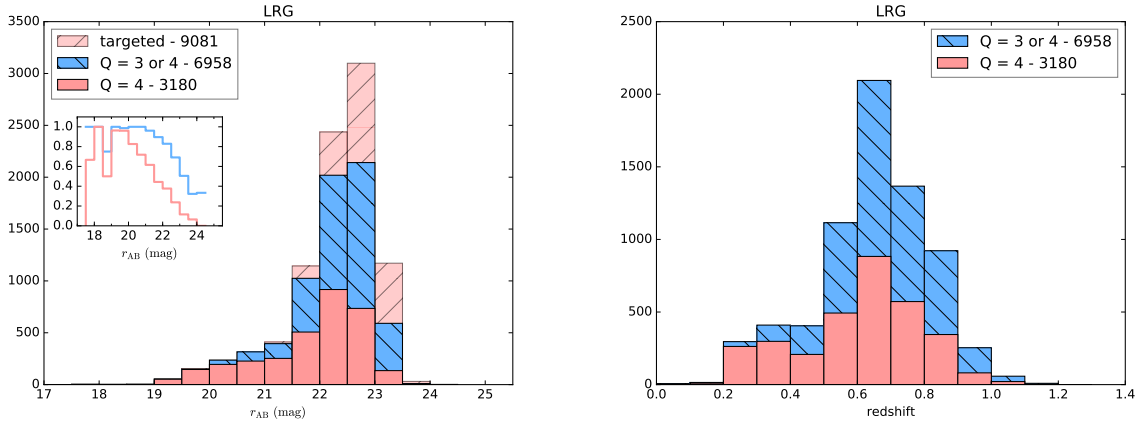


Figure 7. As for Fig. 6 but for Luminous Red Galaxies. LRGs, like redMaGiC galaxies, also tend to have little ongoing star formation, so the spectra are usually devoid of strong emission lines.

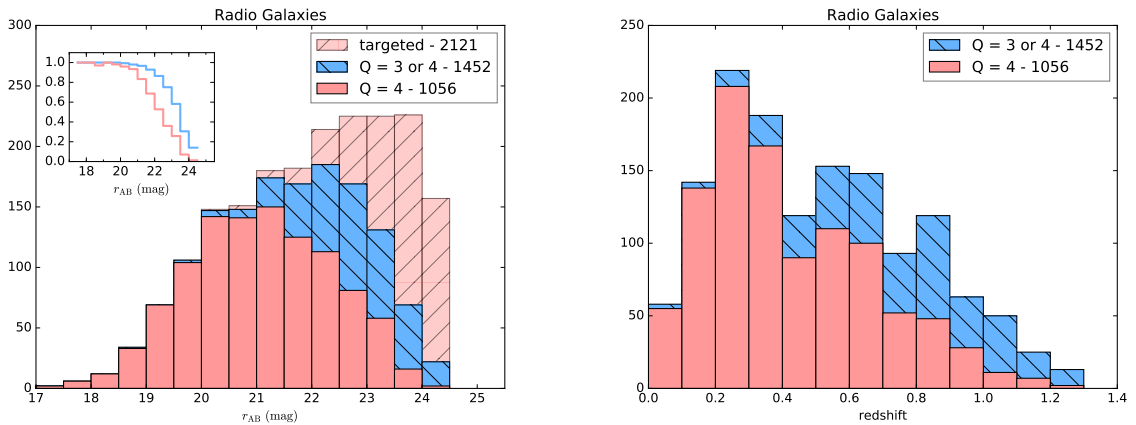


Figure 8. As for Fig. 6 but for radio galaxies.

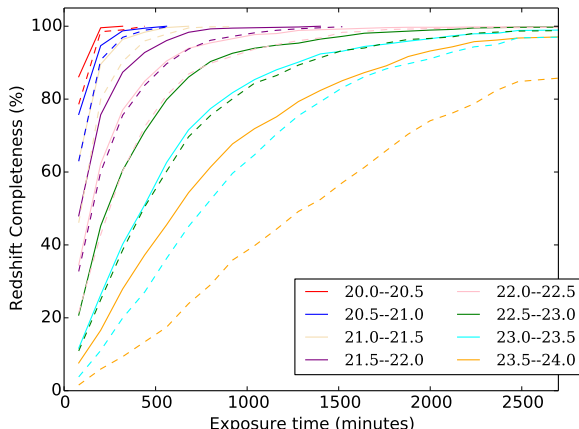


Figure 9. Redshift completeness as a function of exposure time for different magnitude bins and two redshift quality flags. Dashed lines represent redshift quality 4 and solid lines represent redshift quality 3. The lines end once completeness reaches 100%.

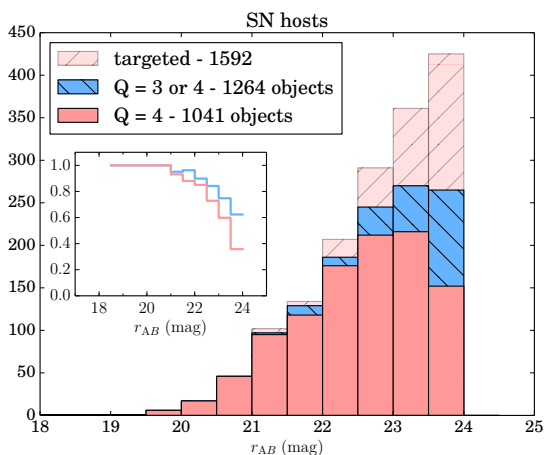


Figure 10. The number of SN hosts with redshifts binned by r_{AB} . The redshift completeness is shown in the insert. Faint SN hosts (i.e. sources classified as `SN_host_faint`) are not included. Throughout the survey, over 7,000 objects were targeted as potential SN hosts. Plotted here are the ones most likely to be SNe Ia. Excluded from this plot are 279 SN hosts that have redshifts from external surveys, as these objects were not targeted. The last bin would reach 95% completeness if we were able to integrate for 40 hours on all sources, as shown in Fig. 9.

tween the DECam CCDs were not covered and so no SNe are found in these gaps. A deficit of SN hosts can be seen in areas where CCDs were not functioning for some fraction of the survey.

Similarly, in the two panels of Fig. 12, we plot the position of LRGs with and without redshifts. The boundary of the DECam chips can no longer be seen, as LRGs were selected from imaging data that was more continuous.

The redshift completeness is shown for SN hosts and LRGs in Fig. 13. The completeness is not uniform for either class of object. In both cases, there is a drop towards the edge of the field. This is not unexpected and is due a drop

in the image quality delivered by the 2dF corrector at the edge of the field and to the accuracy at which the plate can be rotated, which affects fibre near the field edge more than fibres in the centre. Less expected is a gradient from lower left to upper right in the redshift completeness of LRGs

In Fig 14, we show the average number of exposures it takes to get a redshift. There is a general trend of increasing exposure time as one moves from the centre of the field-of-view to the edges. Again, the LRGs show a more significant gradient than the SN hosts.

There are a couple of potential reasons why the gradient is less significant for SN hosts. Firstly, SN hosts are not removed from the target catalogue until a secure, quality 4 redshift is obtained, whereas LRGs are removed from the target catalog once a quality 3 redshift is obtained. This will reduce field-dependent redshift incompleteness. Secondly, the flux lost due to fibre positioning errors is more severe for point sources than extended objects. LRGs due to their higher redshift and redder colours, will be more point like than SN hosts. At the AAT the profile of sources at these redshifts is dominated by the seeing, which will be better in the red than the blue.

The gradient has no impact on the primary OzDES science goals, as the spatial location of the AGN and SNe are not relevant, but the gradient needs to be taken into account if these data were to be used to estimate clustering from statistics such as the 2-point galaxy-galaxy correlation function.

5 THE SECOND OZDES DATA RELEASE

This paper marks the second OzDES release, which will be referred to as OzDES-DR2. OzDES-DR2 differs from OzDES-DR1⁵, the first OzDES data release, in two key aspects.

- The redshift catalogue includes all objects observed during the 6 years of OzDES and the OzDES pre-survey. This includes SN hosts, SNe, and targets for which a redshift could not be obtained.
- The fully processed spectra are provided, including all individual spectra, as well as the stacked spectrum of each object.

In OzDES-DR2, we release 375,000 spectra of almost 39,000 objects, of which three quarters have a redshift with a quality flag of 3 or greater. We release all spectra, including the spectra of objects that do not have a redshift. These data can be used to help secure redshifts for these objects if they are observed in the future with the AAT or other spectroscopic facilities. A full description of the data release is provided in Appendix C.

6 AGN REVERBERATION MAPPING

Second to SN hosts, AGN were the next most frequently observed target type. As shown in Fig. 2, about a quarter of the fibre hours available to OzDES were allocated to the AGN

⁵ The redshift catalogue for OzDES-DR1 is available from <https://datacentral.org.au>.

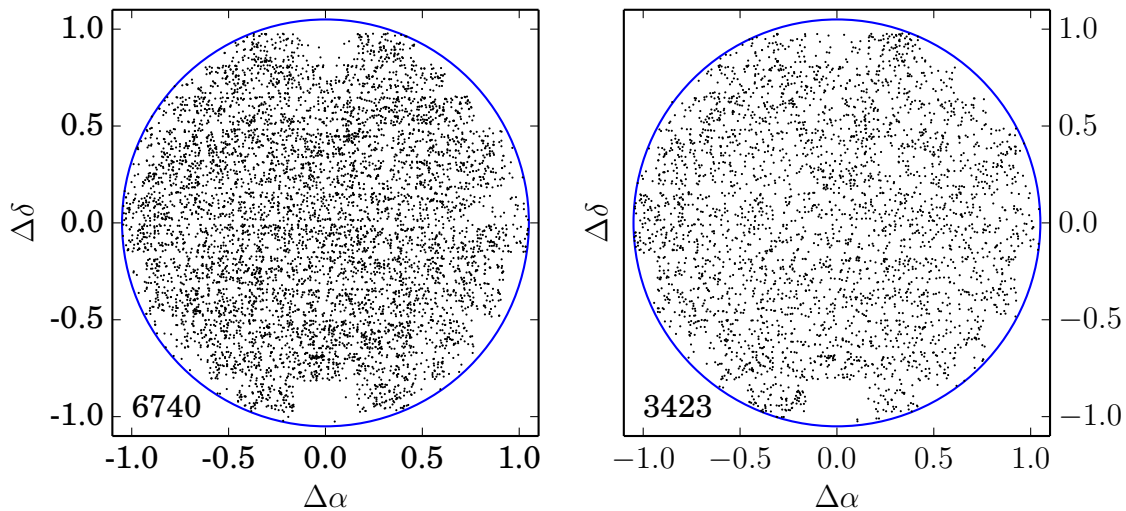


Figure 11. The position, with respect to the centre of the 2dF field plates, of SN hosts with redshifts (left plot) and without redshifts (right plot). The blue circle represents the 2.1 degree diameter of the 2dF field of view. Notable in these images are the chip gaps, the boundary of the mosaic and the two CCD chips at the bottom and top of the mosaic that were not operational for some part of DES. In this plot North is up and East is to the right. Object numbers are given in the lower left corners.

reverberation mapping (RM) programme. The first year of the OzDES RM programme was used to define the sample that was followed in the subsequent five years. Details of the selection are provided in [King et al. \(2015\)](#). In brief, AGN were selected on the basis of the quality of the spectra that were taken in Y1, on the presence of emission lines, while avoiding AGN that occasionally fell in chip gaps, maximising the number of AGN with two reverberating spectral lines in our wavelength range, and ensuring that the AGN selected sampled the redshift interval $0 < z < 4$. Most AGN that pass these cuts have $r_{AB} < 21.2$. From an initial sample of 3331 AGN, 771 AGN were followed during the next five years.

Apart from active transients, AGN in the OzDES RM sample have the highest priority, so the percentage of AGN that were allocated a fibre in any given observation was near to 100%. However, not all fields could be observed during every OzDES run, because of a combination of the finite duration of the observing runs, time lost to poor observing conditions, and the differing priority of the ten DES deep fields. Considering the latter and the bar chart shown in [Fig. 3](#), we expect that the AGN in the two E-region fields and the X3 and C3 fields will have the greatest number of observations, both in terms of the raw number of times they

were observed, and the number of runs in which they were observed.

In [Fig. 15](#), we show the distribution of the number of runs in which AGN were observed. Two plots are shown. In the lower plot, all observations are included, regardless of the observing conditions or the quality of the data. In the upper plot, only data that had a relative zero point greater than -2 (see [Fig. 4](#)) and a quality control (QC) flag of ‘ok’ are included (see Appendix B for a discussion on the QC flag). The vertical solid line shows the number of runs on which OzDES observations were scheduled during the six years that OzDES ran. The two runs from the OzDES pre-survey are excluded, since the AGN RM pre-sample had not been defined at that time. This solid line represents an upper limit to the number of epochs. The median values for the four regions are shown in [Table 4](#). The median values for the X and C regions include X3 and C3, which were the highest priority fields, so we list the median values for the C3 and X3 separately as well. For the E region fields and X3 and C3, the median is close to the 25 epochs that were used in the simulations of the OzDES RM programme in [King et al. \(2015\)](#).

With this number of epochs, simulations in [King et al. \(2015\)](#) predict that one should recover lags for 35-45% of

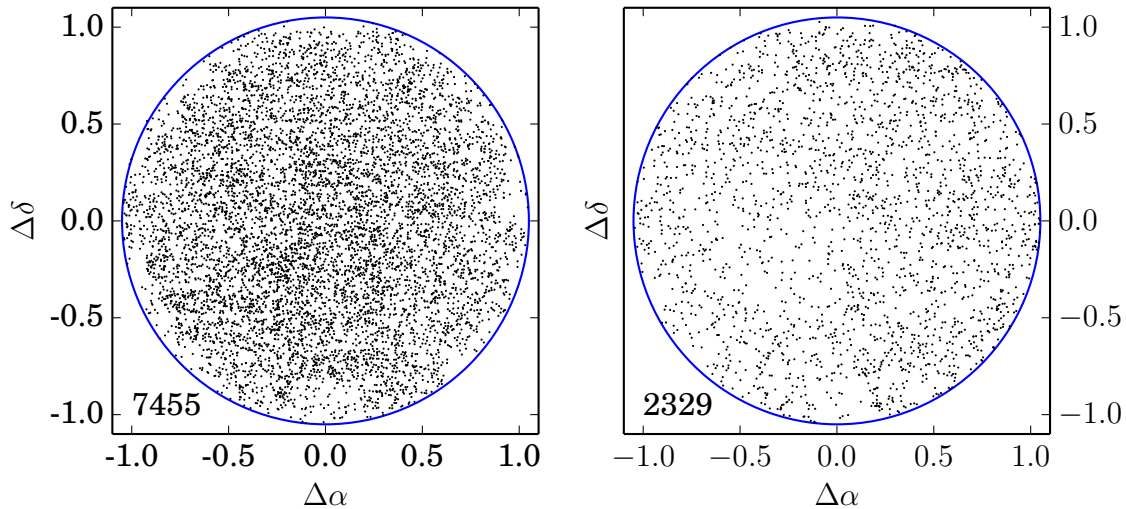


Figure 12. The position, with respect to the centre of the 2dF field plates, of LRGs with redshifts (left plot) and without redshifts (right plot). Gone are the chip gaps. The gap left by the chip at the bottom of the mosaic is still visible. Unlike the plot for SN hosts, there appears to be a gradient, which becomes clearer in Fig. 13. Object numbers are given in the lower left corners.

Table 4. Median number of AGN epochs.

Fields	Epochs
C1, C2, and C3	23
E1 and E2	24
S1 and S2	18
X1, X2, and X3	21
C3	25
X3	25

the AGN depending on the accuracy with which one can calibrate the line flux. The simulations used a 10% accuracy for the line measurements. Hoormann et al. (2019) demonstrate that we are able to achieve this level of accuracy, so we anticipate that we will recover this fraction of lags in the two E fields and the C3 and X3 fields, and a lower fraction in the other six fields. The first two lags, using the CIV line and based on a subset of the OzDES data, were published in Hoormann et al. (2019).

7 SN HOST GALAXIES: A CASE STUDY

As illustrated in Fig. 2, slightly more than one-third of all the fibre hours allocated to targets during the six years of OzDES were allocated to SN hosts. Unlike other targets, SN hosts were observed until a secure redshift (see Sec. 4) was obtained. SN hosts also extend to fainter magnitudes than most targets. Both of these factors meant that the integration time for some hosts are very long, as long as 2 days in a number of cases.

We use SN hosts to examine how well the measured noise changes with time, and explore the implications of de-selecting SN hosts once they reach a $Q = 3$ quality redshift and before they reach the more secure $Q = 4$ quality redshift.

7.1 Dependence on the measured noise with integration time

Ch17 showed that the noise in the spectra evolves with increasing exposure time, t , according to expectations. The OzDES pipeline averages spectra from individual exposures rather than summing them, so the measured noise decrease as \sqrt{t} , or with the square root of the number of exposures if the exposure time is constant. Nearly all OzDES exposures were taken with an integration time of 40 minutes. We repeat the Ch17 analysis, which was done with just three years of

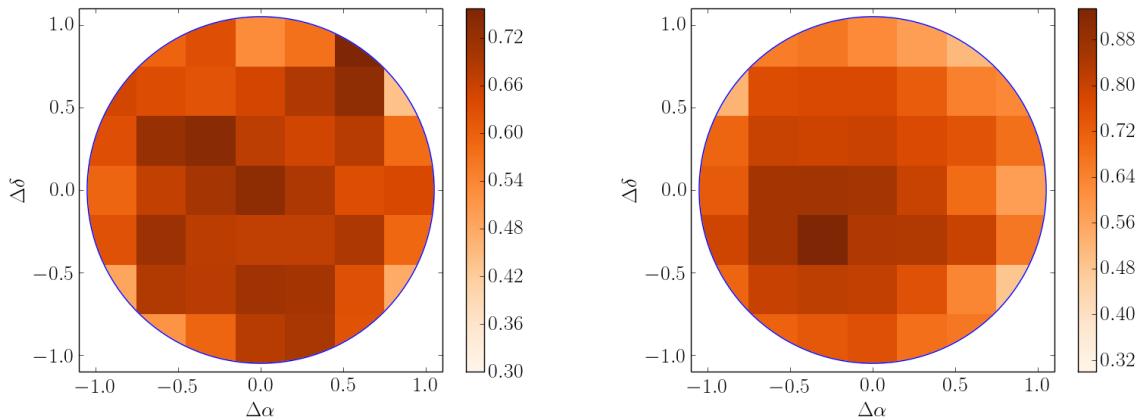


Figure 13. Redshift completeness in bins of 0.3 degrees by 0.3 degrees for SN hosts (left) and LRGs (right). The blue circle represents the 2.1 degree diameter of the 2dF field of view. The decreasing redshift completeness with distance from the centre of 2dF is a consequence of poorer image quality at larger radii. The presence of a gradient in the right plot and not the left one is surprising.

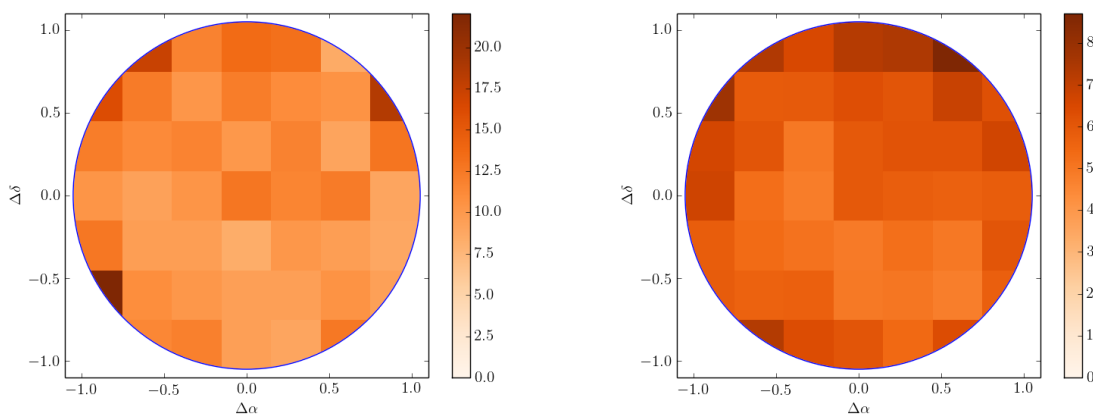


Figure 14. The average number of exposures, in bins of 0.3 degrees by 0.3 degrees, required for a successful redshift measurement for SN hosts (left) and LRGs (right). The blue circle represents the 2.1 degree diameter of the 2dF field of view. Overall, as one approaches edges of the 2dF field of view, it takes more exposures to secure a redshift. The gradient from lower left to upper right seen in Fig. 13 is also seen here and is also stronger for the LRGs.

OzDES data. With six years of data, there are more objects in total and there are more objects with longer integration times.

We do apply some qualitative cuts to the data chosen for this analysis. We exclude spectra that i) were taken in poor conditions (defined here as exposures that have a relative zero point of less than -2 ; see Fig 4), ii) did not pass quality control, and iii) were taken during Y1 and SV. During Y1 and SV, we do not have reliable zero points, and did not take dome flats, so the quality of the processed data is not as good as it was in later years.

With these cuts, we then select SN hosts that have at least 12 exposures. This will tend to select SN hosts that are fainter than average, as redshifts for bright hosts would have typically been obtained with fewer exposures and would have been deselected from the target pool before 12 exposures had been taken. As the selected objects are faint, the signal to noise ratio is quite small, so spectral features do not contribute to the noise.

We then select a wavelength region over which to measure the noise. We start off with a region that is relatively free of night sky lines. The region we select starts at 6610\AA and ends at 6750\AA . We then examine how our results change if we choose a region that contains bright night sky lines.

We then randomise the order in which exposures were taken. The ordering is different for each object. For a given object, we take the first exposure and compute the clipped RMS over the wavelength region noted above. In computing the clipped RMS, we reject the 5 highest and 5 lowest values. Over the wavelength region that we sample, there are about 135 pixels, so the clipping removes less than 10% of the data. The clipping allows us to exclude residuals from cosmic rays that were not perfectly removed during the processing of the data and mitigates the effect of strong spectral features, such as bright emission lines from the object. We experimented with the normalised median absolute deviation instead of the clipped RMS and obtained qualitatively similar results. All

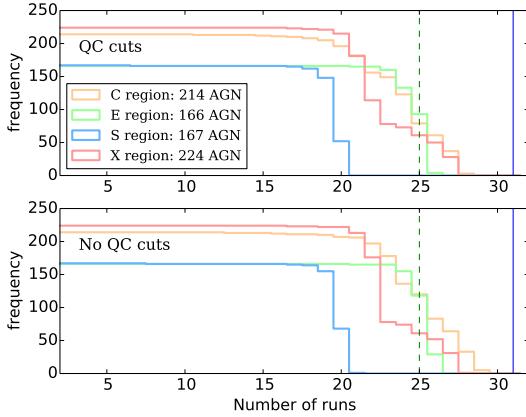


Figure 15. Reverse cumulative histograms of the number of independent spectroscopic epochs in the AGN reverberation mapping programme, plotted without restriction on data quality (lower histogram) and excluding poor quality data (upper histogram). See text for details. The vertical line corresponds to the number of runs on which some OzDES data were taken during the six years OzDES ran. It corresponds to the maximum possible number of epochs. The prediction on the number of recovered lags made in King et al. (2015) were based on obtaining 25 epochs, which is shown as the dotted line.

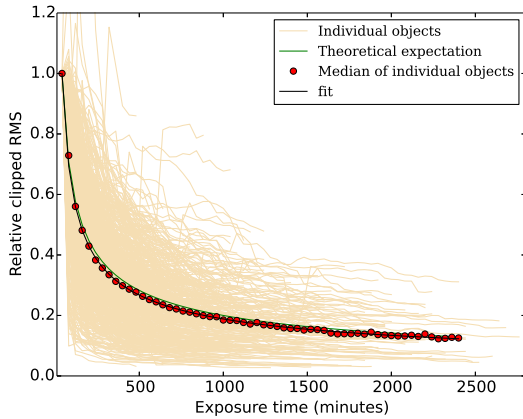


Figure 16. The scaled clipped RMS of 639 individual objects (in light yellow) from the three C fields plotted against exposure time. The median of these curves is plotted as the solid red circles. The theoretical expectation is the green curve and the best fit is the black curve. Note how closely the two match. The clipped RMS is computed over the region 6610-6750 Å.

subsequent measurements of the clipped RMS are scaled by the clipped RMS that is computed in this first exposure.

We then add data from subsequent exposures in the randomised sequence, noting that we average the spectra, rather than summing them, and recompute the clipped RMS. This is done for every SN host. The results are shown in Fig 16.

There is considerable scatter from object to object. This is caused by data of poorer than average quality entering the average at different stages of the sequence. For example, a sequence of exposures in which the first exposure is noisy will result in a much larger than anticipated drop in the relative clipped RMS when data of better quality is included.

Applying stronger quality cuts reduces the scatter. To get an idea of the general trend, we take the median of all objects. These are the red circles in Fig. 16.

We fit the function

$$\frac{n^\alpha + \beta}{1 + \beta} \quad (1)$$

to the red points. The function equals 1 when $n = 1$, and allows for a noise floor. The theoretical expectation for Poisson noise is $\alpha = -0.5$ and $\beta = 0.0$. This is represented with the green curve. The best fit to Eq. 1 is the black curve. The best fit has $\alpha = -0.54 \pm 0.01$ and $\beta = 0.02 \pm 0.01$. The best fit is close to theoretical expectation for Poisson noise, and the floor is low. We find there is no reason to stop observing targets over the range of exposure times explored here. We will explore this point further in the next section.

There are limitations to our approach. Eq. 1 may not be the best functional form, and we have not allowed for the hard limits imposed by theory (i.e. $\beta \geq 0$ and $\alpha \geq -0.5$) when fitting Eq. 1. Indeed, we find $\alpha < -0.5$. It is unclear what causes this.

We varied our selection cuts and repeated the analysis for the other fields to see how robust our results were. We do see some variability in the results, the origins of which are unclear. For the permutations we tried, we found $-0.55 < \alpha < -0.5$ and $0 < \beta < 0.05$, i.e. a slightly steeper dependence than expected from theory and weak evidence for a noise floor. The floor is still less than the contribution from Poisson noise even after a day of exposure. At 1.5 days (which corresponds to 54 40-minute exposures), $54^{-0.5} \sim 0.14$.

We also varied the wavelength region over which the statistics were computed, using both a region free of bright night sky lines (6610-6750 Å) and a region dominated by them (7700-8000 Å). In both cases, a dependence close to $1/\sqrt{n}$ was observed.

7.2 Deselecting targets

It is clear that, for the range of exposure times explored in OzDES (up to two days of integration), the signal-to-noise continues to improve as expected, i.e. with the square root of the exposure time. If OzDES were to continue into the future, then exposure times could grow into many days. A number of questions then arise. How long should one spend on a target before deselecting it for another one if one is unable to get a redshift, and is there a way of knowing if some targets are more likely to result in a redshift than others? Any decision to deselect one source in preference for another needs careful consideration, as such a decision will invariably lead to biases that need to be modelled.

During Y6, when 71% of the available fibre time was spent targeting SN hosts, there were more SN hosts without quality 4 redshifts than 2dF fibres, especially for C3 and X3, the two deepest DES fields. If there were many more fibres than the 392 available with 2dF, and SN hosts were the only targets, then the answer would be to continue observing these targets.

For instruments like 4MOST, which has 4 times as many fibres as 2dF, one could observe more targets longer without running out of fibres. Of course, 4MOST could choose to

push to fainter limits than OzDES did (most OzDES hosts are brighter than $r_{AB} = 24$), in which case it too would have to make a choice about what to keep and what to drop in favour of something that has a higher chance of success.

Ultimately, this decision has to be based on the questions one is trying to address. A SN with a noisy light curve at $z \sim 1$ may not be as valuable as two SNe with high quality light curves at a redshift of a half. We do not address this question in this paper, but leave it to when the DES 5-year photometrically selected sample is defined.

For SN hosts, there is an additional consideration. SN hosts are deselected once they get a $Q = 4$ quality redshift. Some SN hosts obtain a $Q = 3$ quality redshift relatively quickly, but fail to reach the $Q = 4$ quality redshift that is necessary for deselection. This raises an interesting question. How long should one continue observing such targets once the $Q = 3$ quality redshift is obtained?

We examine the number of exposures SN hosts remain at $Q = 3$ once that quality is obtained. We divide these objects into two. Those that eventually obtain a $Q = 4$ quality redshift and those that do not. Almost 85% of objects that move from $Q = 3$ to $Q = 4$ do so within 10 exposures of obtaining $Q = 3$. If one were to deselect targets that fail to reach $Q = 4$ within 10 exposures of obtaining $Q = 3$, then one would free up 6.6% of the fibre hours, which could then be used on other sources, at the cost of obtaining 6.2% fewer $Q = 4$ quality redshifts.

In OzDES, we did not deselect SN hosts before they reached $Q = 4$, no matter how many exposures were being taken. We are now in position to explore how the choice of deselecting SN hosts biases the fitted cosmological parameters. This will be reported in a future paper.

About 35 SN hosts fainter the $r = 23$ fail to obtain a $Q = 4$ quality redshift once a $Q = 3$ quality redshift was obtained, even after continuing to integrate for 30 or more exposures (equivalent to 20 hours of integration). We examined each of these sources, and found that more than half of them show a single bright line, which is assumed to be the [OII] doublet. These objects reach the $Q = 3$ level relatively quickly. The redshift of these sources is around one, meaning that other potential emission lines were outside the spectral range of AAOmega. During run041, we used a redder setting for AAOmega, which allowed us to pick up the [OIII] doublet in some cases. Most of the remaining cases consist of sources with weak features, typically a combination weak [OII], [OIII], and $H\beta$, that only start to become apparent after many hours of integration, but are not yet strong enough for a $Q = 4$ quality redshift to be assigned, even after an additional 20 hours of additional integration.

8 IMPLICATIONS FOR FUTURE SURVEYS

By the early 2020s, LSST will have started surveying the Southern sky and several new multi-object spectroscopic facilities fed by rapidly configurable fibre positioners covering wide fields of view will be operational. Of the planned surveys using these new facilities, the Time-Domain Extragalactic Survey (TiDES, Swann et al. 2019), using 4MOST (de Jong et al. 2019) in combination with LSST, is the one that most closely matches the OzDES concept. A comparison of the OzDES and TiDES surveys is shown in Table 5.

Table 5. A comparison between TiDES/4MOST and OzDES/AAOmega.

Metric	OzDES/AAT	TiDES/4MOST
Telescope, Spectrograph, and Positioner		
M1 Diameter (m)	3.9	4.2
Field of view (sq. deg.)	3.46 (circle)	4.2 (hexagon)
Fibre diameter (")	2.0-2.1	1.45
Number of fibres	392	1624 ^a
Config. time (min)	40 ^b	2
Spectral Resolution	1400-1700	6500
Wavelength range (Å)	3700-8800	3700-9500
Smallest sep. (")	30	15
Survey characteristics		
Survey region ^{c,d}	DES deep fields	LSST DDF and $+5^\circ < \delta < -70^\circ$
Median Seeing (")	1.5	0.8
Usable clear fraction	67%	90%
Duration (years)	5 ^g	5
Transient redshifts	5,000	50,000
AGN monitored	771	700
Live transients	1,450	30,000 ^e
Fibre hours	193,000	250,000
Freq. (per lunation)	1	2 ^f
Observing season	5 months	9 months ^f

^a Fibres feeding the two low-resolution spectrographs.

^b Done in parallel while observing with the opposing plate.

^c The 10 DES deep fields cover 27 sq. degrees.

^d Deep Drilling Fields

^e Mostly from the TiDES wide survey

^f In the deep drilling fields.

^g The SN programme in DES and OzDES lasted 5 seasons.

The survey characteristics of TiDES listed in Table 5 are subject to change as the survey strategy is still evolving.

On practically every metric, TiDES will be superior to OzDES. It will observe at least 10 times more live transients, mostly in the wide survey, it will measure up to 10 times more host galaxy redshifts, and it will monitor AGN at twice the frequency that OzDES did and with shorter seasonal gaps. TiDES will also observe transients over the majority of the southern sky, something that OzDES did not do.

Although the science goals and strategy of TiDES are heavily based upon that of OzDES, practically there are some differences. The first difference between OzDES and TiDES is that TiDES will not drive the pointings of 4MOST. Wherever 4MOST is scheduled to point there will be previously discovered transient targets, a majority of these targets anticipated to be discovered by LSST. If these transients are still bright enough to obtain a spectrum for spectral typing, TiDES will attempt to do so. Otherwise TiDES will target the host galaxy of the transient to obtain a spectroscopic redshift for type Ia SN cosmology.

The second difference between OzDES and TiDES is that TiDES will have two components, a wide field component and a deep field component. The wide field follow-up will cover the majority of the southern sky, where low redshift transients and host galaxies will be targeted with a maximum exposure time of two hours. The second component to TiDES will be spectroscopic follow-up in the deep fields of 4MOST, where higher redshift live transients will be

observed, the AGN reverberation mapping will take place, and where one can repeatedly target SN host galaxies to build up redshifts on the faintest of targets.

The 4MOST Deep fields currently correspond to four pointings, each covering 4.2 sq. degrees and contained within the LSST deep drilling fields (Guiglion et al. 2019). LSST data will provide the high precision photometry and high cadence light curves required for sub-percent precision cosmology. As three of the DES deep fields (ECDFS, XMM-LSS, ELAIS S1) are also covered by the 4MOST and LSST DDFs, this offers the exciting possibility of continued monitoring a subsample of the AGN observed by OzDES over a 15 year baseline, if the time gap between OzDES and TiDES can be covered. The higher cadence and longer observing season will also mean that TiDES will be more sensitive to AGN with shorter lags and less sensitive to seasonal gaps. Therefore TiDES will obtain a higher fraction of lags compared to OzDES.

With more than 4 times the number of fibres, and a field of view that is only slightly larger than 2dF, TiDES on 4MOST will be able to observe SN hosts in the deep fields longer without running out of fibres. Given the higher throughput, broader spectral coverage, and higher spectral resolution of 4MOST compared to AAOmega, and the better observing conditions on Cerro Paranal compared to Siding Spring, TiDES will be able to obtain host galaxy redshifts more quickly and to a fainter magnitude limit than OzDES.

9 SUMMARY

Over six consecutive observing seasons, starting in 2013, OzDES obtained over 375,000 spectra of almost 39,000 sources in the ten DES deep fields. These spectra were obtained with the 2dF fibre positioner and AAOmega spectrograph on the AAT. Over that time, each DES SN field was observed between 18 and 25 times.

In these fields, OzDES has obtained redshifts for almost 7,000 galaxies that hosted a transient, and classified almost 270 transients. The strategy of targeting each field repeatedly has enabled redshifts to be obtained for sources as faint as $r_{AB} \sim 24$. We show that one is not limited by a noise floor, even after two days of exposure, and that one can reach high spectroscopic completeness at these faint limits.

The strategy of targeting each field repeatedly has allowed OzDES to monitor 771 AGN. These data will be used to measure the lag between the continuum from the accretion disk and the broad lines from the broad line emission region. We anticipate that we will recover lags and derive black hole masses for approximately 30% of the AGN we monitored.

The OzDES observing strategy allowed a number of science programmes to run in parallel. Just over 60% of the 192,000 fibre hours were allocated to monitoring AGN and obtaining redshifts for galaxies that hosted a transient. The remaining 40% was used to obtain redshifts for wide variety of targets, such as radio galaxies, cluster galaxies, and galaxies to train photometric redshift algorithms.

OzDES can be used as a template for future surveys that use the next generation of wide-field multi-object fibre-fed spectroscopic facilities to target and monitor sources in the LSST deep drilling fields. These new facilities will have

up to an order of magnitude more fibres than 2dF, and they will enable one to follow a larger number of fainter sources more frequently and for longer.

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APPENDIX B: OBSERVING LOGS FOR Y4, Y5 AND Y6

The observing logs for Y4, Y5, and Y6 are shown in Tables B1, B2, and B3, respectively. They follow the format presented in Yu15 and Ch17. The MaxVis field (Yu et al. 2020) was observed for 155 minutes during run 28 on 25 and 26 December 2017.

Table B1. OzDES fourth year (Y4) observing log for DES deep fields (c.f. Table 2 of Yu15 and Tables A1, A2, and A3 in Ch17).

UT Date	Observing run ^a	Total exposure time (minutes)										Notes for entire run	
		E1	E2	S1	S2	C1	C2	C3 (deep)	X1	X2	X3 (deep)		
2016-08-25	024	80	-	-	-	-	-	-	-	-	-	-	0.5/1.5 clear nights
2016-08-26		-	90	-	-	-	-	-	-	-	-	-	
2016-08-27		-	-	-	-	-	-	-	-	-	-	110	
2016-08-28		-	-	-	-	85	-	-	-	-	-	-	
2016-08-29		180	-	-	-	-	-	-	80	80	-	-	
2016-09-25	025	-	80	-	-	-	20	120	-	-	120	3.5/4.5 clear nights	
2016-09-26		80	-	-	-	30	120	-	120	-	-		
2016-09-27		-	80	40	-	80	-	-	-	120	-		
2016-10-05		-	40	100	120	-	-	-	-	-	-		
2016-10-31	026	-	-	-	-	-	-	80	-	-	120	3.3/3.5 clear nights	
2016-11-01		120	-	120	-	-	-	60	-	-	120		
2016-11-02		-	120	-	120	-	-	-	80	-	-		
2016-11-03		80	-	-	-	120	110	-	-	120	-		
2016-11-25	027	120	-	-	-	40	-	120	-	-	120	5.0/5.0 clear nights	
2016-11-26		-	120	120	-	-	120	-	-	-	-		
2016-11-27		-	-	-	120	120	-	-	120	-	-		
2016-11-28		-	-	120	-	-	68	80	-	120	-		
2016-11-29		120	-	-	-	-	153	-	-	-	120		
2016-12-25	028	-	-	-	-	-	80	80	-	-	120	2.8/4.0 clear nights	
2016-12-26		-	-	80	-	80	-	-	-	80	-		
2016-12-27		-	-	-	80	-	-	80	80	-	-		
All Y4	Total(min)	780	530	580	525	470	671	700	480	440	830		

^a Run numbering includes runs from 2dFLenS and DEVILS projects, so is not necessarily contiguous.

^b Bright object backup programme for poor conditions, not counted toward final total.

APPENDIX C: THE SECOND OZDES DATA RELEASE - OZDES-DR2

The second OzDES data release, OzDES-DR2, consists of the OzDES redshift catalogue and the reduced spectra. The data are available from Data Central⁶.

C1 Redshift Catalogue

The OzDES-DR2 redshift catalogue is a FITS binary table containing almost 38,700 entries, of which almost 30,000 have a redshift with quality flag 3 or greater. The columns in the catalogue are described in Table C1.

The transient type is sourced from the 34 ATels (e.g., Calcino et al. 2018) that OzDES published over 6 years. D’Andrea et al. (2018) provides a description of how transients are classified.

When available, we add comments made by redshifters. Not all sources have comments, as adding a comment was left to the discretion of the redshifter.

Many sources were assigned multiple types. For example, sources that were assigned the LRG type were often assigned as redMaGiC sources and visa-versa. Sources observed during the OzDES pre-survey (runs 001 and 002) were not given an object type.

⁶ <https://datacentral.org.au>

C2 Spectra

All spectra are provided as multi-extension FITS. For targets observed up to and including run 40, there is usually one FITS file per target. If a target was also observed during run 41, which used a redder setting for AAOmega, we provide a second FITS file containing the data from that run combined with data from earlier runs. A small number of targets in run 41 were only observed in that run. In this case, there is just one FITS file per target.

Within each FITS file, there are extensions containing the stacked spectrum and the spectra that went into the stack. Also included are the variances and bad pixel masks for each spectrum. The extensions are described in Table C2. The individual spectra that go into the stack are ordered by date, from the one taken first to the last.

All spectra have been shifted to the heliocentric reference frame.

The keywords in the FITS headers are described in Table C3.

C3 Data Quality

The quality of the data is characterised in a couple of ways. The keywords REDZP and BLUEZP give an indication of the observing conditions at the time the data were taken. They are determined from F stars that were observed at the same time as the target. For the red arm, a zero point that

Table B2. OzDES fifth year (Y5) observing log for DES deep fields (c.f. with Table B1).

UT Date	Observing run ^a	Total exposure time (minutes)										Notes for entire run
		E1	E2	S1	S2	C1	C2	C3 (deep)	X1	X2	X3 (deep)	
2017-08-21	030	80	80	-	-	-	-	80	-	-	80	4.0/4.2 clear nights
2017-08-22		80	80	80	-	-	120	-	-	-	-	
2017-08-23		40	80	-	-	80	-	-	-	80	-	
2017-08-24		80	80	-	120	-	-	-	80	-	-	
2017-08-25		80	80	-	-	-	-	120	-	-	80	
2017-08-26		34	40	-	-	49	-	-	25	-	-	
2017-08-27		-	-	-	80	-	-	-	80	80	-	
2017-10-17	032	-	-	-	-	40	-	75	-	-	-	4.3/6.0 clear nights
2017-10-18		120	-	120	-	-	103	-	-	-	120	
2017-10-21		-	120	-	130	72	-	-	-	120	-	
2017-10-22		93	80	80	-	-	-	80	120	-	-	
2017-10-23		-	105	-	120	-	-	102	-	-	120	
2017-11-16	033	80	-	-	-	60	-	80	-	-	120	2.8/4.5 clear nights
2017-11-19		-	80	-	-	30	120	-	80	80	-	
2017-11-20		-	-	120	120	40	-	80	-	-	-	
2017-12-18	034	-	-	-	-	-	-	100	-	-	80	3.3/4.0 clear nights
2017-12-19		-	-	-	-	80	-	-	-	80	-	
2017-12-21		-	-	-	-	-	107	-	-	-	50	
2017-12-22		80	-	-	-	-	110	-	-	-	-	
2017-12-23		-	80	-	-	-	-	-	100	-	-	
2017-12-24		-	-	70	-	80	-	-	-	-	-	
2017-12-25		-	-	-	87	-	80	-	-	-	-	
All Y5	Total(min)	768	905	470	657	531	640	717	485	440	650	

^a Run numbering includes runs from 2dFLenS and DEVILS projects, so is not necessarily contiguous.

^b Bright object backup programme for poor conditions, not counted toward final total.

is smaller than 31 usually means that conditions were poor, which could be due to seeing or the presence of cloud. It could also be due to an exposure that was cut short.

Most of the 375,000 spectra in the data release have been visually inspected and problematic cases noted in the QC keyword. As this is not 100% foolproof, there might be a number of problematic cases that have been missed. A value for QC that differs from “ok” indicates that there is a problem with the spectrum. The nature of the problem is noted in the value of the QC keyword and is described in Table C4.

Some of the problems are caused by artifacts that affect a limited wavelength range. Others affect the whole spectrum. The region of the spectrum affected by the artifact is noted in the masks. A value of 1 in the mask marks the region of poor quality data.

Examples of what the artifacts look like are available from the [OzDES-DR2 documentation page](#) at Data Central.

APPENDIX D: SOURCE TYPES FOR THE MAXVIS FIELD

OzDES targeted the DES MaxVis field towards the end of Y5. Centred at RA=97.5° and DEC=-58.75°, it is one of the DES standard star fields and is observable during the entire DES observing season. Listed in Table D1 is a description of the target types together with the number observed.

In more detail, stellar-like objects were selected on the basis of their location in $r-i$ vs $g-r$ colour-colour diagram or on the amount of variability that they showed. Targets with more variability were given higher priority. Some of the variable sources are AGN.

The three tail sources were located in a small tail-like feature in the colour-colour diagram near to where the stellar sequence turns up with redder stars. All three sources were galaxies.

QSO candidates were selected on the basis of their colour, with QSOuv representing sources that were red in $r-i$, but blue in $g-r$, QSOHz representing QSOs that were likely to be at high redshift, and QSOc representing sources that were selected by the Cambridge QSO group.

At lower priority were blue stars (BBstar), of which three were white dwarfs and two were AGN. Lowest priority were other randomly selected stars (RNDstars).

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.

Table B3. OzDES sixth year (Y6) observing log for DES deep fields (c.f. with Table B1).

UT Date	Observing run ^a	Total exposure time (minutes)										Notes for entire run	
		E1	E2	S1	S2	C1	C2	C3 (deep)	X1	X2	X3 (deep)		
2018-06-15	036	-	80	-	-	-	-	-	-	-	-	-	0.6/1.5 clear nights
2018-06-24		150	80	-	-	-	-	-	-	-	-	-	
2018-09-04	037	60	120	-	-	-	40	120	-	-	80	1.0/3.5 clear nights	
2018-09-05		117	-	-	-	-	-	-	-	-	-		
2018-09-06		-	-	-	-	-	-	-	-	40	-		
2018-10-06	038	120	-	120	-	-	-	80	-	-	120	3.0/3.0 clear nights	
2018-10-07		-	120	-	120	-	120	-	-	120	-		
2018-10-08		120	-	-	-	120	-	120	120	-	-		
2018-11-08	039	-	120	40	-	-	-	120	-	-	120	2.0/3.0 clear nights	
2018-11-09		120	-	40	-	-	120	-	-	120	-		
2018-12-01	040	80	-	80	-	-	-	80	-	-	80	2.2/3.0 clear nights	
2018-12-02		-	80	-	80	-	120	-	-	80	-		
2018-12-03		74	-	-	-	-	22	-	-	71	-		
2019-01-04	041	-	-	-	-	-	-	120	-	-	80	1.7/3.0 clear nights	
2019-01-06		-	-	-	-	-	-	120	-	-	80		
2019-01-08		-	-	-	-	-	-	120	-	-	80		
2019-01-09		-	-	-	-	-	-	-	-	-	40		
All Y6	Total(min)	841	600	280	200	142	400	880	191	360	680		

^a Run numbering includes runs from 2dFLenS and DEVILS projects, so is not necessarily contiguous.

^b Bright object backup programme for poor conditions, not counted toward final total.

Table C1. Description of the columns in the OzDES-DR2 redshift catalogue.

Name	Description	Units
OzDES ID	Official DES name or an assigned OzDES name if an official DES name is unavailable	...
Alpha J2000	Right Ascension	deg
Delta J2000	Declination	deg
r _{mag}	r_{AB} band magnitude	...
z	Redshift; -9.99 means no redshift	...
qop (Q)	Redshift quality flag	...
Object types	Assigned types	...
Transient type	Type of transient (e.g. SN Ia, SN II, etc.); Assigned to <i>None</i> if no type	...
Comment	Comments from redshifters	...

Table C2. Description of FITS extensions in OzDES-DR2.

Extension	Description
Primary	Stacked spectrum
1	Variance spectrum
2	Mask
3	Spectrum of first epoch
4	Variance spectrum of first epoch
5	Mask for first epoch
⋮	⋮
3i	Spectrum of <i>i</i> th epoch
3i + 1	Variance spectrum of <i>i</i> th epoch
3i + 2	Mask for <i>i</i> th epoch

Table C3. Description of FITS headers in OzDES-DR2.

Primary Header ^a	
DO_HELIO	An internal flag recognised by the redshift code MARZ that ensures that no helio-centric correction is applied by MARZ when set to false
NDATE	The date the data were combined
SOURCE	The original source name
VERCOADD	The version of the coadd pipeline that stacks the reduced data
RMAG	r_{AB} band magnitude in a 2" aperture
TYPE _n	Target type
Z	redshift; -9.99 means no redshift
QOP	redshift quality flag
Individual epochs ^b	
REDZP ^d	The zero point in the r -band. Measured from the spectra of F stars that were observed at the same time as the source
REDSNSF	The sensitivity function used in the red arm
BLUEZP	The zero point in the g -band. Measured from the spectra of F stars that were observed at the same time as the source
BLUESNSF	The sensitivity function used in the blue arm
QC	Quality control flag
VERSION	The version of the pipeline that processes the raw data and produces single epoch reduced spectra
INCOADD	True if the epoch is included in the coadd
PIVOT	The 2dF fibre number
XPLATE	The X position on the plate (microns)
YPLATE	The Y position on the plate (microns)
ORIGTARG	The name of the target used in this observation
SOURCEF	The raw data filename (red arm only)
FLXSCALE	The scale factor that was applied in the co-adding
FLSCTYPE	The method used to compute scale factor
RMAG	r_{AB} in the input target catalogue
Individual epochs ^c	
EXPOSED	Exposure time
DICHROIC	Dichroic used - x6700 for run041, x5700 otherwise
SOURCE	2dF field plate - either plate 0 or plate 1
UTMJD	Modified Julian Date
ZDSTART	Zenith Distance at the start of the exposure
ZDEND	Zenith Distance at the end of the exposure

^a Keywords that are in the FITS standard are not listed.

^b Keywords that are inserted when the data are co-added.

^c Keywords that are part of the header produced at the telescope and keywords. Not all keywords are listed.

^d Zero points are defined as $2.5 \log_{10}(f) + m$, where m is the catalog magnitude of the F star and f is the observed spectrum integrated over the g band for the blue arm and the r band for the red arm.

Table C4. Description of values contained in the QC keyword.

QC value	Description
700nmLED	Contamination from an LED in the gripper arm. This was discovered a few years, into the survey, so it mostly affects early data.
redDip	Most likely associated to charge traps in the old AAOmega red CCD. Affects data taken prior to run007.
doubleDip_1	Most likely associated to charge traps in the old AAOmega red CCD. Affects data taken prior to run007.
doubleDip_2	Most likely associated to charge traps in the old AAOmega red CCD. Affects data taken prior to run007.
doubleHump	Origin unknown.
fib388p0	Issues with fibre 388 on plate 0. Cause unknown.
fib94p1	Issues with fibre 94 on plate 1. Cause unknown.
redfail	Failed reduction in red arm.
redPCAFail	Failed PCA sky subtraction in the red arm.
bluefail	Failed reduction in blue arm.
failedReduction	Failed reduction, both arms.
redPCAFail	Failed PCA sky subtraction in red arm.
brightStarBlue	A bright star in a nearby fibre causes the reduction in the blue arm to fail or contaminates the extracted spectrum in the blue arm.
brightStarRed	As above, but for the red arm.
brightStarAll	As above, but for the entire spectrum.
fringed	Fringes detected.
poorConditions	Poor observing conditions.
blueSkySubFail	Sky subtraction failed in blue arm
blueFeature_1	A blue feature of unknown origin
IRcontamination	Contamination in red arm of unknown origin
failedReduction	Processing failed for unspecified reason

Table D1. Description of target types in the DES MaxVis field.

Type	Number of sources	Description
High Priority targets		
QSOHz	49	High-z quasar candidates
QSOuv	14	uv-excess quasar candidates
QSOc	67	quasar candidates selected from the Cambridge QSO group
VarObj1	94	Variable Object candidate (high priority)
Tail1, Tail2, Tail3		3 star-like objects with colour in the tail-like region
Low Priority targets		
BBstar	72	stars selected on the basis of blue colours in g-r and r-i
VarObj2	55	Variable Object candidate (lower priority)
Lowest Priority targets		
RNDstars	1	Random star-like objects