

Correlating spectral and timing properties in the evolving jet of the microblazar MAXI J1836–194

M. Lucchini ¹★, T. D. Russell ¹★, S. B. Markoff ^{1,2}★, F. Vincentelli ³, D. Gardenier ^{1,4},
C. Ceccobello ⁵ and P. Uttley¹

¹API – Anton Pannekoek Institute for Astronomy, University of Amsterdam, Science Park 904, 1098 XH Amsterdam, the Netherlands

²GRAPPA – Gravitational and Astroparticle Physics Amsterdam, University of Amsterdam, Science Park 904, 1098 XH Amsterdam, the Netherlands

³Department of Physics and Astronomy, University of Southampton, SO17 1BJ, UK

⁴ASTRON – the Netherlands Institute for Radio Astronomy, Oude Hoogeveensedijk 4, 7991 PD, Dwingeloo, the Netherlands

⁵Department of Space, Earth and Environment, Chalmers University of Technology, Onsala Space Observatory, 439 92 Onsala, Sweden

Accepted 2020 December 21. Received 2020 December 18; in original form 2020 August 13

ABSTRACT

During outbursts, the observational properties of black hole X-ray binaries vary on time-scales of days to months. These relatively short time-scales make these systems ideal laboratories to probe the coupling between accreting material and outflowing jets as the accretion rate varies. In particular, the origin of the hard X-ray emission is poorly understood and highly debated. This spectral component, which has a power-law shape, is due to Comptonization of photons near the black hole, but it is unclear whether it originates in the accretion flow itself, or at the base of the jet, or possibly the interface region between them. In this paper, we explore the disc–jet connection by modelling the multiwavelength emission of MAXI J1836–194 during its 2011 outburst. We combine radio through X-ray spectra, X-ray timing information, and a robust joint-fitting method to better isolate the jet’s physical properties. Our results demonstrate that the jet base can produce power-law hard X-ray emission in this system/outburst, provided that its base is fairly compact and that the temperatures of the emitting electrons are subrelativistic. Because of energetic considerations, our model favours mildly pair-loaded jets carrying at least 20 pairs per proton. Finally, we find that the properties of the X-ray power spectrum are correlated with the jet properties, suggesting that an underlying physical process regulates both.

Key words: acceleration of particles – accretion, accretion discs – ISM: jets and outflows – X-rays: binaries – X-rays: individual (MAXI J836–194).

1 INTRODUCTION

Low-mass black hole X-ray binaries are a class of binary systems in which a stellar-mass black hole is accreting mass from a low-mass companion star. Typically, these objects (BHXBs for brevity) are transient sources: they spend most of their lifetime in a faint quiescent state, occasionally shifting into outburst phases lasting a few weeks to a few months (Tetarenko et al. 2016). When this happens, galactic BHXBs become some of the brightest sources in the X-ray sky.

During a full outburst most BHXBs exhibit consistent behaviour, in the form of transitions between spectral states (see Homan & Belloni 2005; Remillard & McClintock 2006 for reviews, and Chen, Shrader & Livio 1997 for a discussion of the variation among outbursts). As the source increases its luminosity while transitioning out of quiescence, its X-ray spectrum is dominated by a hard power-law component, originating in a yet poorly understood ‘corona’ close to the black hole; this is defined as the hard state (HS from now on). ‘As the peak X-ray luminosity of the outburst is approached,

the spectrum softens and becomes increasingly dominated by a blackbody component associated with the accretion disc, while the power law weakens or is completely absent. This is defined as the soft state, while transitional states are referred to as intermediate states (HIMS or SIMS respectively, depending on whether the source is closer to the hard or soft state). As the outburst decays and the luminosity decreases, the source eventually transitions back to the HS and then fades into quiescence. Furthermore, the properties of the X-ray light curve are strongly correlated with spectral properties (e.g. Homan et al. 2001; Homan & Belloni 2005; Heil, Uttley & Klein-Wolt 2015b), with harder states generally displaying more variability than softer ones. X-ray variability therefore provides an independent estimator of the properties of a source during an outburst (e.g. Pottschmidt et al. 2000, 2003; Belloni et al. 2005; Cassatella, Uttley & Maccarone 2012; Heil, Uttley & Klein-Wolt 2015a). Steady compact jets are detected in the HS and (typically) quenched in the soft state; as the source transitions from one to the other, individual knots of plasma are ejected from the system (e.g. Fender, Belloni & Gallo 2004). A significant fraction (≈ 40 per cent) of outbursts ‘fail’ and only exhibit HSs (Tetarenko et al. 2016).

Observationally, the properties of jet and accretion flow appear to be connected. In BHXBs, this takes the form of a correlation (in the HS) between the radio and X-ray luminosities (e.g. Hannikainen

* E-mail: m.lucchini@uva.nl (ML); T.D.Russell@uva.nl (TDR); S.B.Markoff@uva.nl (SBM)

et al. 1998; Corbel et al. 2000, 2003, 2013; Gallo, Fender & Pooley 2003; Gallo et al. 2014) which probe regions in the outer jet or near the black hole, respectively. The radio/X-ray correlation can be extended to accreting supermassive black holes in jetted active galactic nuclei (AGNs) by including a term accounting for the black hole mass; this extension takes the form of a two-dimensional plane in the three-dimensional phase-space connecting radio luminosity, X-ray luminosity and black hole mass (Merloni, Heinz & di Matteo 2003; Falcke, Körding & Markoff 2004; Plotkin et al. 2012). This plane is called the Fundamental Plane of black hole accretion; its existence implies that to first approximation black hole physics are scale-invariant (Heinz & Sunyaev 2003), and what is learnt from one class of systems could potentially be applied to the other. However, the physics driving this connection between accreting and outflowing materials is still poorly understood. Thanks to their quick evolution, BHXBs are ideal laboratories to probe the disc–jet connection, and how it might be changing as a function of accretion rate.

Currently there is significant debate about the exact nature of the coronal power-law component. In general, coronal emission is believed to originate near the black hole due to inverse Comptonization of soft disc photons by a population of hot electrons. Models of the corona can be broadly categorized in three groups, invoking either (a) the innermost regions of a radiatively inefficient accretion flow (RIAF; e.g. Narayan & Yi 1994; Yuan et al. 2007) or (b) a slab corona extending over the accretion disc (e.g. Haardt & Maraschi 1993; Haardt, Maraschi & Ghisellini 1994) or (c) a compact location above the black hole (the so-called lamp-post geometry, e.g. Matt, Perola & Piro 1991; Matt et al. 1992; Martocchia & Matt 1996; Beloborodov 1999) as the location of the Comptonized emission. In the latter scenario, the base of a jet is often invoked as a natural physical realization of the lamp-post (e.g. Markoff, Nowak & Wilms 2005; Maitra et al. 2009; Dauser et al. 2013; Kara & et al. 2019). Additional contributions from non-thermal synchrotron emission originated in the jet have also been invoked (e.g. Markoff, Falcke & Fender 2001).

In jet models, the emitting leptons are typically assumed to be fully relativistic throughout the outflow, starting from the base (e.g. Falcke & Biermann 1995; Markoff et al. 2001; Potter & Cotter 2012; Malzac 2014). Connors et al. (2019) highlighted that this assumption has a large impact on the Comptonization spectra from their launching regions (or in other words, from the ‘lamp-post like’ region), and is in tension with observations if one is to fit the full X-ray spectrum without additional contributions, like non-thermal synchrotron emission from the jet. If the electron distribution is relativistic (its temperature is $T_e \gg 511\text{keV}$ or equivalently the minimum lepton Lorentz factor is $\gamma_{\min} \gg 1$), then fitting the data requires the X-ray emitting regions to be extended and very optically thin ($\tau \leq 0.01$). In this regime, the Comptonization spectrum is produced in a single scattering, rather than a superposition of many scatterings. The resulting X-ray spectrum shows significant curvature, and is unlike a typical power-law continuum produced by many consecutive scatterings in more optically thick ($\tau \approx 0.1$ – 1) media (e.g. Grebenev et al. 1993; Zdziarski et al. 1997; Barret et al. 2000; Joinet et al. 2007). The only way to produce a power-law X-ray emission with these parameters is to fine tune different spectral components, possibly originating in different regions of the system. As a result, past works often invoked a mix of synchrotron and Comptonization from two different emitting regions in order to match the X-ray spectra of BHXBs (e.g. Markoff et al. 2005; Nowak et al. 2011), particularly the hardening above 10 keV: the hard X-rays are due to Comptonization in the jet base, while the soft X-rays are due to synchrotron emission produced downstream in the jet. This

Table 1. X-ray data used in this work. The epochs considered in the spectral analysis have their dates shown in bold font; all the others were only used in the timing analysis.

Instrument	ObsID	Date and time
<i>RXTE/PCA</i>	96371-03-01-00	Aug 31, 11:09–15:54
<i>RXTE/PCA</i>	96371-03-02-00	Sep 01, 13:28–16:06
<i>RXTE/PCA</i>	96371-03-03-00	Sep 02, 11:30–13:21
<i>RXTE/PCA</i>	96371-03-03-01	Sep 04, 07:23–08:41
<i>Swift/XRT</i>	00032085001	Aug 30, 16:47–17:04
<i>RXTE/PCA</i>	96438-01-01-04	Sep 14, 20:12–20:31
<i>Swift/XRT</i>	00032087009	Sep 13, 06:54–21:34
<i>RXTE/PCA</i>	96438-01-02-03	Sep16, 09:13–10:25
<i>RXTE/PCA</i>	96438-01-02-00	Sep 17, 20:13–20:51
<i>Swift/XRT</i>	00032087012	Sep 16, 02:05–02:28
<i>RXTE/PCA</i>	96438-01-03-01	Sep 25, 19:05–19:42
<i>RXTE/PCA</i>	96438-01-03-05	Sep 26, 21:21–21:34
<i>RXTE/PCA</i>	96438-01-03-02	Sep 27, 19:51–20:55
<i>Swift/XRT</i>	00032087017	Sep 26, 04:14–19:16
<i>RXTE/PCA</i>	96438-01-05-02	Oct 11, 18:33–18:57
<i>RXTE/PCA</i>	96438-01-05-05	Oct 12, 13:21–13:55
<i>RXTE/PCA</i>	96438-01-05-06	Oct 13, 19:06–19:26
<i>Swift/XRT</i>	00032087024	Oct 12, 20:06–20:23
<i>RXTE/PCA</i>	96438-01-12-01	Oct 26, 06:18–07:18
<i>RXTE/PCA</i>	96438-01-12-02	Oct 27, 17:45–18:14
<i>RXTE/PCA</i>	96438-01-12-03	Oct 28, 09:51–10:10
<i>Swift/XRT</i>	00032087024	Oct 27, 05:06–11:45

combination also resulted in lower reflection fractions as some of the hardening was absorbed into the continuum. Connors et al. (2019) pointed out that this is inconsistent with the observed hard lags in XRBs (Kotov, Churazov & Gilfanov 2001; Arévalo & Uttley 2006). If instead the synchrotron emission does not extend all the way to the soft X-rays, Connors et al. (2019) showed that the base of the jet can contribute a few to ≈ 50 per cent of the X-ray flux, but the bulk of it has to be produced in a different region (such as the hot accretion flow). In this paper we show that this need not be the case if the electrons in the jet base are in the mildly relativistic regime, in which $T_e \leq 511\text{keV}$, and contains a moderate (≈ 20) number of electron–positron pairs for each proton.

MAXI J1836–194 is a low-mass BHXB that went into outburst during late August 2011; it was quickly identified as a black hole candidate (Miller-Jones et al. 2011; Negoro et al. 2011; Russell et al. 2011; Strohmayer & Smith 2011). Later studies of the source’s X-ray spectra found that the black hole spin is likely high (≈ 0.9), and that the hard X-rays show a very prominent reflection component (Reis et al. 2012; Dong et al. 2020). Rather than undergoing a full outburst, the source reached a hard-intermediate state (HIMS) before going back into the HS and fading into quiescence (Ferrigno et al. 2012). The source was the target of an extensive multiwavelength campaign, which produced some of the best spectral coverage of a BHXB outburst to date (Russell et al. 2013, 2014b). Péault et al. (2019) modelled its evolving jet by combining the radio-through-infrared SEDs with the X-ray power spectra, finding that the jet emission is consistent with an internal shock scenario. MAXI J1836–194 is unique in that its viewing angle ($4^\circ < \theta < 15^\circ$; Russell et al. 2014a) is the lowest in the X-ray binary population to date, and is thus comparable to many jet-dominated AGNs such as M87 or (some) blazars. The low viewing angle, combined with the excellent multiwavelength data, makes it an excellent source to study both the disc–corona–jet connection and scale-invariant models of black hole accretion.

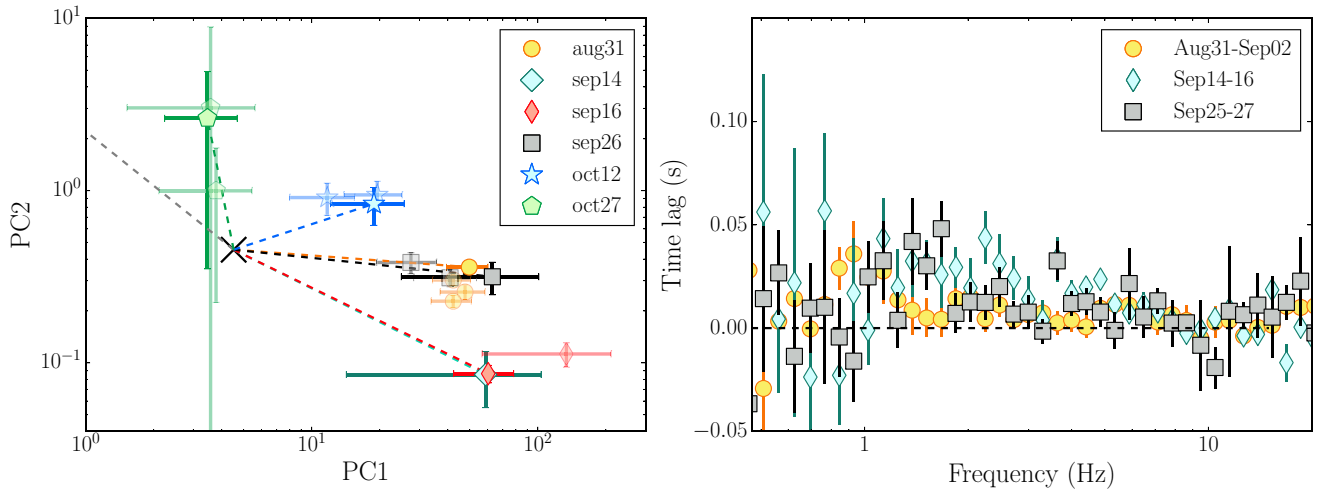


Figure 1. Evolution of X-ray timing properties of the source during the outburst. The left-hand panel shows the evolution of the source power spectra in the (power) colour–colour diagram (Heil et al. 2015a). Opaque points indicate the *RXTE* observation closest to *Swift* coverage; transparent ones indicate all the other epochs reported in Table 1. The right-hand panel shows our tentative ($\approx 2\text{--}3\sigma$) detection of a hard lag between the 3.5–5 and 7–13 keV bands, with hard photons lagging soft ones by $\approx 10^{-2}$ s.

The goal of this paper is to quantify the conditions required for the base of the jet to comprise the X-ray emitting corona, particularly in terms of size, particle temperature and mass content. We do so by modelling the six multiwavelength SEDs presented in Russell et al. (2014b) with the `bljet disc+jet` model (Lucchini et al. 2019a, Paper I from now on), which in itself is an extension of the `agnjet` code (Markoff et al. 2001, 2005; Maitra et al. 2009; Connors et al. 2019). `bljet` was originally developed for modelling AGN SEDs, and is applied to BHXBs here for the first time. Furthermore, we improve the broad-band coverage of the source by analysing additional hard X-ray spectral and timing data gathered by *RXTE*. Finally, we combined timing and spectral information in our modelling with the goal of providing a clearer picture of the coupling between these two observables. The paper is structured as follows. In Section 2 we describe the analysis of the data used, in Section 3 we discuss our full disc+jet multiwavelength model and apply it to the data, in Section 4 we discuss our findings, and in Section 5 we draw our conclusions.

2 DATA ANALYSIS

In this paper we model the six quasi-simultaneous SEDs from Russell et al. (2014b), including additional hard X-ray data provided by *RXTE*. The details of the radio and *Swift*/XRT data reduction are presented in Russell et al. (2014b). The infrared, optical, and UV data were first presented in Russell et al. (2013). We note that in this paper we label epochs based on the date of the *RXTE*/PCA pointings, while Russell et al. (2014b) used the date of the radio observations instead.

2.1 *RXTE* data reduction

We searched the HEASARC archives for *RXTE* observations close to our *Swift*/XRT observations. The resulting ObsIDs are reported in Table 1. We used the standard *RXTE* tools in HEASOFT, version 6.26.1, through the CHROMOS¹ pipeline (Gardenier & Uttley 2018, which also contains the details of the data reduction process) to extract PCA

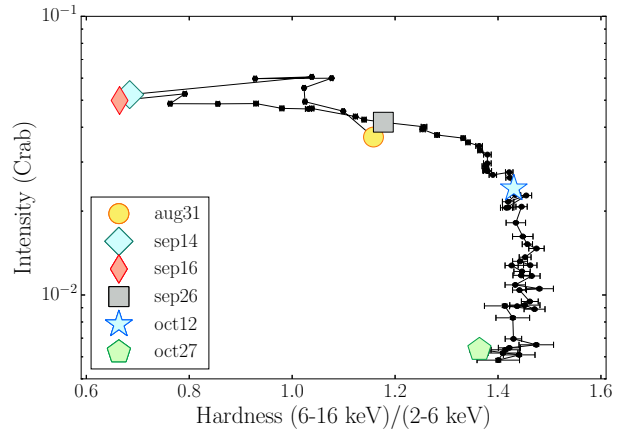


Figure 2. Hardness/intensity diagram of *MAXI* J1836–194 during its outburst; hardness is defined as the ratio between 6–16 and 2–6 keV count rates.

energy spectra and light curves. The spectra were extracted from standard-2 events in the proportional counter unit 2 (PCU2) only. The light curves were extracted from good-xenon events in the 3–13, 3–5.5, and 7–13 keV bands. We jointly modelled the *Swift*/XRT and only the *RXTE*/PCA spectra closest in time to each *Swift* observation, but analysed the light curves of every observation.

2.2 Timing analysis

We investigated the X-ray variability of the source by analysing its *RXTE* light curves using Fourier domain techniques. The CHROMOS pipeline automatically computes the noise-subtracted averaged power spectrum (PSD), correcting for background and dead time detector effects, as well as the power colours of the PSD for each observation chosen. Power colours are the ratio of the integral of the PSD in four bands (band 1: 0.0039–0.031 Hz, band 2: 0.031–0.25 Hz, band 3: 0.25–2.0 Hz, and band 4: 2.0–16.0 Hz); the two power colours are defined as band 3/band 1 and band 2/band 4, respectively, and can be thought of as a way to quantify the shape of the PSD regardless of its normalization. This calculation is identical to that presented in

¹<https://github.com/davidgardenier/chromos>

Gardenier & Uttley (2018). The evolution of the source in the power colour–colour diagram is shown in the left-hand panel of Fig. 1. The power spectral hue, as in Heil et al. (2015a), is defined as the angle between the location of each point in the power colour–colour diagram, and a line angled 45° towards the top left of the panel, shown in Fig. 1 by the grey dashed line. The spectral evolution of the source, presented in Russell et al. (2015), in the hardness–intensity diagram is shown in Fig. 2. The timing and spectral properties of the source are well correlated, with epochs with larger hues being brighter and softer.

We also computed the Fourier time lag of the 7–13 keV light curve with respect to the 3–5.5 keV light curve following the recipe described in Uttley et al. (2014), through the timing software STINGRAY (Huppenkothen et al. 2019). In particular we used 8192 bin per segment and logarithmic binning factor of 1.1. In order to increase the statistics we stacked the light curves in three groups: one with all the observations between August 31 and September 2, a second one with data collected between September 14 and 16, and a last one considering the all observations between September 25 and 27. We note that these epochs have power colours consistent with each other, indicating that the variability properties of the source are relatively similar in the three periods chosen. Results are shown in the right-hand panel of Fig. 1. Even though the statistics are poor, we found evidence ($\approx 2\text{--}3\sigma$) for a hard lag between 1 and 10 Hz for all of the early part of the outburst. In the latter epochs, the statistics are insufficient to detect such a hard lag.

This hard lag is consistent with the standard propagating fluctuations model (e.g. Lyubarskii 1997) and is a first hint that the X-ray emitting region is located close to the black hole, rather than downstream ($\approx 10^3 R_g$) in the jet.

2.3 X-ray spectral fits

All spectral fits in this work are performed in ISIS version 1.6.2–35 (Houck & Denicola 2000). The data are fit first by running the `subplex` least- χ squared algorithm to get close to a good fit, and then by running the `emcee` Markov Chain Monte Carlo algorithm (Foreman-Mackey et al. 2013). We use 20 walkers per free parameter for each chain. All the walkers are initialized uniformly within 1 per cent of the best-fitting values found by `subplex`. We define the chain as having converged after it reaches the point when (a) the acceptance rate of the chain stabilizes and (b) the posterior distribution is unchanged. This results in chain lengths of typically a few thousand; depending on the behaviour of the posterior distribution in each fit, taking the initial burn-in period to be between 30–75 per cent of the chain. We define the value of the best-fitting parameter as the median of the walker distribution after the chain has converged, as we found that this produces a better fit to the data than the peak of the posterior. We define the 1σ uncertainty as the interval in the posterior distribution which contains 68 per cent of the walkers, after excluding the burn-in period.

Swift/XRT spectra were re-binned to a minimum signal-to-noise ratio of 20 per bin in every epoch except on August 31st. In this epoch, due to the lack of statistics, we only used a minimum signal-to-noise ratio of 10. In every *Swift*/XRT spectrum we kept data between 0.5 and 10 keV. *RXTE*/PCA spectra were re-binned to a minimum signal-to-noise ratio of 4.5 per bin in order to use χ^2 statistics. Depending on the quality of the data, we either considered data between 3 and 20 keV (on October 27th) or 3 and 30 keV (every other epoch). In every epoch we added a 1 per cent systematic uncertainty to the *RXTE*/PCA data in order to account for cross-calibration uncertainties. Furthermore, when modelling the *RXTE*

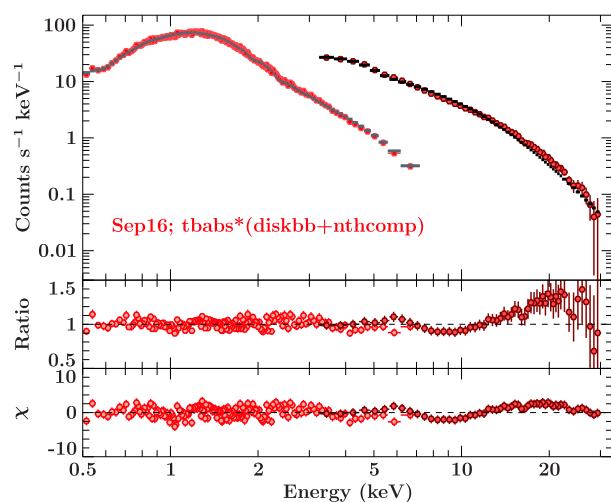


Figure 3. Sep16 *Swift*/XRT (light red) and *RXTE*/PCA (dark red) spectra fitted without a disc reflection contribution. As shown by the residual and ratio plots, the characteristic reflection signatures are very prominent in the *RXTE* band.

spectra we multiplied the data by a constant in order to account for the observations not being strictly simultaneous, as well as any additional cross-calibration uncertainties.

Before applying our physical multiwavelength model to the entire data set, we fit the X-ray spectra alone with phenomenological models in order to gain a better understanding of the broad properties of the system. In particular, we aim to constrain the origin of the coronal X-ray emission. In the framework of a jet-dominated model, two radiative mechanisms are viable candidates to produce X-ray emission: inverse Compton scattering of disc and cyclo-synchrotron photons in the jet base, located close to the black hole, or non-thermal synchrotron emission originating further downstream in the jet. In the former, much stronger reflection features are predicted than in the latter (Markoff & Nowak 2004); while typical sources favour the inverse Compton+reflection scenario (e.g. GX 339–4; Connors et al. 2019), exceptions do exist, such as XTE 1118+480 (Miller et al. 2002; Maitra et al. 2009) and XTE 1550-564 (Russell et al. 2010). To first order, one would expect that a low inclination source like MAXI J1836–194 would be a prime candidate for detection of the non-thermal synchrotron component, as it originates in regions further out in the jets, that are more beamed than the jet base.

We first fit the spectra with an (absorbed) Comptonization continuum (`nthcomp`; Życki, Done & Smith 1999) and a standard accretion disc component (`diskbb`; Makishima et al. 1986), with the exception of the October 27th data for which we only use `nthcomp` (no disc contribution was found in the soft X-rays). In order to account for the non-simultaneity of the *Swift*/XRT and *RXTE*/PCA observations we included a multiplicative calibration constant between the data sets; the final syntax of the model is `constant*tbabs*(nthcomp+diskbb)` or `constant*tbabs*(nthcomp)`, respectively. We freeze the constant to 1 for the *Swift*/XRT data and leave it free for the *RXTE*/PCA data. This model fits the soft X-ray data fairly well, but significant residuals are present around 6 and above 10 keV in every epoch except October 27th; Fig. 3 shows this for the data on September 16th. Both of these are clear signs of reflection, indicating that the source of the X-rays is located close to the black hole and accretion disc.

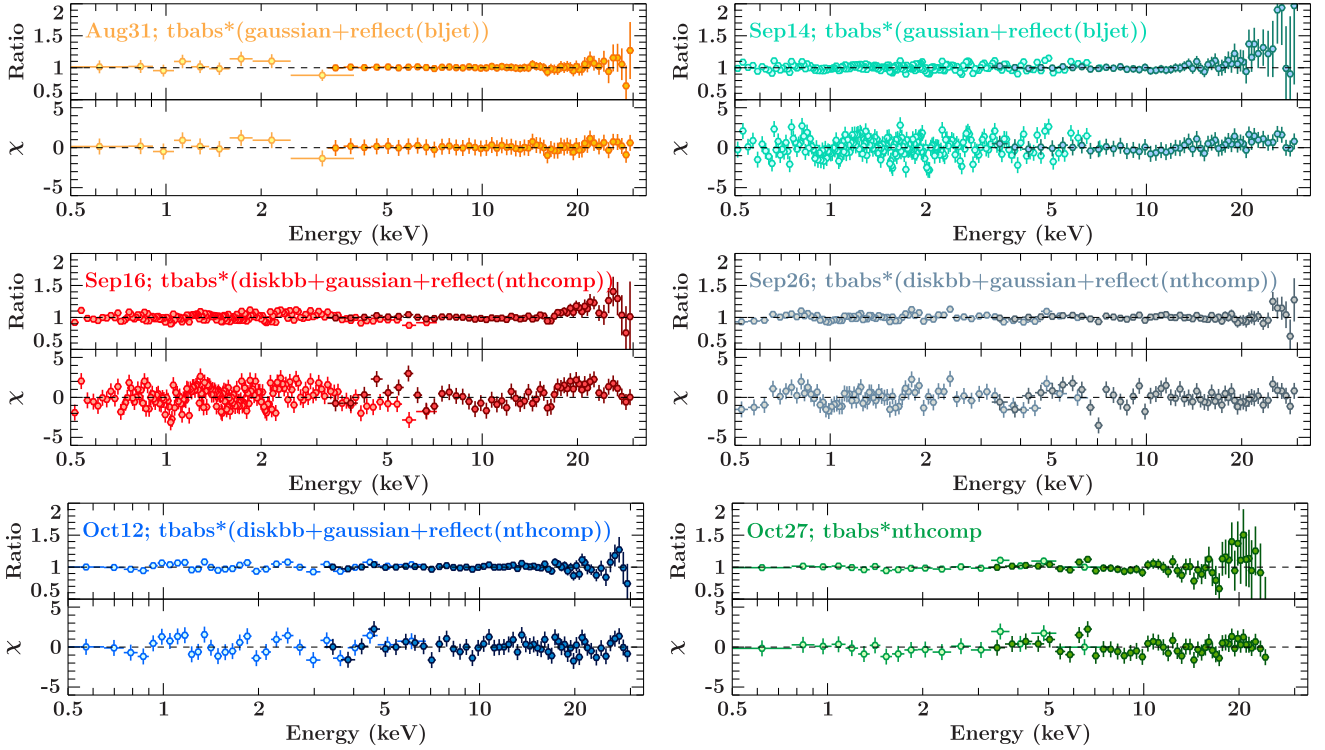


Figure 4. Ratio (data/model) and residual (data-model/error) plots for the X-ray spectra of the source throughout the outburst. On August 31st, October 12th, and October 27th the model is in good agreement with the data. On September 13 and 16, the *Swift*/XRT spectra appear to require a softer power law than the *RXTE*/PCA ones, resulting in residuals above 10 keV. On September 16 and 26 there are systematic residuals between 4 and 7 keV, likely caused by the simple reflection model used here. The χ^2/dof are 52.63/55, 337.98/257, 328.01/224, 121.28/115, 68.68/74, and 44.04/63, respectively.

Table 2. Best-fitting parameters for our phenomenological fits of the *Swift* + *RXTE* spectra.

Epoch	Constant	N_{H} 10^{22} cm^{-2}	Diskbb Norm (10^3)	kT_{disc} (keV)	nthcomp Norm	Γ	relfrac	Gaussian Norm (10^{-4})	Gaussian σ (keV)
Aug 30	$1.15^{+0.07}_{-0.09}$	$0.38^{+0.09}_{-0.07}$	62^{+73}_{-44}	$0.19^{+0.03}_{-0.02}$	$0.27^{+0.02}_{-0.03}$	$2.10^{+0.01}_{-0.02}$	$0.80^{+0.11}_{-0.07}$	$7.2^{+1.8}_{-1.3}$	$0.78^{+0.24}_{-0.19}$
Sep 13	$0.90^{+0.01}_{-0.01}$	$0.36^{+0.01}_{-0.01}$	$5.64^{+0.31}_{-0.27}$	$0.40^{+0.01}_{-0.01}$	$0.77^{+0.02}_{-0.03}$	$2.85^{+0.04}_{-0.05}$	$2.74^{+0.40}_{-0.36}$	$1.8^{+2.4}_{-1.3}$	$1.4^{+2.6}_{-0.9}$
Sep 16	$0.94^{+0.01}_{-0.01}$	$0.396^{+0.005}_{-0.005}$	$8.43^{+0.38}_{-0.35}$	$0.414^{+0.004}_{-0.004}$	$0.27^{+0.01}_{-0.01}$	$2.88^{+0.03}_{-0.03}$	$3.19^{+0.34}_{-0.30}$	$2.1^{+2.4}_{-1.5}$	$1.35^{+1.01}_{-0.49}$
Sep 26	$1.26^{+0.02}_{-0.02}$	$0.36^{+0.02}_{-0.02}$	$11.5^{+4.4}_{-2.7}$	$0.25^{+0.02}_{-0.01}$	$0.32^{+0.02}_{-0.02}$	$2.34^{+0.07}_{-0.06}$	$2.46^{+0.54}_{-0.52}$	$36.9^{+5.9}_{-6.2}$	$3.1^{+0.3}_{-0.4}$
Oct 12	$1.05^{+0.02}_{-0.02}$	$0.35^{+0.03}_{-0.06}$	$9.5^{+6.7}_{-5.6}$	$0.18^{+0.02}_{-0.01}$	$0.12^{+0.01}_{-0.01}$	$1.81^{+0.04}_{-0.03}$	$0.48^{+0.18}_{-0.11}$	$7.4^{+5.6}_{-2.9}$	$1.6^{+0.7}_{-0.5}$
Oct 27	$0.56^{+0.02}_{-0.01}$	$0.31^{+0.02}_{-0.01}$	//	//	$0.050^{+0.002}_{-0.001}$	$1.69^{+0.02}_{-0.02}$	//	//	//

We model these signatures with a Gaussian centred at 6.4 keV and the `reflect` model of Magdziarz & Zdziarski (1995). The syntax for the final X-ray spectral model in epochs that show reflection signatures is `constant*tbabs*(reflect(nthcomp)+diskbb+Gaussian)`. This model provides a fair description of the data; the residuals for each epoch are shown in Fig. 4 and the best-fitting parameters along with uncertainties are reported in Table 2. However, on the spectra of September 13, 16, and 26 residual features remain above 10 keV, and between 4 and 7 keV. We stress that the goal of these phenomenological spectral fits is simply to constrain the location of the X-ray emitting region by quantifying the presence or absences of reflection features. These phenomenological spectral fits favour the inverse Compton scenario, as the reflection features are very strong, particularly near the peak of the outburst, in agreement with other works (Reis et al. 2012; Dong et al. 2020). This picture is further strengthened by our tentative detection of a hard lag in the X-ray light curves.

3 MULTIWAVELENGTH MODEL

Next, we model the six multiwavelength SEDs with an updated version of the multizone jet model `bljet`, the full details of the original model are presented in Paper I. Briefly, `bljet` originated as an extension of the `agnjet` model of Markoff et al. (2005); it is designed to mimic the conservation of energy within the jets similar to ideal MHD, and now confirmed by GRMHD simulations. In this scenario, the jet is launched with a high initial magnetization, and accelerates by converting its Poynting flux into bulk kinetic energy (e.g. McKinney 2006; Chatterjee et al. 2019). Unlike `agnjet`, which is limited to mildly relativistic outflows, the jet in `bljet` can reach arbitrarily high Lorentz factors.

The model assumes that a fraction N_j of the Eddington luminosity is injected at the jet base near the black hole, the resulting internal energy is divided between magnetic fields, cold protons and thermal leptons. This injection occurs in a cylindrical region characterized by an initial radius r_0 , with an aspect ratio $h = R_0/z_0 = 2$ as in papers

1 and 2. The jet then accelerates up to a final speed γ_f by converting its initial magnetic field into bulk kinetic energy; this occurs at a distance z_{acc} , where the magnetization reaches a value $\sigma_f \leq 1$. The jet is assumed to have a parabolic shape in the bulk acceleration region, and a conical shape afterwards, in agreement with VLBI observations of AGNs (e.g. Hada et al. 2016; Mertens et al. 2016). The details of the magnetic-to-kinetic energy conversion and of the rest of the jet dynamics are detailed in Paper I.

At a distance z_{diss} from the black hole, the jet experiences a dissipation region in which particle acceleration begins. From this point onward, a fraction (which initially we take to be 10 per cent, as in Paper I and Lucchini, Krauß & Markoff 2019b) of the leptons is channelled in a non-thermal tail. Following paper I, we set $z_{\text{diss}} = z_{\text{acc}}$ in order to reduce the number of free parameters, but note that it does not always need to be the case (Lucchini et al. 2019b).

3.1 Particle distributions

Previous versions of `agnjet` and `bljet` assumed the particle distributions were always relativistic and thus calculated in energy space, meaning that the leptons are assumed to have an energy E_e and a corresponding Lorentz factor:

$$\gamma_e = E_e/m_e c^2; \quad (1)$$

the lepton particle distribution was defined as $N(E_e)$. This approximation is appropriate in the relativistic regime ($\gamma_e \gg m_e c^2$, implying $T_e \gg 511$ keV), but it is incorrect if the electron temperature is non-relativistic, thus in older versions the minimum temperature allowed was ~ 511 keV.

In this work, we have updated the calculations to be in momentum space in order to treat both relativistic and non-relativistic particle distributions. Similar to Ghisellini, Haardt & Svensson (1998), we describe the leptons starting from their relativistic momentum p_e , which corresponds to a Lorentz factor:

$$\gamma_e(p) = \sqrt{\varrho^2 + 1}, \quad (2)$$

where $\varrho = p/m_e c$ is the electron momentum in units of $m_e c$. Unlike equation (1), this expression is valid in both the relativistic ($p_e \gg m_e c^2$, or $\gamma_e(p) \gg 1$) and non-relativistic ($p_e \leq m_e c^2$, or $\gamma_e(p) \approx 1$) regimes. For the rest of this section, we will drop the subscript e , implicitly assuming that γ and ϱ , as well as the temperature T_e and number density n_e , refer to leptons.

As in older versions of the model, we assume that up to z_{diss} within the jet base, the particle distribution remains thermal and that the jet is isothermal, meaning that adiabatic losses are neglected and therefore T is unchanged. In this inner jet region, the particle distribution is described by the Maxwell–Jüttner distribution:

$$N(\gamma) = N_{\text{th},0} \varrho^2 e^{-\frac{\gamma \varrho}{\theta}}, \quad (3)$$

where $N_{\text{th},0} = n/m_e c^3 \theta K_2(1/\theta)$ is the normalization of the Maxwell–Jüttner so that its integral over particle momenta is n , and $\theta = kT/m_e c^2$ the dimensionless temperature.

Beyond z_{diss} , where we assume non-thermal particle acceleration begins, we calculate the radiating particle distribution by solving the steady-state continuity equation along the length of the jet:

$$N(\varrho) = \frac{\int_{\varrho}^{\infty} Q(\varrho) d\varrho}{\dot{\varrho}_{\text{ad}} + \dot{\varrho}_{\text{rad}}}, \quad (4)$$

where $Q(\varrho)$ is the injection term, $\dot{\gamma}_{\text{rad}}$ is the radiative loss term, and $\dot{\gamma}_{\text{ad}}$ is the adiabatic loss term. The injection term $Q(\varrho)$ is assumed to be a mixed distribution, composed by a Maxwell–Jüttner thermal

distribution plus an exponentially cut-off power-law tail:

$$Q(\varrho) = \begin{cases} \frac{(1 - f_{\text{pl}}) N_{\text{th},0}}{t_{\text{inj}}} \varrho^2 e^{-\frac{\gamma \varrho}{\theta}}, & \varrho < \varrho_{\text{min,pl}} \\ \frac{(1 - f_{\text{pl}}) N_{\text{th},0}}{t_{\text{inj}}} \varrho^2 e^{-\frac{\gamma \varrho}{\theta}} + \frac{f_{\text{pl}} N_{\text{th},0}}{t_{\text{inj}}} \varrho^{-s} e^{-\frac{\gamma \varrho}{\varrho_{\text{max,pl}}}}, & \varrho \geq \varrho_{\text{min,pl}} \end{cases} \quad (5)$$

where f_{pl} is the fraction of particles channelled into the power law, s is the slope of the injected power law, $N_{\text{th},0}$ is the normalization of the thermal particles as in equation (3), $N_{\text{th},0} = (1 - s)/(\varrho_{\text{max}}^{1-s} - \varrho_{\text{min}}^{1-s})$ is the normalization of the power-law tail. The two distributions are normalized such that the total lepton number density n is always $n = n_{\text{th}} + n_{\text{pl}}$. The injection time is defined as $t_{\text{inj}} = r(z)/c$, where $r(z)$ is the radius of the jet and c the speed of light. $\varrho_{\text{min,pl}} = \langle \varrho_{\text{th}}(\theta) \rangle$ is the minimum of the non-thermal particle distribution, which is always assumed to be equal to the average dimensionless momentum of the thermal distribution for a given dimensionless temperature θ . $\varrho_{\text{max,pl}}$ is the maximum particle momentum. The corresponding Lorentz factor from which it is calculated is derived identically to paper I; for completeness this derivation is reported below. From each distribution in ϱ space, the corresponding distribution in Lorentz factor space is always calculated as

$$N(\gamma) = N(\varrho) \frac{d\varrho}{d\gamma}. \quad (6)$$

The radiative loss term is defined as in Ghisellini et al. (1998):

$$\dot{\varrho}_{\text{rad}} = \frac{4\sigma_T c U_{\text{rad}}}{3m_e c^2} \varrho \gamma, \quad (7)$$

where σ_T is the Thomson cross-section. In this paper, the radiative energy density U_{rad} includes only the synchrotron term, such that $U_{\text{rad}} = U_b$. This is because at the distances over which particle acceleration occurs ($\approx 10^3$ – $10^6 R_g$) radiative cooling is dominated by cyclo-synchrotron, and the disc contribution is negligible.

Similarly, the adiabatic loss term is defined as

$$\dot{\varrho}_{\text{ad}} = \frac{\beta_{\text{eff}} c}{r} \varrho, \quad (8)$$

where β_{eff} is the effective expansion speed of the jet. If the jet is purely isothermal then adiabatic losses are assumed to be entirely balanced by some acceleration mechanism, and thus $\dot{\gamma}_{\text{ad}} = 0$ and only radiative losses are present. When this happens, the observed slope of the steady-state non-thermal particle distribution will be $p = s + 1$ for every electron Lorentz factor γ . If the jet is purely adiabatic rather than isothermal, then no reacceleration is present and β_{eff} is the true expansion speed of the jet; this may or may not dominate over radiative cooling, depending on the particle energy. In practice, the true value of β_{eff} is somewhere between these two extremes, and allows one to fine-tune the break in the particle distribution as we did in Paper I (in that work, $f_b = \beta(z)/\beta_{\text{eff}}$, where $\beta(z)$ is the jet speed along the z -axis).

The calculation of the maximum electron energy is unchanged from previous versions. We define the particle acceleration time-scale as

$$t_{\text{acc}} = \frac{4\gamma m_e c}{3f_{\text{sc}} e B(z)}, \quad (9)$$

where e is the charge of the electron, $B(z)$ the magnetic field strength along the jet, and f_{sc} is a free parameter (described in Jokipii 1987) that quantifies the efficiency of particle acceleration, originally in terms of a relative shock velocity and ratio to scattering mean-free

path, but now this is grouped into a single parameter. The maximum energy of the injected particles $\gamma_{\max, \text{pl}}$ is found by solving

$$t_{\text{acc}}^{-1} = t_{\text{ad}}^{-1} + t_{\text{rad}}^{-1}, \quad (10)$$

where t_{ad} and t_{rad} are the time-scales derived from equations (7) and (8), respectively. This results in a maximum injected Lorentz factor:

$$\gamma_{\max, \text{pl}}(z) = \frac{-3m_e c^2 \beta_{\text{eff}}}{8\sigma_t U_{\text{rad}}(z)r(z)} + \frac{1}{2} \sqrt{\left(\frac{-3m_e c^2 \beta_{\text{eff}}}{4\sigma_t U_{\text{rad}}(z)r(z)}\right)^2 + \frac{3f_{\text{sc}} e B(z)}{4\sigma_t U_{\text{rad}}}}. \quad (11)$$

3.2 The role of pairs

The second update to `bljet` presented in this paper consists of a more thorough calculation of the pair content of the jets; previously, `bljet` only considered jets which contain one proton per electron. Abandoning the assumption of one pair per electron requires a change in the equipartition conditions of the plasma injected at the jet base, compared to [Paper I](#). Similarly to [Paper I](#), we define the power injected in the jet as

$$N_j = 2\gamma_0 \beta_0 c \pi R_0^2 (U_{b,0} + U_{p,0} + U_{e,0}), \quad (12)$$

where the factor 2 accounts for the launching of two jets, γ_0 is the initial jet Lorentz factor, β_0 the initial jet speed in units of c , R_0 the radius of the jet base, $U_{b,0} = B^2/8\pi$ the energy density in magnetic fields in the jet base, $U_{p,0} = n_p m_p c^2$ the energy density of the (cold) protons injected in the jet, $U_{e,0} = n_e \langle \gamma \rangle m_e c^2$ the energy density of the injected (hot) electrons. We define the pair content of the jet as $\eta = n_e/n_p$. As in [Paper I](#), we wish to describe how these three components are related to each other in order to fully describe the energy budget of the jet. For convenience, in our code we compute η as a function of the initial magnetization σ_0 and initial plasma- β parameter (defined below), although we stress that this does not imply a physical causality between these numbers. We define the initial magnetization σ_0 as

$$\sigma_0 = \frac{U_{b,0} + P_{b,0}}{U_{p,0} + U_{e,0} + P_{e,0}} = \frac{2U_{b,0}}{U_{p,0} + \Gamma_{\text{ad}} U_{e,0}}, \quad (13)$$

where Γ_{ad} is the adiabatic index of the hot electrons; for simplicity we always take $\Gamma_{\text{ad}} = 4/3$. As long as the electrons do not dominate the energy budget, this approximation (and in general, the electron population) has negligible impact on the jet dynamics. The plasma- β parameter at the jet base is:

$$\beta_{e,0} = \frac{U_{e,0}}{U_{b,0}}. \quad (14)$$

The pair content η is determined by the values of σ_0 and $\beta_{e,0}$, as well as the electron average Lorentz factor $\langle \gamma \rangle$ and does not depend on any other quantities like black hole mass or injected power. This can be shown by solving equation (14) for $U_{b,0}$ and solving equation (13) for η :

$$\sigma_0 = \frac{2U_{e,0}/\beta_{e,0}}{U_{p,0} + \Gamma_{\text{ad}} U_{e,0}} = \frac{2\eta \langle \gamma \rangle n_p m_e c^2}{\beta_{e,0} (n_p m_p c^2 + \Gamma_{\text{ad}} \eta \langle \gamma \rangle n_p m_e c^2)}, \quad (15)$$

which gives

$$\eta(\beta_{e,0}, \sigma_0) = \frac{\beta_{e,0} \sigma_0}{\langle \gamma \rangle (2 - \Gamma_{\text{ad}} \beta_{e,0} \sigma_0)} \frac{m_p}{m_e}. \quad (16)$$

This expression has two critical values, beyond which the combination of σ_0 and $\beta_{e,0}$ results in a charged, unphysical jet. The first critical value occurs when

$$\beta_{e,0} \geq 2/(\Gamma_{\text{ad}} \sigma_0); \quad (17)$$

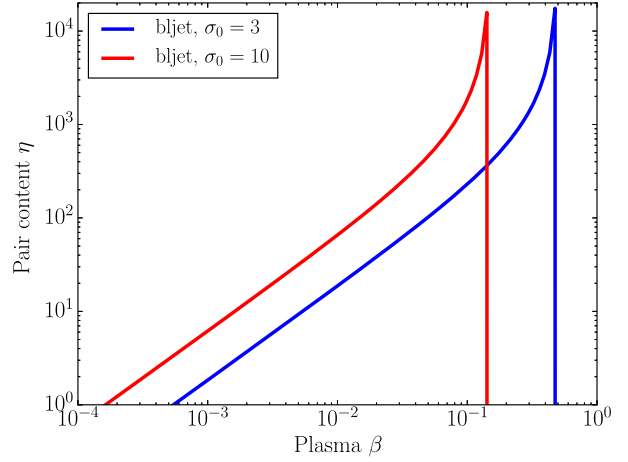


Figure 5. Jet pair content as a function of initial plasma- β $\beta_{e,0}$, for both mildly and highly magnetized jets ($\sigma_0 = 3$ and 10 , respectively), for $T_e = 135$ keV, corresponding to $\langle \gamma \rangle \approx 1.5$.

in this case equation (16) forces the pair content of the jet to be a negative value. The second critical value is

$$\beta_{e,0} = \frac{2\langle \gamma \rangle}{\sigma_0 (m_p/m_e + \langle \gamma \rangle \Gamma_{\text{ad}} \sigma_0)}; \quad (18)$$

in this case, the jet contains exactly one electron per proton, and no pair content is necessary. Within these two critical values, the pair content of the jet for a given initial magnetization is explicitly determined by $\beta_{e,0}$. This is shown in [Fig. 5](#); the magnetization essentially sets the normalization of the relation between η and $\beta_{e,0}$. We note that in the near-relativistic regime which this paper focuses on, $\langle \gamma \rangle \approx 1-3$ and therefore the impact of the electron temperature on the pair content is far smaller than that of $\beta_{e,0}$ and σ_0 .

3.3 Optically thick spectral shapes

The final update to `bljet` presented in this paper is a simple parametrization designed to allow more flexibility in fitting optically thick spectra from compact jets.

The radio spectrum of *MAXI* J1836–194 was highly inverted rather than slightly inverted/flat throughout its outburst ([Russell et al. 2015](#)), with spectral indexes ranging from $\alpha = 0.19$ to $\alpha = 0.70$ (where $F(\nu) \propto \nu^\alpha$). This behaviour cannot be captured by the simple assumption made in the last version of `bljet` of a conical, isothermal jet, which produces a strictly flat radio spectrum ($F(\nu) \propto \nu^0$, [Blandford & Königl 1979](#)), or in the quasi-isothermal treatment in `agnjet` (see, e.g. [Crumley et al. 2017](#)). The optically thick slope is set by the details of the balance between particle acceleration and cooling. `bljet` does not capture these details self-consistently, as this would require a full Fokker–Planck treatment linking particle acceleration and cooling to the jet geometry and dynamics.; therefore, we introduced a phenomenological free parameter in the particle distribution to gradually reduce the emissivity along the jet z -axis. We produce inverted radio spectra by changing the fraction of electrons channelled in the non-thermal tail along the length of the jet:

$$n_{\text{nth}}(z) = n_{\text{nth},0} \left(\frac{\log_{10}(z_{\text{diss}})}{\log_{10}(z)} \right)^{f_{\text{pl}}}, \quad (19)$$

where $n_{\text{nth},0} = 0.1$ is the fraction of particles channelled into the non-thermal tail at z_{diss} , z is the distance from the black hole, and f_{pl} is a free parameter which we fit to the data in order to match the optically

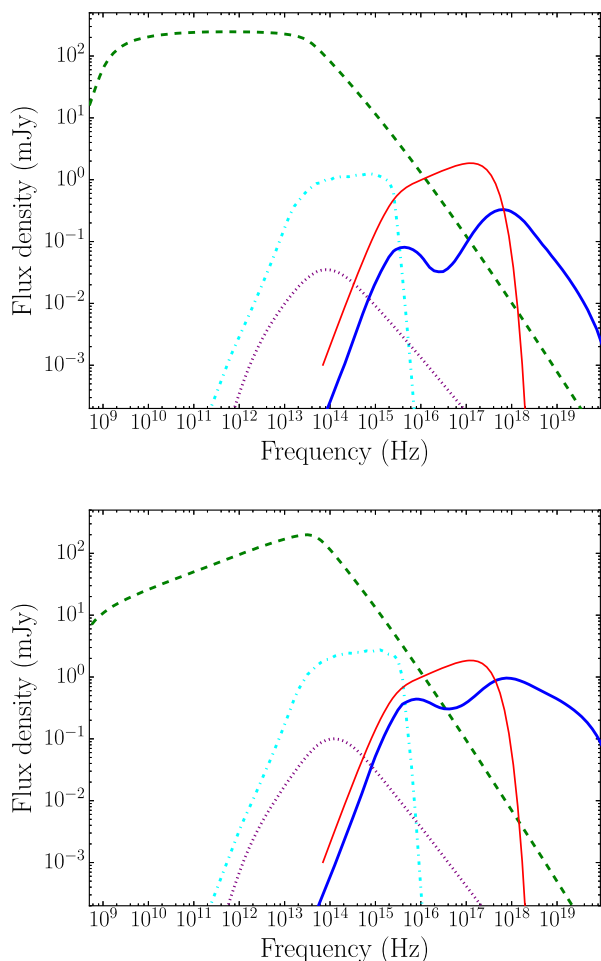


Figure 6. Effect of the f_{pl} parameter on the radio emission. The two SEDs show the difference between jets with different powers ($N_j = 4 \times 10^{-2} L_{\text{Edd}}$ and $N_j = 8 \times 10^{-2} L_{\text{Edd}}$, respectively) and different values of f_{pl} ($f_{\text{pl}} = 0$ and $f_{\text{pl}} = 10$, respectively). Both SEDs assume $M_{\text{bh}} = 10 M_{\odot}$, $D = 7$ kpc, $T_e = 100$ keV, $r_0 = 10 R_g$, $z_{\text{diss}} = 10^4 R_g$, $L_{\text{disc}} = 4 \times 10^{-2} L_{\text{Edd}}$. The dashed green and dotted purple lines represent non-thermal synchrotron and inverse-Compton emission, respectively; the dotted cyan and continuous thick blue lines represent thermal cyclo-synchrotron and inverse-Compton emission from the jet base, respectively; the continuous thin red line represents the emission from the disc.

thick spectral shape. Higher values of f_{pl} suppress the number density of non-thermal particles, mimicking the effect of additional adiabatic cooling and/or reduced particle acceleration, that can hopefully guide the implementation of more physical treatments in future work.

A limitation of this phenomenological approach comes from estimating the jet power by modelling exclusively the optically thick radio flux densities. This is highlighted in Fig. 6. Suppressing the emission from the outer jet in order to invert the radio spectrum can lead to a counter-intuitive regime: due to the shape of the optically thick spectrum, a less powerful, flat-spectrum jet can produce more radio emission than a more powerful, inverted-spectrum one. A similar behaviour would also emerge from a more self-consistent treatment, in which the radio spectrum is inverted with an additional loss term in the FP equation. Either treatment essentially quantifies how much of the jet initial power is converted in radio emission. The goal of this paper is to first quantify this behaviour, before implementing a more self-consistent treatment in future works.

3.4 Model parameters

The focus of this work is to probe the connection between the accretion disc, jet base, and outer compact jet as the source evolves through its outburst; as such, we froze parameters that can be estimated by observations (such as black hole mass), do not greatly impact the SED (such as the location of the cooling break or the outer disc radius), or introduce large amounts of degeneracy (such as the pair content). Furthermore, we fixed the value of N_{H} to that found in Section 2.3.

We assume a black hole mass of $10 M_{\odot}$, a distance of 7 kpc, and a viewing angle of 10° (Russell et al. 2014a, b).

X-ray binary jets are believed to be only mildly relativistic, thus we take $\gamma_f = 3$ (Fender et al. 2004). Constraining the value of σ_f is harder due to modelling degeneracies, particularly in the case of X-ray binary jets in which the non-thermal inverse Compton emission of the jets is not detected. Theoretical models however predict that the magnetization will continue to decrease as the jet accelerates, until the jet either reaches equipartition or is slightly matter-dominated (e.g. Komissarov et al. 2009; Tchekhovskoy, McKinney & Narayan 2009; Ceccobello et al. 2018; Chatterjee et al. 2019). With some exceptions, blazar jets also appear to either be close to equipartition or be somewhat matter dominated (e.g. Ghisellini et al. 2014; Paper I). Therefore, we fix $\sigma_f = 0.1$ in all of our fits.

Similarly to Lucchini et al. (2019b), but unlike Paper I, we find that no additional heating of the electrons is necessary to match the data. Therefore, we fix $f_{\text{heat}} = 1$. In practice, this means that the minimum Lorentz factor of the non-thermal electron distribution is $\gamma_{\text{min,pl}} \approx 1$. The maximum and break energies of the electron distribution are not well constrained. As such, we take $f_{\text{sc}} = 0.1$ and $\beta_{\text{eff}} = 0.1$. This results in $\gamma_{\text{max,pl}} \approx 10^6$ along the jet, varying by a factor of \approx a few with the distance from the black hole. Physically, this means we do not make any a priori assumption on the radiative mechanism responsible for the X-ray emission, and allow the non-thermal synchrotron emission to extend above the *RXTE/PCA* band.

The pair content has a large effect on the SED, changing its normalization and the relative importance of Comptonization. Unfortunately, this effect is almost entirely degenerate with the effect of the injected jet power. As such, we set a pair content of ≈ 20 pairs per proton (changing slightly depending on temperature), giving $\beta_{e,0} = 0.02085$. This results in jet powers roughly on the order of the accretion luminosity; on the other hand, taking one electron per proton requires unreasonable values of the injected jet power. We discuss this choice further in Section 4.

Assuming a standard Shakura–Sunyaev temperature profile $T(R) \propto R^{-3/4}$, the emission from the outer disc is subdominant at every wavelength, hence we always take $R_{\text{out}} = 10^5 R_g$. However, on September 13 and 16 (near the peak of the outburst) we found in preliminary fits that the spectral shape and luminosity, particularly in the UV, cannot be matched by cyclo-synchrotron emission from the jet base. Therefore, we include an additional blackbody peaking at optical/UV wavelengths, likely due to disc irradiation. This is analogous to the models in Russell et al. (2014b) and Péault et al. (2019). In every other epoch this contribution is negligible compared to the jet, and thus we choose to neglect it to reduce the number of free parameters.

In summary, our disc+jet model has up to 13 free parameters: 6 relating to the jet, 4 to the accretion flow, and 3 to the reflection component. These are: the injected jet power N_j , the jet base radius r_0 , the temperature of the electrons at the base T_e (which in our model can be constrained from the optical cyclo/synchrotron bump as well as the X-ray spectrum), the location of particle acceleration z_{diss} , the slope of the accelerated power law s , the (radio) spectral

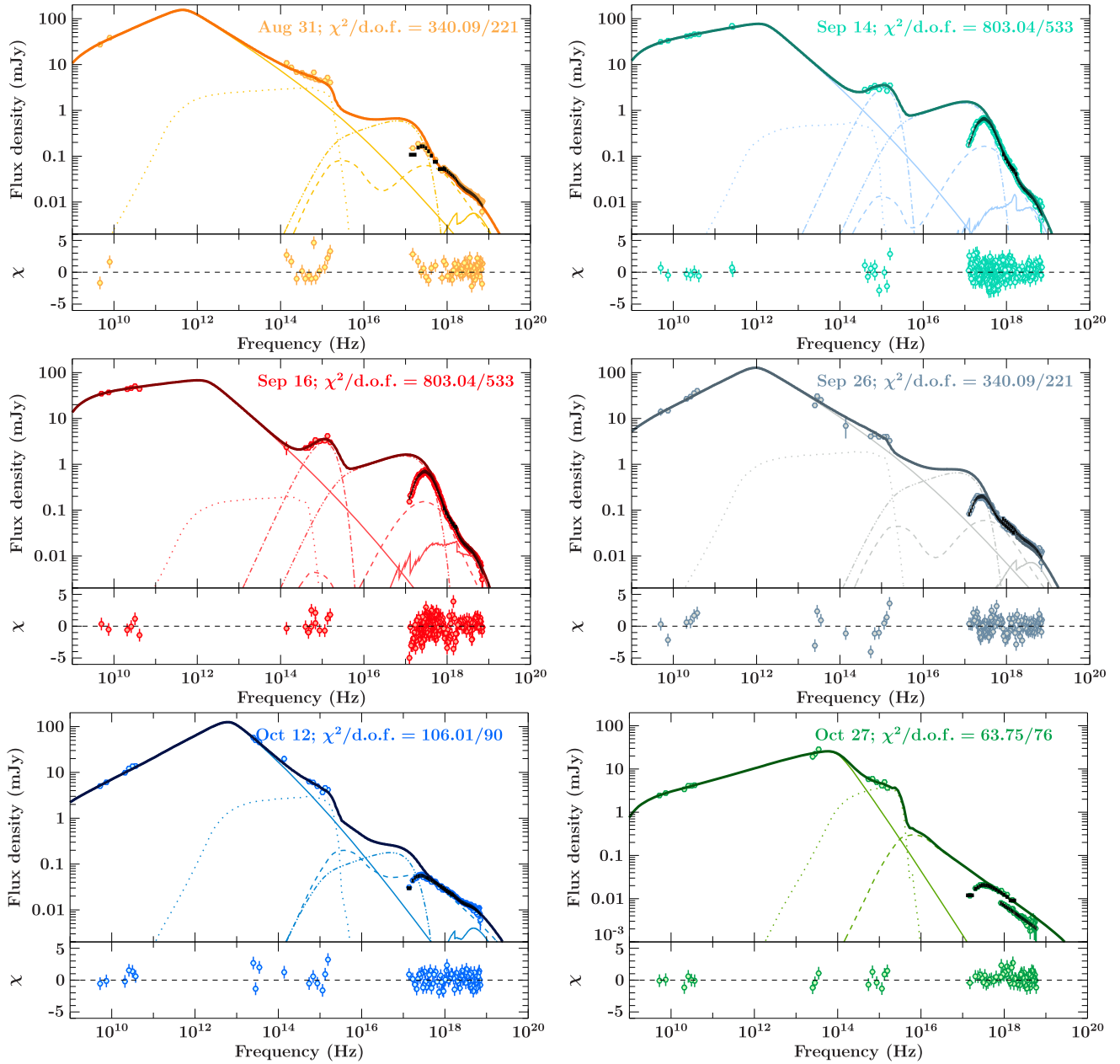


Figure 7. Best fit of each SED; August 30 and September 26, and September 13 and 16 have been fitted jointly, while the October epochs are fit individually. The thick continuous dark line represents the total model, the thin continuous line represents the non-thermal synchrotron emission, the thin dashed line the inverse-Compton emission from the jet base, the thin dotted line represents the thermal cyclo-synchrotron emission from the jet base, the dot-dashed and double dot-dashed lines represent the optical blackbody excess and accretion disc, the triple dot-dashed line represents reflection.

shape parameter f_{pl} , the disc luminosity L_{disc} , the disc inner radius r_{in} , the blackbody luminosity L_{bb} , the blackbody temperature T_{bb} , the reflection fraction, and the normalization and width of the Gaussian. The additional blackbody is only present on September 13 and 16, and the accretion disc and reflection components are absent on October 27. The full list of free and fixed parameters is shown in Table 3.

3.5 Joint fits

We initially modelled each SED independently; while the model reproduced the data very well, we found significant degeneracies in the posterior distributions of the parameters. In order to address these

issues we combined timing and spectral information, and performed joint fits of SEDs that had similar power-colours, using an approach similar to Connors et al. (2019). As shown in Fig. 1, these are September 13/14 and 16, around the peak of the outburst, and August 30/31 and September 26, close to the HS to HIMS and HIMS to HS transitions, respectively. For each of these epochs we take the geometrical parameters of the X-ray emitting regions to be the same; in the case of our model, this means tying the disc truncation radius r_{in} and the jet base radius r_0 . The best-fitting results for the joint (on August 30 and September 26, and September 13 and 16) and individual (October 12 and October 27) fits are reported in Table 4, and all the SEDs are shown in Fig. 7. In every epoch we find that the model is in excellent agreement with the data.

Table 3. Summary of the parameters of `bljet`. The parameters in the top block are frozen, as they can be estimated a priori or do not impact the SED. The parameters in the bottom block, shown in bold font, are kept free and fitted to the data.

Fixed Parameter	Description
$M_{\text{bh}} = 10 M_{\odot}$	Mass of the black hole
$D = 7 \text{ kpc}$	Distance to the source
$\theta = 10^{\circ}$	Source inclination
$\gamma_f = 3$	Terminal jet Lorentz factor after bulk acceleration
$\sigma_f = 0.1$	Final jet magnetization after bulk acceleration
$f_{\text{sc}} = 0.1$	Particle acceleration efficiency
$\beta_{\text{eff}} = 0.1$	Adiabatic cooling efficiency
$\beta_{e,0} = 0.02085$	Initial plasma- β at the jet base; sets the pair content. Corresponds to ≈ 20 pairs per proton with $\gamma_f = 3$, $T_e \approx 100 \text{ keV}$ and $\sigma_0 \approx 2$
$f_{\text{heat}} = 1$	Shock heating parameter
$R_{\text{out}} = 10^5 R_g$	Disc outer radius
Fitted Parameter	Description
$N_j (L_{\text{Edd}})$	Power injected in the jet base
$r_0 (R_g)$	Jet base radius
$z_{\text{diss}} (R_g)$	Location where particle acceleration in the jet starts
$T_e (\text{keV})$	Electron temperature at the base of the jet
f_{pl}	Additional loss parameter to set the radio spectral index
s	Slope of the injected non-thermal particle distribution
$L_{\text{disk}} (L_{\text{Edd}})$	Disc luminosity
$r_{\text{in}} (R_g)$	Disc truncation radius
$T_{\text{bb}} (\text{K})$	Optical excess blackbody temperature
$L_{\text{bb}} (\text{ergs}^{-1})$	Optical excess blackbody luminosity

4 DISCUSSION

The main result of our modelling is that in every epoch the full broad-band SED can be fitted by the disc+jet emission, without the need to invoke contributions from an additional coronal region as in Connors et al. (2019). The reason for our better fits is our model’s new ability to treat sub- and near-relativistic temperatures for the electrons at the jet base. This demonstrates that the coronal emission in these observations can in principle be closely related to the base of the jets, provided that the emitting particle distribution remains in the mildly relativistic regime as might be expected from the inner hot accretion flow, and the optical depth is large enough to produce hard Comptonization spectra. In all of our models we find that $\tau \approx 0.1$ – 1 (depending on the values of r_0 and N_j), resulting in $\Gamma \approx 1.7$ – 2.8 . Most importantly, we find that epochs with similar X-ray power spectra can be fit with similar geometries for the emitting region.

During most epochs of the outburst, the main radiative mechanism responsible for the X-ray emission is inverse-Compton scattering of disc photons in the jet base; however, this slowly evolves during the latter two epochs of the outburst, which are harder and fainter than the first four. For October 12 we find that inverse Comptonization of cyclo-synchrotron photons produced in the jet base (SSC) accounts for ≈ 10 per cent of the coronal continuum; by October 27 the disc is not detected in the soft X-ray anymore, and as its luminosity increases SSC becomes the main channel of Comptonization. Near the peak of the outburst in mid-September, we find very strong reflection features, which dominate over the continuum; in other epochs (again, with the exception of October 27) reflection is still present, but much weaker. We find the extent to which reflection dominates over the continuum at energies $\geq 10 \text{ keV}$ does not depend on the continuum model chosen, and is observed both in the phenomenological fits using `nthcomp`, and in the broad-band fits using `bljet`. This is in

agreement with the more sophisticated reflection modelling in Reis et al. (2012) and Dong et al. (2020).

In the optical/ultraviolet bands we also find that the radiative mechanism changes during the outburst. During the HSs in October, and the HS/HIMS transitions states on August 30 and September 26, we find that the data can be represented fairly well by cyclo-synchrotron emission originating in the jet base. However, this mechanism underpredicts the UV flux of the source during the two HIMS states at the peak of the outburst. Furthermore, during these epochs the optical/UV data showed an upturn, becoming brighter at higher frequencies, which our model cannot easily reproduce. This is in agreement with Russell et al. (2014b).

Finally, our fits seem to indicate that if a disc component is detected, its innermost radius is fairly constant ($R_{\text{in}} \approx 12 - 15 R_g$). However, this is not a robust claim and should be taken with care. Russell et al. (2014b), using the `diskir` model, found that during the HIMS epochs the truncation radius varies with hardness between $R_{\text{in}} \approx 6$ and $20 R_g$. Because `diskir` self-consistently accounts for disc irradiation, which our model does not, this is a more reliable estimate of the disc truncation radius.

4.1 An evolving jet during the outburst

Fig. 8 shows the evolution of the jet dynamical properties, defined by the jet initial radius r_0 and the particle acceleration distance z_{diss} , throughout the outburst, as a function of power-spectral hue (defined as the angle between the coloured and blue-grey lines in the left panel of Fig. 1). Intriguingly, both parameters appear to be correlated with the hue (although we note that in case of z_{diss} , the apparent correlation is weak and largely driven by including the Oct 27th dataset), which is a tracer of X-ray variability properties. We stress that the hue captures the shape of the power spectrum (Heil et al. 2015a), rather than a simply ratio of power at high versus low frequency: values near 0° represent a flat power spectrum, while values closer to 180° correspond to a more peaked power spectrum. This makes the hue a useful tool to categorize spectral states, but means it is not straightforward to deduce information about the characteristic size of the emitting region.

The changes in jet base radius can be understood in terms of optical depth. As the outburst progresses from the HS to the HIMS, the power-law continuum softens. As the temperature in the jet base is only changing by a factor of ≈ 2 at most (see Table 4), one natural way to soften the Comptonization spectrum is to reduce the optical depth by increasing the size of the emitting region. Such a trend is in broad agreement with recent GRMHD simulations of thin discs. Liska et al. (2019) found that compared to the standard thick disc scenario, a thin disc can still launch and collimate a powerful jet; however, compared to a thick ADAF-type disc, a thin disc exerts a lower pressure on the jet edge, leading to a wider outflow. A change in the initial jet collimation would naturally lead to the behaviour we found in our fits, where harder, fainter states with low hue (presumably close to a higher scale-height ADAF-type disc) favour smaller jet bases, while bright intermediate states with large hue require more extended emitting regions. We also note that Dong et al. (2020), who modelled the reflection spectra of the source near the peak of the outburst, favour an extended corona over a compact lamp-post geometry, in agreement with our scenario.

Interestingly, the trend we observe of the corona/jet launching becoming wider (and more optically thin) appears to be unlike what is observed in other sources. Kara & et al. (2019) found that the reverberation lags in MAXI J1820+070 become shorter as the source transitioned towards the soft state, while the iron line remained

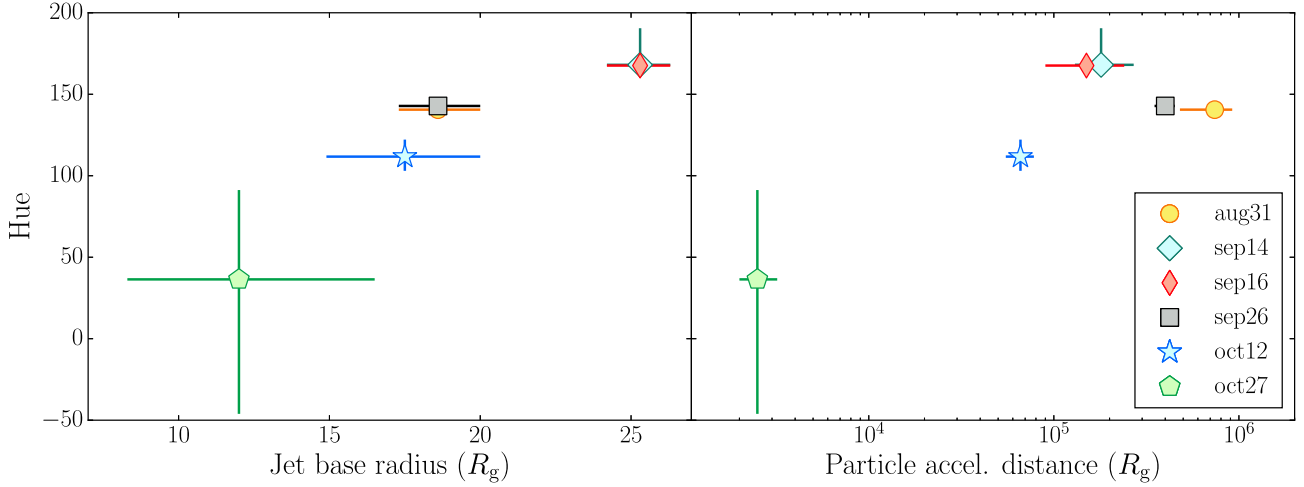


Figure 8. Evolution in the jet base radius (left panel) and location of the start of particle acceleration (right panel) with power spectral hue, from our broadband fits.

Table 4. Summary of the best-fitting parameters for each fit (individual for the October data, jointly for the other epochs), using `bljet` for the continuum.

Parameter	Aug 30	Sep 13	Sep 16	Sep 26	Oct12	Oct 27
N_j ($10^{-2} L_{\text{Edd}}$)	$4.3^{+0.3}_{-0.4}$	$4.2^{+0.3}_{-0.2}$	$3.6^{+0.3}_{-0.3}$	$3.7^{+0.3}_{-0.3}$	$7.7^{+1.1}_{-1.1}$	$6.2^{+2.4}_{-2.2}$
r_0 (R_g) ^a	$18.6^{+1.4}_{-1.3}$	$25.3^{+1.0}_{-1.1}$	$25.3^{+1.0}_{-1.1}$	$18.6^{+1.4}_{-1.3}$	$17.5^{+2.5}_{-2.6}$	$12.0^{+4.5}_{-3.7}$
z_{diss} (R_g)	$7.4^{+1.8}_{-2.6} \cdot 10^5$	$1.8^{+0.9}_{-0.5} \cdot 10^5$	$1.5^{+0.9}_{-0.6} \cdot 10^5$	$4.0^{+0.5}_{-0.5} \cdot 10^5$	$6.6^{+1.2}_{-1.1} \cdot 10^4$	$2.5^{+0.7}_{-0.5} \cdot 10^3$
T_e (keV)	126^{+4}_{-4}	70^{+4}_{-5}	52^{+10}_{-7}	114^{+7}_{-7}	101^{+5}_{-3}	126^{+18}_{-8}
f_{pl}	$15.3^{+1.5}_{-2.9}$	$5.0^{+0.7}_{-0.7}$	$3.4^{+0.9}_{-1.0}$	$19.5^{+1.0}_{-0.7}$	$18.1^{+1.3}_{-1.3}$	$7.5^{+0.5}_{-0.5}$
s	$2.15^{+0.08}_{-0.06}$	$2.55^{+0.10}_{-0.05}$	$2.55^{+0.06}_{-0.07}$	$2.10^{+0.04}_{-0.04}$	$2.45^{+0.06}_{-0.07}$	$3.0^{+0.2}_{-0.2}$
L_{disc} ($10^{-2} L_{\text{Edd}}$)	$0.64^{+0.04}_{-0.03}$	$2.27^{+0.02}_{-0.02}$	$2.44^{+0.04}_{-0.03}$	$0.74^{+0.02}_{-0.01}$	$0.15^{+0.01}_{-0.01}$	//
r_{in} (R_g) ^a	$15.6^{+0.5}_{-0.4}$	$14.0^{+0.2}_{-0.1}$	$14.0^{+0.2}_{-0.1}$	$15.6^{+0.5}_{-0.4}$	$12.8^{+1.9}_{-1.6}$	//
L_{bb} ($10^{36} \text{ ergs}^{-1}$) ^a	//	$1.0^{+0.2}_{-0.1}$	$1.0^{+0.2}_{-0.1}$	//	//	//
T_{bb} (10^4 K) ^a	//	$1.9^{+0.2}_{-0.1}$	$1.9^{+0.2}_{-0.1}$	//	//	//
relfrac	$1.05^{+0.11}_{-0.10}$	$1.6^{+0.2}_{-0.2}$	$3.8^{+0.7}_{-0.8}$	$1.4^{+0.2}_{-0.1}$	$0.6^{+0.1}_{-0.1}$	//
Gaussian norm (10^{-4})	$2.8^{+2.2}_{-1.6}$	$9.7^{+3.5}_{-2.9}$	22^{+7}_{-9}	25^{+10}_{-10}	$12.0^{+6.0}_{-5.0}$	//
Gaussian σ (keV)	$1.0^{+1.6}_{-0.7}$	$0.4^{+0.3}_{-0.2}$	$2.1^{+0.4}_{-0.4}$	$4.6^{+1.3}_{-0.9}$	$2.0^{+0.8}_{-0.6}$	//

^aTied over Aug30/Sep26 and/or Sep13/Sep16.

unchanged (and broad); they interpret this behaviour as the corona contracting during the state transition. Vincentelli et al. (2019) found that in the HS of GX 339–4 the IR power spectra evolved more slowly than the X-ray power spectra, suggesting that the size of the jet launching region remains constant as the source evolves during the HS. However, we (also) notice that these results were found (with different methods and) during different phases of the outbursts, therefore a quantitative comparison between these estimates is not straightforward.

The (weaker) correlation between the X-ray variability properties and the particle acceleration location z_{acc} is less intuitive. We also note that due the relatively poor constraints available on the break frequency on September 14 and 16 (Russell et al. 2014b), this second result should be interpreted with additional care. However, the existence of such a correlation is not entirely unexpected, as the disc and jet are coupled to each other. From a purely hydrodynamical point of view, one way to explain such a correlation is to assume that propagating fluctuations in the accretion flow, which set the shape of the power spectrum (Ingram & Done 2011; Ingram & van der Klis 2013; Rapisarda, Ingram & van der Klis 2014; Rapisarda et al.

2016), also set the details of the internal shocks powering particle acceleration (Malzac 2013; Péault et al. 2019). This could naturally lead to the observed correlation. However, such a scenario does not account for the role that magnetic fields play in determining the dynamics of accretion flows and outflows. From a magneto-hydrodynamics point of view, the location of particle acceleration may be associated with the modified fast point (MFP) of the jet, which can be understood as the location beyond which the outflow ceases to be causally connected with the black hole (e.g. Vlahakis & Königl 2004). In simplified, ideal MHD treatments, the location of the MFP can be uniquely determined when the conditions at the base of the jet are specified, and small changes e.g. in the magnetic field configuration at the jet launching point can change the MFP by orders of magnitude (e.g. Ferreira & Pelletier 1995; Polko, Meier & Markoff 2010, 2013, 2014). In this sense, a change in the physical conditions at the base of the jet and/or innermost regions of the accretion flow, like the magnetic field configuration, could drive large changes in both the location of the MFP and the properties of the turbulence driving accretion, producing the correlations between timing and spectral parameters we found. Regardless of origin, this

new empirical trend provides an interesting benchmark for more complicated numerical studies of dissipation (e.g. Chatterjee et al. 2019). A more explicit comparison is currently not possible, as both simulations and semi-analytical GRMHD have yet to fully explore their parameter spaces, due to the computational costs involved.

Finally, our fits suggest that the jet power in the source is roughly constant and on the order of $\approx 10^{-2} L_{\text{Edd}}$, and that the SEDs in every epoch can be fitted with a moderate terminal Lorentz factor of $\gamma \approx 3$. These findings are in contrast with Péault et al. (2019), whose model fitted to the same data requires a wide range in both jet powers ($\approx 10^{-3}$ to a few $10^{-1} L_{\text{Edd}}$) and Lorentz factor (1.05–17). However, we note that in both models jet speed and power are highly degenerate both with each other and with additional parameters (e.g. opening angle, pair content, efficiency of particle acceleration, as shown in fig. 6 in Péault et al. 2019 and discussed in the following section). These degeneracies mean that the best-fitting values of the jet powers should only be considered as order of magnitude estimates. Otherwise, our treatment of the jet is complementary to that of Péault et al. (2019): in their model, the mechanism responsible for particle acceleration is treated in more detail; in ours, we include inverse Compton emission from the jet base and couple it to the full compact jet. Nevertheless, both in our model and theirs, several parameters need to vary between epochs in order to correctly reproduce the data.

4.2 Pair loading in the jet

As mentioned in Section 3.4, in every fit we fixed the initial plasma- β parameter $\beta_{e,0} = 0.02085$, resulting in a roughly constant pair content of ≈ 20 pairs per proton. We chose this target ratio based on a comparison of the energetics of the jet and disc emission. We note that this is an ad-hoc assumption we made purely to reduce the model’s degeneracy; however, it has important consequences for the energetic balance of inflowing versus outflowing material.

In the framework of a leptonic model, assuming one cold proton per radiating electron leads to the highest possible jet power for a given observed jet luminosity. Such an assumption leads to reasonable energy budgets in canonical blazars; for example, Ghisellini et al. (2014) found that the jet power is on the order of, but slightly higher than, the disc luminosity. A large pair content in the jets in Flat Spectrum Radio Quasars is also disfavoured because of the effect that radiation pressure would have on the outflow dynamics (Ghisellini & Tavecchio 2010).

In our modelling we found, however, that assuming a low pair content in the jet ($\eta \leq 20$) leads to uncomfortably high energy requirements in the jet. Since the jet emission of this source is not significantly different from other XRBs, a similar conclusion would also apply to the XRB population as a whole. This discrepancy in required power is illustrated in Fig. 9, which shows the jet+disc SEDs for a constant disc luminosity $L_{\text{disc}} = 4 \times 10^{-2} L_{\text{Edd}}$, while varying the jet power and matter content. If the jet carries ≈ 20 pairs per electron, the jet power required to produce a ‘typical’ SED (in terms of radio and X-ray luminosities) is $N_j = 4 \times 10^{-2} L_{\text{Edd}}$, as shown in the top left panel; this value is comparable to the disc luminosity, as inferred for blazar jets, and comfortably sub-Eddington as expected for HS XRBs. Instead, if the jet carries one proton per electron, as shown in the bottom right panel, the required jet power to produce a typical BHXB SED is $N_j \approx 4 \cdot 10^{-1} L_{\text{Edd}}$, 20 times higher than the disc luminosity.

This degeneracy can be understood by writing the jet power for a fixed lepton number density, as a function of pair content η . By

combining equations (12), (14), and (16) we get:

$$N_j = 2\gamma_0\beta_0c\pi R_0^2 U_{e,0} \left(1 + \frac{\sigma_0 m_p + \Gamma_{\text{ad}} \sigma_0 \eta \langle \gamma \rangle m_e}{2\eta \langle \gamma \rangle m_e} + \frac{m_p}{\eta \langle \gamma \rangle m_e} \right). \quad (20)$$

The term in brackets is a monotonically decreasing function of η , implying that very similar SEDs can be produced by reducing η and increasing N_j , regardless of whether it is the protons or leptons that carry the bulk of the kinetic energy.

Jet powers in excess of the accretion power can be achieved through the Blandford–Znajek process (Blandford & Znajek 1977); however, GRMHD simulations find that the efficiency of such a process is on the order of a few at most (e.g. Tchekhovskoy, Narayan & McKinney 2011), far lower than the (conservative) factor ≈ 20 estimated above. Despite the large uncertainties arising from directly comparing GRMHD simulations to observations, to date this is the best estimate of the efficiency of the Blandford–Znajek process. These energetic considerations imply one of two things: either black hole X-ray binaries are far more efficient at launching jets than blazars are, or the composition of the plasma of the outflows in the two classes of sources is different, with XRBs carrying more pairs than AGN.

One possible explanation is that the energy densities near the black hole scale inversely with black hole mass, similarly to the peak temperature of an accretion disc, for a fixed Eddington accretion power. Such an increase could lead to an increased efficiency in pair-production processes near the black hole. This scaling can be shown by once again inverting equation (12) to solve for the energy density:

$$U_{\text{tot}} = \frac{N_j}{\pi \gamma_0 \beta_0 c r_0^2} = 1.4 \times 10^{17} \left(\frac{N_j}{1.26 \times 10^{38} \text{ erg s}^{-1}} \right) \left(\frac{1.5 \times 10^5 \text{ cm}}{r_0} \right)^2. \quad (21)$$

Taking $N_j = 10^{-2} L_{\text{Edd}}$ and $r_0 = 10 R_g$, resulting in $U_{\text{tot}} = 1.4 \times 10^{12} \text{ erg cm}^{-3}$ for a $10 M_\odot$ black hole, and $U_{\text{tot}} = 1.4 \times 10^4 \text{ erg cm}^{-3}$ for a $10^9 M_\odot$ black hole. Therefore, for a given energy budget *in dimensionless Eddington units*, the energy density near a stellar-mass black hole is typically ~ 8 orders of magnitude larger than that near a large supermassive black hole. Higher energy densities *in physical units* would naturally lead to an increase of pair-production processes near the black hole (e.g. Neronov & Aharonian 2007; Mościbrodzka et al. 2011; Broderick & Tchekhovskoy 2015). This in turn could naturally result in XRB jets being more pair-loaded than AGN jets.

A useful tool to quantify the importance of pair production is the compactness-temperature diagram (e.g. Fabian et al. 2015, and references therein). The dimensionless compactness l is defined as

$$l = \frac{L}{R} \frac{\sigma_1}{m_e c^3}, \quad (22)$$

where L is the luminosity and R is the radius of the emitting region. In the $l - \theta$ (where θ is the electron temperature in units of 511 keV), regions of high temperature and/or high compactness are forbidden due to pair production becoming a runaway process. The exact location of the pair production limit depends on the geometry of the emitting region (Svensson 1984, 1996; Stern et al. 1995). For our fits, we computed L by integrating the total model luminosity in the 0.1–300 keV range, excluding the non-thermal synchrotron emission which originates away from the base of the jet/corona. The resulting $l - \theta$ plot is shown in Fig. 10. All of our fits lie close and to the left of the pair-production region, in excellent agreement with the results

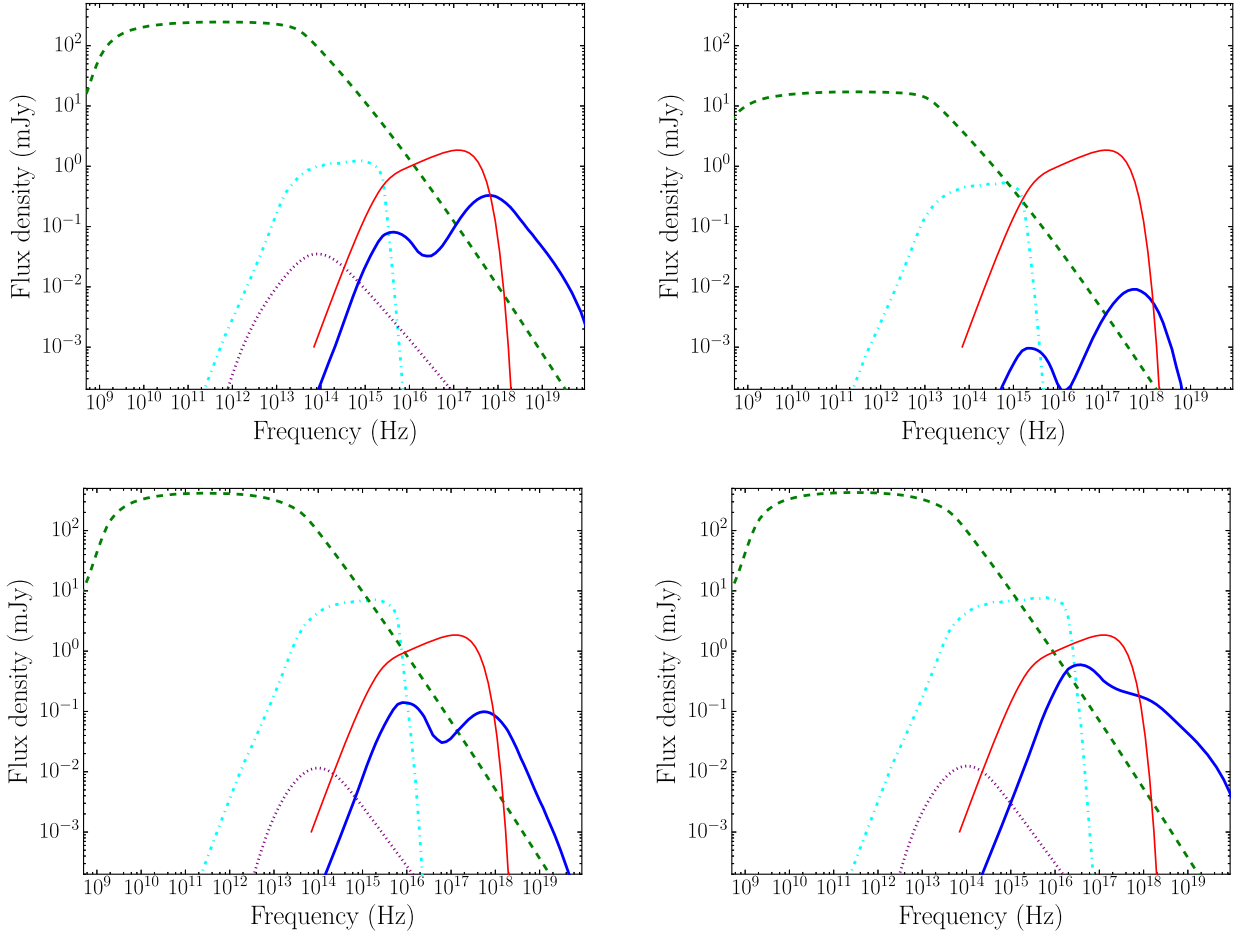


Figure 9. Example jet SEDs for different jet powers and pair contents, demonstrating the model degeneracy between jet power, pair content and jet base geometry. For the top left SED, $N_j = 4 \times 10^{-2} L_{\text{Edd}}$, $r_0 = 10 R_g$, and 20 pairs per proton are injected in the jet. The top right SED uses the same jet power and jet base radius, but one pair per proton. The bottom left SED uses one proton per electron, $N_j = 4 \times 10^{-1} L_{\text{Edd}}$, and $r_0 = 10 R_g$. Finally, the bottom right SED uses $N_j = 4 \times 10^{-1} L_{\text{Edd}}$, and $r_0 = 4 R_g$. The dashed green and dotted purple lines represent non-thermal synchrotron or inverse-Compton emission; the dotted cyan and continuous thick blue lines represent thermal synchrotron or inverse-Compton emission from the jet base; the continuous thin red line the emission from the disc.

of Fabian et al. (2015), suggesting that pair production near the black hole could be regulating the matter content injected in the jet. Older versions of our jet model instead tended to return fits in the runaway region, as discussed by Malzac, Belmont & Fabian (2009), further highlighting the importance of treating quasi-relativistic particles in the corona correctly.

4.3 Implications of a jet-dominated model

Our proposed scenario for the evolution of the broad-band SED of jet-dominated sources makes three predictions:

(1) In the harder states, if the cyclo/synchrotron emission of the jet base is brighter than both the companion star and accretion flow, the optical and X-ray emission should be strongly correlated, and both should vary on similar time-scales, as both originate in the same region near the black hole. As the source softens and disc irradiation dominates over the cyclo-synchrotron emission of the jet base/corona, the optical should become less variable, and uncorrelated with the hard X-ray power law.

(2) In the intermediate states, the seed photons for inverse Compton scattering originate in the disc, while in harder states, they originate in the jet base itself. In the first regime, the jet base/corona should react

to changes both in the accretion luminosity (which affects the seed photons) and accretion rate, which should impact the mass loading of the jet. In the second case, only mass loading is important, as seed photons are produced locally. These two scenarios would lead to different pivoting behaviour of the X-ray power law, leading in turn to a change in the shape of the reverberation lags (Mastroserio, Ingram & van der Klis 2018).

(3) In states closer to quiescence, as the jet power drops, the optical depth should also drop. This can be shown by estimating the dependence of the optical depth in the jet base as a function of jet power and initial jet radius:

$$\tau = n_e r_0 \sigma_t \quad (23)$$

where n_e is the electron number density in the jet base, r_0 its radius, and σ_t is the Thomson cross-section. c the jet power starting from equation (12):

$$\begin{aligned} N_j &= 2\pi r_0^2 \gamma_0 \beta_0 c (U_e + U_b + U_p) \\ &= 2\pi r_0^2 \gamma_0 \beta_0 n_e m_e \langle \gamma \rangle c^3 \left(1 + \frac{1}{\beta_{e,0}} + \frac{1}{\kappa_{p,0}} \right) \end{aligned} \quad (24)$$

where $\langle \gamma \rangle \approx a$ few is the average Lorentz factor of the electrons,

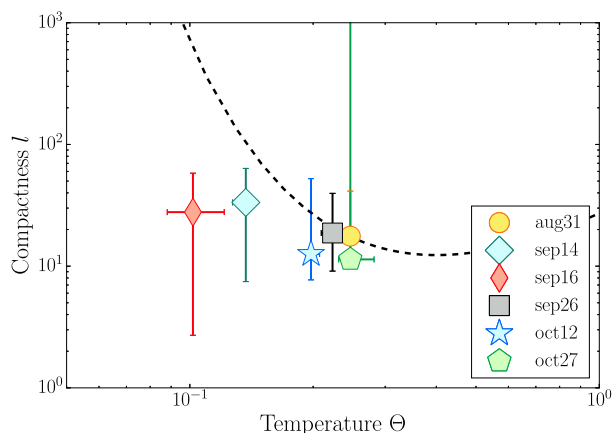


Figure 10. Compactness versus dimensionless temperature for the corona, derived from our multiwavelength fits. The dashed line indicates the pair runaway threshold estimated in Svensson (1984); each epoch lies close to it, but in the allowed region, implying that some amount of pair production is likely to be occurring in the jet base/corona.

$\beta_{e,0} = U_e/U_b$, and $\kappa_{p,0} = U_p/U_e$. By making the simplifying assumption that $\langle \gamma \rangle \approx 1$, $\beta_{e,0} = U_e/U_b$, and $\kappa_{p,0} = U_p/U_e$ do not change by factors of more than a few as the jet power varies, we find:

$$n_e = \frac{N_j}{2\pi r_0^2 \gamma_0 \beta_0 m_e \langle \gamma \rangle c^3 \left(1 + \frac{1}{\beta_{e,0}} + \frac{1}{\kappa_{p,0}}\right)} \propto \frac{N_j}{r_0^2} \quad (25)$$

and therefore the optical depth scales as

$$\tau \propto n_e r_0 \propto \frac{N_j}{r_0}. \quad (26)$$

The size of the X-ray emitting region in the jet base is highly unlikely to ever become smaller than $\approx 1 R_g$, so as the jet power drops at some point the outflow will become increasingly more optically thin, leading to softer Comptonization spectra. While this can be offset by an increase in temperature, the spectrum of an optically thin, hot corona is not a power law, but carries some intrinsic curvature. If a power-law spectrum is to be maintained down to quiescence, therefore, an additional component should become the dominant source of hard X-ray photons – for example, non-thermal synchrotron emission from downstream in the jets. This will also have an impact on the variability of the X-ray continuum, as both the radiative mechanism and emitting region will have changed dramatically from a typical bright hard or hard-intermediate state. Finally, we note that similar conclusions could be drawn in a scenario in which the X-ray power law originates in the RIAF, by replacing N_j in equation (24) with the accretion rate \dot{M} .

To summarize, our jet model predicts that both the characteristics of the X-ray variability should show distinct patterns at different stages of the outburst, tracing the radiative mechanism responsible for the emission in different wavelengths. Furthermore, correlated X-ray and optical variability could be used to discriminate between the latter originating in the accretion flow, companion star or jet base.

5 CONCLUSION

In this paper we have presented an extension of the `bljet` multiwavelength, steady-state jet+disc model, and modelled six quasi-simultaneous, multiwavelength SEDs of the BHBX MAXI J1836–194. `bljet` was initially designed to model the SEDs of blazar jets (Lucchini et al. 2019a), starting from the `agnjet` model

(Markoff et al. 2001, 2005) which instead was limited to low-luminosity AGNs and BHBXBs. Here, we have shown that the `bljet` formalism can be used to fit the broad-band SED of hard and HIMS states in X-ray binaries. Furthermore, we combined spectral and timing information in order to provide a more self-consistent picture of the system throughout its outburst.

We found that in a regime in which the electrons are mildly or non-relativistic at the base of the jet ($T_e \approx 100$ keV), this region of the system can produce a power-law coronal continuum by inverse Compton upscattering either disc-photons (during bright hard and intermediate states), or local cyclo-synchrotron photons (during more faint HSs). We find that the corona/jet launching region is relatively compact, with radii of $\approx 8 - 25 R_g$; we do not require any additional spectral component to account for the power-law X-ray continuum emission. In our model, if the outflow carries a mild pair content (≈ 20 per proton) we find that the jet power is on the order of, but slightly higher than, the accretion luminosity. A proton-heavy jet carrying no pairs and one proton per electron can in principle reproduce the data, but in this case the jet power required is about 20 times higher than the accretion luminosity; these energetic constraints therefore favour the pair-loaded scenario.

Our fits indicate that the geometrical properties of the jet are evolving throughout the outburst and are correlated with the shape of the X-ray power spectrum. We find that during softer HIMS states the jet base widens, while the location of particle acceleration in the jet moves downstream, away from the black hole. Vice versa, harder states require more compact jet launching regions and particle acceleration occurring closer to the black hole. We propose that this is caused by a change in the physical conditions at the base of the jet, which cause a variation in both the X-ray variability and jet structure. Such a change in the jet launching conditions could arise, for example, in differences in the collimation environment near the black hole. Recent GRMHD simulations (e.g. Chatterjee et al. 2019; Liska et al. 2019) are in broad agreement with this picture.

Our work shows that combining physically motivated models, advanced fitting techniques and different observables is a powerful tool to better understand the poorly understood physics of the disc-jet coupling in accreting black holes. In particular, the correlations between timing and spectral properties we have outlined require a time-dependent model to be investigated in depth. This will be the subject of future works.

ACKNOWLEDGEMENTS

We thank the anonymous referee for useful comments that greatly improved the quality of the manuscript. ML and SM are thankful for support from an NWO (Netherlands Organisation for Scientific Research) VICI award, grant number 639.043.513. FV acknowledges support from STFC under grant ST/R000638/1. TDR acknowledges support from an NWO Veni Fellowship, grant number 639.041.646. This research has made use of ISIS functions (ISIScripts) provided by ECAP/Remeis observatory and MIT (<http://www.sternwarte.uni-erlangen.de/isis/>).

DATA AVAILABILITY

All data in this paper are publicly available. The radio data were published and tabulated in Russell et al. (2014a,b, 2015). The Infrared, optical, and UV data were published in Russell et al. (2013). The X-ray data are publicly available from HEASARC (<https://heasarc.gsfc.nasa.gov/>). A reproduction package is available at DOI: 10.5281/zenodo.4384809.

REFERENCES

- Arévalo P., Uttley P., 2006, *MNRAS*, 367, 801
- Barret D., Olive J. F., Boirin L., Done C., Skinner G. K., Grindlay J. E., 2000, *ApJ*, 533, 329
- Belloni T., Homan J., Casella P., van der Klis M., Nespoli E., Lewin W. H. G., Miller J. M., Méndez M., 2005, *A&A*, 440, 207
- Beloborodov A. M., 1999, *ApJ*, 510, L123
- Blandford R. D., Königl A., 1979, *ApJ*, 232, 34
- Blandford R. D., Znajek R. L., 1977, *MNRAS*, 179, 433
- Blumenthal G. R., Gould R. J., 1970, *Rev. Mod. Phys.*, 42, 237
- Broderick A. E., Tchekhovskoy A., 2015, *ApJ*, 809, 97
- Cassatella P., Uttley P., Maccarone T. J., 2012, *MNRAS*, 427, 2985
- Ceccobello C., Cavecchi Y., Heemskerk M. H. M., Markoff S., Polko P., Meier D., 2018, *MNRAS*, 473, 4417
- Chatterjee K., Liska M., Tchekhovskoy A., Markoff S. B., 2019, *MNRAS*, 490, 2200
- Chen W., Shrader C. R., Livio M., 1997, *ApJ*, 491, 312
- Connors R. M. T. et al., 2019, *MNRAS*, 485, 3696
- Corbel S., Fender R. P., Tzioumis A. K., Nowak M., McIntyre V., Durouchoux P., Sood R., 2000, *A&A*, 359, 251
- Corbel S., Nowak M. A., Fender R. P., Tzioumis A. K., Markoff S., 2003, *A&A*, 400, 1007
- Corbel S., Coriat M., Brocksopp C., Tzioumis A. K., Fender R. P., Tomsick J. A., Buxton M. M., Bailyn C. D., 2013, *MNRAS*, 428, 2500
- Crumley P., Ceccobello C., Connors R. M. T., Cavecchi Y., 2017, *A&A*, 601, A87
- Dauser T., Garcia J., Wilms J., Böck M., Brenneman L. W., Falanga M., Fukumura K., Reynolds C. S., 2013, *MNRAS*, 430, 1694
- Dong Y., García J. A., Liu Z., Zhao X., Zheng X., Gou L., 2020, *MNRAS*, 493, 2178
- Fabian A. C., Lohfink A., Kara E., Parker M. L., Vasudevan R., Reynolds C. S., 2015, *MNRAS*, 451, 4375
- Falcke H., Biermann P. L., 1995, *A&A*, 293, 665
- Falcke H., Körding E., Markoff S., 2004, *A&A*, 414, 895
- Fender R. P., Belloni T. M., Gallo E., 2004, *MNRAS*, 355, 1105
- Ferreira J., Pelletier G., 1995, *A&A*, 295, 807
- Ferrigno C., Bozzo E., Del Santo M., Capitanio F., 2012, *A&A*, 537, L7
- Foreman-Mackey D., Hogg D. W., Lang D., Goodman J., 2013, *PASP*, 125, 306
- Gallo E., Fender R. P., Pooley G. G., 2003, *MNRAS*, 344, 60
- Gallo E. et al., 2014, *MNRAS*, 445, 290
- Gardenier D. W., Uttley P., 2018, *MNRAS*, 481, 3761
- Ghisellini G., Tavecchio F., 2010, *MNRAS*, 409, L79
- Ghisellini G., Haardt F., Svensson R., 1998, *MNRAS*, 297, 348
- Ghisellini G., Tavecchio F., Maraschi L., Celotti A., Sbarrato T., 2014, *Nature*, 515, 376
- Grebenev S. et al., 1993, *A&AS*, 97, 281
- Haardt F., Maraschi L., 1993, *ApJ*, 413, 507
- Haardt F., Maraschi L., Ghisellini G., 1994, *ApJ*, 432, L95
- Hada K. et al., 2016, *ApJ*, 817, 131
- Hannikainen D. C., Hunstead R. W., Campbell-Wilson D., Sood R. K., 1998, *A&A*, 337, 460
- Heil L. M., Uttley P., Klein-Wolt M., 2015a, *MNRAS*, 448, 3339
- Heil L. M., Uttley P., Klein-Wolt M., 2015b, *MNRAS*, 448, 3348
- Heinz S., Sunyaev R. A., 2003, *MNRAS*, 343, L59
- Homan J., Belloni T., 2005, *Ap&SS*, 300, 107
- Homan J., Wijnands R., van der Klis M., Belloni T., van Paradijs J., Klein-Wolt M., Fender R., Méndez M., 2001, *ApJS*, 132, 377
- Houck J. C., Denicola L. A., 2000, in Manset N., Veillet C., Crabtree D., eds, ASP Conf. Ser. Vol. 216, *Astronomical Data Analysis Software and Systems IX*. Astron. Soc. Pac., San Francisco, p. 591
- Huppenkothen D. et al., 2019, *ApJ*, 881, 39
- Ingram A., Done C., 2011, *MNRAS*, 415, 2323
- Ingram A., van der Klis M., 2013, *MNRAS*, 434, 1476
- Joinet A., Jourdain E., Malzac J., Roques J. P., Corbel S., Rodriguez J., Kalemci E., 2007, *ApJ*, 657, 400
- Jokipii J. R., 1987, *ApJ*, 313, 842
- Kara E. et al., 2019, *Nature*, 565, 198
- Komisarov S. S., Vlahakis N., Königl A., Barkov M. V., 2009, *MNRAS*, 394, 1182
- Kotov O., Churazov E., Gilfanov M., 2001, *MNRAS*, 327, 799
- Kubota A., Tanaka Y., Makishima K., Ueda Y., Dotani T., Inoue H., Yamaoka K., 1998, *PASJ*, 50, 667
- Lind K. R., Blandford R. D., 1985, *ApJ*, 295, 358
- Liska M., Tchekhovskoy A., Ingram A., van der Klis M., 2019, *MNRAS*, 487, 550
- Lucchini M., Markoff S., Crumley P., Krauß F., Connors R. M. T., 2019a, *MNRAS*, 482, 4798 (Paper I)
- Lucchini M., Krauß F., Markoff S., 2019b, *MNRAS*, 489, 1633
- Lyubarskii Y. E., 1997, *MNRAS*, 292, 679
- McKinney J. C., 2006, *MNRAS*, 368, 1561
- Magdziarz P., Zdziarski A. A., 1995, *MNRAS*, 273, 837
- Maitra D., Markoff S., Brocksopp C., Noble M., Nowak M., Wilms J., 2009, *MNRAS*, 398, 1638
- Makishima K., Maejima Y., Mitsuda K., Bradt H. V., Remillard R. A., Tuohy I. R., Hoshi R., Nakagawa M., 1986, *ApJ*, 308, 635
- Malzac J., 2013, *MNRAS*, 429, L20
- Malzac J., 2014, *MNRAS*, 443, 299
- Malzac J., Belmont R., Fabian A. C., 2009, *MNRAS*, 400, 1512
- Markoff S., Nowak M. A., 2004, *ApJ*, 609, 972
- Markoff S., Falcke H., Fender R., 2001, *A&A*, 372, L25
- Markoff S., Nowak M. A., Wilms J., 2005, *ApJ*, 635, 1203
- Martocchia A., Matt G., 1996, *MNRAS*, 282, L53
- Mastroserio G., Ingram A., van der Klis M., 2018, *MNRAS*, 475, 4027
- Matt G., Perola G. C., Piro L., 1991, *A&A*, 247, 25
- Matt G., Perola G. C., Piro L., Stella L., 1992, *A&A*, 257, 63
- Merloni A., Heinz S., di Matteo T., 2003, *MNRAS*, 345, 1057
- Mertens F., Lobanov A. P., Walker R. C., Hardee P. E., 2016, *A&A*, 595, A54
- Miller J. M., Ballantyne D. R., Fabian A. C., Lewin W. H. G., 2002, *MNRAS*, 335, 865
- Miller-Jones J. C. A., Sivakoff G. R., Rupen M., Altamirano D., 2011, *Astron. Telegram*, 3628, 1
- Mościbrodzka M., Gammie C. F., Dolence J. C., Shiokawa H., 2011, *ApJ*, 735, 9
- Narayan R., Yi I., 1994, *ApJ*, 428, L13
- Negoro H. et al., 2011, *Astron. Telegram*, 3611, 1
- Neronov A., Aharonian F. A., 2007, *ApJ*, 671, 85
- Nowak M. A. et al., 2011, *ApJ*, 728, 13
- Péault M. et al., 2019, *MNRAS*, 482, 2447
- Plotkin R. M., Markoff S., Kelly B. C., Körding E., Anderson S. F., 2012, *MNRAS*, 419, 267
- Polko P., Meier D. L., Markoff S., 2010, *ApJ*, 723, 1343
- Polko P., Meier D. L., Markoff S., 2013, *MNRAS*, 428, 587
- Polko P., Meier D. L., Markoff S., 2014, *MNRAS*, 438, 959
- Potter W. J., Cotter G., 2012, *MNRAS*, 423, 756
- Pottschmidt K., Wilms J., Nowak M. A., Heindl W. A., Smith D. M., Staubert R., 2000, *A&A*, 357, L17
- Pottschmidt K. et al., 2003, *A&A*, 407, 1039
- Rapisarda S., Ingram A., van der Klis M., 2014, *MNRAS*, 440, 2882
- Rapisarda S., Ingram A., Kalamkar M., van der Klis M., 2016, *MNRAS*, 462, 4078
- Reis R. C., Miller J. M., Reynolds M. T., Fabian A. C., Walton D. J., 2012, *ApJ*, 751, 34
- Remillard R. A., McClintock J. E., 2006, *A&A*, 44, 49
- Russell D. M., Maitra D., Dunn R. J. H., Markoff S., 2010, *MNRAS*, 405, 1759
- Russell D. M. et al., 2011, *Astron. Telegram*, 3689, 1
- Russell D. M. et al., 2013, *ApJ*, 768, L35
- Russell T. D., Soria R., Motch C., Pakull M. W., Torres M. A. P., Curran P. A., Jonker P. G., Miller-Jones J. C. A., 2014a, *MNRAS*, 439, 1381
- Russell T. D., Soria R., Miller-Jones J. C. A., Curran P. A., Markoff S., Russell D. M., Sivakoff G. R., 2014b, *MNRAS*, 439, 1390
- Russell T. D. et al., 2015, *MNRAS*, 450, 1745
- Shakura N. I., Sunyaev R. A., 1973, *A&A*, 500, 33

- Stern B. E., Poutanen J., Svensson R., Sikora M., Begelman M. C., 1995, *ApJ*, 449, L13
- Strohmer T. E., Smith E. A., 2011, *Astron. Telegram*, 3618, 1
- Svensson R., 1984, *MNRAS*, 209, 175
- Svensson R., 1996, *A&AS*, 120, 475
- Tchekhovskoy A., McKinney J. C., Narayan R., 2009, *ApJ*, 699, 1789
- Tchekhovskoy A., Narayan R., McKinney J. C., 2011, *MNRAS*, 418, L79
- Tetarenko B. E., Sivakoff G. R., Heinke C. O., Gladstone J. C., 2016, *ApJS*, 222, 15
- Uttley P., Cackett E. M., Fabian A. C., Kara E., Wilkins D. R., 2014, *A&AR*, 22, 72
- Vincentelli F. M. et al., 2019, *ApJ*, 887, L19
- Vlahakis N., Königl A., 2004, *ApJ*, 605, 636
- Yuan F., Zdziarski A. A., Xue Y., Wu X.-B., 2007, *ApJ*, 659, 541
- Zdziarski A. A., Johnson W. N., Poutanen J., Magdziarz P., Gierlinski M., 1997, in Winkler C., Courvoisier T. J. L., Durouchoux P., eds, *ESA Special Publication*, Vol. 382, *The Transparent Universe*. ESA, Noordwijk, p. 373
- Zycki P. T., Done C., Smith D. A., 1999, *MNRAS*, 309, 561

APPENDIX: RADIATION CALCULATIONS

We include a truncated Shakura–Sunyaev disc (Shakura & Sunyaev 1973) using the simplified boundary condition (Kubota et al. 1998) to calculate the normalization. The disc is parametrized by an inner truncation radius r_{in} and luminosity L_{d} ; the temperature therefore is

$$T_{\text{in}} = \left(\frac{L_{\text{d}}}{2\sigma_{\text{sb}}r_{\text{in}}^2} \right)^{1/4}, \quad (\text{A1})$$

where σ_{sb} is the Stefan–Boltzmann constant. We take the disc height at each radius r to be

$$H = \max(0.1, L_{\text{d}}/L_{\text{Edd}}) \times r, \quad (\text{A2})$$

where L_{Edd} is the black hole’s Eddington luminosity. This results in a thin disc with $H/R = 0.1$ at sub-Eddington luminosities and a puffed up thicker disc at higher luminosity. We note that for the scope of this work, the disc is always in the thin regime. The only difference in the treatment of the disc is that previous versions of `agnjet` and `bljet` considered a fixed height of the disc, while here we fix H/R .

Cyclo-synchrotron and inverse-Compton emission in `bljet` are computed following Blumenthal & Gould (1970), with some small changes designed to account for emission in the cyclotron regime. In order to account for the cyclo-synchrotron emission of both mildly and highly relativistic electrons, we use the phenomenological emissivity formula of Ghisellini et al. (1998):

$$j'_{\text{c}}(v', \gamma) = \frac{4\rho^2}{3} \frac{\sigma_{\text{T}}cU_{\text{b}}}{v_1} \frac{2}{1+3\rho^2} \exp \left[\frac{2(1-v'/v'_1)}{1+3\rho^2} \right] \quad (\text{A3})$$

for electrons with $\gamma \leq 2$, and the full synchrotron emissivity

$$j'_{\text{s}}(v', \gamma) = \frac{\sqrt{3}e^3 B \sin \theta}{m_e c^2} \frac{v'}{v'_s} \int_{v'/v'_s}^{\infty} K_{5/3}(y) dy \quad (\text{A4})$$

otherwise. In both equations, $q = p/m_e c^2$ is the electron momentum in units of $m_e c$. Previous versions of `agnjet` and `bljet` only included the emissivity in equation (A4). The emissivity integrated over the particle distribution is

$$j'(v') = \int_{\gamma_{\text{min}}}^{\gamma_{\text{max}}} N(\gamma) j'_{\text{c,s}}(v', \gamma) d\gamma, \quad (\text{A5})$$

where $j'_{\text{c,s}}(v', \gamma)$ is the pitch-angle average of equations (A3) and

(A4). The absorption coefficient is

$$\alpha'(v') = \int_{\gamma_{\text{min}}}^{\gamma_{\text{max}}} \frac{N(\gamma)}{\gamma^2} \frac{d}{d\gamma} (\gamma^2 j'_{\text{c,s}}(v', \gamma)) d\gamma \quad (\text{A6})$$

as in Blumenthal & Gould (1970). The comoving cyclo-synchrotron luminosity per unit frequency then is

$$L'_s(v') = \frac{16\pi H(z)R(z)v'^2}{c^2} \frac{j'(v')}{\alpha'(v')} (1 - e^{-\tau'(v')}), \quad (\text{A7})$$

where $R(z)$ the radius of the jet at a distance z from the black hole, $H(z)$ is the height of each cylindrical slice of the jet, and

$$\tau'_s(v') = \frac{\pi}{2\alpha'(v')R(z)\delta \sin(\theta)} \quad (\text{A8})$$

is the cyclo-synchrotron optical depth, which includes skin depth and viewing angle effects. Finally, the luminosity is boosted in the observer frame with the standard transformations (for a continual jet flow; Lind & Blandford 1985):

$$L_s(v) = \delta^2 L'_{\text{c,s}}(v'); \quad v = \delta v'. \quad (\text{A9})$$

Similarly to cyclo-synchrotron emission, we follow the treatment of Blumenthal & Gould (1970) to calculate the inverse-Compton emission along the jet, which includes the full Klein–Nishina regime. This is unchanged from previous versions of the model, and is only reported here for completeness. Given an electron with Lorentz factor γ and a photon differential density (number of photons per unit volume and per unit of initial photon energy) $N(\epsilon_0)$, the scattered photon spectrum in the comoving frame of the jet is

$$\frac{dN}{dt d\epsilon_1} = \frac{2\pi r_e^2 m_e c^3 N(\epsilon_0) d\epsilon_0}{\gamma \epsilon_0} F'_{\text{IC}}(q, \Gamma_e), \quad (\text{A10})$$

where ϵ_1 is the final photon energy, ϵ_0 is the initial photon energy, r_e is the classical radius of the electron, γ is the electron’s Lorentz factor, and

$$F(q, \Gamma_e) = 2q \ln q + (1 + 2q)(1 - q) + \frac{(\Gamma_e q)^2 (1 - q)}{2(1 + \Gamma_e q)}, \quad (\text{A11})$$

where $\Gamma_e = 4\epsilon_0 \gamma / m_e c^2$ regulates whether the scattering happens in the Thomson or Klein–Nishina regime ($\Gamma_e \ll 1$ and $\Gamma_e \gg 1$, respectively), $q = E_1 / (\Gamma_e (1 - E_1))$ accounts for the photon energy gain from ϵ_0 to ϵ_1 , and $E_1 = \epsilon_1 / \gamma m_e c^2$ is the final photon energy in units of the initial electron energy. The total spectrum is found by integrating over the particle and seed photon distributions:

$$\frac{dN_{\text{tot}}}{dt d\epsilon_1} = \int \int N(\gamma) \frac{dN}{dt d\epsilon_1} d\gamma d\epsilon_0, \quad (\text{A12})$$

which has units of total number of scatterings, per unit of outgoing photon energy, volume, and time. Finally, the comoving luminosity is obtained by multiplying equation (A12) by the emitting volume times $h\epsilon_1$, where h is the Planck constant, in order to obtain a specific luminosity ($\text{erg s}^{-1} \text{Hz}^{-1}$). The observed spectrum is then calculated with the same Doppler boost as the cyclo-synchrotron one (equation A9).

For inverse-Compton scattering, both disc and cyclo-synchrotron photons produced within the jet are considered. The seed cyclo-synchrotron photon distribution is calculated by converting equation (A7) into units of number of photons per unit volume and energy:

$$n(\epsilon_0) = \frac{L'_s(v')}{c h^2 v' \pi R(z)^2}. \quad (\text{A13})$$

The disc seed photon distribution is calculated by integrating the disc

temperature profile along the disc radial profile:

$$n_{\text{disc}}(\nu, \delta) = \int_{\alpha_{\text{min}}}^{\alpha_{\text{max}}} \frac{4\pi \nu^2(\delta)}{hc^3(e^{h\nu(\delta)/k_b T(\alpha, \delta)} - 1)} d\alpha, \quad (\text{A14})$$

where the extremes of the integral are $\alpha_{\text{min}} = \arctan(r_{\text{in}}/z)$, $\alpha_{\text{max}} = \arctan(R_{\text{out}}/(z - HR_{\text{out}}/2))$ if $z < HR_{\text{out}}/2$ and $\alpha_{\text{max}} = \pi/2 \arctan(R_{\text{out}}/(z - HR_{\text{out}}/2))$ otherwise. In this way, α includes the entire disc viewing angle, from innermost to outermost regions. r_{in} and R_{out} are the inner and outer disc radii, H is the disc aspect ratio defined in equation (A2), z is the distance of the jet segment from the black hole, $\nu(\delta)$ is the seed photon frequency in the co-moving frame of the jet, and $T(\alpha, \delta)$ is the observed disc temperature in the comoving frame of the jet, assuming a standard Shakura–Sunyaev disc temperature profile.

To calculate multiple scatterings, the IC/SSC spectrum is then passed as the input spectrum back in equation (A12)

$$\frac{dN_{\text{tot}}}{dt d\epsilon_1} = \sum_i \int \int N(\gamma) \frac{N_i}{dt d\epsilon_1} d\gamma d\epsilon_1 \quad (\text{A15})$$

to calculate the total spectrum which is made by the sum of i -scatterings. We fix the maximum number of iterations/scattering to 15 to maintain a short run time per fit and, at the same time, confidently include all the scattering orders that could be relevant even in the regime of large optical depths.

This paper has been typeset from a $\text{\TeX}/\text{\LaTeX}$ file prepared by the author.