

Millihertz X-ray variability during the 2019 outburst of black hole candidate *Swift* J1357.2 – 0933

Aru Beri,^{1,2*} Vishal Gaur,¹ Phil Charles,² David R. A. Williams^{1,3}, Jahanvi^{1,4}, John A. Paice^{1,2,3}, Poshak Gandhi^{1,2}, Diego Altamirano,² Rob Fender,^{5,6} David A. Green^{1,7} and David Titterton⁷

¹Indian Institute of Science Education and Research (IISER) Mohali, Punjab 140306, India

²Physics & Astronomy, University of Southampton, Southampton, Hampshire SO17 1BJ, UK

³Jodrell Bank Centre for Astrophysics, School of Physics and Astronomy, The University of Manchester, Manchester M13 9PL, UK

⁴GAPHE, University of Liege, Liege 4000, Belgium

⁵Astrophysics, Department of Physics, University of Oxford, Keble Road, Oxford OX1 3RH, UK

⁶Department of Astronomy, University of Cape Town, Private Bag X3, Rondebosch 7701, South Africa

⁷Astrophysics Group, Cavendish Laboratory, 19 J.J. Thomson Avenue, Cambridge CB3 0HE, UK

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ABSTRACT

Swift J1357.2 – 0933 is a black-hole candidate X-ray transient, which underwent its third outburst in 2019, during which several multiwavelength observations were carried out. Here, we report results from the *Neil Gehrels Swift* and *NICER* observatories and radio data from *AMI*. For the first time, millihertz quasi-periodic X-ray oscillations with frequencies varying between ~ 1 and 5 mHz were found in *NICER* observations and a similar feature was also detected in one *Swift*–XRT data set. Our spectral analysis indicate that the maximum value of the measured X-ray flux is much lower compared to the peak values observed during the 2011 and 2017 outbursts. This value is ~ 100 times lower than found with *MAXI* on MJD 58558 much (~ 68 d) earlier in the outburst, suggesting that the *Swift* and *NICER* fluxes belong to the declining phase of the 2019 outburst. An additional soft component was detected in the XRT observation with the highest flux level, but at a relatively low $L_X \sim 3 \times 10^{34}$ ($d/6$ kpc)² erg s⁻¹, and which we fitted with a disc component at a temperature of ~ 0.17 keV. The optical/UV magnitudes obtained from *Swift*–UVOT showed a correlation with X-ray observations, indicating X-ray reprocessing to be the plausible origin of the optical and UV emission. However, the source was not significantly detected in the radio band. There are currently a number of models that could explain this millihertz-frequency X-ray variability; not least of which involves an X-ray component to the curious dips that, so far, have only been observed in the optical.

Key words: accretion, accretion discs – black hole physics – X-rays: binaries – X-rays: individual: *Swift* J1357.2 – 0933.

1 INTRODUCTION

Low-mass X-ray binaries (LMXB) are composed of a black hole (BH) or neutron star (NS) accreting from a companion star of mass less than a solar mass (M_\odot ; see e.g. Lewin & van der Klis 2006). These include a significant number of transient systems that show an abrupt increase in X-ray luminosity of several orders of magnitude over a few days followed by a decay on a time-scale of a few weeks, a month, or several months (e.g. Frank, King & Lasota 1987).

Swift J1357.2 – 0933 (hereafter J1357) is a transient LMXB discovered by *Swift* Burst Alert Telescope (BAT) during its 2011 outburst (Krimm et al. 2011). It is believed to host a BH of mass $\geq 9 M_\odot$ (Corral-Santana et al. 2016) in a close binary orbit with an orbital period (P_{orb}) of 2.8 h (Corral-Santana et al. 2013), although this has not yet been confirmed. Soon after its discovery, Rau, Greiner & Filgas (2011) reported a detection of the optical counterpart (at $r' = 16.30$) using GROND images, while SDSS archive images revealed the presence of a pre-outburst counterpart

with $r' = 21.96$. Subsequent time-resolved photometry of J1357 in quiescence (Shahbaz et al. 2013; Russell et al. 2018) showed substantial variability, but no periodicity that might be associated with eclipses, dips, or ellipsoidal modulation due to the donor. Consequently, even in quiescence, the donor is not the dominant emitter at optical/IR wavelengths, making it difficult to constrain J1357's distance.

Time-resolved optical spectroscopy (Torres et al. 2015) revealed the presence of prominent broad double-peaked H α emission with no hint of any late-type spectral features, such as TiO bands. Russell et al. (2018) found substantial variability in the 6-yr optical/IR (OIR) light curve of J1357 during quiescence, indicating a substantial and continuing disc-emitting component. This study suggested that J1357 lies at a greater distance, as was also discussed by Shahbaz et al. (2013) who used the outburst amplitude– P_{orb} relation, extending the possible distance out to ≥ 6.3 kpc. More recently, Charles et al. (2019) argued similarly, thus increasing earlier luminosity estimates by $\sim \times 40$, and making J1357 a member of the luminous, Galactic X-ray binaries instead of ‘very faint X-ray transients’ (or VFXTs).

* E-mail: aru.beri8@gmail.com

J1357 has proven to be a highly enigmatic source and several observed properties of J1357 cannot yet be explained by ‘standard’ LMXB models. Based on the presence of the optical dipping whose period evolves during the outburst, Corral-Santana et al. (2013) proposed that there exists a warped disc, including a thick inner torus, which is seen at a high inclination $i > 70^\circ$. However, such a geometry has difficulties explaining the lack of eclipses by the donor or any X-ray dipping, as well as the lack of X-ray reflection features (see Paice et al. 2019; Beri et al. 2019a for details). An updated multicomponent model of J1357 to account for these properties has been proposed by Paice et al. (2019), which requires a truncated accretion disc, with an inner disc radius much greater than the radius of the innermost stable circular orbit (R_{ISCO}). They suggest that the presence of an extended X-ray corona in between the BH and accretion disc could explain the lack of X-ray reflection features and absence of X-ray dips. They also propose a jet region (synchrotron emission) near the BH observed at a high inclination and sporadically occluded by the accretion disc’s vertical extensions, causing significant dips in the red band optical light curves. However, these authors also noted that this model could not explain all the source properties, such as why the perturbations in the accretion disc move outwards during outburst.

In X-rays, there has only been one detection of a quasi-periodic oscillation (QPO) from J1357, and that was a low-frequency (~ 6 mHz) QPO seen at the start of the 2011 outburst (Armas Padilla et al. 2014) by the *Rossini X-ray Timing Explorer (RXTE)*. During the 2017 outburst of J1357, high-speed multiwavelength photometry was performed, revealing the lack of any X-ray response to the peculiar optical dips. The lack of X-ray dips in the light curves were explained as being due to the presence of a truncated disc with an extended X-ray corona. However, the detailed accretion geometry remains poorly understood, and so further high time-resolution X-ray studies are clearly required.

Such an opportunity presented itself early in 2019 when J1357 underwent its third outburst (Gandhi et al. 2019; Russell et al. 2019; van Velzen, Bellm & van Roestel 2019; Beri et al. 2019b). Here, we report on observations with the *Neil Gehrels Swift* observatory, X-ray data from the Neutron star Interior Composition Explorer (*NICER*) and radio data from the Arcminute Microkelvin Imager (*AMI*). In Section 2, we give observational and data reduction details, while spectral results are presented in Section 3. Timing analysis and results are presented in Section 4. Section 5 gives details on results from UV/Optical and radio observations. We discuss our new results in Section 6.

2 OBSERVATIONS AND DATA ANALYSIS

2.1 Swift

We found that weak X-ray activity was detected in *MAXI* around MJD 58 558 (see Fig. 1), and estimated the corresponding unabsorbed flux value using the *WEBPIMMS HEASARC* tool which translates to $L_X \sim 3 \times 10^{36}$ erg s^{-1} for $d = 6$ kpc. This was also noted by Gandhi et al. (2019) and Russell et al. (2019), and these authors suggested that the 2019 outburst started around MJD 58562.

We have used data from the X-ray Telescope (XRT) and the Ultraviolet and Optical Telescope (UVOT) instruments on-board the *Neil Gehrels Swift* Observatory (Gehrels et al. 2004). The XRT is sensitive in the 0.2–10 keV energy range and has an effective area of 100 cm^2 (see Burrows et al. 2005, for details) while UVOT covers UV and optical bands (170–600 nm) (Romig et al. 2004).

We obtained a total of 15 observations of J1357 between 2019 May 27 and August 8 totalling ~ 18 ks of exposure time (see Table 1 for

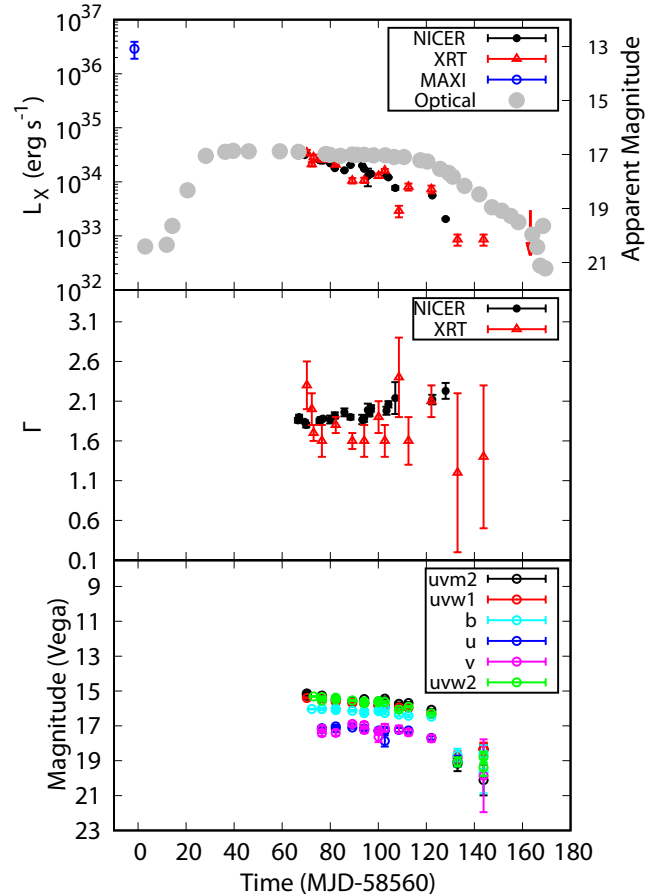


Figure 1. Overview of the J1357 2019 outburst. The (*upper panel*) shows L_X (for $d = 6$ kpc) using data from *Swift*–XRT and *NICER* (0.5–10 keV) and an outline (grey discs) of the ground-based r -band monitoring from Russell et al. (2019). The (*middle panel*) shows the evolution of photon index, Γ , and the (*bottom panel*) shows UV/optical magnitudes (Vega system). Time is in days relative to MJD 58560, the estimated start date of the outburst.

details). The first two XRT observations were in the windowed timing (WT) mode, while the rest were in photon counting (PC) mode. The online tools provided by the UK Swift Science Data Centre¹ (Evans et al. 2009) were used to obtain light curves and spectra from these data.

UVOT observations were taken in image mode, the majority using the six available filters (v , b , u , $uvw1$, $uvm2$, $uvw2$). Light curves in each filter were created using the UVOTMAGHIST tool that uses UVOTSOURCE to perform aperture photometry on all sky images (in each filter) available for an individual observation. Source magnitudes were computed in the Vega system using a 5 arcsec radius circle centred on J1357, together with a neighbouring source-free 10 arcsec radius circle for background correction. All flux values and magnitudes were corrected for Galactic extinction in this direction ($E(B - V) = 0.04$) as done by Armas Padilla et al. (2013) and Beri et al. (2019a).

2.2 NICER

The *NICER* X-ray Timing Instrument (Gendreau et al. 2016) has 56 concentrators, each coupled to a silicon drift detector housed in

¹<http://www.swift.ac.uk/>

Table 1. Log of the *Swift* observations and XRT spectral results during the 2019 outburst of J1357.

Obs-ID	Time MJD	Mode	Count rate (count s ⁻¹)	Exp time (ks)	Γ	$F_{X, \text{unabs}}$ (10 ⁻¹² erg cm ⁻² s ⁻¹)	L_X^b (10 ³² erg s ⁻¹)	cstat/dof
31918084 ^a	58630.2533	WT	0.22 ± 0.02	1.3	2.3 ± 0.3	8.2 ± 1.0	347 ± 42	51/64
31918085	58632.2957	WT	0.11 ± 0.01	1.5	2.0 ± 0.2	4.9 ± 0.4	212 ± 19	160/155
88872001	58633.0559	PC	0.16 ± 0.01	1.1	1.7 ± 0.1	6.6 ± 0.4	284 ± 34	97/133
31918086	58636.5232	PC	0.15 ± 0.01	1.6	1.6 ± 0.2	6.0 ± 0.3	258 ± 25	144/166
31918087	58642.2938	PC	0.14 ± 0.01	1.6	1.8 ± 0.1	4.8 ± 0.3	207 ± 12	133/144
31918088	58649.1678	PC	0.11 ± 0.01	1.4	1.6 ± 0.1	2.5 ± 0.3	105 ± 12	89/110
31918089	58654.1139	PC	0.08 ± 0.01	0.55	1.6 ± 0.2	2.4 ± 0.2	107 ± 12	38/41
31918090	58660.1249	PC	0.10 ± 0.0086	1.6	1.9 ± 0.2	3.0 ± 0.2	129 ± 12	87/97
31918091	58662.6706	PC	0.06 ± 0.01	0.7	1.6 ± 0.2	3.8 ± 0.6	161 ± 12	19/38
31918092	58668.5556	PC	0.05 ± 0.01	0.6	2.4 ± 0.5	0.7 ± 0.2	29 ± 7	10/21
31918093	58672.5342	PC	0.07 ± 0.01	1.0	1.6 ± 0.3	1.9 ± 0.5	81 ± 12	28/66
31918094	58682.1041	PC	0.06 ± 0.006	1.4	2.1 ± 0.2	1.7 ± 0.2	73 ± 12	43/70
31918095	58692.9878	PC	0.008 ± 0.003	1.3	1.2 ± 1.0	0.2 ± 0.1	9 ± 2	4/7
31918096	58703.8096	PC	0.006 ± 0.002	1.5	1.4 ± 0.9	0.2 ± 0.1	9 ± 2	5/8
31918097	58712.3114	PC	0.0035 ^c	1.3	2.5	<0.102 ^a	<4.4 ^a	–

Notes. ^aShows a sharp mHz peak in the PDS at a 95 per cent significance level.

^bCalculated from the 0.5–10 keV unabsorbed flux ($F_{X, \text{unabs}}$) for a distance of 6.0 kpc.

^c95 per cent confidence upper level limit count rates using the prescription given by Gehrels (1986). We estimated the corresponding unabsorbed flux upper limits using the WEBPIMMS HEASARC tool.

a Focal Plane Module (FPM). At the time of J1357 observations, 52 of the 56 FPMs were functional, providing an effective area of ~ 1750 cm² in the 0.2–12 keV band. J1357 was observed with *NICER* between May 23 and 2019 August 20 (see Table 3), and the data were reprocessed using *NICERDAS* version 8c, distributed as part of *HEASOFT* version 6.29c. The cleaned event files were generated using the *nicerl2*² task in *HEASOFT* version 6.29c using the default filter criteria which includes limiting analysis to time intervals with a pointing offset <54 arcsec, a bright Earth limb angle >30°, a dark Earth limb angle <15°, and ignoring data collected during the passage through the South Atlantic Anomaly. Standard screening procedures also include rejecting all time intervals where the rate of saturating particle events (overshoots) is greater than 1 counts s⁻¹ detector⁻¹. These cleaned events were then filtered using the *nifpmsel* tool to remove data from two noisier detectors (Detectors 14 and 34). The final cleaned events (from all MPU) were then processed using *XSELECT* to obtain scientific products. The background spectra were generated using the *nibackgen3C50* tool (Remillard et al. 2022). Standard tools (*nicerrmf* and *nicerarf*) were used to create response and auxiliary files for the spectral analysis. In order to investigate the presence of any artificial features such as dips in the X-ray light curves, we also examined data following the procedure recommended by the *NICER* team.³

2.3 Arcminute Microkelvin Imager

Radio observations with the *Arcminute Microkelvin Imager Large Array* (AMI-LA; Zwart et al. 2008; Hickish et al. 2018, hereafter AMI) were triggered after the source went into outburst. AMI is an 8-element radio array located in Cambridge, UK, operating at a central frequency of 15.5 GHz with a 5 GHz bandwidth. A single observation of J1357 was performed on 2019 May 26 (MJD 58629.84) for a total of 4 h. The observation used six of the eight 13 m AMI antennas, interleaving 100 s scans of the phase calibrator NVSS J140412 – 001324 with 10 min scans of the target field.

The data were flux calibrated using 3C286 and were reduced with a custom pipeline *REDUCE_DC* (Perrott et al. 2015) and imaged using the Common Astronomical Software Applications (CASA).⁴ Unfortunately, NVSS J135720 – 093003 is <3 arcmin away from J1357 and within the primary beam of *AMI*. It was removed from the *uv* data using the CASA task *uvsub* and a final image was obtained of the field, with an rms-noise limit of 70 μ Jy beam⁻¹. J1357 is not detected above a 3 σ rms noise (210 μ Jy beam⁻¹) threshold.

3 X-RAY SPECTROSCOPY

3.1 X-ray spectral analysis

We performed spectral analysis using *XSPEC* 12.12.0 (Arnaud 1996). Owing to the limited *Swift*-XRT statistics, we grouped spectra using the *ftools* task *grppha* so as to have at least one count per bin. *W*-statistics (background subtracted Cash statistics) were employed while performing the X-ray spectral fitting (Wachter, Leach & Kellogg 1979). Spectra obtained with *NICER* were grouped into bins with a minimum of 25 photons and χ^2 statistics were applied. Interstellar absorption was included in all our spectral fits, employing Wilms, Allen & McCray (2000) abundances and Verner et al. (1996) photoelectric cross-sections. *TBABS* was used to model the hydrogen column density (N_{H}) with a value fixed to 1.2×10^{20} cm⁻² as obtained by Armas Padilla et al. (2014) from the *XMM-Newton* high-resolution X-ray spectra. The X-ray spectral fitting was performed over ranges of 0.5–10 keV for PC mode and 0.7–10 keV for WT mode, due to the presence of low-energy residuals in the latter.⁵ For the *NICER* spectra, emission from the source was detectable over a narrow energy range in all cases (0.4–2.5 keV), and so our spectral fits are confined to this range. All the fluxes reported are unabsorbed 0.5–10 keV fluxes obtained using the convolution model ‘*CFLUX*’. Unless explicitly mentioned, we quote all errors at 1 σ confidence level.

²<https://heasarc.gsfc.nasa.gov/lheasoft/ftools/headas/nicerl2.html>

³https://heasarc.gsfc.nasa.gov/docs/nicer/analysis_threads/iss_obstruction/

⁴<https://casa.nrao.edu/>

⁵<http://www.swift.ac.uk/analysis/xrt/>

Table 2. Spectral parameters of J1357 for an absorbed disc-blackbody and power-law model from *Swift*-XRT data. We have assumed a distance of 6.0 kpc for estimating the diskbb radius.

Obs ID	kT (keV)	Diskbb Radius (km)	Γ	Power-law norm	cstat/dof
31918084	0.17 ± 0.06	25_{-17}^{+45}	0.4 ± 0.6	$(3.1_{-1}^{+3}) \times 10^{-4}$	31/62

Table 3. Log of the *NICER* observations during the 2019 outburst of J1357.

Obs-ID	Time (MJD)	Exp-time (ks)	Count rate (count s ⁻¹)	Γ	$F_{X, \text{unabs}}$ (10^{-12} erg cm ⁻² s(10^{32} erg s ⁻¹))	L_X^a	χ^2/dof
2200730101 (Obs 1)	58626.57	1.9	3.47 ± 0.05	1.86 ± 0.03	7.25 ± 0.10	312.3 ± 4.2	143/130
2200730102 (Obs 2)	58627.02	1.8	3.65 ± 0.05	1.86 ± 0.03	7.49 ± 0.11	322.4 ± 4.8	116/124
2200730103 (Obs 3)	58629.41	6.4	3.30 ± 0.03	1.84 ± 0.02	7.28 ± 0.05	313.7 ± 2.3	157.8/157
2200730104 (Obs 4)	58629.99	1.5	3.63 ± 0.05	1.80 ± 0.03	7.89 ± 0.12	339.6 ± 5.0	118/119
2200730105 (Obs 5)	58635.53	2.8	2.65 ± 0.04	1.85 ± 0.03	5.84 ± 0.07	251.3 ± 3.1	132/141
2200730106 (Obs 6)	58636.05	3.3	2.66 ± 0.03	1.87 ± 0.03	5.82 ± 0.07	250.7 ± 2.9	123/147
2200730107 (Obs 7)	58637.01	5.4	2.60 ± 0.03	1.88 ± 0.02	5.75 ± 0.05	247.5 ± 2.3	151/155
2200730108 (Obs 8)	58639.98	1.4	2.15 ± 0.04	1.87 ± 0.05	5.12 ± 0.11	220.5 ± 4.7	789/810
2200730109 (Obs 9)	58641.91	2.2	2.15 ± 0.04	1.92 ± 0.04	4.18 ± 0.08	179.9 ± 3.3	105/107
2200730110 (Obs 10)	58645.90	1.3	1.89 ± 0.05	1.96 ± 0.05	3.79 ± 0.08	163.4 ± 3.5	862/910
2200730111 (Obs 11)	58648.42	2.8	2.19 ± 0.03	1.90 ± 0.03	4.75 ± 0.06	204.4 ± 2.8	151/128
2200730112 (Obs 12)	58653.32	0.84	1.87 ± 0.06	1.87 ± 0.06	4.66 ± 0.11	200.7 ± 4.9	650/650
2200730113 (Obs 13)	58654.09	2.5	1.65 ± 0.03	1.88 ± 0.04	4.05 ± 0.06	174.3 ± 2.8	106/113
2200730114 (Obs 14)	58655.64	0.6	1.31 ± 0.06	1.99 ± 0.08	3.00 ± 0.11	129.1 ± 4.6	417/350
2200730115 (Obs 15)	58656.35	3.3	1.44 ± 0.03	1.95 ± 0.04	3.25 ± 0.05	139.8 ± 2.2	141/120
2200730116 (Obs 16)	58656.99	3.2	1.68 ± 0.03	2.01 ± 0.04	3.28 ± 0.06	141.3 ± 2.4	92/106
2200730117 (Obs 17)	58663.38	2.4	1.50 ± 0.03	1.98 ± 0.05	3.01 ± 0.06	129.6 ± 2.6	87/100
2200730118 (Obs 18)	58664.22	3.0	1.48 ± 0.03	2.06 ± 0.04	2.81 ± 0.05	121.0 ± 2.1	102/117
2200730119 (Obs 19)	58667.06	0.25	0.97 ± 0.08	2.14 ± 0.19	1.79 ± 0.13	77.2 ± 5.8	9.2/11
2200730120 (Obs 20)	58682.46	2.5	0.80 ± 0.02	2.12 ± 0.06	1.31 ± 0.03	56.5 ± 1.4	78/73
2200730121 (Obs 21)	58688.01	3.5	0.06 ± 0.01	2.23 ± 0.10	0.48 ± 0.02	20.5 ± 0.9	72/68

Note. ^aCalculated from the 0.5–10 keV unabsorbed flux ($F_{X, \text{unabs}}$) for a distance of 6.0 kpc.

3.2 X-ray spectral results

The top panel of Fig. 1 shows the L_X (0.5–10 keV) from *Swift*-XRT and *NICER*. The maximum value of unabsorbed flux measured is $\sim 8 \times 10^{-12}$ erg cm⁻² s⁻¹ which corresponds to L_X of about 3.4×10^{34} (d/6 kpc)² erg s⁻¹, lower than previous outbursts, likely as a result of not observing until well after the outburst had started and declines throughout the *Swift* and *NICER* observations (Tables 1 and 3). We have also replotted the published optical outburst profile from Russell et al. (2019). Using an absorbed power-law model, we find that the power-law index (Γ) evolves over the outburst (Fig. 1, middle panel) with Γ between ~ 1.5 and ~ 2.5 during these observations, indicating the presence of soft X-ray spectra.

In one of the XRT spectra (ObsID 31918084) with higher statistics, we found that using an absorbed power-law model showed some excess in the spectral residuals (see Fig. 2). Therefore, we added a *diskbb* component to the absorbed power-law model and found significant improvement in the spectral fit, giving a disc temperature of 0.17 keV (see Table 2). To verify the statistical significance of adding a disc component, we performed a *ftest* to determine the probability of chance improvement (PCI), which was very low (4×10^{-7}), indicating it is reasonable to add the *diskbb* component in the X-ray spectra. However, we would like to add a caveat that given limited statistics, it was difficult to constrain all fit parameters. To further evaluate the chance probability of improvement by adding the extra *diskbb* component, we simulated 100 000 data sets using *simftest* in XSPEC. The evaluated chance probability is less than

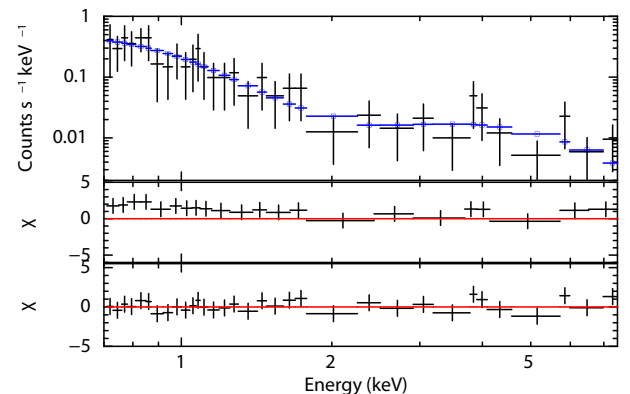


Figure 2. Top panel: Best-fitting *Swift*-XRT spectra of J1357, obtained using an absorbed *diskbb* and power law. The lower panels show residuals (χ) to the fits of (middle panel) a simple absorbed power-law (which reveals a low-energy excess), and (bottom panel) the best-fitting model.

10^{-6} , indicating that the presence of an additional disc component is significant.

The unabsorbed 0.5–10 keV flux during this XRT observation is $\sim 8 \times 10^{-12}$ erg cm⁻² s⁻¹, to which the thermal component contributes 50 per cent, i.e. an inferred $L_X \sim 3.4 \times 10^{34}$ erg s⁻¹. The spectral residuals did not indicate the presence of a neutral iron K_α line at 6.4 keV, which is consistent with the previous study carried out by Beri et al. (2019b) using *Swift* and *NuSTAR* data. The non-

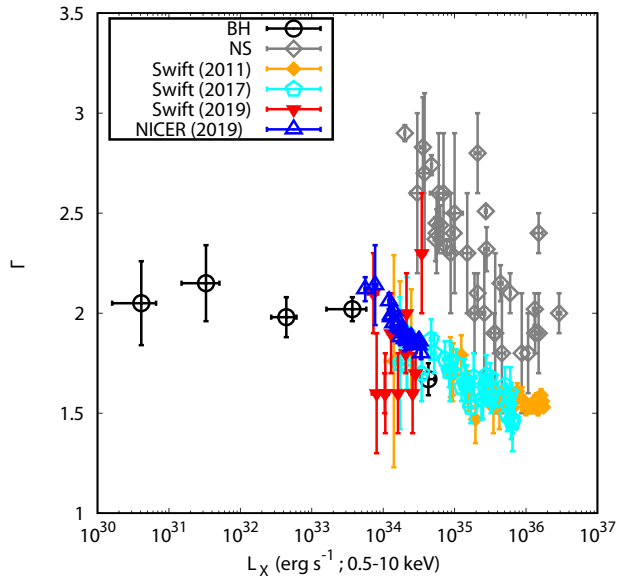


Figure 3. Comparison of the spectral (photon index Γ) – X-ray luminosity (L_X , 0.5–10 keV) behaviour of J1357 (coloured points, see legend, assuming $d = 6$ kpc) with other low-level accreting LMXBs as compiled by Wijnands et al. (2015) for both BH (solid black points) and NS (grey points) transients. The earlier J1357 values are from Armas Padilla et al. (2013) and Beri et al. (2019a) for the 2011 and 2017 outbursts, respectively.

detection of the iron line was also confirmed from our simulations of 10 000 data sets that returned the null-hypothesis probability of about 0.98.

In Fig. 3, we compare J1357’s spectral properties as a function of L_X for all three outbursts, together with those from the Wijnands et al. (2015) survey of low-level accreting NS and BH transients at $L_X < 10^{36}$ erg s $^{-1}$. NS LMXBs are significantly softer than BH systems below an L_X (0.5–10 keV) of 10^{35} erg s $^{-1}$. During all three outbursts of J1357, the power-law index showed similar behaviour, clearly following the general trend of the BH sample. However, an interesting point to note is that the 2019 outburst showed higher values of Γ compared to the earlier outbursts of J1357, indicating softer X-ray spectra.

4 TIMING RESULTS

We investigated the X-ray light curves of J1357 to search for the presence of dip-like features similar to those observed in the optical (e.g. Corral-Santana et al. 2013). The X-ray light curves were extracted with a bin size of 100 ms in the 0.2–12 keV energy range, with average count rates for each observation given in Table 3.

We created a power density spectrum (PDS) using data of Obs 9 with the FTOOL task ‘powspec’. Obs 9 showed the most flux variability while similar features were also seen with other *NICER* observations, albeit during brief snapshots (~ 500 s). The Obs 9 light curve with 20 s bins (0.2–12 keV) was used to search for signals in the lower frequency range. The PDS showed the presence of excess power around ~ 3 mHz (Fig. 4), similar to that seen by *RXTE* during J1357’s 2011 outburst. In order to measure the characteristic frequency and full width at half-maximum (FWHM), we fitted the Obs 9 PDS with a single, zero-centred Lorentzian, obtaining ν_0 and FWHM of 0.0028 ± 0.0001 and 0.0007 ± 0.0003 Hz. The quality factor (Q) is $\sim 4.0 \pm 1.6$ and the root mean square (rms) amplitude

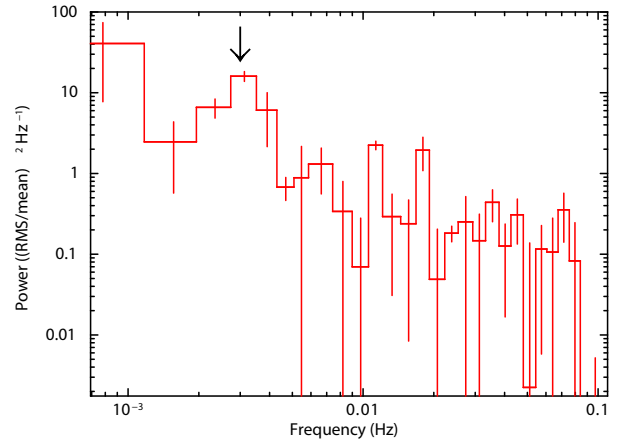


Figure 4. PDS obtained using the *NICER* data of Obs 9, showing the presence of a millihertz QPO (marked with an arrow) with a rms fractional variability of ~ 4 per cent (0.2–12 keV).

is $4.8^{+0.5}_{-1.3}$ per cent. We also observed two additional sharp peaks in the PDS around ~ 0.011 and 0.014 Hz.

In order to further investigate all these peaks in the PDS we used independent methods such as Lomb–Scargle periodogram (Lomb 1976; Scargle 1982; Horne & Baliunas 1986) and CLEAN (Roberts, Lehar & Dreher 1987) as implemented in the PERIOD programme distributed with STARLINK SOFTWARE COLLECTION⁶ (Currie et al. 2014). The CLEAN algorithm is a powerful tool that basically deconvolves the spectral window from the discrete Fourier power spectrum (or dirty spectrum) and produces a CLEAN spectrum, which is largely free of the many effects of spectral leakage. A sharp peak was observed in all the periodograms, and a consistent value of QPO frequency was obtained using this method (Fig. 5). The QPO frequency varied between 1 and 5 mHz during these observations. However, we did not observe any sharp features around 11 and 14 mHz as seen in Fig. 4.

The significance of these features was determined using a Fisher randomization test (Linnell Nemeč & Nemeč 1985), which consists of calculating the periodogram of a new, randomized time-series. This randomization of data and periodogram calculation loop is then performed for a large number of permutations. The False Alarm Probabilities (FAP) were calculated to indicate the significance of any peak detected, and were found to be zero (95 per cent confidence interval 0.0 to 0.01) for those shown in Fig. 5 and the values are listed in Table 4. The lower significance value implies that the obtained period is correct. A large value of FAP would suggest that any detected periodicities were unlikely to be real. Although all X-ray light curves showed the presence of variability, a PDS as shown in Fig. 4 could not be created for Obs 19 with *NICER* due to its short duration.

We also evaluated the significance of the QPO against the red noise using a Monte Carlo simulation (see e.g. Benlloch et al. 2001; Vaughan 2005). This method involves determination of the shape of the PDS of each observation by fitting a power-law model (as shown in Fig. A1). Thereafter, we generated 10^5 trial light curves from a best-fitting power index of PDS using PYTHON libraries of Stingray⁷

⁶<http://starlink.eao.hawaii.edu/starlink>

⁷Stingray is a PYTHON package for X-ray astronomy, and is available at <https://github.com/StingraySoftware/stingray>.

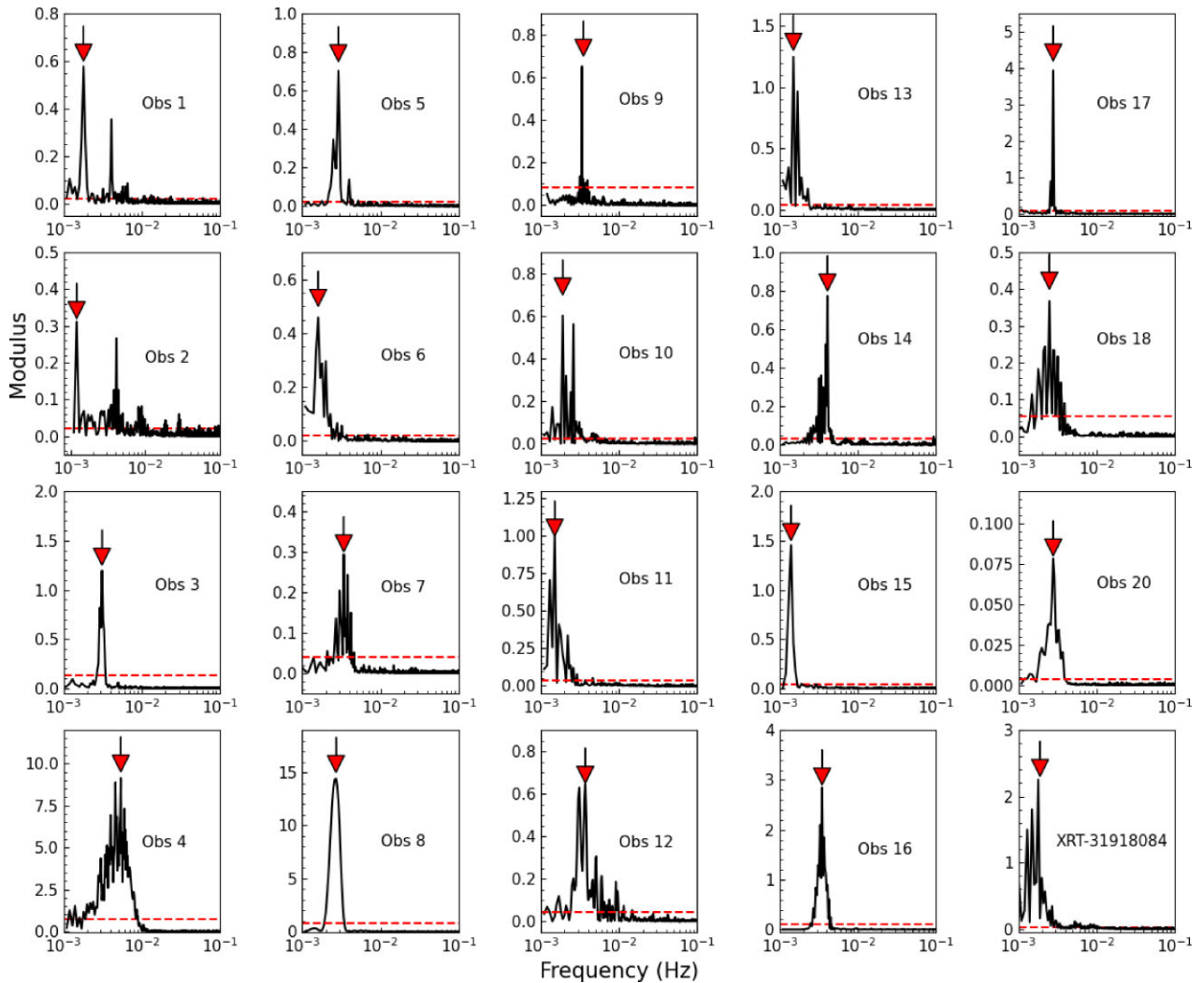


Figure 5. The plots show CLEAN power density spectra. The arrows indicating sharp peaks in the millihertz frequency range. We have included only those observations during which the false alarm probabilities for the period detection lie below 0.01 with 95 per cent confidence. The dashed line (red) is the 5σ significance level.

version 1.1 (Huppenkothen et al. 2019a; Huppenkothen et al. 2019b; Bachetti et al. 2022). We have used duration, mean count rate, and variance as that of the observed light curve. An example of simulated red-noise light curve is shown in Fig. A2. We calculated corresponding power spectra searching for peaks in the 0.001–0.1 Hz range. The 3σ confidence limits on the maximum power are calculated, including the uncertainties in the red noise model. The chance probability of occurrence of the observed signal is obtained by counting the number of trial time series with powers equal to or exceeding the observed power in the frequency range of 0.001–0.009 Hz. Our simulation results are shown in Fig. A3. The detection significance for the observations showing the QPO is given in Table 4. The detection significance for the observations showing the QPO is well above 3σ for all observations except Obs 7, Obs 13, Obs 18, and XRT observation.

Fig. 6 shows the evolution of X-ray QPOs compared to optical QPOs observed during previous outbursts of this source. Unlike optical QPOs, X-ray QPOs did not show a monotonic trend of dip frequency decreasing as the outburst declines.

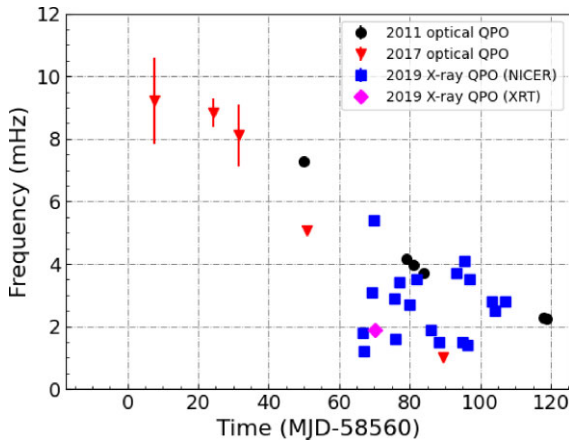
5 UV/OPTICAL AND RADIO OBSERVATIONS OF J1357

We also explored the multiwavelength variability of J1357 during the 2019 outburst. The bottom panel of Fig. 1 shows that the UV/optical magnitudes display a continuous decrease in brightness (from magnitude ~ 15 to ~ 19), similar to that observed in X-rays. There is a correlation between X-ray and UV/optical fluxes (Fig. 7). We performed a Monte-Carlo simulation to compute the 1σ confidence interval (shaded region in Fig. 7). 10 000 iterations were performed and at each step, 1000 data points were generated. We used the covariance matrix to determine confidence intervals of each random small subset of data, and this step was repeated for 10 000 iterations to obtain an average value of the confidence interval. The best-fitting correlation slopes (β) are given in Table 5. We also observe that β increases at shorter wavelengths, consistent with previous reports on this source (Fig. 8).

We compared the correlation between the UV/optical and X-ray fluxes against correlations observed for three emission processes: X-ray reprocessing in the disc, the viscously heated disc, and jet

Table 4. Peaks observed in the PDS shown in Fig. 5 are listed below.

Observation	Frequency (mHz)	Significance (σ)
Obs 1	1.8 ± 0.1	12
Obs 2	1.2 ± 0.1	8
Obs 3	3.1 ± 0.1	12
Obs 4	5.4 ± 0.1	3.9
Obs 5	2.9 ± 0.1	4.4
Obs 6	1.6 ± 0.1	12
Obs 7	3.4 ± 0.1	0.45
Obs 8	2.7 ± 0.1	14
Obs 9	3.5 ± 0.1	12
Obs 10	1.9 ± 0.1	5
Obs 11	1.5 ± 0.1	6
Obs 12	3.7 ± 0.1	7
Obs 13	1.5 ± 0.1	2.1
Obs 14	4.1 ± 0.1	7
Obs 15	1.4 ± 0.1	4.3
Obs 16	3.5 ± 0.1	6.7
Obs 17	2.8 ± 0.1	10
Obs 18	2.5 ± 0.1	0.4
Obs 20	2.8 ± 0.1	4.3
XRT-31918084	1.8 ± 0.1	2.7

**Figure 6.** Evolution of J1357’s QPO frequency with time during the 2019 outburst (blue squares – *NICER*; magenta diamond – *Swift*), compared with the optical QPO behaviour in previous outbursts (see legend) from Corral-Santana et al. (2013) and Paice et al. (2019).

emission. For X-ray reprocessing, we adopt the theoretical model between the optical and X-ray luminosities given by van Paradijs & McClintock (1994). According to this model optical luminosity of an X-ray reprocessing accretion disc varies as $L_{\text{opt}} \propto L_X^{0.5} a$, where a is the orbital separation of the system given by $3.5 \times 10^{10} (M_{\text{BH}})^{1/3} (1 + q)^{1/3} (P_{\text{hr}})^{2/3}$ (Frank, King & Raine 2002). The values of BH mass (M_{BH}), the mass ratio of the companion star to the compact object, $q = M_c/M_{\text{BH}}$ and the orbital period (P_{hr}) were taken from Casares (2016). For the viscously heated disc and jet emission, we have used the following relations: $L_{\text{opt}} \propto L_X^{0.25}$ and $L_{\text{opt}} \propto L_X^{0.7}$, respectively (see Russell et al. 2006). The best-fitting correlation slopes (Table 5) show that for all the UVOT bands the values lie closer to the model for X-ray reprocessing. However, for the V-band emission, there can also be a contribution from a viscously heated disc around a BH.

We did not detect the source at radio wavelengths with *AMI*, our observations only yielding 3σ upper limits of $210 \mu\text{Jy beam}^{-1}$.

6 DISCUSSION

6.1 Spectral behaviour

Swift-XRT and *NICER* observations made during the 2019 outburst of J1357 were well-fit using a simple absorbed power-law consistent with the 2011 and 2017 outbursts of J1357 (Armas Padilla et al. 2013; Beri et al. 2019a). We found that an additional disc component was needed to obtain the best fit in one of the XRT observations with a higher signal-to-noise ratio. During the 2011 outburst of J1357, Armas Padilla et al. (2014) did find disc emission in their *XMM-Newton* spectrum with a disc temperature of ~ 0.2 keV, similar to that measured here. The low value of disc temperature observed during the outburst in 2011 was believed to be due to J1357 being in its low hard spectral state. In fact, during the 2017 outburst, models fitted to *NuSTAR* observations did not require any disc contribution. The non-detection of the disc component was explained as being due to the presence of a cool disc (see Beri et al. 2019a for details). Moreover, we did not find the presence of any reflection features in the form of an iron line in the *NICER* spectra. This is again consistent with previous reports on this source (see e.g. Beri et al. 2019a).

The exponential (viscous) decay time-scale measured during earlier outbursts was found to be ~ 64 d (Tetarenko et al. 2018). Moreover, Fig. 3 shows that values of photon index are comparable to those observed during the decline of the 2011 and 2017 outbursts (Armas Padilla et al. 2013; Beri et al. 2019a). During the 2019 outburst of J1357, our *NICER* and *Swift*-XRT observations revealed the maximum value of unabsorbed flux to be $\sim 8 \times 10^{-12} \text{ erg cm}^{-2} \text{ s}^{-1}$ which corresponds to L_X of about $3.4 \times 10^{34} (\text{d}/6 \text{ kpc})^2 \text{ erg s}^{-1}$. This flux is almost 18 times lower than that observed in 2017 (Beri et al. 2019a) and 55 times lower than measured in 2011 (Armas Padilla et al. 2013). All this indicates that *NICER* and *Swift* observations were made during the declining phase of the outburst.

6.2 Timing behaviour

Thanks to the excellent fast timing capability of *NICER*, for the first time we were able to probe the X-ray variability in J1357. Structures in the X-ray light curves were found to vary on time-scales between ~ 200 and 700 s with a strong energy dependence.

The *NICER* PDS for Obs 9 (Fig. 4) showed the presence of a QPO at around 3 mHz with a Q factor of ~ 4 . Early in the 2011 outburst of J1357, *RXTE* found a similar QPO at ~ 6 mHz with $Q \sim 3$ (Armas Padilla et al. 2014). The presence of a sharp peak at corresponding frequencies was also observed in our Lomb-Scargle and CLEAN periodograms (see Fig. 5). One of our most intriguing results is that we found peaks in the millihertz frequency range that were consistent with those seen in the optical during the decline phase of the 2011 and 2017 outbursts (Corral-Santana et al. 2013; Paice et al. 2019).

It could be possible that the origin of X-ray QPOs is linked to that for optical dips, as they were also observed at similar frequencies during the 2019 outburst (Jimenez-Ibarra et al. 2019). During the previous outbursts of J1357, a monotonic decrease of the optical dip frequency as the outburst progresses has been observed; however, this behaviour is not evident in our X-ray data (Fig. 6). The interpretation of this optical dip behaviour was that the inner edge of some obscuring material was moving outwards through the disc during the decline. Thus, we cannot rule out other possibilities as to the origin of the X-ray QPO. Millihertz X-ray QPOs observed in J1357 are believed to resemble ‘1 Hz QPOs’ seen in dipping neutron star systems (see Armas Padilla et al. 2014), and it is therefore possible that they share a common origin. Thus, millihertz X-ray QPOs

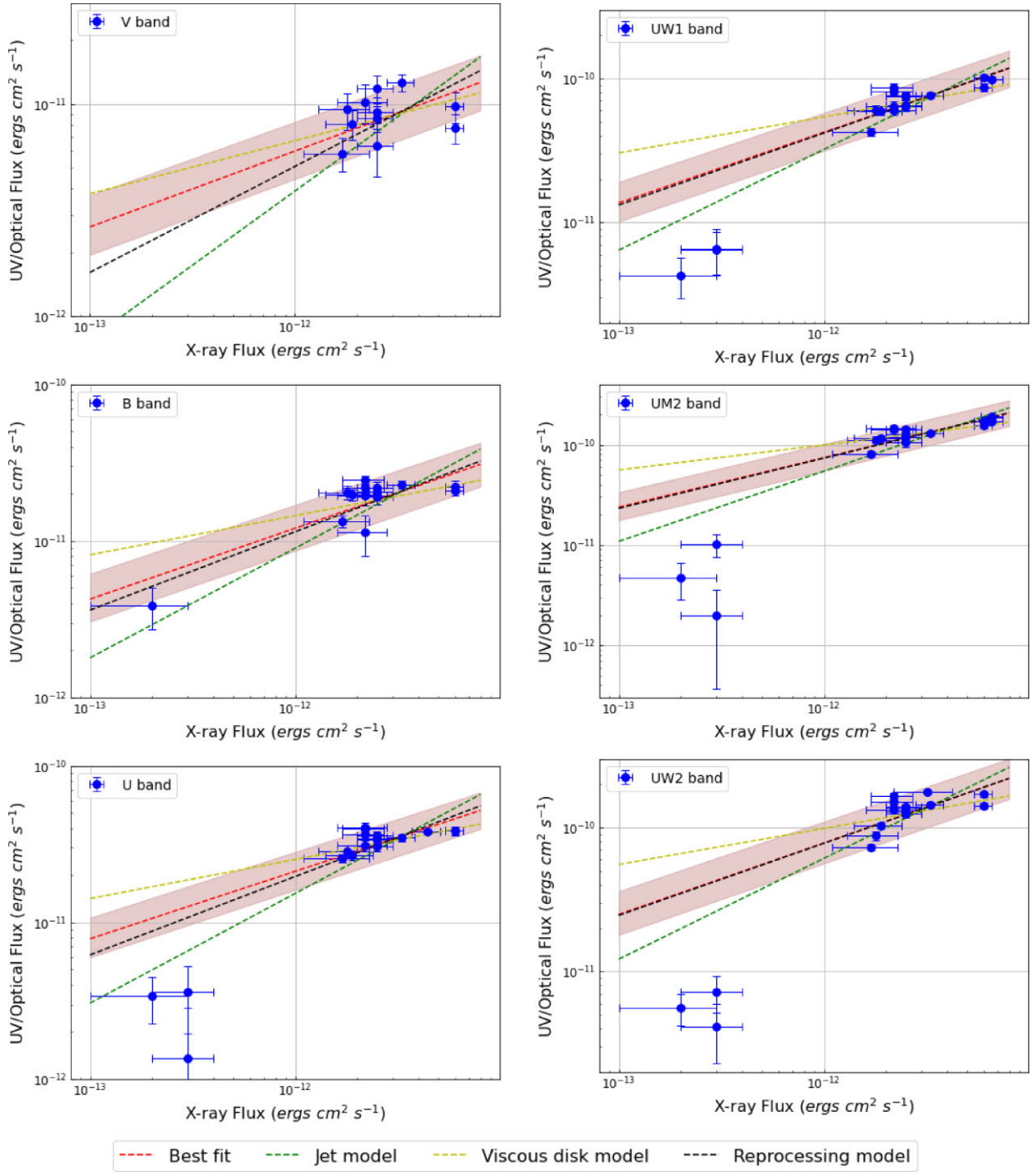


Figure 7. Correlation slopes (β) between the UVOT bands and the unabsorbed X-ray flux in the 0.5–10 keV energy band, with models for X-ray reprocessing, viscously heated disc and jet components. The shaded region represents the 1σ confidence interval (refer to the text for more details).

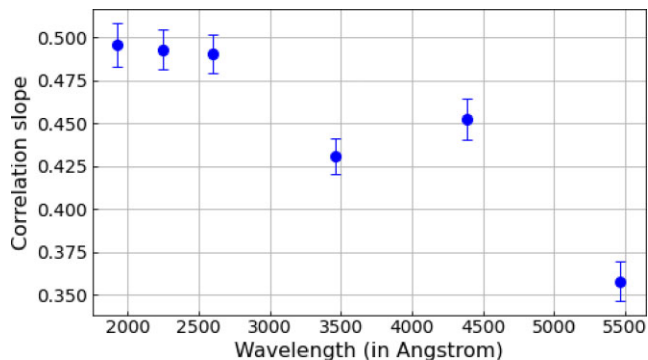
observed in J1357 could be due to a structure in the inner disc which quasi-periodically obscures the inner region (Jonker, van der Klis & Wijnands 1999) or due to relativistic Lense–Thirring precession of the inner accretion disc (Homan 2012). For the first time millihertz X-ray QPOs have been significantly detected during the decline phase of the J1357 outburst, ruling out the earlier suggestion that the X-ray QPOs observed in J1357 are part of the new class of BH QPOs with frequencies in the millihertz range seen only at the start of an outburst (Altamirano & Strohmayer 2012).

6.3 UV/Optical and X-ray correlation

The decay profile after the outburst peak is well explained using the disc-instability model including irradiation (Dubus et al. 1999; Dubus, Hameury & Lasota 2001) and such profiles have been observed in a number of BH-LMXB outburst light-curves, including outburst light-curves of J1357 (Tetarenko et al. 2018). Therefore, as reported in Beri et al. (2019a), we compared the correlation between the UV/optical and X-ray fluxes against correlations observed for

Table 5. Correlation slope (β) between UV/optical and X-ray fluxes during 2019 outburst of J1357.

Filter (wavelength; Å)	Best fit slopes (β)
V (5468)	0.36 ± 0.01
B (4392)	0.45 ± 0.01
U (3465)	0.43 ± 0.01
UVW1 (2600)	0.49 ± 0.01
UVM2 (2246)	0.49 ± 0.01
UVW2 (1928)	0.49 ± 0.01

**Figure 8.** Best-fitting correlation slopes (β) between the UVOT bands and the 0.5–10 keV X-ray flux.

three emission processes: X-ray reprocessing in the disc, the viscously heated disc, and jet emission (Fig. 7). For the V band, the best-fitting β value was found to deviate from the predicted values for the reprocessing and jet models (Fig. 8). It might be possible that intrinsic thermal emission from the viscously heated outer accretion disc contributes significant light in the optical (Frank et al. 2002). However, for the other bands the β values were more consistent with the predicted values for the reprocessing model, favouring the idea suggested by van Paradijs & McClintock (1994) that for smaller accretion discs (i.e. smaller P_{orb}) a higher value of the average surface temperature of the disc is expected (presumably as it is closer to the compact object and irradiating source). Therefore, one would expect to find a larger fraction of the reprocessed emission in the UV band.

7 CONCLUSIONS

In this work, we have analysed multiwavelength data of J1357 during its 2019 outburst. We summarize our findings as follows:

(i) Our X-ray spectral analysis suggests that our *NICER* and *Swift*-XRT observations were made during the declining phase of the outburst. The maximum X-ray flux we observed is almost 18 times lower than the peak seen in 2017 (Beri et al. 2019a) and 55 times lower than that in 2011 (Armas Padilla et al. 2013). Moreover, our spectral index measurements are comparable to those observed during the decay phase observations in 2011 and 2017. One of the XRT observations also revealed the presence of a cool accretion disc.

(ii) X-ray QPOs in the millihertz frequency range have been detected for the first time throughout the outburst of J1357. This is in contrast to that observed in the 2011 outburst, during which 6 mHz QPO was detected only during the first *RXTE* observation and was not present during any of the later *RXTE* observations or during

the *XMM*–*Newton* observation which was taken 3 d after the first *RXTE* observation (Armas Padilla et al. 2014). A number of models could explain this X-ray variability, not least of which involves an X-ray component to the curious dips that, so far, have chiefly been seen only in the optical. Therefore, future X-ray observations of such outbursts would help to answer these questions.

(iii) Our optical/UV and X-ray correlation study indicates a significant contribution of the X-ray reprocessing to the optical and UV emission. However, this is in contrast to what has been observed during the 2011 and 2017 outbursts of J1357 where Armas Padilla et al. (2013) and Beri et al. (2019a) observed UV/optical emission to be dominated by the viscously heated disc, with little or no contribution from X-ray irradiation.

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This work made use of the NUMPY (van der Walt, Colbert & Varoquaux 2011), MATPLOTLIB (Hunter 2007), SCIPY (Jones, Oliphant & Peterson 2001), PANDAS (Wes McKinney 2010), ASTROPY (Astropy Collaboration et al. 2013), STINGRAY (Huppenkothen et al. 2019a) PYTHON packages.

DATA AVAILABILITY

The data underlying this article are publicly available from the High Energy Astrophysics Science Archive Research Center (HEASARC), provided by NASA’s Goddard Space Flight Center.

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APPENDIX: LIGHT-CURVE SIMULATION RESULTS

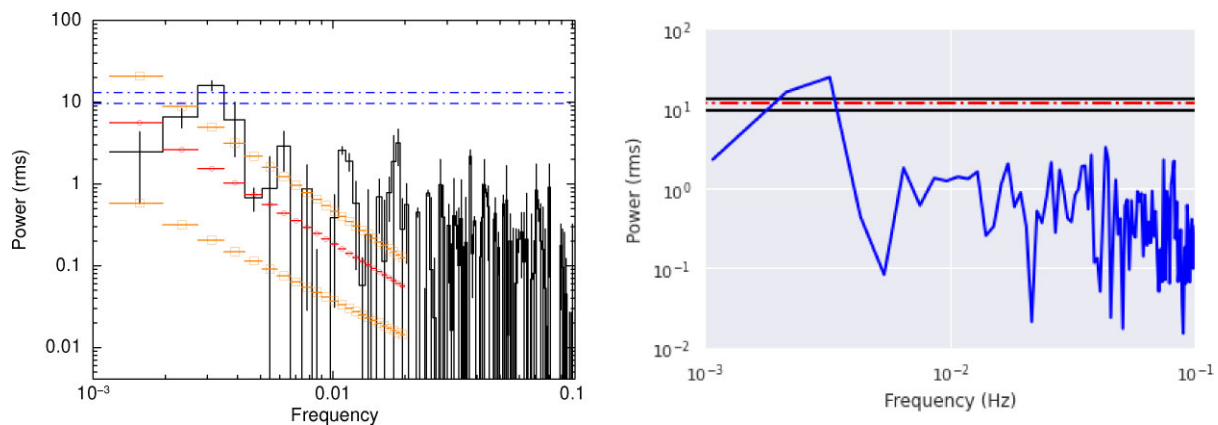


Figure A1. The left plot shows the PDS created using ‘*powspec*’ tool (as shown in Fig. 4), orange squares are the uncertainties in the red noise model while the best fitting photon index is shown in red circles. These values were used to simulate light curves, and blue dashed lines is the maximum power observed at 3σ confidence using the observed range of photon index. On the right, we show an example PDS created using light curve of Obs 9, using *stingray* library ‘*powerspectrum*’. In order to compute significance of observed QPO peak, we have generated PDS of simulated light curves using this tool. The dotted dashed lines in red indicate the maximum power observed at 3σ confidence in the simulated light curves, using the best-fitting photon index and solid black lines are calculated, including the uncertainties in the red noise model.

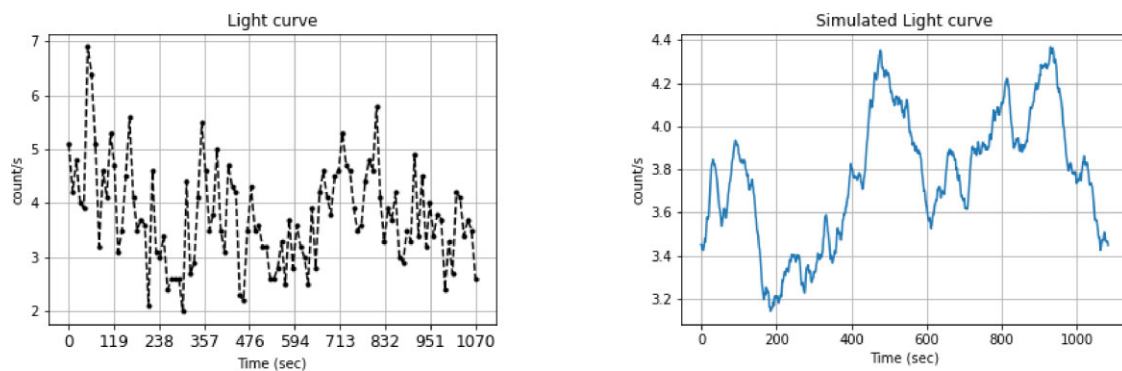


Figure A2. The left plot is the light curve of one of the observations (Obs 8), while the right plot is the simulated red-noise light curve obtained using the best-fitting parameters of the observed light curve.

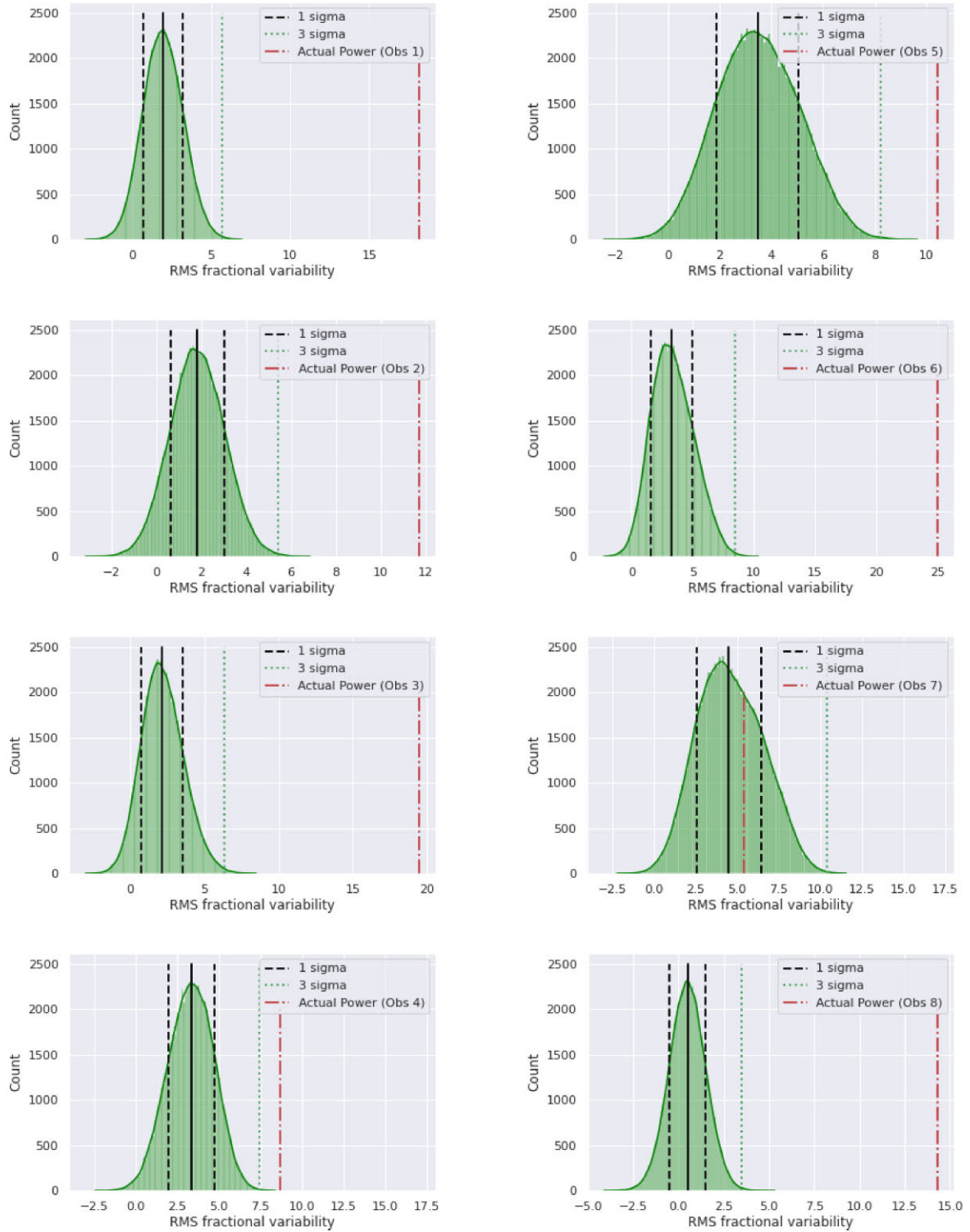
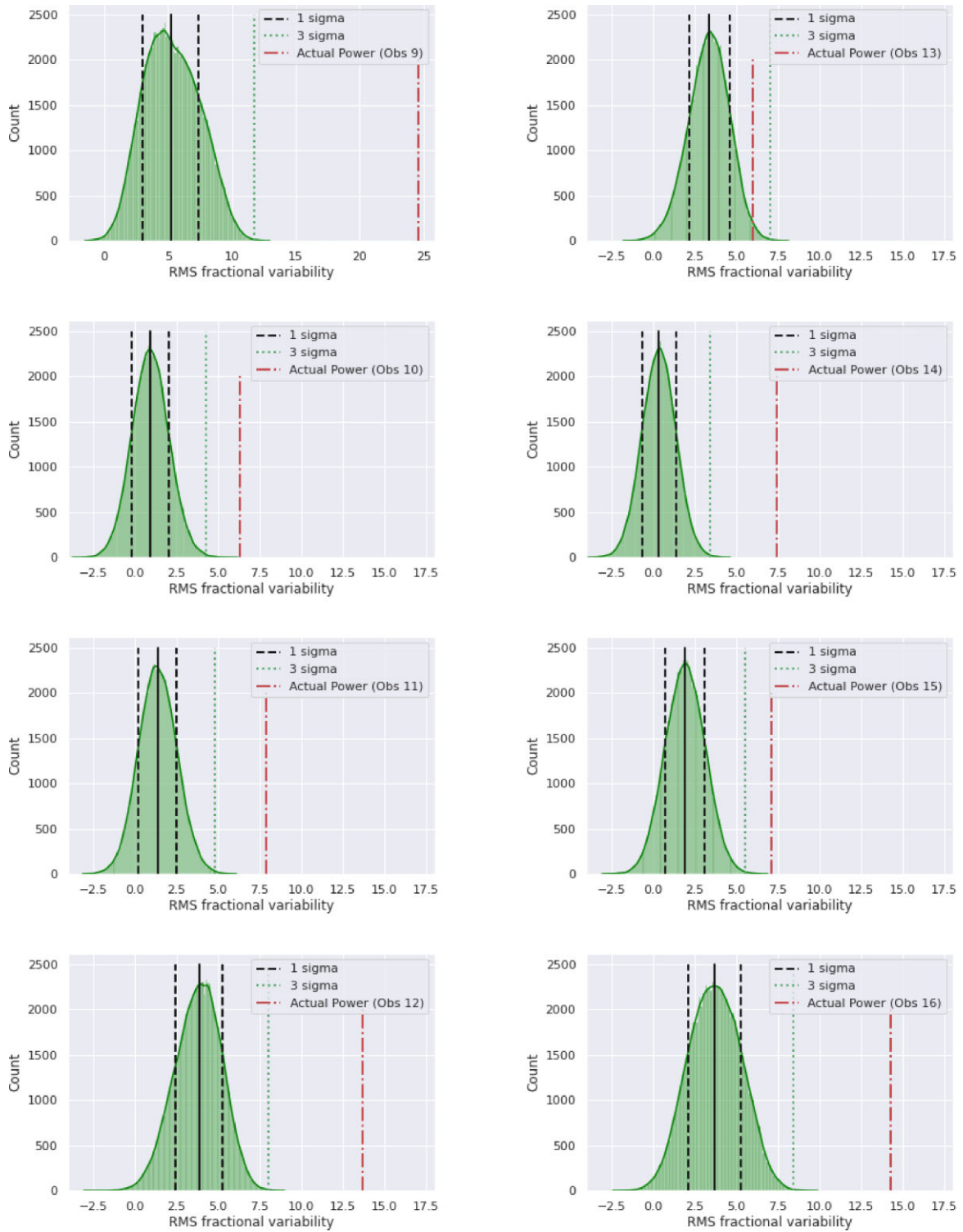


Figure A3. Results from Monte-Carlo simulations for testing significance of the QPO peak in the periodogram. For each simulation, we compared power in the frequency range of 0.001–0.009 Hz to that observed in the PDS of real light curve. We have used *Powerspectrum* of Stingray package to create power density spectra normalized to squared fractional rms. The red dashed lines represent actual power observed in real light curves.

Figure A3. *continued*

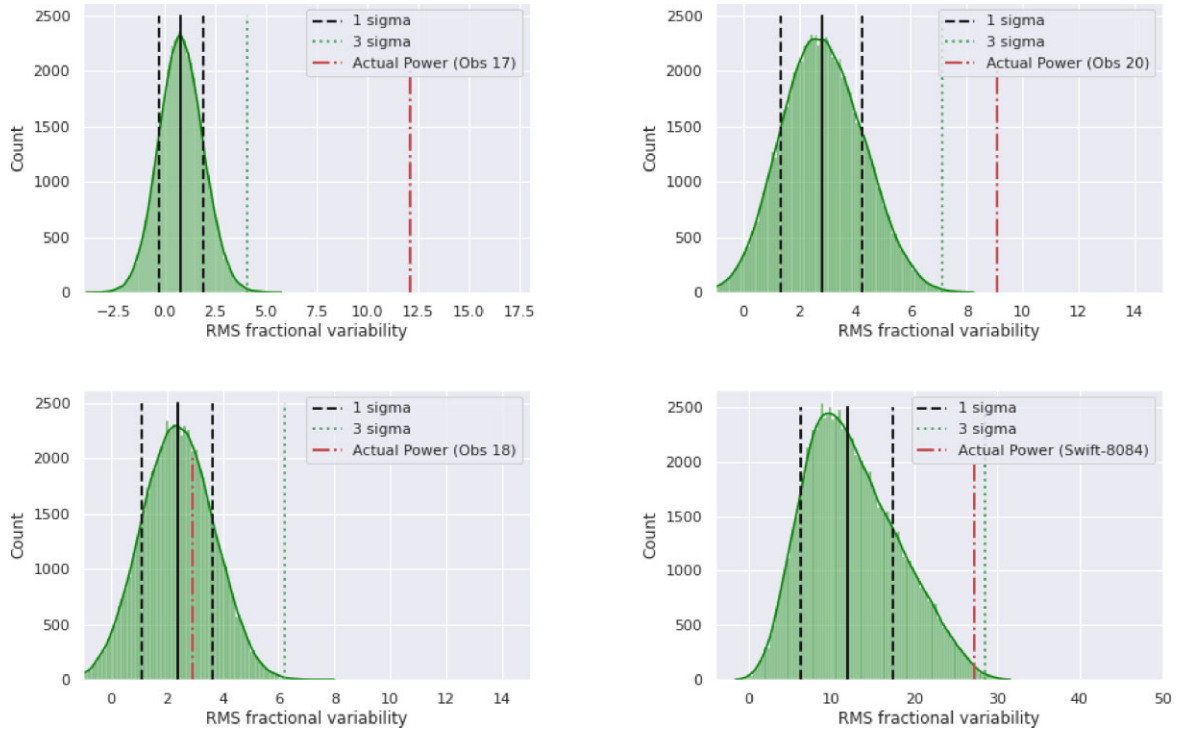


Figure A3. continued

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