Non-Abelian Domain Walls and Gravitational Waves

Bowen Fu,^a Stephen F. King,^b Luca Marsili,^c Silvia Pascoli, d,e Jessica Turner f and Ye-Ling Zhou^g

- c Instituto de Fìsica Corpuscular (IFIC), Universitat de València, Parc Cientific UV, C/ Catedratico Jose Beltran 2, E-46980 Paterna, Spain
- d DIFA, University of Bologna, via Irnerio 46, 40126 Bologna, Italy
- e INFN, Sezione di Bologna, viale Berti Pichat 6/2, 40127 Bologna, Italy
- f Institute for Particle Physics Phenomenology, Department of Physics, Durham University, Durham DH1 3LE, U.K.
- ^gSchool of Fundamental Physics and Mathematical Sciences, Hangzhou Institute for Advanced Study, UCAS, Hangzhou 310024, China

E-mail: fubowen@mail.neu.edu.cn, s.f.king@soton.ac.uk, luca.marsili@ific.uv.es, silvia.pascoli@unibo.it, jessica.turner@durham.ac.uk, zhouyeling@ucas.ac.cn

Abstract: We investigate the properties of domain walls arising from non-Abelian discrete symmetries, which we refer to as non-Abelian domain walls. We focus on S_4 , one of the most commonly used groups in lepton flavour mixing models. The spontaneous breaking of S_4 leads to distinct vacua preserving a residual Z_2 or Z_3 symmetry. Five types of domain walls are found, labelled as SI, SII, TI, TII, and TIII, respectively, the former two separating Z_2 vacua and the latter three separating Z_3 vacua. We highlight that SI, TI and TIII may be unstable for some regions of the parameter space and decay to stable domain walls. Stable domain walls can collapse and release gravitational radiation for a suitable size of explicit symmetry breaking. A symmetry-breaking scale of order 100 TeV may explain the recent discovery of nanohertz gravitational waves by PTA experiments. For the first time, we investigate the properties of these domain walls, which we obtain numerically with semi-analytical formulas applied to compute the tension and thickness across a wide range of parameter space. We estimate the resulting gravitational wave spectrum and find that, thanks to their rich vacuum structure, non-Abelian domain walls manifest in a very interesting and complex phenomenology.

KEYWORDS: non-Abelian flavour symmetry, domain wall, gravitational waves

^aKey Laboratory of Cosmology and Astrophysics (Liaoning) \mathcal{B} College of Sciences, Northeastern University, Shenyang 110819, China

 b School of Physics and Astronomy, University of Southampton, Southampton, SO17 1BJ, U.K.

Contents

1 Introduction

Since the discovery of neutrino oscillations $[1–5]$ $[1–5]$, the origin of neutrino masses and leptonic mixing remains the most concrete evidence of physics beyond the Standard Model. Compared with the CKM quark mixing, leptonic mixing is significantly larger. To explain the leptonic mixing pattern, the most studied approach is that of flavour symmetries (or family symmetries), among which non-Abelian discrete symmetries are the most popular candidate [\[6](#page-26-1)[–9\]](#page-26-2). The scale at which these flavour symmetries break is typically assumed to be very high making them largely inaccessible at current collider experiments.

Since their discovery, gravitational waves (GW) have become a powerful probe of new physics at high scales [\[10\]](#page-26-3). Moreover, recently there has been the discovery of a stochastic background of nanohertz gravitational waves by pulsar timing array experiments $[11-15]$ $[11-15]$ which may be of a cosmological origin [\[16\]](#page-26-6). Among the cosmological sources of gravitational waves, topological defects $[17-19]$ $[17-19]$ including cosmic strings $[20, 21]$ $[20, 21]$ $[20, 21]$, domain walls $[22-25]$ $[22-25]$, and hybrid defects [\[26,](#page-27-2) [27\]](#page-27-3), provide a strong connection between high-scale new physics and the gravitational wave observations [\[28](#page-27-4)[–31\]](#page-27-5). It has been highlighted that GWs from collapsing domain walls provide a good fit to the data of the recent NANOGrav Pulsar Time Array (PTA) measurement [\[32\]](#page-27-6). Several studies have explored this signal, with most focusing on Z_2 domain walls or axion domain walls $[33-37]$ $[33-37]$. General discussions on the classification of domain walls from the cyclic group Z_N (for $N > 2$) and the consequent GW signals are presented in [\[38,](#page-27-9) [39\]](#page-27-10). It was pointed out that non-Abelian discrete symmetries can also lead to domain walls and thus observable gravitational waves [\[40\]](#page-28-0). However, the analysis of domain walls of Ref. [\[40\]](#page-28-0) was qualitative as it was based on the results for the Z_2 symmetry. The domain wall properties from non-Abelian discrete symmetry breaking and the resulting GW signal are still unexplored.

In this work, we perform a detailed analysis of the domain walls from non-Abelian discrete symmetry breaking. We consider the group to be the octahedral group S_4 , which is a subgroup of $SO(3)$ and one of the most popular groups used in addressing the origin of lepton flavour mixing. By analysing the vacuum structure of a general renormalisable scalar potential respecting this symmetry, we find that there are two different types of vacua, one preserving Z_2 and the other preserving Z_3 symmetries. For each type of vacuum, there are different types of domain walls (five types in total) that depend on the topology of the vacuum manifold. We point out that some domain walls are unstable in part of the parameter space, as their energies are high enough to enable them to decay to lower-energy domain walls. We also estimate the gravitational wave produced by stable non-Abelian domain walls, in the presence of suitable explicit symmetry breaking. Due to different types of domain walls, the gravitational wave spectrum has a richer structure than the simple Z_2 case.

The paper is organised as follows. We take S_4 as a representative example and discuss the vacuum structure based on the most generic renormalisable potential in Sec. [2.](#page-2-0) Then we present the domain wall solutions for each type of vacuum in Sec. [3.](#page-4-0) In Sec. [4,](#page-13-0) we clarify the stability of the domain wall solutions obtained in Sec. [3.](#page-4-0) Finally, we estimate the gravitational wave signal produced by the non-Abelian domain walls in Sec. [5](#page-18-0) and summarise in Sec. [6.](#page-21-0)

2 Flavon potential and vacuum alignment

We begin our discussion with a brief introduction to the octahedral group S_4 . S_4 is the group of even permutations of four objects. It contains 24 elements, with three generators S, T and U satisfying $S^2 = T^3 = (ST)^3 = U^2 = (SU)^2 = (TU)^2 = (STU)^4 = 1$. There are five irreducible representations (irrep), the trivial one-dimensional (1d) irreps 1, another 1d sign representation $1'$, one two-dimensional $(2d)$ irrep 2, and two three-dimensional $(3d)$ irreps 3 and 3'. The S , T and U generators of S_4 in 3d irreps are given by

$$
T = \begin{pmatrix} 0 & 0 & 1 \\ 1 & 0 & 0 \\ 0 & 1 & 0 \end{pmatrix}, \quad S = \begin{pmatrix} 1 & 0 & 0 \\ 0 & -1 & 0 \\ 0 & 0 & -1 \end{pmatrix}, \quad U = \pm \begin{pmatrix} 1 & 0 & 0 \\ 0 & 0 & 1 \\ 0 & 1 & 0 \end{pmatrix}, \tag{2.1}
$$

where the \pm sign of U is for 3 and 3', respectively. Here, we have applied the 3d representation in the Ma-Rajasekaran basis of S and T [\[41\]](#page-28-1). Compared with the Altarelli-Feruglio basis [\[42\]](#page-28-2), which is widely used in flavour model building, it is easier for us to find the vacua of ϕ in this basis and the physics remains equivalent.

We introduce a real flavon triplet $\phi = (\phi_1, \phi_2, \phi_3)^T$. By arranging the flavon as a **3'** of S_4 , the most general renormalisable flavon potential is

$$
V(\phi) = -\frac{\mu^2}{2}I_1 + \frac{g_1}{4}I_1^2 + \frac{g_2}{2}I_2, \qquad (2.2)
$$

where

$$
I_1 = \phi_1^2 + \phi_2^2 + \phi_3^2,
$$

\n
$$
I_2 = \phi_1^2 \phi_2^2 + \phi_2^2 \phi_3^2 + \phi_3^2 \phi_1^2.
$$
\n(2.3)

To ensure the potential positive, $g_1 > 0$ and $g_2 > -4g_1$ are required. We provide details of the potential construction in Appendix [A.](#page-22-0) Note that if ϕ is arranged as a 3 of S_4 , a cubic term $\phi_1\phi_2\phi_3$ is allowed in the potential. Its effect on the vacuum structure and domain wall properties can be found in Appendix [C.](#page-24-0)

A necessary condition for the vacuum of ϕ is $\frac{\partial V(\phi)}{\partial \phi_i} = 0$. The solutions satisfying this condition can be divided into three classes according to the corresponding values of $V(\phi)$:

$$
(1) \quad v_m \in \left\{ \begin{pmatrix} 1 \\ 0 \\ 0 \end{pmatrix}, \begin{pmatrix} 0 \\ 1 \\ 0 \end{pmatrix}, \begin{pmatrix} 0 \\ 0 \\ 1 \end{pmatrix}, \begin{pmatrix} -1 \\ 0 \\ 0 \end{pmatrix}, \begin{pmatrix} 0 \\ -1 \\ 0 \end{pmatrix}, \begin{pmatrix} 0 \\ 0 \\ -1 \end{pmatrix} \right\} v, \qquad \text{for } m = 1, 2, \dots 6;
$$
\n
$$
(2) \quad u_n = \left\{ \begin{pmatrix} 1 \\ 1 \\ 1 \end{pmatrix}, \begin{pmatrix} -1 \\ 1 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ -1 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ 1 \\ -1 \end{pmatrix}, \begin{pmatrix} -1 \\ -1 \\ -1 \end{pmatrix}, \begin{pmatrix} 1 \\ -1 \\ -1 \end{pmatrix}, \begin{pmatrix} -1 \\ 1 \\ -1 \end{pmatrix}, \begin{pmatrix} -1 \\ 1 \\ -1 \end{pmatrix} \right\} u,
$$
\n
$$
(3) \quad s_l \in \left\{ \begin{pmatrix} 0 \\ 1 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ 0 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ 1 \\ 0 \end{pmatrix}, \begin{pmatrix} 0 \\ 1 \\ -1 \end{pmatrix}, \begin{pmatrix} -1 \\ 0 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ -1 \\ 0 \end{pmatrix}, \begin{pmatrix} 1 \\ -1 \\ 0 \end{pmatrix} \right\} s,
$$
\n
$$
(4) \quad \text{for } l = 1, 2, \dots, 12,
$$
\n
$$
(5) \quad \text{for } l = 1, 2, \dots, 12,
$$
\n
$$
(6) \quad \text{for } l = 1, 2, \dots, 12,
$$
\n
$$
(7) \quad \text{for } l = 1, 2, \dots, 12,
$$
\n
$$
(8) \quad \text{for } l = 1, 2, \dots, 12,
$$
\n
$$
(9) \quad \text{for } l = 1, 2, \dots, 12,
$$
\n
$$
(10) \quad \text{for } l = 1, 2, \dots, 12,
$$
\n
$$
(2.4)
$$

where

$$
v = \frac{\mu}{\sqrt{g_1}}, \qquad V(v_m) = -\frac{\mu^4}{4g_1},
$$

\n
$$
u = \frac{\mu}{\sqrt{3g_1 + 2g_2}}, \qquad V(u_n) = -\frac{3\mu^4}{4(3g_1 + 2g_2)},
$$

\n
$$
s = \frac{\mu}{\sqrt{2g_1 + g_2}}, \qquad V(s_l) = -\frac{\mu^4}{4g_1 + 2g_2}.
$$
\n(2.5)

The first class of solutions, v_m , preserves a different Z_2 symmetry. For example, $(1,0,0)T$ is invariant under a Z_2 transformation generated by S, and $(0, 1, 0)T$ is invariant under a Z_2 transformation generated by TST2. Similarly, each u_n in the second class of solutions preserves a different Z_3 symmetry, e.g., $(1,1,1)^T$ invariant in a Z_3 generated by T and $(-1, 1, 1)^T$ is invariant under a Z_3 transformation generated by STS. For convenience, v_m and u_m will also be denoted as Z_2 -preserving and Z_3 -preserving vacua in the remainder of the paper, respectively.

To guarantee that some of the above solutions are vacua, $V(\phi)$ must be a local minimum at these solutions, corresponding to the requirement of a positive-definite second derivative of $V(\phi)$. More specifically, the matrix M_{ϕ}^2 defined in the following should be positive-definite at these solutions:

$$
(M_{\phi}^2)_{ij} = \frac{\partial^2 V(\phi)}{\partial \phi_i \partial \phi_j} \Big|_{\langle \phi \rangle} . \tag{2.6}
$$

In general, M_{ϕ}^2 is a 3×3 real symmetric matrix which can be diagonalised through $W^T M_{\phi}^2 W = \text{diag}\{m_1^2, m_2^2, m_3^2\}$, with m_i^2 the eigenvalues of M_{ϕ}^2 . In the above three classes of solutions, we find:

(1)
$$
m_1^2 = 2g_1v^2
$$
, $m_2^2 = m_3^2 = g_2v^2$;
\n(2) $m_1^2 = 2(3g_1 + 2g_2)u^2$, $m_2^2 = m_3^2 = -2g_2u^2$;
\n(3) $m_1^2 = 2(2g_1 + g_2)s^2$, $m_2^2 = -2m_3^2 = -2g_2s^2$, (2.7)

at v_m , u_n and s_l , respectively. Although m_1^2 is always positive in all solutions, m_2^2 or m_3^2 may be positive or negative, depending on the sign of the coefficient g_2 . The third solution class is less interesting since m_2^2 and m_3^2 always take opposite signs at s_l . Thus, s_l is always an unstable saddle point of $V(\phi)$ and cannot be a vacuum. For the first two classes of solutions, if $g_2 > 0$, m_2^2 and m_3^2 are positive at v_m and negative at u_n . Therefore, u_n is an unstable saddle point, and $V(\phi)$ can only take a local minimum value (thus, the global minimum value) at v_m . As a result, v_m is the only choice of the ϕ vacuum expectation value (VEV). On the contrary, if $g_2 < 0$, u_n is the only choice of the ϕ VEV. All VEVs for $g_2 > 0$ and $g_2 < 0$ have been geometrically shown in Figs. [1](#page-5-0) and [2,](#page-5-1) respectively, in the three-dimensional field space up to a mass-dimension normalisation factor.

3 Domain wall solutions

A domain wall (DW) is characterised by a classical and stable one-dimensional scalar field profile in three-dimensional space. The equation satisfied by the scalar is obtained from the scalar Equation of Motion (EOM), i.e., the Euler-Lagrange equation of the scalar, which is given by

$$
\left(\frac{d^2}{dt^2} - \nabla^2\right)\phi_i + \frac{\partial V(\phi)}{\partial \phi_i} = 0\,,\tag{3.1}
$$

where i denotes the components of the scalar ϕ . A DW refers to a stable solution of the scalar field and thus is time-independent. Defining the direction perpendicular to the DW

Figure 1: Z_2 -preserving vacua of S_4 in the three-dimensional flavon field space and two topologically different types of DWs. A straight line connecting to two vacua should not be understood as the path in the field space, but just represent the topology of the DW separating these two vacua. SI and SII DWs are given in the left and right panels, respectively.

Figure 2: Z_3 -preserving vacua of S_4 and three topologically different types of DWs. TI, TII and TIII DWs are given in the left, middle and right panels.

surface as the z-axis, the profile variation is only z-dependent. Thus the equation of motion for the scalar field is simplified to

$$
\frac{d^2\phi_i(z)}{dz^2} = \frac{\partial V(\phi)}{\partial \phi_i}.
$$
\n(3.2)

Given a theory where a discrete symmetry is spontaneously broken, a series of vacua (e.g., v_1, v_2, \ldots as obtained above) can be generated. These vacua are disconnected from each other in the field space. By fixing two disconnected vacua v_i and v_j $(i \neq j)$ along the z -axis, a DW can form between them. If the theory provides more than two vacua, there would be multiple choices of pairs of vacua, and different DWs can be generated.

In this section, we derive all DW solutions between any degenerate vacua discussed in Sec. [2.](#page-2-0) Following Eq. [\(3.2\)](#page-5-2), we refer to the spatial coordinate perpendicular to the wall's surface as z. Two different vacua are fixed at $z = \pm \infty$ as two different boundary conditions.

The solution $\phi_i(z)$ varies continuously from one vacuum, (e.g., v_1 at $z = +\infty$) to another vacuum (e.g., v_2 at $z = -\infty$). At any z coordinate between two boundaries, $\phi_i(z)$ does not take a vacuum value, and thus, there must be some energy stored in the scalar field. The total energy density stored in all scalars compared with the vacuum is a sum of the gradient energy and the potential

$$
\rho(z) = \sum_{i} \frac{1}{2} [\phi_i'(z)]^2 + \Delta V(\phi(z)),
$$
\n(3.3)

where the potential is shifted such that the vacuum energy is zero, i.e., $\Delta V(\phi(z))$ = $V(\phi(z)) - V(\phi(\overline{z} \to \infty))$. Integrating along z gives the tension of the wall, i.e., the energy per unit area of the wall,

$$
\sigma = \int_{-\infty}^{+\infty} dz \,\rho(z). \tag{3.4}
$$

The DW has a thickness δ representing the length scale over which most of the energy is localised. Following the discussion in [\[39\]](#page-27-10), we define the thickness δ as^{[1](#page-6-1)}

$$
\delta = \int_{z_0 - \delta/2}^{z_0 + \delta/2} dz \, \rho(z) \approx 64.38\% \times \sigma \,. \tag{3.6}
$$

This definition applies to the case where $\rho(z)$ has only a single peak centred at $z = z_0$, which refers to the core of the DW. Without loss of generality, we set $z_0 = 0$ when solving scalar profiles perpendicular to the DW.

3.1 Domain walls separating Z_2 -preserving vacua: S-type

We first discuss DWs formed in separating Z_2 -preserving vacua v_m . It is straightforward to prove that each vacua is invariant under the transformation of S or its conjugate transformations. Due to this, we denote these kinds of DWs in general as S-type DWs. It is convenient to normalise the fields and the z-coordinate as

$$
\overline{\phi}_i = \frac{\sqrt{g_1}}{\mu} \phi_i = \frac{\phi_i}{v}, \quad \overline{z} = \mu z, \qquad (3.7)
$$

and define $\beta = g_2/g_1$. Then, the EOM simplifies to

$$
\frac{d^2\overline{\phi}_i(\overline{z})}{d\overline{z}^2} = \overline{\phi}_i[-1 + \overline{\phi}_1^2 + \overline{\phi}_2^2 + \overline{\phi}_3^2 + \beta(\overline{\phi}_j^2 + \overline{\phi}_k^2)],\tag{3.8}
$$

¹In the special Z_2 DW case, the width δ is defined as the factor appearing in the hyperbolic tangent function of the profile $\propto \tanh(z/\delta)$ [\[17\]](#page-26-7). This leads to the condition

$$
\int_{-\delta/2}^{\delta/2} dz \,\rho(z) = \frac{1}{2} \tanh \frac{1}{2} \left(3 - \tanh^2 \frac{1}{2} \right) \times \sigma \approx 64.38\% \times \sigma ,\qquad (3.5)
$$

where σ is the wall tension. We apply this condition as a generalised definition of the wall thickness for walls despite a hyperbolic profile.

which includes only one free parameter β . The DW tension is given by $\sigma = \mu v^2 \overline{\sigma}$ with dimensionless tension

$$
\overline{\sigma} = \int_{-\infty}^{+\infty} d\overline{z} \left\{ \frac{1}{2} \left[\overline{\phi}_1^2(\overline{z}) + \overline{\phi}_2^2(\overline{z}) + \overline{\phi}_3^2(\overline{z}) \right] + \Delta \overline{V}(\overline{\phi}(\overline{z})) \right\},\tag{3.9}
$$

where $\Delta \overline{V} = \Delta V / (\mu^2 v^2)$. The thickness of the DW can be parameterised into a dimensionless parameter as $\overline{\delta} = \mu \delta$. The S-type DWs are further classified into two topologically different types of DWs, SI and SII. The topology of SI and SII DWs are shown in Fig. [1](#page-5-0) and the profiles of the scalars in the field space for typical values of β are shown in Fig. [3.](#page-8-0)

SI domain walls

The first type of DWs includes those separating v_1 and v_4 , v_2 and v_5 , and v_3 and v_6 . We follow the convention in [\[38\]](#page-27-9) and denote them as

$$
\boxed{v_1 \mid v_4}, \quad \boxed{v_2 \mid v_5}, \quad \boxed{v_3 \mid v_6}, \tag{3.10}
$$

respectively. $|v_i| v_j = |v_j| v_i |$ (for $i \neq j$) because exchanging v_i and v_j on the two boundaries gives a physically equivalent DW.

To make a quantitative analysis, we take $\boxed{v_1 \mid v_4}$ as an example. In this case, we find a DW solution that can be solved explicitly. As shown in the left panel of Fig. [3](#page-8-0) and more clearly as a projection in the left panel of Fig. [4,](#page-8-0) the solution is a straight line between v_1 and v_4 in the field space. When $\overline{\phi}_2$ and $\overline{\phi}_3$ are both fixed to be 0, the EOM of $\overline{\phi}_1$ and the boundary conditions are given by

$$
\overline{\phi}_1''(\overline{z}) = \overline{\phi}_1[-1 + \overline{\phi}_1^2],
$$

\n
$$
\overline{\phi}_1|_{\overline{z} \to +\infty} = +1, \quad \overline{\phi}_1|_{\overline{z} \to -\infty} = -1.
$$
\n(3.11)

This is simply the DW formed from the spontaneous symmetry breaking of Z_2 , which has been widely discussed in the literature. We refer to the latter case as the Z_2 DW. The scalar field follows the profile

$$
\overline{\phi}_1(\overline{z}) = \tanh\left(\frac{\overline{z}}{\sqrt{2}}\right),\tag{3.12}
$$

which is also shown in the left panel of Fig. [5.](#page-8-0) The denominator in the Hyperbolic function in Eq. (3.12) denotes the dimensionless thickness of the wall, i.e. $\delta = \sqrt{2}$. The dimensionless tension can be straightforwardly calculated:

$$
\overline{\sigma}_{\rm SI} = \int_{-\infty}^{+\infty} d\overline{z} \left\{ \frac{1}{2} \overline{\phi}_1^2(\overline{z}) + \frac{1}{4} \left[\overline{\phi}_1^2(\overline{z}) - 1 \right]^2 \right\} = \frac{2\sqrt{2}}{3} . \tag{3.13}
$$

With the dimensionless thickness and tension, one can express the physical tension and thickness in terms of physical parameters as

$$
\sigma_{\rm SI} = \frac{2}{3}\sqrt{2g_1}v^3 = \frac{2}{3}m_1v^2 \,, \quad \delta_{\rm SI} = \sqrt{\frac{2}{g_1v^2}} = \frac{2}{m_1} \,, \tag{3.14}
$$

where m_1 is the scalar mass given in Eq. [\(2.7\)](#page-4-1). The result is independent of β and consistent with the classical one for Z_2 DW [\[17\]](#page-26-7).

Figure 3: Profiles of scalars in the field space for SI DW $\boxed{v_1 \mid v_4}$ (left panel) and SII DW $\overline{v_1 \mid v_2}$ (right panel). The green, orange and blue curves refer to $\beta = 0.1, 2, 10$, respectively.

Figure 4: Projection of the paths in Fig. [3](#page-8-0) in $\overline{\phi}_1 - \overline{\phi}_2$ plane at $\overline{\phi}_3 = 0$.

Figure 5: Profiles of scalars along the spatial coordinate \overline{z} for SI DW $|v_1|$ $|v_4|$ (left) and SII DW $\boxed{v_1 \mid v_2}$ (right). $\overline{\phi}_2 = \overline{\phi}_3 = 0$ along the path for SI DW $\boxed{v_1 \mid v_4}$, while $\overline{\phi}_3 = 0$ along the path for SII DW $\boxed{v_1 \mid v_2}$. The green, orange and red curves refer to $\beta = 0.1, 2, 10$, respectively.

SII domain walls

The vacua combinations of the second type DWs are all listed below,

$$
[v_1] v_2], [v_1] v_3], [v_1] v_5], [v_1] v_6], [v_4] v_2], [v_4] v_3], [v_4] v_5], [v_4] v_6], (3.15)
$$

which are presented visually in the right panel of Fig. [\(1\)](#page-5-0). We take $|v_1|$ $|v_2|$ as an example to solve the scalar profiles perpendicular to the DW. At both v_1 and $\overline{v_2, \phi_3} = 0$ is satisfied, and the reflection symmetry $\phi_3 \leftrightarrow -\phi_3$ is satisfied on the boundaries. We remind the reader that the EOM in Eq. [\(3.8\)](#page-6-2) also satisfies this reflection symmetry. Thus, we expect there is the DW solution obtained by fixing $\phi_3 \rightarrow 0$. Then, the EOM and boundary conditions are simplified to be

$$
\overline{\phi}_{1}^{\prime\prime}(\overline{z}) = \overline{\phi}_{1}[-1 + \overline{\phi}_{1}^{2} + \overline{\phi}_{2}^{2} + \beta \overline{\phi}_{2}^{2}],\n\overline{\phi}_{2}^{\prime\prime}(\overline{z}) = \overline{\phi}_{2}[-1 + \overline{\phi}_{1}^{2} + \overline{\phi}_{2}^{2} + \beta \overline{\phi}_{1}^{2}],\n\begin{pmatrix} \overline{\phi}_{1} \\ \overline{\phi}_{2} \end{pmatrix}\Big|_{\overline{z} \to +\infty} = \begin{pmatrix} 1 \\ 0 \end{pmatrix}, \quad \begin{pmatrix} \overline{\phi}_{1} \\ \overline{\phi}_{2} \end{pmatrix}\Big|_{\overline{z} \to -\infty} = \begin{pmatrix} 0 \\ 1 \end{pmatrix}.
$$
\n(3.16)

The tension simplifies to

$$
\overline{\sigma}_{\text{SII}} = \int_{-\infty}^{+\infty} d\overline{z} \left\{ \frac{1}{2} \left[\overline{\phi}_1^{\prime 2}(\overline{z}) + \overline{\phi}_2^{\prime 2}(\overline{z}) \right] + \frac{1}{4} \left[\overline{\phi}_1^2(\overline{z}) + \overline{\phi}_2^2(\overline{z}) - 1 \right]^2 + \frac{\beta}{2} \overline{\phi}_1^2(\overline{z}) \overline{\phi}_2^2(\overline{z}) \right\} . (3.17)
$$

With the above parametrisation, scalar profiles and the dimensionless DW tension depend only on a single parameter, $\beta \equiv g_2/g_1$. Given any positive value of β , we can calculate the scalar profiles in the domain wall solution and the corresponding dimensionless tension and thickness numerically. The path of the scalars in the field space is shown in the right panel of Fig. [3](#page-8-0) and as a projection in the right panel of Fig. [4.](#page-8-0) The domain wall solution is a straight line between v_1 and v_2 in the field space when $\beta = 2$. We note that an enhanced symmetry for $\beta = 2$ occurs since the potential of Eq. [\(2.2\)](#page-3-0) becomes only a function of a single coupling. We expand this point further in Appendix [B.](#page-23-0) For $\beta < 2$, ϕ_2 is a concave function of ϕ_1 and the path curves away from the origin; for $\beta > 2$, in contrast, ϕ_2 is a convex function of ϕ_1 and the path moves toward the origin. The profiles of ϕ_2 along \overline{z} for $\beta = 0.1, 2, 10$ are shown in the right panel of Fig. [5,](#page-8-0) illustrating an increase in the dimensionless thickness as β decreases.

To study the properties of the domain walls in detail, we perform a scan on β from 10^{-3} to 10^3 and obtain the dependence of dimensionless tension and thickness upon β , which is shown as the blue curves in Fig. [6.](#page-10-1) While the dimensionless tension of SI is independent of β as computed in Eq. [\(3.13\)](#page-7-1), that of SII increases as β becomes larger, as shown in the left panel of Fig. [6.](#page-10-1) Similarly, the dimensionless thickness of SI remains a constant for all β, but that of SII shows a negative correlation first and then a positive one with β. The result can be semi-analytically expressed as

$$
\overline{\sigma}_{\rm SII}(\beta) \approx \frac{2\sqrt{2}}{3} \frac{1}{1 + 1.875\beta^{-1/2}} \left[1 + 0.5 \frac{\beta^{1/2}}{1 + 2\beta} \right],
$$

\n
$$
\overline{\delta}_{\rm SII}(\beta) \approx \frac{1.52}{\beta^{0.50}} + \frac{0.66\beta^{0.79}}{\beta^{0.70} + 1}.
$$
\n(3.18)

Figure 6: Dimensionless tension (left) and thickness (right) of DWs SI and SII as functions **Figure 6.** Dimensionless tension (left) and thickness (Figure of DWs
of β. For DW SI, $\overline{\sigma}_{SI} = 2\sqrt{2}/3$ and $\overline{\delta}_{SI} = \sqrt{2}$ are independent of β.

with relative error less than 6% for $10^{-4} < \beta < 10^3$.

3.2 Domain walls separating Z_3 -preserving vacua: T-type

The DWs separating Z_3 -preserving vacua have three topologically different types, which are labelled as TI, TII and TIII, respectively. The label "T" refers to the vacua preserving residual Z_3 symmetries either generated by T or its conjugate elements. For convenience, we perform the following normalisation for the fields and the coordinate

$$
\tilde{\phi}_i = \frac{\sqrt{3g_1 + 2g_2}}{\mu} \phi_i = \frac{\phi_i}{u}, \quad \bar{z} = \mu z.
$$
\n(3.19)

Under such a normalisation, the EOM can be simplified into

$$
\frac{d^2 \tilde{\phi}_i(\bar{z})}{d\bar{z}^2} = \tilde{\phi}_i \Big[-1 + \frac{1}{3+2\beta} (\tilde{\phi}_1^2 + \tilde{\phi}_2^2 + \tilde{\phi}_3^2) + \frac{\beta}{3+2\beta} (\tilde{\phi}_j^2 + \tilde{\phi}_k^2) \Big].
$$
 (3.20)

The DW tension is given by $\sigma = \mu u^2 \tilde{\sigma}$ with

$$
\tilde{\sigma} = \int_{-\infty}^{+\infty} d\overline{z} \left\{ \frac{1}{2} \left[\tilde{\phi}'_1^2(\overline{z}) + \tilde{\phi}'_2^2(\overline{z}) + \tilde{\phi}'_3^2(\overline{z}) \right] + \Delta \tilde{V}(\tilde{\phi}(\overline{z})) \right\},\tag{3.21}
$$

where $\Delta \tilde{V} = \Delta V / (\mu^2 u^2)$. The DW thickness is given by $\tilde{\delta} = \mu \delta$. In particular, the EOMs become decoupled for each $\tilde{\phi}_i$ when $\beta = -1$, i.e. $d^2 \tilde{\phi}_i(\overline{z})/d\overline{z}^2 = \tilde{\phi}_i[-1 + \tilde{\phi}_i^2/(3 + 2\beta)]$. As a result, there is no correlation between the field when $\beta = -1$ and the paths in field space are straight lines. The topology of TI, TII and TIII DWs are shown in Fig. [2.](#page-5-1) Typical examples of scalar profiles in the field space are shown in Fig. [7](#page-11-0) with $\beta = -0.1, -1, -1.4$, with projections in Fig. [8.](#page-11-0) Those profiles along the spatial coordinate are shown in Fig. [9.](#page-11-0) By varying the value of β , we obtain the β -dependent tensions and thicknesses of each type of DW, as shown in Fig. [10.](#page-14-1) We discuss each type of DWs in more detail in the following sections.

Figure 7: Profiles of scalars in the field space for TI DW $|u_1| |u_5|$ (left panel), TII DW $|u_1|$ $|u_4|$ (middle panel) and TIII DW $|u_1|$ $u_8|$ (right panel). The green, orange and red curves refer to $-\beta = 0.1, 1, 1.4$, respectively.

Figure 8: Projection of the paths in Fig. [7](#page-11-0) in $\tilde{\phi}_1 - \tilde{\phi}_3$ with $\tilde{\phi}_1 = \tilde{\phi}_2$.

Figure 9: Profiles of scalars along the spatial coordinate \overline{z} for TI DW $|u_1|$ $u_5|$ (left panel), TII DW $\boxed{u_1 \mid u_2}$ (middle panel) and TIII DW $\boxed{u_1 \mid u_8}$ (right panel). There are $\tilde{\phi}_1 = \tilde{\phi}_2 =$ $\tilde{\phi}_3$, $\tilde{\phi}_2 = \overline{\tilde{\phi}_3}$, and $\tilde{\phi}_1 = \tilde{\phi}_2$ along the path for the three cases, respectively. The green, orange and red curves refer to $-\beta = 0.1, 1, 1.4$, respectively.

TI domain walls

Similar to the SI domain walls, the TI domain walls are also equivalent to the Z_2 domain walls, including

$$
\boxed{u_1 \mid u_5}, \quad \boxed{u_2 \mid u_6}, \quad \boxed{u_3 \mid u_7}, \quad \boxed{u_4 \mid u_8}.
$$
 (3.22)

Again, we choose $|u_1| u_5$ as an example for detailed discussion. As visualised in the left panel of Fig. [7](#page-11-0) and Fig. [8,](#page-11-0) the path of domain wall $|u_1| \, |u_5|$ in the field space is a straight line between u_1 and u_5 . The EOMs and the boundary conditions for scalar fields across the DW are given by

$$
\tilde{\phi}_1''(\overline{z}) = \tilde{\phi}_1[-1 + \tilde{\phi}_1^2],
$$

\n
$$
\tilde{\phi}_1|_{\overline{z} \to +\infty} = 1, \quad \tilde{\phi}_1|_{\overline{z} \to -\infty} = -1,
$$
\n(3.23)

with $\tilde{\phi}_2 = \tilde{\phi}_3 = \tilde{\phi}_1$ satisfied. The EOM for $\tilde{\phi}_1$ and the boundary conditions are equivalent to Eq. [\(3.11\)](#page-7-2) and hence the profile of $\tilde{\phi}_1$ is in the same form as Eq. [\(3.12\)](#page-7-0) and the dimensionless Eq. (5.11) and hence the prome of φ_1 is in the same form as Eq. (5.12) and the dimensionless thickness of the wall is $\sqrt{2}$. The dimensionless tension includes contribution from all three fields $\tilde{\phi}_1, \tilde{\phi}_2, \tilde{\phi}_3$, with $2\sqrt{2}/3$ from each field equally. Thus the total energy stored in the DW is given by 3×2 $\sqrt{2}/3 = 2\sqrt{2}$. We summarise the dimensionless tension and thickness as follows,

$$
\tilde{\sigma}_{\rm TI} = 2\sqrt{2}, \quad \tilde{\delta}_{\rm TI} = \sqrt{2}.
$$
\n(3.24)

TII domain walls

The second type domain walls include

$$
\boxed{u_1 \mid u_2}, \quad \boxed{u_1 \mid u_3}, \quad \boxed{u_1 \mid u_4}, \quad \boxed{u_2 \mid u_7}, \quad \boxed{u_2 \mid u_8}, \quad \boxed{u_3 \mid u_6},
$$
\n
$$
\boxed{u_3 \mid u_8}, \quad \boxed{u_4 \mid u_6}, \quad \boxed{u_4 \mid u_7}, \quad \boxed{u_5 \mid u_6}, \quad \boxed{u_5 \mid u_7}, \quad \boxed{u_5 \mid u_8}.
$$
\n(3.25)

We take the DW $\boxed{u_1 \mid u_4}$ as an example. In both u_1 and u_4 , $\tilde{\phi}_1$ and $\tilde{\phi}_2$ take the same value. Since there is no source to break the permutation symmetry between $\tilde{\phi}_1$ and $\tilde{\phi}_2$, we expect that $\tilde{\phi}_1 = \tilde{\phi}_2$ for \bar{z} from $-\infty$ to ∞ . In this case, EOM and boundary conditions for scalar profiles are simplified to

$$
\tilde{\phi}_1''(\overline{z}) = \tilde{\phi}_1 \Big[-1 + \frac{2 + 2\beta}{3 + 2\beta} \tilde{\phi}_1^2 + \frac{1}{3 + 2\beta} \tilde{\phi}_3^2 \Big],
$$

\n
$$
\tilde{\phi}_3''(\overline{z}) = \tilde{\phi}_3 \Big[-1 + \frac{2 + 2\beta}{3 + 2\beta} \tilde{\phi}_1^2 + \frac{1}{3 + 2\beta} \tilde{\phi}_3^2 \Big],
$$

\n
$$
\left(\tilde{\phi}_1 \atop \tilde{\phi}_3 \right) \Big|_{\overline{z} \to +\infty} = \begin{pmatrix} 1 \\ 1 \end{pmatrix}, \quad \left(\tilde{\phi}_1 \atop \tilde{\phi}_3 \end{pmatrix} \Big|_{\overline{z} \to -\infty} = \begin{pmatrix} 1 \\ -1 \end{pmatrix}.
$$
\n(3.26)

By fixing $-\beta = 0.1, 1, 1.4$ respectively, we obtain solutions for scalar profiles in the field space in the middle panel of Fig. [7](#page-11-0) and Fig. [8,](#page-11-0) and those along the \overline{z} coordinate in the middle panel of Fig. [9.](#page-11-0) As mentioned, the path of the scalar profiles in the field space is a straight line between u_1 and u_4 when $\beta = -1$, which corresponds to the yellow line. The distance between the path and the origin positively depends on β .

Following the numerical results of β -dependent dimensionless tension and thickness as the yellow lines in Fig. [10,](#page-14-1) we find that the dimensionless tension decreases as β decreases while the dimensionless thickness increases. The following analytical formulas can fit the numerical results very well with relative error less than 3%,

$$
\tilde{\sigma}_{\text{TH}}(\beta) = \frac{0.77(-\beta)^{0.5}}{(1.5+\beta)^{0.25}}, \qquad \tilde{\delta}_{\text{TH}}(\beta) = \frac{2.12(-\beta)^{-0.5}}{0.48 + (-\beta)^{1.35}}.
$$
\n(3.27)

TIII domain walls

The third type of domain walls include

u¹ u⁸ , u¹ u⁶ , u¹ u⁷ , u⁶ u⁷ , u⁶ u⁸ , u⁷ u⁸ , u² u³ , u² u⁴ , u² u⁵ , u³ u⁴ , u³ u⁵ , u⁴ u⁵ . (3.28)

We take $\boxed{u_1 \mid u_8}$ as an example to solve for the DW solution. Similar to the TII case, there is no source to break the permutation symmetry between $\tilde{\phi}_1$ and $\tilde{\phi}_2$ and thus, we expect $\tilde{\phi}_1 = \tilde{\phi}_2$ for any \bar{z} . Then, the EOM and boundary conditions for $\tilde{\phi}_1$ and $\tilde{\phi}_3$ are given by:

$$
\tilde{\phi}_1''(\overline{z}) = \tilde{\phi}_1 \Big[-1 + \frac{2+\beta}{3+2\beta} \tilde{\phi}_1^2 + \frac{1+\beta}{3+2\beta} \tilde{\phi}_3^2 \Big],
$$

\n
$$
\tilde{\phi}_3''(\overline{z}) = \tilde{\phi}_3 \Big[-1 + \frac{2+\beta}{3+2\beta} \tilde{\phi}_1^2 + \frac{1+\beta}{3+2\beta} \tilde{\phi}_3^2 \Big],
$$

\n
$$
\begin{pmatrix} \tilde{\phi}_1 \\ \tilde{\phi}_3 \end{pmatrix} \Big|_{\overline{z} \to +\infty} = \begin{pmatrix} 1 \\ 1 \end{pmatrix}, \quad \begin{pmatrix} \tilde{\phi}_1 \\ \tilde{\phi}_3 \end{pmatrix} \Big|_{\overline{z} \to -\infty} = \begin{pmatrix} -1 \\ 1 \end{pmatrix}.
$$
\n(3.29)

The scalar profiles in the field space obtained by numerical computation are shown in the right panel of Fig. [7](#page-11-0) and Fig. [8.](#page-11-0) Again, the path is straight between u_1 and u_8 for $\beta = -1$ and moves away from the origin as β increases. The scalar profiles along the \overline{z} coordinate in the right panel of Fig. [9.](#page-11-0) The dependence of dimensionless tension and thickness on the value of β in the TIII case is qualitatively similar to the TII case, which is shown by the green lines in Fig. [10.](#page-14-1) The following analytical formulas fit the numerical results with relative errors within 3%,

$$
\tilde{\sigma}_{\text{TIII}}(\beta) = \frac{2.06(-\beta)^{0.5}}{1 + 0.09(-\beta)^{0.6}},
$$

$$
\tilde{\delta}_{\text{TIII}}(\beta) = \frac{1.62(-\beta)^{-0.5}}{1 + 0.125(-\beta)^{0.6}}.
$$
\n(3.30)

4 Stability of the domain wall solutions

Different DWs can be generated through the usual Kibble mechanism during phase transitions in the early universe. However, not all of them are necessarily stable. Given two vacua as boundary conditions and the generic static differential equations of scalars given

Figure 10: Dimensionless tension (left) and thickness (right) of DWs TI, TII, and TIII as **Figure 10.** Dimensionless tension (left) and thickness (right) of DWS 11, 1 functions of β . For DW TI, $\tilde{\sigma}_{TI} = 2\sqrt{2}$ and $\tilde{\delta}_{TI} = \sqrt{2}$ are independent of β .

in Eq. [\(3.2\)](#page-5-2), there may be other DW solutions beyond those in Sec. [3.](#page-4-0) These solutions, as shown below, could consist of several other DWs. If the energy stored in a DW solution, i.e., the tension, is not the lowest compared to other DW solutions, it can decay. These unstable solutions with higher tensions can evolve into DWs with the lowest tension, i.e. the ground state. In this section, we determine the stability of S-type and T-type DWs separately.

4.1 S-type domain walls

In the last section, we summarised that there are two obviously distinct types of DWs, SI and SII, for the DWs separating two Z_2 -preserving vacua. To check the stability of the SI DW, we focus on the vacuum transition from v_1 to v_4 . Given the EOM in Eq. [\(3.2\)](#page-5-2) and boundaries v_1 and v_4 at $z \to \pm \infty$, we will check if the SI DW solution, in which v_1 and v_4 are directly connected by a straight line in the field space, is the only solution. We have found several solutions, as shown in Fig. [11:](#page-15-0)

- 1. The most obvious solution is the SI DW solution in Sec. [3.1,](#page-7-3) e.g., $|v_1| v_4$, referring to the straight path connecting v_1 and v_4 in the field space. We remind the reader that this solution is the Z_2 domain wall solution with the tension independent of the Lagrangian parameters g_1 and g_2 .
- 2. The second solution, shown as the purple curve in Fig. [11,](#page-15-0) is a path indirectly between v_1 and v_4 with the third vacuum, e.g. v_2 , as transit point. This path is a combination of two SII DWs v_1 v_2 and v_2 v_4 and we can denote the combination as v_1 v_2 v_4 for simplicity. Note that the transit point is an arbitrary choice for any vacuum of v_2, v_3, v_5 and v_6 and such paths in the field space constitute a class of paths which minimise the energy density of the domain wall for $\beta < 2$.
- 3. A third solution is the straight line connection between v_1 and v_4 without crossing any other vacuum. However, the path in the field space is different from that of the SI solution but lies in the 45° (or 135° , 225° , 315°) direction when projected to the $\overline{\phi}_2 - \overline{\phi}_3$ plane, shown as orange curve in the left panel of Fig. [11.](#page-15-0)

Figure 11: Left: Three different types of possible domain solutions from v_1 to v_4 when $\beta = 0.3$. Right: The corresponding tension of the three different types of solutions for different β values.

Figure 12: The energy density of the DW along the z direction for the three different solutions in Fig. [11](#page-15-0) with the same conventions.

Tensions for different solutions as functions of the free parameter β are presented in the right panel of Fig. [11.](#page-15-0) The tension of the first solution, which is independent of β , is given by a constant $\overline{\sigma}_{SI} = \frac{2\sqrt{2}}{3}$ $\frac{\sqrt{2}}{3}$. Tensions obtained in the other solutions are β -dependent. For β < 2, tensions in these solutions are smaller than in the first solution, and the tension of the second solution is the smallest. In this case, $\beta < 2$, $\bar{\sigma}_{SI} > 2\bar{\sigma}_{SI}$ and thus the SI DW is unstable. For $\beta > 2$, the tension of the second solution is the largest, and the third solution gives tension equal to that of the first solution. In this case, $\beta > 2$, $\bar{\sigma}_{SI} < 2\bar{\sigma}_{SI}$ and thus the SI DW is stable.

When $\beta = 2$, the tensions of all three solutions are equal, which can be proved analytically. Notice that the second solution is equivalent to two SII DWs, i.e. $|v_1| v_2$ and v_2 v_4 . Proving the tension of the second solution to be 2 √ 2/3 is equivalent to proving that the tension of a SII DW at $\beta = 2$ is $\sqrt{2}/3$, i.e., $\overline{\sigma}(\beta = 2) = \sqrt{2}/3$. With $\beta = 2$, it is

convenient to make the parameterisation $\phi_{\pm} = \phi_1 \pm \phi_2$ so that Eq. [\(3.16\)](#page-9-0) becomes

$$
\overline{\phi}_{\pm}''(\overline{z}) = \overline{\phi}_{\pm}[-1 + \overline{\phi}_{\pm}^{2}], \n\overline{\phi}_{\pm}|_{\overline{z} \to +\infty} = 1, \quad \overline{\phi}_{\pm}|_{\overline{z} \to -\infty} = \pm 1.
$$
\n(4.1)

 $\overline{\phi}_+$ in this case has a constant solution $\overline{\phi}_+(\overline{z}) = 1$ and $\overline{\phi}_-$ obeys the Z_2 DW solution, $\overline{\phi}_{-}(\overline{z}) = \tanh(\frac{\overline{z}}{\sqrt{2}})$. Substituting this solution into Eq. [\(3.17\)](#page-9-1), the tension is

$$
\overline{\sigma}_{\rm SII}(\beta = 2) = \frac{\sqrt{2}}{3} = \frac{1}{2}\overline{\sigma}_{\rm SI}.
$$
\n(4.2)

The SII DW does not have the stability problem as it always gives the smallest tension for vacuum transition between v_1 and v_2 for any positive value of β , which is obviously seen from the topology.

We show the dimensionless energy density distribution along the z direction of the three different solutions between v_1 and v_4 in Fig. [12](#page-15-0) with $\beta = 0.3$, where

$$
\overline{\rho}(\overline{z}) = \sum_{i} \frac{1}{2} \left[\overline{\phi}'_i(\overline{z}) \right]^2 + \Delta \overline{V}(\overline{\phi}(\overline{z})), \tag{4.3}
$$

is the dimensionless energy density normalised using Eq. [\(3.3\)](#page-6-3). The blue curve refers to the energy density profile of SI DW $|v_1| v_4$. As anticipated we observe that the energy density peaks at $\overline{z} = 0$ as falls off at $|\overline{z}| \sim 1/0.3$. The purple curve shows the energy density profile of $|v_1| v_2 | v_4|$ and we see that there are two peaks (i.e. two domain walls) with the energy density peaking around $\overline{z} \sim \pm 10$ with magnitude significantly smaller than the SI DW scenario. The orange curve refers to the third solution where there is a flatter and broader domain wall than the SI DW case.

The first and second solutions give tensions $\sigma_{SI} = 2\sqrt{2}/3$ and $2\sigma_{SII}(\beta = 0.3) \approx 0.5$, respectively, representing the largest and smallest energy unit area stored in the scalar fields. In this case, the SI DW is unstable and decays to two SII DWs. We anticipate that during this decay the difference in energy of the higher and lower tension domain walls is released as gravitational radiation or possible light scalar (flavon) degrees of freedom.

We have numerically checked that given an initial SI path in the field space with a small perturbation to $\bar{\phi}_2$ and $\bar{\phi}_3$, the path automatically deviates from the initial one during the deformation iteration and eventually stabilised at the path $|v_1|$ $|v_2|$ $|v_4|$ or any other equivalent ones $|v_1| v_{3,5,6} | v_4$. We developed code to numerically calculate domain wall solutions using a method similar to that of CosmoTransitions [\[43\]](#page-28-3), solving the equations of motion for scalar field configurations. While CosmoTransitions focuses on solving the equations of motion for first-order phase transitions by finding bounce solutions and computing the Euclidean action, which determines the tunnelling probability between vacuum, our approach uses a similar methodology but for domain wall solutions.

4.2 T-type domain walls

For β < 0, we arrive at three T-type DWs. The TII DW is always stable as seen from the topology. We first check the stability of the TI DW. There can be three different types of solutions between u_1 and u_5 , as shown in the left panel of Fig. [13](#page-17-0) in field space.

Figure 13: Left: Three different types of possible domain solutions from u_1 to u_5 when $\beta = -0.1$. Right: Tension of different possible DW solutions from u_1 to u_5 .

- 1. The TI DW which is a blue straight line in the left panel of Fig. [13](#page-17-0) that runs between u_1 and u_5 in the field space and is labelled as $|u_1|$ u_5
- 2. The combination of a TII DW and a TIII DW, which passes one of the other vacua in the field space, e.g. $|u_1| u_4$ and $|u_4| u_5$, or equivalently $|u_1| u_4 | u_5$. This is shown as a grey line in the left panel of Fig. [13.](#page-17-0)
- 3. The combination of three TII DWs, which consists of three edges of the cube, e.g. $u_1 | u_2 |$, $| u_2 | u_7 |$ and $| u_7 | u_5 |$, or equivalently, $| u_1 | u_2 | u_7 | u_5 |$. This is shown as a green line in the left panel of Fig. [13.](#page-17-0)

The tensions of the three different solutions are shown in the right panel of Fig. [13](#page-17-0) for various β. For −β > 1, the TI DW, referred to as solution 1 here, has the lowest energy and is, therefore, energetically stable. The TI DW is unstable for $-\beta < 1$ as solution 3 has the lowest energy. We then check the stability of the TIII DW. Taking u_2 and u_5 as boundary conditions. We find two types of solutions.

- 1. One is the path directly connecting between u_2 and u_5 , which is equivalent to that between u_4 and u_5 . See $|u_4|$ u_5 in Fig. [13.](#page-17-0)
- 2. The other is the path from u_1 to u_2 to u_8 . This is equivalent to the path from u_2 to u_7 to u_5 , i.e., $|u_2| u_7 | u_5$ in Fig. [13.](#page-17-0)

By comparing tensions in these two solutions, we see that the TIII DW is stable when $-\beta$ > 1. Finally, we conclude that both TI and TIII DWs are stable for $-\beta$ > 1 and unstable for $-\beta < 1$.

5 Gravitational waves from non-Abelian domain walls

In general, stable DWs can be problematic as their energy would come to dominate the total energy density of the universe. This problem can be solved if the residual discrete symmetry is softly broken by energy biases between the vacua. The energy biases can be expressed as

$$
V_{ij}^{\text{bias}} = \epsilon_{ij} v^4 \,, \tag{5.1}
$$

where i, j refers to the scalar field indices with v the VEV of the field in the absence of softly symmetry breaking. In addition to leading to the decay of domain walls, the bias term can help correct leading order mixing patterns which describe flavour mixing in the lepton sector [\[40,](#page-28-0) [44,](#page-28-4) [45\]](#page-28-5). These biases make the DWs unstable and annihilate at a temperature $T_{\rm ann}$, generating gravitational waves. The spectrum of the GW for frequency f and at time t is given by

$$
\Omega h^2(f,t) = \frac{h^2}{\rho_c(t)} \left(\frac{d\rho_{\rm gw}(t)}{d\ln f} \right) \,. \tag{5.2}
$$

For a typical Z_2 DW, the GW spectrum is characterised by a peak at the frequency

$$
f_{\text{peak}} \simeq a \left(t_{\text{ann}} \right) H \left(t_{\text{ann}} \right) \simeq 1 \times 10^{-4} \text{ Hz} \left(\frac{g_{*} \left(T_{\text{ann}} \right)}{10} \right)^{1/6} \left(\frac{T_{\text{ann}}}{\text{TeV}} \right)
$$

$$
\simeq 3 \times 10^{3} \text{ Hz} \left(\frac{10}{g_{*} \left(T_{\text{ann}} \right)} \right)^{1/12} \left(\frac{V_{\text{bias}}}{\sigma \text{ TeV}} \right)^{1/2}, \tag{5.3}
$$

with amplitude

$$
\Omega_{\rm gw}h^2\big|_{\rm peak} \simeq 0.7 \times 10^{-37} \left(\frac{10}{g_*(T_{\rm ann})}\right)^{4/3} \left(\frac{\sigma}{\rm TeV} \frac{\sigma}{T_{\rm ann}^2}\right)^2
$$
\n
$$
\simeq 0.9 \times 10^{-67} \left(\frac{10}{g_*(T_{\rm ann})}\right)^{1/3} \left(\frac{\sigma}{\rm TeV}^3\right)^4 \left(\frac{\rm TeV}^4}{V_{\rm bias}}\right)^2,
$$
\n(5.4)

where the present critical density $\rho_c(t_0) = 10.53h^2$ keV/cm³ and effective degrees of freedom $g_{s*}(t_0) = 3.91$ is applied. On each side of peak, we use the power-law approximation obtained from numerical simulation, which are $\Omega_{\rm gw} h^2(f) \propto f^3$ for $f < f_{\rm peak}$ and $\Omega_{\rm gw}h^2(f) \propto f^{-1}$ for $f > f_{\rm peak}$ [\[25\]](#page-27-1). In the non-Abelian DWs discussed in this work, some solutions are unstable even in the absence of any biases. The evolution of those unstable DW solutions can change the kinetic history of the DW network and thus produce special gravitational waves signal. An interesting example is the decay of an SI DW into a pair of SII DWs, during which extra energy is released, partly into the kinetic energy of the walls and gravitational radiation. The details of those complex processes are beyond the scope of this work. Here, we focus on the case where all the wall solutions are stable, namely $\beta > 2$ for the Z_2 preserving case and $\beta < -1$ for the Z_3 preserving case.

5.1 Gravitational waves emission by non-Abelian domain walls

For the S-type DWs, the SII DWs are four times more likely to be produced than SI DWs during the phase transition. We assume that the ratio of number density to be constant

		$\begin{bmatrix} \epsilon_{12}^v & \epsilon_{13}^v & \epsilon_{14}^v & \epsilon_{15}^v & \epsilon_{16}^v \end{bmatrix} \begin{bmatrix} \epsilon_{12}^u & \epsilon_{13}^u & \epsilon_{14}^u & \epsilon_{15}^u & \epsilon_{16}^u & \epsilon_{17}^u & \epsilon_{18}^u \end{bmatrix}$			
		$\begin{bmatrix} 2\hat{\epsilon} & 3\hat{\epsilon} & \hat{\epsilon} & 4\hat{\epsilon} & 5\hat{\epsilon} \end{bmatrix} \begin{bmatrix} 2\hat{\epsilon} & 4\hat{\epsilon} & 6\hat{\epsilon} & \hat{\epsilon} & 3\hat{\epsilon} & 5\hat{\epsilon} & 7\hat{\epsilon} \end{bmatrix}$			

Table 1: Choice of the bias for the difference between the vacua of the Z_2 (left) and Z_3 (right) preserving scenario in terms of $\hat{\epsilon}$, which is a free parameter which we set when discussing testability.

Figure 14: The black line shows the gravitational waves spectrum due to overlapping spectra for the S-type DWs (left) with $\beta = 10$, and T-type DWs (right) with $\beta = -1.2$. The frequencies and the spectrum are normalised to $\hat{f} = f_{\text{peak}}^{Z_2}$, the peak frequency of the contribution from SI/TI DWs, and its correspondent amplitude $\hat{\Omega}h^2$. The values of five possible biases for the Z_2 case and seven possible biases for the Z_3 case are given in Tab. [1.](#page-19-0)

during the subsequent evolution and we have n_{SI} : $n_{\text{SI}} = 1 : 4$. To be consistent with cosmology, biases are introduced to make the DWs unstable and decay. From the vacuum structure in Fig. [1,](#page-5-0) at first glance there are 15 biases between the six vacua: three for SI DWs and 12 for SII DWs. However, only five of them are linearly independent. Here, as a example, v_1 is chosen to be the lowest vacuum and the other five vacua are determined by the biases between them and v_1 . The biases for the benchmark case are summarised in Tab. [1.](#page-19-0)

The left panel of Fig. [14](#page-19-1) shows the resulting GW amplitude as a function of the wavenumber. While the biases are free parameters, the tensions of the DWs can be fixed by the potential as discussed in Sec[.3.](#page-4-0) The consequence of different possible tensions is similar to what has been observed in Ref. $[40]$ but more significant. In Ref. $[40]$, the origin of having different peaks is the difference between the biases while the tension is unique. Here, since there are different types of DWs, these DWs with different tensions contribute to various peaks in the complete spectrum. This additional effect leads to a more pronounced multipeak structure than the one caused by different biases. In particular, if $\Omega^{\text{SI}} \sim \Omega^{\text{SI}}$, the spectrum does not evolve as f^{-1} in a significant interval of frequencies after the first peak, which is more likely to be detected.

For the T-type DWs, the ratio of the number densities of different vacua is n_{TI} : n_{TI} . $n_{\text{THI}} = 1:3:3$. Again, u_1 is chosen to be the lowest energy vacuum and the other seven

	$\sigma^{1/3}/\text{TeV}$	$V_{\text{bias}}^{1/4}/\text{TeV}$
BP1	3×10^2	$10^{-3.75}$
BP2	3×10^5	10^{1}

Table 2: The values for the tension of SI and TI DWs and the smallest bias of the two benchmark points.

vacua are determined by the biases between them and u_1 . The biases for the benchmark case are summarised in Tab. [1.](#page-19-0) Following the same procedure as the S-type DWs, we find a similar behaviour, where there is a deviation from f^{-1} power-law spectrum after the peak, as shown in the right panel of Fig. [14.](#page-19-1)

Finally, it is important to recall that the discussion above holds only when $\beta > 2$ for the S-type DWs and $\beta < -1$ for the T-type DWs. If β is smaller (larger) for the S-type (T-type) DWs, there would be unstable DWs which may decay much earlier than the annihilation time of the stable ones. Such a process may generate additional signals at higher frequencies and provide a different GW spectrum.

5.2 Testability with gravitational wave observatories

Given the recent discovery of a stochastic gravitational wave background by Pulsar Time Array (PTA) observatories and planned next-generation earth and space GW observatories, it is interesting to see how the deviation from the ordinary Z_2 spectrum can be detected. Here, we provide a qualitative analysis regarding the possibility of such GW observatories to detect this deviation. Two free parameters can affect the gravitational wave spectrum produced by the DW network: the wall tensions and the biases. As discussed in Sec. [3,](#page-4-0) the DWs tension can be determined by β up to an overall scale. In the following discussion, we fix the overall scale of the tensions by fixing the tension of SI and TI DW in the Z_2 - and Z_3 -preserving cases, respectively. For the biases, we follow the ratio given in Tab. [1,](#page-19-0) with the smallest bias as a free parameter.

To provide a more concrete discussion, we choose two benchmark values for the tensions of SI and TI DWs and the smallest biases, as shown in Tab. [2.](#page-20-1) The estimated gravitational wave spectrum is shown in Fig. [15.](#page-21-1) The peak of the first benchmark point is in the nHz range and the spectrum is consistent with the signal observed by PTA, as shown in Fig. [15.](#page-21-1) At the higher frequencies spanned by PTAs, the multi-peak structure shows up instead of the power-law structure, changing the slope of the spectrum after the peak. According to Ref. [\[32\]](#page-27-6), the annihilation of the network happens at around 100 MeV. For the other benchmark point, the peak falls within the frequency range that is detectable for nextgeneration space-based gravitational wave detectors, such as LISA and Taiji. Moreover, the deviation from the ordinary Z_2 spectrum around the peak frequencies can also be observed. Nevertheless, more quantitative discussion about the GW spectrum would require a numerical simulation of the DW dynamics and GW production, which will be addressed in subsequent works.

Figure 15: Gravitational waves spectrum due to overlapping spectra for the S-type DWs (top) with $\beta = 10$, and T-type DWs (bottom) with $\beta = -1.2$. The solid and dashed lines correspond to the benchmark points BP1 and BP2 in Tab. [2,](#page-20-1) respectively.

6 Conclusion

In this work, we have performed the first detailed analysis of the properties of domain walls (DWs) emerging from the spontaneous breaking of non-Abelian discrete symmetries. Choosing S_4 as an example, we have discussed the vacuum structure from a generic renormalisable potential of a triplet flavon ϕ . Depending on the potential parameters, this group provides two types of vacua, six Z_2 -preserving and eight Z_3 -preserving vacua. From these vacua, we found five types of DWs, labelled as SI, SII and TI, TII, TIII, respectively. Here, SI and SII refer to two types DWs separating two Z_2 -preserving, and TI, TII and TIII refer to three types of DWs separating two Z_3 -preserving vacua.

Given the different types of DWs, we explored their properties by computing the profiles, the tension and the thickness for each of them with respect to β (where $\beta = q_2/q_1$), which is the only free parameter of S_4 after the redefinition of the scalar fields and the space coordinate. Noticing that the DW between some pairs of vacua can be unstable, we have discussed the stability of the various solutions, and we find the tension plays a crucial role. For $|\beta| \ll 1$ the difference between the tensions of the different DWs can be more than an order of magnitude. In that case, the DWs with higher tension will evolve into the ones with lower tension. The stable DW solutions are the SII type for the Z_2 -preserving vacua and the TII type for the Z_3 -preserving vacua.

Finally, we have estimated the gravitational wave signature for non-Abelian DWs and qualitatively compared it with the signal from a network of Z_2 DWs. We have found a multi-peak behaviour which is more enhanced with respect to what has been discussed in Ref. [\[40\]](#page-28-0) because there are DW solutions with differing tensions which may be hierarchical. When all the solutions are stable, i.e. $\beta > 2$ for the S-type DWs and $\beta < -1$ for the T-type DWs, the gravitational wave signals produced by the DWs with different tensions reach their peaks at various frequencies, which is a distinctive feature of DWs generated by non-Abelian discrete symmetries.

In summary, the DWs from non-Abelian discrete symmetry breaking have a richer structure than the Z_2 scenario in many aspects. This could lead to interesting phenomena regarding the evolution of the DWs, such as solution instabilities that could affect the production of gravitational wave signals and may have other possible cosmological consequences. The study of DWs from non-Abelian discrete symmetries allows a precise connection between flavour physics and gravitational wave phenomenology and our work opens a new pathway to test flavour models.

Acknowledgement

We want to thank Y. C. Wu for the helpful discussion. J.T. is supported by the STFC under Grant No. ST/X003167/1. This work was partially supported by European Union's Horizon 2020 Research, Innovation Programme under Marie Sklodowska-Curie grant agreement HIDDeN European ITN project (H2020-MSCA-ITN-2019//860881-HIDDeN), National Natural Science Foundation of China (NSFC) under Grant Nos. 12205064, 12347103, Zhejiang Provincial Natural Science Foundation of China under Grant No. LDQ24A050002 and STFC Consolidated Grant ST/T000775/1.

A Derivation of renormalisable potentials in S_4

The renormalisable potential $V(\phi)$ in Eq. [\(A.4\)](#page-23-1) applies to S_4 in the Ma-Rajasekaran basis [\[41\]](#page-28-1). It is the general form of renormalisable potentials invariant in these discrete groups. Below we show how to derive it in S_4 .

Given two triplets $a = (a_1, a_2, a_3)^T$ and $b = (b_1, b_2, b_3)^T$ of S_4 , In S_4 , there are five irreps 1, 1' 2, 3 and 3'. The product of two $3^{(\prime)}$ -plets follows the decomposition $3 \times 3 =$ $3' \times 3' = 1 + 2 + 3 + 3'$, with

$$
(ab)1 = a1b1 + a2b2 + a3b3,(ab)2 = (a1b1 + \omega a2b2 + \omega2 a3b3, a1b1 + \omega2 a2b2 + \omega a3b3)T,(ab)3 = (a2b3 + a3b2, a3b1 + a1b3, a1b2 + a2b1)T,(ab)3 = (a2b3 - a3b2, a3b1 - a1b3, a1b2 - a2b1)T.
$$
(A.1)

The product of 3 and 3' follows the decomposition $3 \times 3' = 1' + 2 + 3' + 3$. Here no trivial singlet is given. Constructions of $(ab)_1$ and $(ab)_2$ follows those in $(ab)_1$ and $(ab)_2$ in Eq. [\(A.1\)](#page-23-2), respectively. Constructions of $(ab)_3$ and $(ab)_3$ ['] are reversed compared with those in Eq. $(A.1)$.

If ϕ is arranged to be a triplet 3', the most general S_4 -invariant potential is given by

$$
V(\phi) = -\frac{1}{2}\mu^2(\phi\phi)_1 + \frac{1}{4}\left[f_1((\phi\phi)_1)^2 + \frac{1}{2}f_2((\phi\phi)_2(\phi\phi)_2)_1 + f_3((\phi\phi)_3(\phi\phi)_3)_1\right],
$$
 (A.2)

where μ^2 and $f_{1,2,3}$ are real parameters. Then we fully recover Eq. [\(2.2\)](#page-3-0) with $g_1 = f_1 + f_2$, $g_2 = 2f_3 - \frac{2}{3}$ $\frac{2}{3}f_2$. If ϕ is arranged to be the triplet 3, where will also be a cubic term ($(\phi \phi)_{\mathbf{3}} \phi$ ₁ = 6 $\phi_1 \phi_2 \phi_3$.

Including the cubic term actually gives the most general A_4 -invariant renormalisable potential for a triplet. A_4 has four irreps, 1, 1', 1'', and 3. Kronecker products of two triplet follows the decomposition $3 \times 3 = 1 + 1' + 1'' + 3_S + 3_A$, where the subscripts s and ^A stand for the symmetric and anti-symmetric parts, respectively,

$$
(ab)1 = a1b1 + a2b2 + a3b3,
$$

\n
$$
(ab)1' = a1b1 + \omega a2b2 + \omega2 a3b3,
$$

\n
$$
(ab)1'' = a1b1 + \omega2 a2b2 + \omega a3b3,
$$

\n
$$
(ab)3s = (a2b3 + a3b2, a3b1 + a1b3, a1b2 + a2b1)T,
$$

\n
$$
(ab)3s = (a2b3 - a3b2, a3b1 - a1b3, a1b2 - a2b1)T.
$$
 (A.3)

Following the representation theory of A_4 , one can construct the general A_4 -invariant renormalisable potential as

$$
V(\phi) = -\frac{1}{2}\mu^2(\phi\phi)_1 + \frac{1}{4}\left[f_1((\phi\phi)_1)^2 + f_2(\phi\phi)_1/(\phi\phi)_1 + f_3((\phi\phi)_3/(\phi\phi)_3)\right] + \frac{A}{6}((\phi\phi)_3/\phi)_3\phi)_1,
$$
\n(A.4)

where A is real. This is almost the same except the last cubic term. In A_4 , this term can be forbidden by introducing a parity Z_2^P symmetry, $\phi \leftrightarrow -\phi$.

B Domain wall properties at $\beta = 2$

In the special case at $\beta = 2$, i.e., $g_2 = 2g_1$ in the potential. The domain wall can be analytically solved and we provide the of derivation below. We take the DW of $|v_1|$ $|v_2|$ as an example. Since ϕ_3 vanishes in the solution, it is enough for us to consider the effective potential of only ϕ_1 and ϕ_2 , i.e.,

$$
V(\phi_1, \phi_2) = -\frac{\mu^2}{2}(\phi_1^2 + \phi_2^2) + \frac{g_1}{4}[(\phi_1^2 + \phi_2^2)^2 + 4\phi_1^2 \phi_2^2].
$$
 (B.1)

By taking the transformation $\phi_{\pm} = \frac{1}{\sqrt{2}}$ $\frac{1}{2}(\phi_1 \pm \phi_2)$, we obtain $V(\phi_1, \phi_2) = V(\phi_+) + V(\phi_-)$ with $V(\phi_+)$ given by

$$
V(\phi_{\pm}) = -\frac{\mu^2}{2}\phi_{\pm}^2 + \frac{g_1}{2}\phi_{\pm}^4.
$$
 (B.2)

In this parameterisation, ϕ_+ and ϕ_- decouple in the potential. And the potential is invariant under two seperate Z_2 symmetries Z_2^+ : $\phi_+ \leftrightarrow -\phi_+$ and Z_2^- : $\phi_- \leftrightarrow -\phi_-$. Boundary conditions for $\boxed{v_1 \mid v_2}$ are equivalent to

$$
\phi_{\pm}|_{z \to \infty} = \frac{v}{\sqrt{2}}, \quad \phi_{\pm}|_{z \to -\infty} = \pm \frac{v}{\sqrt{2}} \tag{B.3}
$$

This suggests $\phi_+(z)$ takes the VEV $\frac{v}{\sqrt{2}}$ along the z-direction from $+\infty$ to $-\infty$ and $\phi_-(z)$ just follows the Z_2 domain wall solution $\phi_-(z) = \frac{v}{\sqrt{2}} \tanh(\frac{\mu z}{\sqrt{2}})$. The tension and thickness of the domain wall are obtained straightforwardly,

$$
\sigma_{\text{SII}}(\beta = 2) = \mu \frac{v^2}{2} \frac{2\sqrt{2}}{3}, \quad \delta_{\text{SII}}(\beta = 2) = \frac{\sqrt{2}}{\mu}.
$$
 (B.4)

Following the parametrisation in the Sec. [3.1,](#page-6-0) we obtain the dimensionless tension and thickness:

$$
\bar{\sigma}_{\text{SII}}(\beta = 2) = \frac{\sqrt{2}}{3} = \frac{1}{2}\bar{\sigma}_{\text{SI}}, \quad \bar{\delta}_{\text{SI}}(\beta = 2) = \sqrt{2} = \bar{\delta}_{\text{SI}}.
$$
\n(B.5)

We have proved that at $\beta = 2$, the energy stored in an SII DW is half that of a SI DW.

C Effect of the cubic term in the potential

By arranging the flavon as a 3 of S_4 or considering the potential in A_4 , we have to include a cubic term in the potential. Then the full renormalisable potential is written to be

$$
V(\phi) = -\frac{\mu^2}{2}I_1 + \frac{g_1}{4}I_1^2 + \frac{g_2}{2}I_2 + A\phi_1\phi_2\phi_3.
$$
 (C.1)

We discuss the vacuum structure and DW property induced by the last term.

Including this term, the six Z_2 -preserving vacua do not be modified, but the eight Z_3 -preserving solutions u_n are further split into to two groups,

$$
u_{1,6,7,8} = \left\{ \begin{pmatrix} 1 \\ 1 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ -1 \\ -1 \end{pmatrix}, \begin{pmatrix} -1 \\ 1 \\ -1 \end{pmatrix}, \begin{pmatrix} -1 \\ -1 \\ 1 \end{pmatrix} \right\} u_{-}
$$
(C.2)

$$
u_{2,3,4,5} = \left\{ \begin{pmatrix} -1 \\ 1 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ -1 \\ 1 \end{pmatrix}, \begin{pmatrix} 1 \\ 1 \\ -1 \end{pmatrix}, \begin{pmatrix} -1 \\ -1 \\ -1 \end{pmatrix} \right\} u_+, \tag{C.3}
$$

where

$$
u_{\mp} = \frac{\mu}{\sqrt{3g_1 + 2g_2}} \left(\sqrt{1 + a^2} \mp a \right)
$$
 (C.4)

and $a = \frac{2A}{\sqrt{3a}}$ $\frac{2A}{\mu\sqrt{3g_1+2g_2}}$. The potential at these solutions are split into two values

$$
V_{\mp} = -\frac{3\mu^4}{4(3g_1 + 2g_2)} \left(1 + \frac{2}{3}a^2 \mp \frac{2}{3}a\sqrt{a^2 + 1} \right) \left(1 + 2a^2 \mp 2a\sqrt{a^2 + 1} \right) ,\qquad\text{(C.5)}
$$

where the \mp signs are for $u_{1,6,7,8}$ and $u_{2,3,4,5}$, respectively. In the case of negative A (i.e., $a < 0$, $V₋ < V₊$, and thus $V₋$ is the global minimum. $u_{1,6,7,8}$ and $u_{2,3,4,5}$ are not equivalent any more and the former four solutions represent the true vacua, while the latter solutions become metastable false vacua. On the contrary, for positive A (i.e., $a > 0$), the vacua properties are flipped. In either cases, the flavon masses are modified to

$$
m_{1,\mp}^2 = 2(3g_1 + 2g_2)u_{\mp}^2 \left(1 + a^2 \pm a\sqrt{a^2 + 1}\right),
$$

\n
$$
m_{2,\mp}^2 = m_{3,\mp}^2 = -2g_2u_{\mp}^2 \left(1 + (a^2 \pm a\sqrt{a^2 + 1})(4 + 6g_1/g_2)\right).
$$
\n(C.6)

Due to inclusion of the cubic term, the vacua structure become more complicated and we expect that DWs can have more complicated structures compared with those in the maintext. We will study this scenario in detail elsewhere. Below, we give brief comments on the influence of the cubic term on DWs.

For S-type DWs, where two Z_2 -preserving vacua have been fixed by boundary conditions, the coefficient A of the cubic term does not contribute to the EOM of the scalar. Thus all S-type DWs are independent of the cubic term.

For T-type DWs, the cubic term can contribute to their properties but at different levels for different types. For TIII DWs, the cubic term contributes to both the vacua expectation value and flavon EOM. Thus, it modifies the DWs just quantitatively. Modifications to TI and TII DWs are more significant. A TI or TII DW separates two vacua in two different sets. We take $|u_1| \cdot |u_2|$ with $A < 0$ as an example. From the discussion in the end of Sec. [2,](#page-2-0) it is clear that the vacuum energy $V_-\,$ of the true vacua u_1 on one side is lower than the vacuum energy V_+ of the false vacua u_2 on the other side. Such a term, if it is not too large, can be treated as a bias. As discussed in section [5,](#page-18-0) a bias makes DWs collapse at a certain temperature and generate gravitational waves. Here, a non-zero cubic term can generate a pre-existing bias term to the vacua. If this term is larger than the bias from soft symmetry breaking, the TI and TII DWs, as well as half of the TIII DWs, may collapse at a higher temperature, namely, earlier than the other DWs. Its collapsing also generates GWs, but the signal is weak due to the large bias.

References

[1] KAMIOKANDE-II collaboration, Results from one thousand days of real time directional solar neutrino data, [Phys. Rev. Lett.](https://doi.org/10.1103/PhysRevLett.65.1297) 65 (1990) 1297.

- [2] B.T. Cleveland, T. Daily, R. Davis, Jr., J.R. Distel, K. Lande, C.K. Lee et al., Measurement of the solar electron neutrino flux with the Homestake chlorine detector, [Astrophys. J.](https://doi.org/10.1086/305343) 496 [\(1998\) 505.](https://doi.org/10.1086/305343)
- [3] GALLEX collaboration, GALLEX solar neutrino observations: Results for GALLEX IV, [Phys. Lett. B](https://doi.org/10.1016/S0370-2693(98)01579-2) 447 (1999) 127.
- [4] SAGE collaboration, Measurement of the solar neutrino capture rate with gallium metal, Phys. Rev. C 60 [\(1999\) 055801](https://doi.org/10.1103/PhysRevC.60.055801) [[astro-ph/9907113](https://arxiv.org/abs/astro-ph/9907113)].
- [5] Super-Kamiokande collaboration, Solar B-8 and hep neutrino measurements from 1258 days of Super-Kamiokande data, [Phys. Rev. Lett.](https://doi.org/10.1103/PhysRevLett.86.5651) 86 (2001) 5651 [[hep-ex/0103032](https://arxiv.org/abs/hep-ex/0103032)].
- [6] G. Altarelli and F. Feruglio, Discrete Flavor Symmetries and Models of Neutrino Mixing, [Rev. Mod. Phys.](https://doi.org/10.1103/RevModPhys.82.2701) 82 (2010) 2701 [[1002.0211](https://arxiv.org/abs/1002.0211)].
- [7] S.F. King and C. Luhn, Neutrino Mass and Mixing with Discrete Symmetry, [Rept. Prog.](https://doi.org/10.1088/0034-4885/76/5/056201) Phys. 76 [\(2013\) 056201](https://doi.org/10.1088/0034-4885/76/5/056201) [[1301.1340](https://arxiv.org/abs/1301.1340)].
- [8] S.F. King, Unified Models of Neutrinos, Flavour and CP Violation, [Prog. Part. Nucl. Phys.](https://doi.org/10.1016/j.ppnp.2017.01.003) 94 [\(2017\) 217](https://doi.org/10.1016/j.ppnp.2017.01.003) [[1701.04413](https://arxiv.org/abs/1701.04413)].
- [9] Z.-z. Xing, Flavor structures of charged fermions and massive neutrinos, [Phys. Rept.](https://doi.org/10.1016/j.physrep.2020.02.001) 854 [\(2020\) 1](https://doi.org/10.1016/j.physrep.2020.02.001) [[1909.09610](https://arxiv.org/abs/1909.09610)].
- [10] LIGO Scientific, Virgo collaboration, Observation of Gravitational Waves from a Binary Black Hole Merger, [Phys. Rev. Lett.](https://doi.org/10.1103/PhysRevLett.116.061102) 116 (2016) 061102 [[1602.03837](https://arxiv.org/abs/1602.03837)].
- [11] NANOGrav collaboration, The NANOGrav 15 yr Data Set: Evidence for a Gravitational-wave Background, [Astrophys. J. Lett.](https://doi.org/10.3847/2041-8213/acdac6) 951 (2023) L8 [[2306.16213](https://arxiv.org/abs/2306.16213)].
- [12] EPTA, InPTA: collaboration, The second data release from the European Pulsar Timing Array - III. Search for gravitational wave signals, [Astron. Astrophys.](https://doi.org/10.1051/0004-6361/202346844) 678 (2023) A50 [[2306.16214](https://arxiv.org/abs/2306.16214)].
- [13] EPTA collaboration, The second data release from the European Pulsar Timing Array: V. Implications for massive black holes, dark matter and the early Universe, [2306.16227](https://arxiv.org/abs/2306.16227).
- [14] D.J. Reardon et al., Search for an Isotropic Gravitational-wave Background with the Parkes Pulsar Timing Array, [Astrophys. J. Lett.](https://doi.org/10.3847/2041-8213/acdd02) 951 (2023) L6 [[2306.16215](https://arxiv.org/abs/2306.16215)].
- [15] H. Xu et al., Searching for the Nano-Hertz Stochastic Gravitational Wave Background with the Chinese Pulsar Timing Array Data Release I, [Res. Astron. Astrophys.](https://doi.org/10.1088/1674-4527/acdfa5) 23 (2023) 075024 [[2306.16216](https://arxiv.org/abs/2306.16216)].
- [16] E. Madge, E. Morgante, C. Puchades-Ibáñez, N. Ramberg, W. Ratzinger, S. Schenk et al., Primordial gravitational waves in the nano-Hertz regime and PTA data — towards solving the GW inverse problem, JHEP 10 [\(2023\) 171](https://doi.org/10.1007/JHEP10(2023)171) [[2306.14856](https://arxiv.org/abs/2306.14856)].
- [17] A. Vilenkin, Cosmic Strings and Domain Walls, [Phys. Rept.](https://doi.org/10.1016/0370-1573(85)90033-X) 121 (1985) 263.
- [18] T.W.B. Kibble, Topology of Cosmic Domains and Strings, J. Phys. A 9 [\(1976\) 1387.](https://doi.org/10.1088/0305-4470/9/8/029)
- [19] J. Preskill, MAGNETIC MONOPOLES, [Ann. Rev. Nucl. Part. Sci.](https://doi.org/10.1146/annurev.ns.34.120184.002333) 34 (1984) 461.
- [20] P. Auclair et al., Probing the gravitational wave background from cosmic strings with LISA, JCAP 04 [\(2020\) 034](https://doi.org/10.1088/1475-7516/2020/04/034) [[1909.00819](https://arxiv.org/abs/1909.00819)].
- [21] Y. Gouttenoire, G. Servant and P. Simakachorn, Beyond the Standard Models with Cosmic Strings, JCAP 07 [\(2020\) 032](https://doi.org/10.1088/1475-7516/2020/07/032) [[1912.02569](https://arxiv.org/abs/1912.02569)].
- [22] T. Hiramatsu, M. Kawasaki and K. Saikawa, Gravitational Waves from Collapsing Domain Walls, JCAP 05 [\(2010\) 032](https://doi.org/10.1088/1475-7516/2010/05/032) [[1002.1555](https://arxiv.org/abs/1002.1555)].
- [23] M. Kawasaki and K. Saikawa, Study of gravitational radiation from cosmic domain walls, JCAP 09 [\(2011\) 008](https://doi.org/10.1088/1475-7516/2011/09/008) [[1102.5628](https://arxiv.org/abs/1102.5628)].
- [24] T. Hiramatsu, M. Kawasaki and K. Saikawa, On the estimation of gravitational wave spectrum from cosmic domain walls, JCAP 02 [\(2014\) 031](https://doi.org/10.1088/1475-7516/2014/02/031) [[1309.5001](https://arxiv.org/abs/1309.5001)].
- [25] K. Saikawa, A review of gravitational waves from cosmic domain walls, Universe 3 [\(2017\) 40](https://doi.org/10.3390/universe3020040) [[1703.02576](https://arxiv.org/abs/1703.02576)].
- [26] W. Buchmuller, V. Domcke and K. Schmitz, Stochastic gravitational-wave background from metastable cosmic strings, JCAP 12 [\(2021\) 006](https://doi.org/10.1088/1475-7516/2021/12/006) [[2107.04578](https://arxiv.org/abs/2107.04578)].
- [27] D.I. Dunsky, A. Ghoshal, H. Murayama, Y. Sakakihara and G. White, GUTs, hybrid topological defects, and gravitational waves, Phys. Rev. D 106 [\(2022\) 075030](https://doi.org/10.1103/PhysRevD.106.075030) [[2111.08750](https://arxiv.org/abs/2111.08750)].
- [28] S.F. King, S. Pascoli, J. Turner and Y.-L. Zhou, Gravitational Waves and Proton Decay: Complementary Windows into Grand Unified Theories, [Phys. Rev. Lett.](https://doi.org/10.1103/PhysRevLett.126.021802) 126 (2021) 021802 [[2005.13549](https://arxiv.org/abs/2005.13549)].
- [29] S.F. King, S. Pascoli, J. Turner and Y.-L. Zhou, Confronting SO(10) GUTs with proton decay and gravitational waves, JHEP 10 [\(2021\) 225](https://doi.org/10.1007/JHEP10(2021)225) [[2106.15634](https://arxiv.org/abs/2106.15634)].
- [30] B. Fu, S.F. King, L. Marsili, S. Pascoli, J. Turner and Y.-L. Zhou, A predictive and testable unified theory of fermion masses, mixing and leptogenesis, JHEP 11 [\(2022\) 072](https://doi.org/10.1007/JHEP11(2022)072) [[2209.00021](https://arxiv.org/abs/2209.00021)].
- [31] B. Fu, S.F. King, L. Marsili, S. Pascoli, J. Turner and Y.-L. Zhou, Testing realistic SO(10) SUSY GUTs with proton decay and gravitational waves, Phys. Rev. D 109 [\(2024\) 055025](https://doi.org/10.1103/PhysRevD.109.055025) [[2308.05799](https://arxiv.org/abs/2308.05799)].
- [32] NANOGrav collaboration, The NANOGrav 15 yr Data Set: Search for Signals from New Physics, [Astrophys. J. Lett.](https://doi.org/10.3847/2041-8213/acdc91) 951 (2023) L11 [[2306.16219](https://arxiv.org/abs/2306.16219)].
- [33] N. Kitajima, J. Lee, K. Murai, F. Takahashi and W. Yin, Gravitational waves from domain wall collapse, and application to nanohertz signals with QCD-coupled axions, [Phys. Lett. B](https://doi.org/10.1016/j.physletb.2024.138586) 851 [\(2024\) 138586](https://doi.org/10.1016/j.physletb.2024.138586) [[2306.17146](https://arxiv.org/abs/2306.17146)].
- [34] Y. Bai, T.-K. Chen and M. Korwar, QCD-collapsed domain walls: QCD phase transition and gravitational wave spectroscopy, JHEP 12 [\(2023\) 194](https://doi.org/10.1007/JHEP12(2023)194) [[2306.17160](https://arxiv.org/abs/2306.17160)].
- [35] E. Babichev, D. Gorbunov, S. Ramazanov, R. Samanta and A. Vikman, NANOGrav spectral index $\gamma = 3$ from melting domain walls, Phys. Rev. D 108 [\(2023\) 123529](https://doi.org/10.1103/PhysRevD.108.123529) [[2307.04582](https://arxiv.org/abs/2307.04582)].
- [36] Z. Zhang, C. Cai, Y.-H. Su, S. Wang, Z.-H. Yu and H.-H. Zhang, Nano-Hertz gravitational waves from collapsing domain walls associated with freeze-in dark matter in light of pulsar timing array observations, Phys. Rev. D 108 [\(2023\) 095037](https://doi.org/10.1103/PhysRevD.108.095037) [[2307.11495](https://arxiv.org/abs/2307.11495)].
- [37] S. Blasi, A. Mariotti, A. Rase and A. Sevrin, Axionic domain walls at Pulsar Timing Arrays: QCD bias and particle friction, JHEP 11 [\(2023\) 169](https://doi.org/10.1007/JHEP11(2023)169) [[2306.17830](https://arxiv.org/abs/2306.17830)].
- [38] Y. Wu, K.-P. Xie and Y.-L. Zhou, Classification of Abelian domain walls, [Phys. Rev. D](https://doi.org/10.1103/PhysRevD.106.075019) 106 [\(2022\) 075019](https://doi.org/10.1103/PhysRevD.106.075019) [[2205.11529](https://arxiv.org/abs/2205.11529)].
- [39] Y. Wu, K.-P. Xie and Y.-L. Zhou, Collapsing domain walls beyond Z2, [Phys. Rev. D](https://doi.org/10.1103/PhysRevD.105.095013) 105 (2022) 095013 [[2204.04374](https://arxiv.org/abs/2204.04374)].
- [40] G.B. Gelmini, S. Pascoli, E. Vitagliano and Y.-L. Zhou, Gravitational wave signatures from discrete flavor symmetries, JCAP 02 [\(2021\) 032](https://doi.org/10.1088/1475-7516/2021/02/032) [[2009.01903](https://arxiv.org/abs/2009.01903)].
- [41] E. Ma and G. Rajasekaran, Softly broken A(4) symmetry for nearly degenerate neutrino masses, Phys. Rev. D 64 [\(2001\) 113012](https://doi.org/10.1103/PhysRevD.64.113012) [[hep-ph/0106291](https://arxiv.org/abs/hep-ph/0106291)].
- [42] G. Altarelli and F. Feruglio, Tri-bimaximal neutrino mixing, A(4) and the modular symmetry, [Nucl. Phys. B](https://doi.org/10.1016/j.nuclphysb.2006.02.015) 741 (2006) 215 [[hep-ph/0512103](https://arxiv.org/abs/hep-ph/0512103)].
- [43] C.L. Wainwright, CosmoTransitions: Computing Cosmological Phase Transition Temperatures and Bubble Profiles with Multiple Fields, [Comput. Phys. Commun.](https://doi.org/10.1016/j.cpc.2012.04.004) 183 [\(2012\) 2006](https://doi.org/10.1016/j.cpc.2012.04.004) [[1109.4189](https://arxiv.org/abs/1109.4189)].
- [44] F. Riva, Low-Scale Leptogenesis and the Domain Wall Problem in Models with Discrete Flavor Symmetries, [Phys. Lett. B](https://doi.org/10.1016/j.physletb.2010.05.073) 690 (2010) 443 [[1004.1177](https://arxiv.org/abs/1004.1177)].
- [45] J. Preskill, S.P. Trivedi, F. Wilczek and M.B. Wise, Cosmology and broken discrete symmetry, [Nucl. Phys. B](https://doi.org/10.1016/0550-3213(91)90241-O) 363 (1991) 207.