

Simultaneous retrieval of orbital phase resolved *JWST*/MIRI emission spectra of the hot Jupiter WASP-43b: evidence of water, ammonia, and carbon monoxide

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ABSTRACT

Spectroscopic phase curves of hot Jupiters measure their emission spectra at multiple orbital phases, thus enabling detailed characterization of their atmospheres. Precise constraints on the atmospheric composition of these exoplanets offer insights into their formation and evolution. We analyse four phase-resolved emission spectra of the hot Jupiter WASP-43b, generated from a phase curve observed with the Mid-Infrared Instrument/Low Resolution Spectrometer onboard the *JWST*, to retrieve its atmospheric properties. Using a parametric 2D temperature model and assuming a chemically homogeneous atmosphere within the observed pressure region, we simultaneously fit the four spectra to constrain the abundances of atmospheric constituents, thereby yielding more precise constraints than previous work that analysed each spectrum independently. Our analysis reveals statistically significant evidence of NH_3 (4σ) in a hot Jupiter's emission spectra for the first time, along with evidence of H_2O (6.5σ), CO (3.1σ), and a non-detection of CH_4 . With our abundance constraints, we tentatively estimate the metallicity of WASP-43b at $0.6\text{--}6.5\times$ solar and its C/O ratio at $0.6\text{--}0.9$. Our findings offer vital insights into the atmospheric conditions and formation history of WASP-43b by simultaneously constraining the abundances of carbon, oxygen, and nitrogen-bearing species.

Key words: radiative transfer – methods: numerical – planets and satellites: atmospheres – planets and satellites: individual: WASP-43b.

1 INTRODUCTION

The discovery of hot Jupiters stands as a testament to the diverse outcomes of planetary formation processes. Understanding the mech-

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anisms that give rise to their close-in orbits is crucial for advancing our knowledge of planetary formation, thus necessitating detailed characterization of their physical properties. Fortunately, because of their large atmospheric scale heights and favourable planet-to-host-star flux ratios, hot Jupiters are excellent targets for spectroscopic atmospheric characterization (e.g. Sing et al. 2016; Fisher & Heng 2018; Irwin et al. 2020; Mansfield et al. 2021; Ahrer et al. 2023). Precise constraints on hot Jupiters' atmospheric properties, such as molecular composition and thermal structure, further our knowledge of atmospheric physics under extreme conditions and grant us insight into their formation and migration pathways (Madhusudhan, Amin & Kennedy 2014; Mordasini et al. 2016; Madhusudhan et al. 2017; Dawson & Johnson 2018; Cridland et al. 2019b; Chachan et al. 2023). In recent years, phase-resolved emission spectroscopy, or 'spectroscopic phase curves', where the disc-integrated emission spectrum of a transiting exoplanet is observed at multiple orbital phases, has gained attention as a powerful observation strategy (Stevenson et al. 2014; Kreidberg et al. 2018; Irwin et al. 2020; Mikal-Evans et al. 2022). These phase curves are valuable data sets as they allow us to probe the longitudinal variation of atmospheric properties across a planet and mitigate the degeneracy between thermal structure and molecular abundances in atmospheric retrievals (Feng et al. 2016; Blecic, Dobbs-Dixon & Greene 2017; Taylor et al. 2020).

In this work, we retrieve the atmospheric properties of the transiting hot Jupiter WASP-43b (Hellier et al. 2011) using a phase curve observed with the Mid-Infrared Instrument/Low Resolution Spectrometer (MIRI/LRS) onboard the *JWST*. WASP-43b orbits its K7 host star in just 19.2 h (Gillon et al. 2012), making it a prime target for phase curve observations. The planet is presumably tidally locked due to the proximity to its host star, allowing us to convert the orbital phase to the central longitude of the observed hemisphere. Using the parametric atmospheric model of Yang, Irwin & Barstow (2023), we constrain the molecular abundances in the atmosphere of WASP-43b by simultaneously fitting four emission spectra generated at different orbital phases, with the assumption that atmospheric dynamics effectively homogenize the chemical composition in the pressure range probed by the observation (Cooper & Showman 2006; Agúndez et al. 2014; Drummond et al. 2018a, b; Mendonça et al. 2018b; Venot et al. 2020; Baeyens et al. 2021). By analysing the four spectra together, we are able to boost the signal-to-noise ratio and better constrain the molecular abundances than the previous analysis that fit each spectrum separately (Bell et al. 2024). We confirm previous detection of H_2O using the same data set (Bell et al. 2024) at higher statistical significance and additionally find statistically significant evidence of NH_3 and CO . Our work adds to a wealth of studies on WASP-43b, spanning telescopic observations (Czesla et al. 2013; Blecic et al. 2014; Kreidberg et al. 2014; Stevenson et al. 2014, 2017; Weaver et al. 2020; Fraine et al. 2021; Scandariato et al. 2022; Lesjak et al. 2023; Murphy et al. 2023; Bell et al. 2024), atmospheric retrievals (Changeat & Al-Refaie 2020; Feng, Line & Fortney 2020; Irwin et al. 2020; Cubillos et al. 2021; Chubb & Min 2022; Dobbs-Dixon & Blecic 2022; Taylor & Parmentier 2023; Yang et al. 2023) and atmospheric modelling (Kataria et al. 2015; Keating & Cowan 2017; Mendonça et al. 2018a, b; Carone et al. 2020; Helling et al. 2020; Venot et al. 2020; Schneider et al. 2022; Teinturier et al. 2024). Our simultaneous abundance constraints on water, methane, and ammonia can enable further studies on both the atmospheric chemistry and the formation history of WASP-43b (Fortney et al. 2020; Line et al. 2021; Ohno & Fortney 2023a, b).

This paper is structured as follows: In the methodology section (Section 2), we first outline how the emission spectra are generated from the *JWST* MIRI/LRS phase curve in 2.1, then intro-

duce our parametric atmospheric model in 2.2, our radiative transfer routine in 2.3, and the set-up of our atmospheric retrievals in 2.4. We present our results in Section 3, including the spectral fits in 3.1, the abundance constraints in 3.2 and the thermal structure constraints in 3.3. We compare our results to past *Hubble Space Telescope* (*HST*) and *Spitzer* observations in 3.4 and briefly discuss the implications of our results on the formation history of WASP-43b in 3.5. We explore the implications of our retrieved atmospheric properties with 1D chemical modelling in 3.6. We detail the limitations of the assumptions made in our analysis in Section 4 before concluding in Section 5.

2 METHODOLOGY

We describe the *JWST* observation in 2.1, followed by an outline of our atmospheric retrieval procedure, which is divided into three parts. First, we model the atmosphere of WASP-43b using the global parametric atmospheric model from Yang et al. (2023) based on a radiative-equilibrium temperature–pressure profile and assuming uniform molecular abundances (2.2). Secondly, we use a correlated- k radiative transfer pipeline to generate phase-resolved disc-integrated emission spectra from a given atmospheric model (2.3). Thirdly, we use a Bayesian parameter inference algorithm to constrain the parameters of our atmospheric model given the observed spectra and calculate the significance of the molecular detections from Bayesian evidence (2.4).

2.1 Observation

We analyse four disc-integrated emission spectra of WASP-43b derived from a phase curve observation with *JWST* MIRI in Low Resolution Spectroscopy slitless mode collected between 2022 December 1 and 2 (*JWST* Transiting Exoplanet Community Early Release Science Programme, *JWST*-ERS-1366; Stevenson et al. 2016; Bean et al. 2018; Bell et al. 2024). The original observation spans 5–12 μm and contains a full phase curve with two eclipses and one transit, lasting 26.5 h at a cadence of 10.34 s. However, we only fit the observation in the 5–10.5 μm region due to unresolved systematic effects in the 10.6–11.8 μm region (see 2.1.1). The spectra we analyse are the final fiducial spectra presented in Bell et al. (2024), which are the mean average of the spectra reduced using four different data reduction pipelines (Eureka! v1, Eureka! v2, TEATRO, SPARTA). The uncertainties are set as the mean uncertainty per wavelength bin, and then adding in quadrature the root mean square between the individual reductions and the mean spectrum, as was done in Bell et al. (2024).

2.1.1 'Shadowed region effect'

The original MIRI/LRS observation spanned 5–12 μm . However, the observation in the 10.6–11.8 μm region was affected by the so-called shadowed region effect of the MIRI instrument (Bell et al. 2024), which has not yet been robustly separated from the astrophysical phase variations, making the observation in this wavelength range unreliable for this particular observation (not all MIRI/LRS observations are affected, see Bell et al. 2024). Therefore, we exclude this wavelength region from our retrievals, in line with the retrieval analysis in Bell et al. (2024). This wavelength region contains strong NH_3 features; thus, successfully removing the systematic effect in the future will allow us to more confidently constrain the NH_3 abundance of WASP-43b and other hot Jupiters.

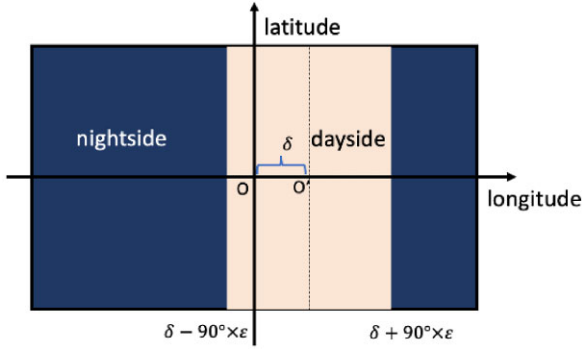


Figure 1. Schematics of our parametric 2D temperature model as defined by equation (1). The model divides the atmosphere into a dayside region and a nightside region, each modelled with a representative TP profile. The dayside central longitude and the dayside width are allowed to vary. ‘O’ marks the substellar point and ‘O’ marks the centre of the dayside region. While temperature is constant with longitude on isobars in the nightside region, we can parametrize the variation of temperature with longitude on the dayside. Note that the temperature is constant with latitude and only varies with pressure and longitude (we interpret the retrieved thermal structure as a latitudinal average, see 2.2).

2.2 Parametric atmospheric model

Our atmospheric model combines a parametric 2D temperature model (model 4 from Yang et al. 2023) with a simple chemistry model that assumes uniform gas volume mixing ratios (VMRs) in the atmospheric region probed by the observation. The temperature model is ‘2D’ in the sense that temperature only varies with pressure and longitude and is held constant with latitude. Of course, physically, we do not expect the temperature to be constant with latitude, but should instead fall off towards the poles. We interpret the thermal structure retrieved using this 2D model as a latitudinally averaged thermal structure, as Yang et al. (2023) found that the temperature structure retrieved from synthetic phase curve data simulated from a general circulation model (GCM) closely resembles the meridional mean thermal structure of the GCM calculated using $\cos(\text{latitude})$ as the weight. We choose not to model the latitudinal temperature variation since we are much more sensitive to longitudinal variation than latitudinal variation with our phase curve observation (the central longitude during the observation changes, whereas the central latitude is held constant). We note that the dayside latitudinal variation can be constrained with the eclipse mapping data of WASP-43b (Hammond et al. 2024), discussed in 4.1.

The 2D temperature model divides the atmosphere into a dayside region and a nightside region that are each described with a representative temperature–pressure (TP) profile (see Fig. 1). The width and the central longitude of the dayside region are free parameters, modelling the effects of atmospheric circulation on heat redistribution. These parameters can also mimic the impact of a cloud distribution across the atmosphere, which may cause an apparent hot spot offset (Parmentier, Showman & Fortney 2020). We allow the TP profile to vary with longitude on the dayside while keeping the temperature constant with longitude on isobars on the nightside, inspired by GCM studies which show approximately constant-with-longitude meridional mean thermal structures on the nightsides of hot Jupiters (e.g. fig. 10, Kataria et al. 2015; fig. 4, Mendonça et al. 2018a; fig. 7, Yang et al. 2023; fig. 2, Teinturier et al. 2024). Mathematically, the temperature T in our model atmosphere at pressure P , longitude

Λ , and latitude Φ is given by

$$T(P, \Lambda, \Phi) = \begin{cases} T_{\text{night}}(P) & \text{if } \Lambda > \delta + 90^\circ \times \epsilon \text{ or } \Lambda < \delta - 90^\circ \times \epsilon, \\ T_{\text{night}}(P) + (T_{\text{day}}(P) - T_{\text{night}}(P)) \cos^n\left(\frac{\Lambda - \delta}{\epsilon}\right) & \text{if } \delta + 90^\circ \times \epsilon \geq \Lambda \geq \delta - 90^\circ \times \epsilon, \end{cases} \quad (1)$$

where $T_{\text{day}}(P)$ and $T_{\text{night}}(P)$ are respectively the representative TP profiles for the dayside region and the nightside region, δ is the longitudinal deviation of the centre of the dayside region from the substellar point, and ϵ is a parameter governing the width of the dayside region. Our coordinate system is defined such that the substellar point (where the star would be perceived to be directly overhead) is at 0 degree longitude, and the antistellar point is at -180 degree longitude. The dayside region in our model is bound by the meridians $\Lambda = \delta - 90^\circ \times \epsilon$ and $\Lambda = \delta + 90^\circ \times \epsilon$, and the rest of the model is the nightside region (see Fig. 1). We reiterate that the ‘dayside region’ in this model can deviate from the illuminated region of the planet. The exponent n of the cosine term in equation (1) additionally prescribes how strongly temperatures vary with longitude on isobars on the dayside.

Following Yang et al. (2023), we use the 1D radiative equilibrium temperature profile given by equation (29) of Guillot (2010) to set $T_{\text{day}}(P)$ and $T_{\text{night}}(P)$. While the derivation of this TP profile assumes radiative equilibrium, we find in Yang et al. (2023) that it is flexible enough to approximate the temperature profiles calculated by a GCM in the limited pressure range currently probed by low-resolution emission spectroscopy. The Guillot (2010) profile is given by

$$T^4(P) = \frac{3T_{\text{int}}^4}{4} \left(\frac{2}{3} + \tau \right) + \frac{3T_{\text{irr}}^4}{4} f \left[\frac{2}{3} + \frac{1}{\gamma\sqrt{3}} + \left(\frac{\gamma}{\sqrt{3}} - \frac{1}{\gamma\sqrt{3}} \right) e^{-\gamma\tau\sqrt{3}} \right], \quad (2)$$

where τ is the infrared optical depth given by

$$\tau(P) = \frac{\kappa_{\text{th}} P}{g}. \quad (3)$$

This TP profile contains four free parameters: κ_{th} is the mean infrared opacity, γ is the ratio between the mean visible and mean infrared opacities, T_{int} is the internal heat flux, and f is a catch-all parameter of order unity that models the effects of albedo and the redistribution of stellar flux due to atmospheric circulation. We assume a negligible change in gravity g in the pressure range probed by the observation for the computation of the TP profile,¹ so that τ is linear in P . T_{irr} is the irradiation temperature defined by

$$T_{\text{irr}} = \left(\frac{R_{\text{star}}}{a} \right)^{1/2} T_{\text{star}}, \quad (4)$$

where a is the orbital semimajor axis, and R_{star} and T_{star} are the host star radius and temperature, respectively. In our model, the TP profile parameters for $T_{\text{day}}(P)$ and $T_{\text{night}}(P)$ are set independently. In summary, our temperature model has eleven free parameters: eight parameters for the two representative TP profiles, plus three parameters that characterize the dayside longitudinal thermal structure (see Table 1).

For the chemistry component of our atmospheric model, we assume the atmosphere is H_2/He dominated, with a fixed Jupiter-like H_2 to He VMR ratio at 86:14 (Von Zahn & Hunten 1996; Niemann et al. 1998). Our model includes five spectrally active molecules whose VMRs are free parameters: H_2O , CO_2 , CO , CH_4 , and NH_3 . These molecules are expected at potentially detectable abundance in

¹For our best-fitting model shown in Section 3, the change in g from the highest pressure level to the lowest pressure level is about 5 per cent.

Table 1. Free parameters of our atmospheric model and their priors. Compared to Yang et al. (2023), we additionally include NH_3 , an error inflation parameter b , and extend our priors for the $\log_{10}$ VMRs from $(-8, -2)$ to $(-10, -2)$ to account for the generally higher opacities in the MIRI bandpass compared to the *HST*/Wide Field Camera 3 (WFC3) bandpass modelled in Yang et al. (2023).

Parameter	Description	Prior	Unit
δ	Dayside longitudinal offset	$U(-45, 45)$	degree
ε	Dayside longitudinal width scaling	$U(0.5, 1.2)$	dimensionless
n	Dayside longitudinal variation exponent	$U(0, 2)$	dimensionless
$\log \kappa_{\text{th,day}}$	$\log_{10}$ mean infrared opacity (dayside)	$U(-4, 2)$	$\text{m}^2 \text{kg}^{-1}$
$\log \gamma_{\text{day}}$	$\log_{10}$ ratio of visible and infrared opacities (dayside)	$U(-4, 1)$	dimensionless
$\log f_{\text{day}}$	$\log_{10}$ heat redistribution parameter (dayside)	$U(-4, 1)$	dimensionless
$T_{\text{int,day}}$	Internal heat flux temperature (dayside)	$U(100, 1000)$	Kelvin
$\log \kappa_{\text{th,night}}$	$\log_{10}$ mean infrared opacity (nightside)	$U(-4, 2)$	$\text{m}^2 \text{kg}^{-1}$
$\log \gamma_{\text{night}}$	$\log_{10}$ ratio of visible and infrared opacities (nightside)	$U(-4, 1)$	dimensionless
$\log f_{\text{night}}$	$\log_{10}$ heat redistribution parameter (nightside)	$U(-4, 1)$	dimensionless
$T_{\text{int,night}}$	Internal heat flux temperature (nightside)	$U(100, 1000)$	Kelvin
b	Error inflation parameter	$U(1, 5)$	dimensionless
$\log \text{VMR}_{\text{H}_2\text{O}}$	$\log_{10}$ volume mixing ratio of H_2O	$U(-10, -2)$	dimensionless
$\log \text{VMR}_{\text{CO}_2}$	$\log_{10}$ volume mixing ratio of CO_2	$U(-10, -2)$	dimensionless
$\log \text{VMR}_{\text{CO}}$	$\log_{10}$ volume mixing ratio of CO	$U(-10, -2)$	dimensionless
$\log \text{VMR}_{\text{CH}_4}$	$\log_{10}$ volume mixing ratio of CH_4	$U(-10, -2)$	dimensionless
$\log \text{VMR}_{\text{NH}_3}$	$\log_{10}$ volume mixing ratio of NH_3	$U(-10, -2)$	dimensionless

the atmosphere of WASP-43b from chemical modelling (Mendonça et al. 2018b; fig. 5, Venot et al. 2020; fig. 18, Baeyens et al. 2021) and given their significant opacities in the MIRI wavelengths (Fig. 2). We assume that the VMRs of the spectrally active molecules are constant with altitude, longitude, and latitude based on previous modelling works (e.g. Cooper & Showman 2006; Agúndez et al. 2014; Baeyens et al. 2021), which suggest that atmospheric circulation should effectively homogenize the molecular abundances in hot Jupiters in the pressure range ($10\text{--}10^{-3}$ bar) typically probed by low-resolution emission spectroscopy.

The fixed stellar and planetary parameters used in this study are listed in Table 2, and the free parameters of our model are listed in Table 1.

2.3 Radiative transfer

We use the correlated- k method for radiative transfer (Lacis & Oinas 1991), following Irwin et al. (2008). The molecular opacity data are taken from the ExoMolOP database (Chubb et al. 2021). We generate channel-integrated k -tables at the resolution of the spectra (constant $0.5 \mu\text{m}$ bins) from the original ExoMolOP k -tables, which were computed at a resolving power of $R = 1000$. We model the molecular absorption due to the five spectrally active molecules in our model and the collision-induced absorption due to $\text{H}_2\text{--H}_2$ pairs and $\text{H}_2\text{--He}$ pairs. The full list of opacity data is given in Table 3. We calculate the disc-integrated emission spectra at arbitrary orbital phase using the method of Irwin et al. (2020) with five zenith angle quadratures, which takes into account the variation of emission angle and optical path across the atmosphere. The disc-integrated spectrum at a given orbital phase is computed assuming a circular, tidally locked and edge-on orbit, using the parameters in Table 2. The spectral calculations are implemented by the NEMESISPY package² following Yang et al. (2023), which is a Python development of the Fortran NEMESIS library (Irwin et al. 2008).

2.4 Retrieval set-up

We perform retrievals on the emission spectra of WASP-43b at four different orbital phases simultaneously. Following Bell et al. (2024), the spectra are binned to eleven $0.5 \mu\text{m}$ wavelengths bins in the $5\text{--}10.5 \mu\text{m}$ wavelength range, at orbital phase 0 (nightside), 0.25 (observer facing the evening terminator), 0.5 (dayside), and 0.75 (observer facing the morning terminator). We fit the four spectra simultaneously using spectra calculated from our atmospheric model described in 2.2 via the radiative transfer procedure described in 2.3. We convert planetary emission to planet-to-star flux ratio using a stellar spectrum computed using a PHOENIX stellar model (Allard & Hauschildt 1995; Hauschildt, Allard & Baron 1999; Husser et al. 2013), generated assuming an effective temperature of 4300 K, a surface gravity of $\log(g) = 4.50$, and a solar metallicity.

The atmospheric model is defined from 20 to 10^{-5} bar, on 40 points equally spaced in log pressure. We retrieve the eleven free temperature model parameters, the five molecular VMRs, and an error inflation parameter described in 2.4.1. The priors of these parameters are listed in Table 1. We find the posterior distribution of these parameters given the data using the Nested Sampling algorithm (Feroz & Hobson 2008) as implemented by the PYMULTINEST software package (Buchner et al. 2014), using 1000 sampling live points. For each molecule, we assess the significance of its detection via Bayesian model comparison by performing five additional retrievals, each omitting one of the five molecules in our fiducial model. The detection significance of a molecule can be calculated by computing the Bayes factor of the fiducial model over the model without the particular molecule, following Trotta (2008) and Benneke & Seager (2013).

2.4.1 Error inflation

We include an error inflation parameter to account for underestimated measurement and forward modelling uncertainties, following the approach used in, for example, Line et al. (2015) and Bell et al. (2024). We also refer the reader to Irwin et al. (2008), where the forward modelling error is treated in the optimal estimation framework. The error inflation is implemented by modifying the likelihood

²<https://pypi.org/project/nemesispy>

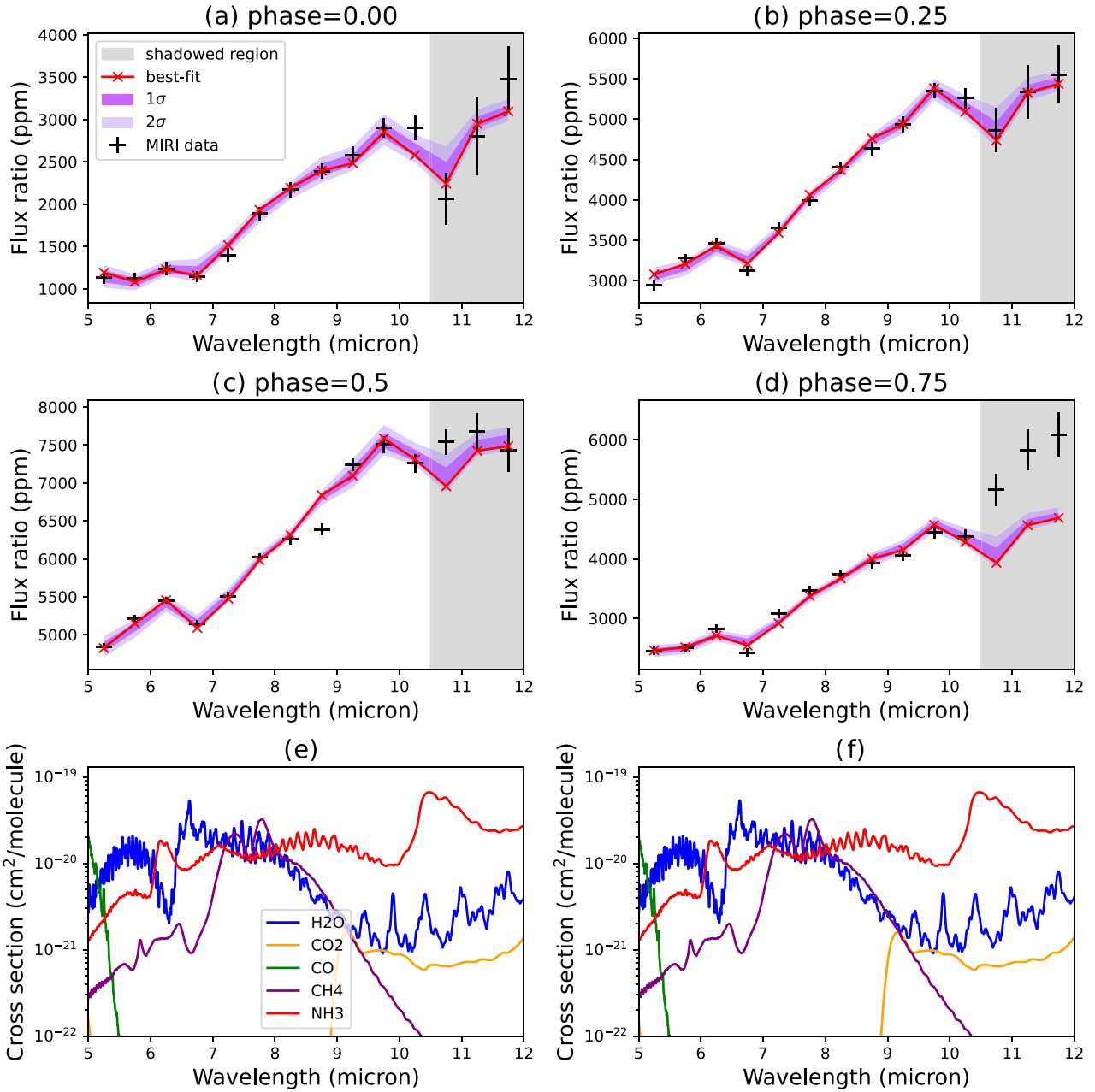


Figure 2. (a)–(d): Best-fitting model spectra at the four orbital phases calculated from the maximum a posteriori parameters. We plot the 1σ and 2σ central credible intervals with dark and light purple shading, respectively. The data in the grey shaded region (10.5–12 μm) are affected by the ‘shadowed region’ systematics and are not included in our retrievals (see 2.1.1). Interestingly, our best-fitting model, when extended to 12 μm , is consistent with the current reduction of the ‘shadowed region’ data for all phases except at phase 0.75 (see 3.1 for why we think phase 0.75 is worst affected by the ‘shadowed region effect’). Since NH_3 has strong spectral features in the ‘shadowed region’, a reliable reduction of the ‘shadowed region effect’ in the future can refine the abundance constraints on NH_3 . Note the absorption feature at the 8.75 μm bin at phase 0.5, which our model cannot explain. (e)–(f): Cross-sections of the spectrally active molecules included in our retrievals, computed at a spectral resolution of $R=1000$ at 1 bar pressure and 1500 K temperature. (f) is identical to (e) except for the omission of legend for ease of reference. The spectral features between ~ 5 and ~ 8 μm are indicative of H_2O molecules, whereas the absence of spectral features between ~ 7 and ~ 9 μm allows us to rule out CH_4 at high abundance. While NH_3 has strong opacities throughout the MIRI/LRS bandpass, its retrieved abundance is about two orders of magnitude lower than H_2O . This means that the spectral features of H_2O would dominate over those of NH_3 in the wavelength regions where H_2O has higher or similar opacity compared to NH_3 . The notable exception is the wavelength region greater than ~ 8 μm , where the H_2O opacity is generally much lower than the NH_3 opacity. The constraints on CO are driven by the 5.25 μm bin; the absence of CO would increase the planetary flux at the 5.25 μm bin relative to the 5.75 μm bin. The constraints on CO_2 are driven by the data in 9–10.5 μm region.

Table 2. WASP-43b stellar and planetary parameters used in this study.

Parameter	Value	Reference
T_{star}	4300 K	Bell et al. (2024)
R_{star}	$0.667 R_{\odot}$	Gillon et al. (2012)
$[\text{Fe}/\text{H}]_{\text{star}}$	-0.01 dex	Gillon et al. (2012)
a	0.0153 au	Gillon et al. (2012)
M_{plt}	$2.034 M_{\text{Jup}}$	Gillon et al. (2012)
R_{plt}	$1.036 R_{\text{Jup}}$	Gillon et al. (2012)

Table 3. Opacity data used to calculate emission spectra in this study. Apart from the He and H_2 collision-induced absorption opacity, all data are downloaded from the ExoMolOP database (Chubb et al. 2021).

Molecule	Opacity data
H_2O	Polyansky et al. (2018)
CO_2	Yurchenko et al. (2020)
CO	Li et al. (2015)
CH_4	Yurchenko et al. (2017)
NH_3	Coles, Yurchenko & Tennyson (2019)
He	Borysow, Frommhold & Moraldi (1989) & Borysow & Frommhold (1989)
H_2	Borysow et al. (1989) & Borysow & Frommhold (1989)

term in Bayesian inference with a constant-with-wavelength multiplicative factor. The log likelihood function with error inflation is given by

$$\ln L(\mathbf{D}|\boldsymbol{\theta}) = -\frac{1}{2} \sum_{i=1}^n \frac{(D_i - F_i(\boldsymbol{\theta}))^2}{s_i^2} - \frac{1}{2} \ln(2\pi s_i^2), \quad (5)$$

where $\mathbf{D}$ is the set of point data points in the observed spectra, $F_i(\boldsymbol{\theta})$ is the i th data point in the model spectra given model parameters $\boldsymbol{\theta}$, and s_i is defined as

$$s_i = b * \sigma_i, \quad (6)$$

where σ_i is the uncertainty for the i th data point, and b is a free error inflation parameter between 1 and 5.

3 RESULTS

We describe the results of our retrievals in this section. In 3.1, we show that our parametric atmospheric model can explain the observation. In 3.2, we show that there is statistically significant evidence of H_2O , NH_3 , CO , weak evidence of CO_2 , and no detectable level of CH_4 . In 3.3, we plot the transmission weighting function of our best-fitting model atmosphere and the best-fitting thermal structure. We compare our results to past *HST* + *Spitzer* observations in 3.4 and briefly discuss the implication of our results on the formation history of WASP-43b in 3.5. We investigate the implications of our retrieved abundances and thermal structure on the atmospheric chemistry in 3.6. The posterior distribution of all our model parameters is shown in the appendix (A1), alongside a comparison of our retrieved H_2O abundance to a list of previous studies (B1).

3.1 Spectral fit

We plot the best-fitting model spectra against the observed emission spectra in Fig. 2. The best-fitting model spectra are calculated from the maximum a posteriori parameters (which in our case are the same as the maximum likelihood parameters), and we propagate the $1\sigma/2\sigma$ central credible intervals to show the spread of our posterior

distribution. Our best-fitting model can fit almost all observed data points within the original uninflated error bars in Bell et al. (2024), which is also reflected by the fact that the error inflation parameter we retrieve is only around 1.65 (see 2.4.1 and A1). One notable exception is the $8.75 \mu\text{m}$ bin at phase 0.5, which cannot be accounted for by any of the opacity sources in our atmospheric model (see 4.3 for more discussion). For ease of reference, we include plots of the cross-sections of the molecules included in our retrievals in the MIRI/LRS wavelength range in Fig. 2. Note that the absorption due to each molecule is also weighted by its VMR. Chemical models predict that the VMRs of H_2O and CO in the atmosphere of WASP-43b are at least two orders of magnitude higher than the VMRs of CO_2 , CH_4 , and NH_3 (e.g. Venot et al. 2020), so we expect H_2O opacity to dominate in the $5\text{--}8 \mu\text{m}$ region. On the other hand, NH_3 opacity should be comparable to the H_2O opacity in the $8\text{--}10.5 \mu\text{m}$ region, allowing us to constrain the abundance of NH_3 . For the carbon-bearing molecules, as we can see from panel (e) of Fig. 2, the constraints on CO could only come from the $5.25 \mu\text{m}$ bin, whereas the constraints on CO_2 would be driven by the $9\text{--}10.5 \mu\text{m}$ region. We note that a more robust determination of the relative contribution of different wavelength bins to the molecular detections can be achieved by a leave-one-out cross-validation analysis, as detailed in Welbanks et al. (2023).

We reiterate that our retrievals exclude the data in the ‘shadowed region’, which is shaded in grey in Fig. 2. Therefore, we do not fit for the wavelength region longer than $10.5 \mu\text{m}$, so our retrieval analysis excludes most of the strong NH_3 features centred around $10.5 \mu\text{m}$. Interestingly, our best-fitting model, when extended to the $10.5\text{--}12 \mu\text{m}$ ‘shadowed region’, seems to be consistent with the current reduction of the ‘shadowed region’ data for three out of four orbital phases (phases 0, 0.25, and 0.5). Because the severity of the ‘shadowed region effect’ (see 2.1.1) decreased as a function of time since the beginning of the MIRI observation, we expect phases 0.5 and 0.75 to be more severely affected as the observation started just before a secondary eclipse (where phase = 0.5). However, since phase 0.5 was observed again at the end of the observation, we may expect it to be less affected than phase 0.75, which was only observed once. On the other hand, we would expect phases 0 and 0.25 to be least affected since they were observed later when much of the ‘shadowed region effect’ had decayed away. Therefore, if we assume that our best-fitting model is representative of the true atmosphere of WASP-43b and the ‘shadowed region’ data is more reliable at phases 0, 0.25, and 0.5 than at phase 0.75, then we expect our best-fitting model to be more consistent with the ‘shadowed region’ data at phases 0, 0.25, and 0.5 than at phase 0.75. As shown in Fig. 2, our best-fitting model agrees well with the data in the ‘shadowed region’ for phase 0 and phase 0.25, and slightly less so at phase 0.5. In contrast, our best-fitting model does not agree with the ‘shadowed region’ data at phase 0.75. Considering the time dependence of the ‘shadowed region effect’, the data in the ‘shadowed region’ is consistent with our best-fitting model and the presence of NH_3 . The important caveat is that the poorly fit systematic noise could unexpectedly impact the data reduction in Bell et al. (2024), so we do not want to draw any definitive conclusions from the affected data until the systematics issue is fully resolved. Further observations of WASP-43b using other JWST instruments at shorter wavelengths such as the scheduled Near Infrared Spectrograph (NIRSpec) observation (Guaranteed Time Observations 1224, PI: S. Birkmann), re-observations of WASP-43b using MIRI after the (still unknown) cause of the ‘shadowed region effect’ has been mitigated, or a robust reduction of the ‘shadowed region effect’ can all be used to confirm our NH_3 signal and tighten the constraints on its abundance.

3.2 Abundance constraints

We show the retrieved posterior distribution of the molecular VMRs in Fig. 3. We mark our best-fitting model with red solid lines and a solar-metallicity GCM-based chemistry model in orange solid lines for reference (R. Baeyens, private communication). Using the Bayesian evidence ratio test described in 2.4, we find statistically significant evidence of H_2O at 6.5σ , NH_3 at 4σ , and CO at 3.1σ . We find weak evidence of CO_2 (2.5σ) and no detectable level of CH_4 (95 per cent VMR upper limit at $10^{-6.39}$). We note that our analysis is one of the first studies to report evidence of NH_3 in the atmosphere of an exoplanet (MacDonald & Madhusudhan 2017a, b; Giacobbe et al. 2021; Dyrek et al. 2024), and to our knowledge, the first to report evidence of NH_3 by analysing emission spectra. We list the detection significance and the abundance constraints (given as 1σ central credible intervals) in Table 4.

As shown in Fig. 3, the retrieved H_2O and CO abundances are broadly consistent with a solar metallicity atmosphere. We constrain the metallicity and C/O ratio of WASP-43b using the posterior distribution of the molecular VMRs in Fig. 4, taking the reference solar abundances from Lodders (2010). Note that our metallicity estimate is calculated taking into account all of the molecules included in our retrievals, as listed in Table 4, thus taking into account O, C, and N elemental abundances. We tentatively estimate the metallicity of WASP-43b at $1.6^{+4.9}_{-1.0} \times$ solar and its C/O ratio at $0.8^{+0.1}_{-0.2}$. We find that our retrieved abundance constraints, and the inferred metallicity and C/O ratio are very similar to the results of Lesjak et al. (2023), which were based on high-resolution dayside emission spectra observed with Very Large Telescope/CRYogenic InfraRed Echelle Spectrograph+ (VLT/CRIRES+), thus providing validation to our results with an independent observational method. We discuss the implications of our retrieved abundances on the formation of WASP-43b in Section 3.5.

We caution the reader that our metallicity and C/O ratio estimations are only tentative for the following reasons. First, our CO constraints come only from the shortest $5.25 \mu\text{m}$ wavelength bin (Fig. 2). Since the shortest 1–2 wavelength bins are much more strongly illuminated, they may suffer more from poorly modelled non-linearity effects (Argyriou et al. 2023) in the data reduction stage of Bell et al. (2024). Furthermore, the detection significance of the carbon-bearing molecules are relatively low, as the MIRI/LRS wavelengths are not optimal for constraining the abundances of CO and CO_2 . The scheduled NIRSpec observation of WASP-43b (GTO 1224, PI: S. Birkmann) covering the $2.9\text{--}5.2 \mu\text{m}$ region will more robustly constrain the abundances of both CO and CO_2 , which will lead to a more accurate estimation of C/O ratio. We further note that cloud formation can also bias our estimated metallicity and C/O ratio. As pointed out by Lodders & Fegley (2002) and Woitke et al. (2018), the formation of silicate clouds can take oxygen away from gas-phase molecules. Therefore, if there is a significant level of silicate cloud formation on the nightside of WASP-43b, then our C/O ratio estimation would be an overestimate and our metallicity estimation would be an underestimate.

We now compare our abundance constraints to those in Bell et al. (2024) in Fig. 5, which were obtained by performing retrievals on the spectrum at each orbital phase separately and using a dilution parameter to account for spatial inhomogeneity in thermal structure (Taylor et al. 2020). We confirm the detection of H_2O at a much higher significance level (6.5σ) than the analysis in Bell et al. (2024), which reported significance ranging from 1.8σ to 4.1σ depending on the orbital phase and the retrieval pipeline. Our constraints are much more precise than those in Bell et al. (2024) and are most consistent

with their HYDRA retrieval pipeline results. For the retrievals in Bell et al. (2024) that included H_2O abundance as a free parameter, we note that all three retrieval pipelines show results consistent with our model assumption that the H_2O abundance is constant across all orbital phases. We also confirm the non-detection of CH_4 in Bell et al. (2024) at a 95 per cent upper limit of $10^{-6.39}$. In addition to tightening the abundance constraints reported in Bell et al. (2024), we are able to report statistically significant evidence of two additional molecules, namely NH_3 and CO. Our results demonstrate the power of phase-resolved spectroscopy in characterizing the chemistry of hot Jupiters in detail, as well as the importance of a global model in interpreting phase curves.

3.3 Thermal structure

We plot the transmission weighting function of our best-fitting model at the substellar point and the nightside in Figs 6 and 7, respectively. We see that our retrieval is primarily sensitive to pressures inside the 1 bar to 1 millibar pressure range. We plot the retrieved best-fitting temperature profile as a function of pressure and longitude in the 1 bar to 1 millibar pressure range in Fig. 8, as the profile outside this pressure range is not directly constrained by the data and is rather set by our parametrization. The retrieved temperature profile should be interpreted as a latitudinally averaged thermal structure: Yang et al. (2023) find that the retrieved thermal structure from GCM-simulated phase curves closely resemble the latitudinally averaged GCM thermal structure using $\cos(\text{latitude})$ as the weight.³ This weighting scheme is also commonly used in GCM studies to calculate meridionally averaged temperature structure (e.g. Kataria et al. 2015; Mendonça et al. 2018a; Carone et al. 2020).

The best-fitting thermal structure, as shown in Fig. 8, consists of a hot dayside region and a cooler nightside region, where the centre of the hot dayside region is moderately offset to the east. The eastward shift of our retrieved thermal structure is due to the fact that the peak of the phase curve occurred before the secondary eclipse (Bell et al. 2024), and similar phase curve maximum offsets have been observed in numerous hot Jupiters (e.g. Parmentier & Crossfield 2017, fig. 4). The eastward shift has been predicted and explained by GCM studies (Showman & Guillot 2002; Showman et al. 2009; Showman & Polvani 2011), which proposed that strong equatorial super-rotating (eastward) jets would develop in the atmospheres of hot Jupiters and act to transport heat eastwards, so the hottest region of the atmosphere is displaced eastward of the substellar point. This offset is also seen in numerous GCM simulations specific to the planetary parameters of WASP-43b (Kataria et al. 2015; Mendonça et al. 2018a, b; Carone et al. 2020; Schneider et al. 2022; Teinturier et al. 2024). We retrieve an eastward hotspot offset of $8.68^{+0.96}_{-1.14}$ degrees, consistent with the $7.34^{+0.38}_{-0.38}$ degrees eastward offset in Bell et al. (2024).⁴ We note that our best-fitting model in Fig. 8 resembles some GCM simulations of WASP-43b, for example, fig. 10 in Kataria et al. (2015) and fig. 4 in Mendonça et al. (2018a).

We plot the retrieved representative nightside and dayside TP profiles T_{night} and T_{day} in Fig. 9. We find large day–night temperature

³In Irwin et al. (2020), the authors also found that the retrieved thermal structure from GCM-simulated phase curves using a different method closely resembles the latitudinally averaged GCM thermal structure using $\cos(\text{latitude})$ as the weight.

⁴We note that the offset of the phase curve maximum does not necessarily correspond to the longitudinal offset of the hottest atmospheric region if there are clouds, latitudinal thermal variation or more complex longitudinal thermal variation than permitted by our model (Parmentier et al. 2020).

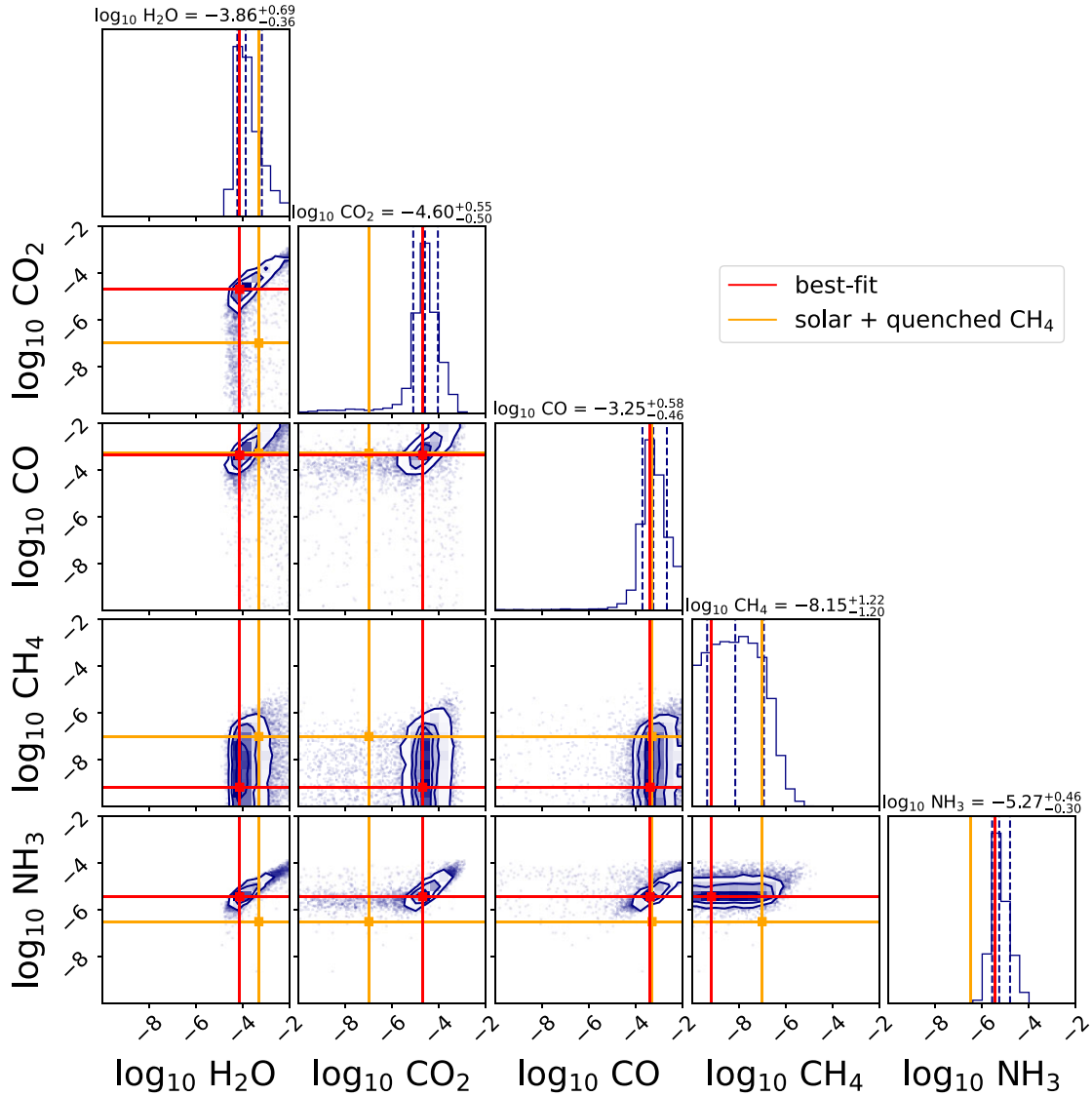


Figure 3. Posterior distribution of the molecular abundances. The histograms on the diagonal show the marginalized posterior distributions of each molecular VMR, where the blue dashed lines give the 16 percent, 50 percent, and 84 percent percentiles (the percentile values are given at the top of each panel). The off-diagonal plots show the joint posterior distributions of any pairs of parameters. The red solid lines mark our retrieved best-fitting parameters (maximum a posteriori parameters). The orange solid lines mark the expected $\log_{10}$ VMR of the molecules in a solar metallicity atmosphere, taking horizontal and vertical mixing into account: -3.3 for H_2O and CO , -7 for CO_2 and CH_4 , and -6.5 for NH_3 (Venot et al. 2020; Baeyens et al. 2021). While the abundances of H_2O , CO , and CO_2 in the chemical models of Baeyens et al. (2021) and Venot et al. (2020) are approximately constant with altitude, there are some vertical variations in the abundances of CH_4 and NH_3 in the pressure region probed by our observation. We take the ‘solar abundances’ of CH_4 and NH_3 to be roughly the values at 1 bar.

Table 4. Detection significance and abundance constraints of molecules included in our retrieval model. For all molecules apart from CH_4 , we quote the medians of the marginalized posterior distributions, with the uncertainties showing the 1σ central credible intervals. For CH_4 , we quote the 95 percent upper limit.

Molecule	Detection significance	$\log_{10}(\text{VMR})$
H_2O	6.5σ	$-3.86^{+0.69}_{-0.36}$
CO_2	2.5σ	$-4.60^{+0.55}_{-0.50}$
CO	3.1σ	$-3.25^{+0.58}_{-0.46}$
CH_4	NA	<-6.39
NH_3	4σ	$-5.27^{+0.46}_{-0.30}$

contrasts of at least 700 K at all pressure levels. As shown by Fig. 9, our retrieved temperature profiles are broadly consistent with the TP profiles retrieved in Bell et al. (2024, see their fig. 4). While we do not model clouds in our retrievals, past studies (e.g. Burningham et al. 2017; Mollière et al. 2020) have shown that flexible TP profiles can mimic the spectral effects of clouds by becoming more isothermal. Our retrieved large day–night temperature contrast and relatively isothermal nightside TP profile are consistent with the presence of clouds on the nightside, which echo the results of Bell et al. (2024) that cloudy GCMs can match the nightside emission of WASP-43b better than cloudless GCMs. However, we cannot rule out or confirm the presence of nightside clouds directly with our model.

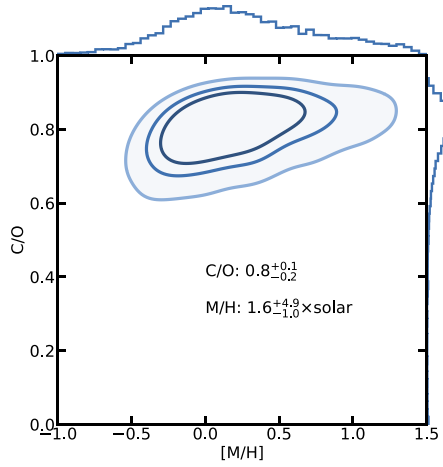


Figure 4. Joint distribution of metallicity and C/O ratio derived from the posterior distribution of all molecular VMRs in our model.

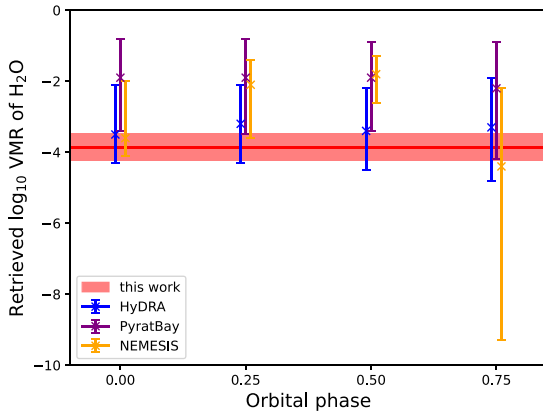


Figure 5. Comparison of the retrieved constraints on H_2O VMR in this work to those retrieved in Bell et al. (2024). The horizontal red solid line shows the median of the marginalized posterior distribution, and the red shaded region around it shows the 1σ central credible interval. We only include the retrievals in which the abundance of H_2O was a free parameter.

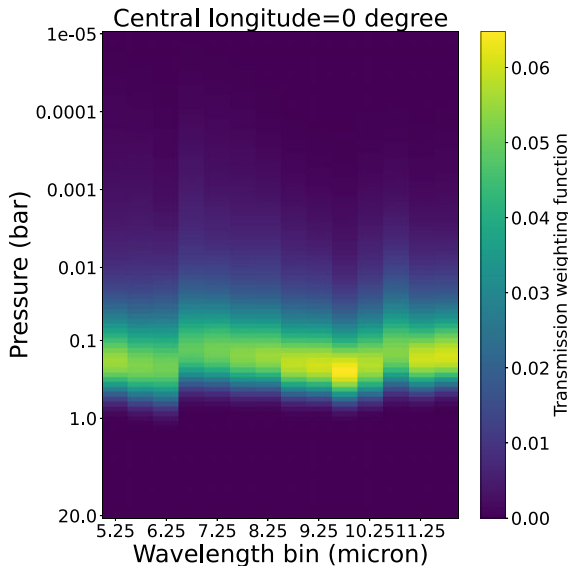


Figure 6. Transmission weighting function of the best-fitting retrieved model on the nightside.

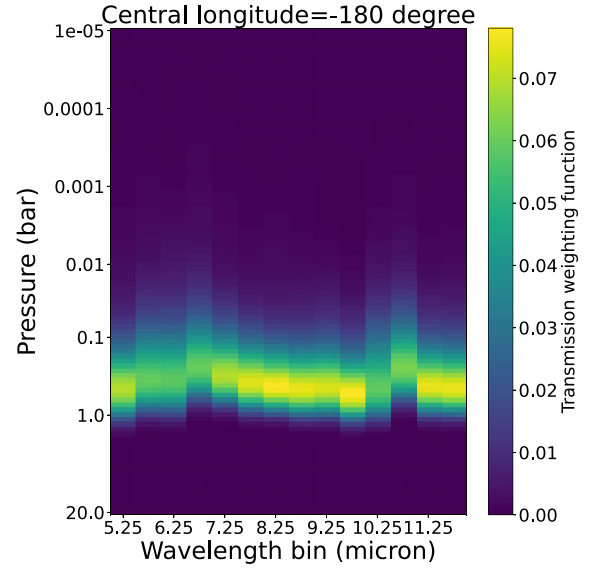


Figure 7. Transmission weighting function of the best-fitting retrieved model on the nightside.

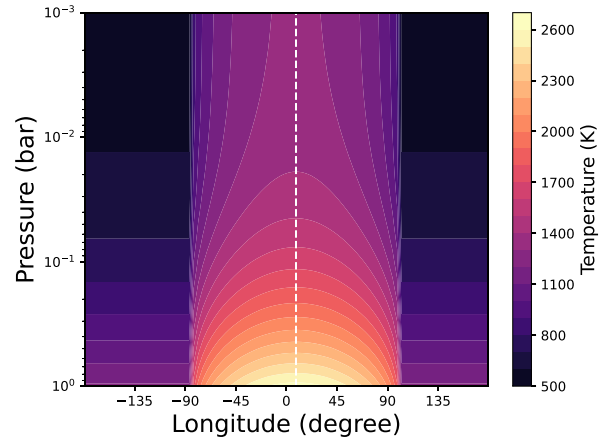


Figure 8. Best-fitting 2D thermal structure calculated from the maximum a posteriori parameters. The white vertical dashed line marks the position of the hotspot offset.

3.4 Comparison with *HST*/WFC3 and *Spitzer* observations

We compare the constraints on chemical abundances retrieved in this work to the abundance constraints retrieved from the *HST*/WFC3 + *Spitzer*/IRAC data set (Stevenson et al. 2017) in a previous study (Yang et al. 2023) in Fig. 10. The abundance constraints on H_2O retrieved from the MIRI/LRS data set are entirely consistent with those retrieved from the *HST*/WFC3 + *Spitzer*/IRAC data set. This agreement boosts our confidence in the H_2O detection in the atmosphere of WASP-43b. Moreover, as shown in Yang et al. (2023) and Irwin et al. (2020), the *HST*/WFC3 observation probes deeper in the atmosphere (to pressure levels as high as ~ 10 bar) than the MIRI/LRS observation analysed in this work. The consistent H_2O abundance retrieved from the two data sets reaffirms the assumption that the abundance of H_2O is constant with pressure in the observable pressure levels of hot Jupiters. We compare our retrieved H_2O abundance to more previous studies in Appendix B.

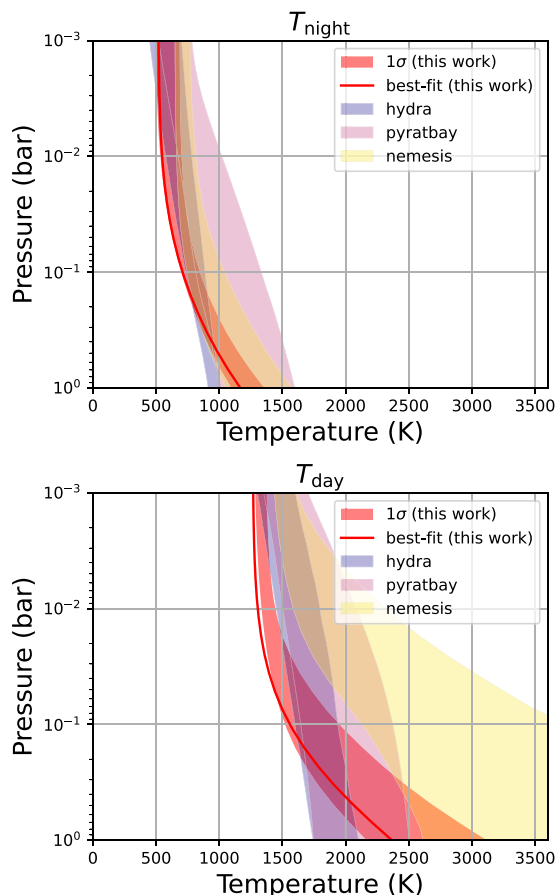


Figure 9. Comparison of the retrieved representative dayside/nightside TP profiles in this work to the retrieved dayside/nightside TP profiles in Bell et al. (2024). The best-fitting model (red line) is calculated from our maximum a posteriori parameters, whereas the 1σ credible intervals (red shaded areas) are propagated from the 1σ central credible intervals of our posterior distribution. The blue, purple, yellow shaded regions represent the 1σ credible intervals of the TP profiles retrieved using three different retrieval pipelines in fig. 4 of Bell et al. (2024). We only include the retrievals in which the chemical abundances were free parameters.

3.5 Implication on formation

The composition of the protoplanetary discs in which planets form is a function of time and location, so the final composition of the planets may be used to constrain their formation and migration history. As discussed in 3.2, our results tentatively suggest that the atmosphere of WASP-43b has a metallicity of $0.6\text{--}6.5\times$ solar, with a supersolar C/O ratio of $0.8^{+0.1}_{-0.2}$. We now discuss the implications of these estimates on the formation history of WASP-43b. The retrieved metallicity is consistent with the mass–metallicity relation in which more massive planets are expected to have lower metallicities (e.g. fig. 4, Kreidberg et al. 2014; fig. 2, Thorngren & Fortney 2019; fig. 12, Cridland et al. 2019b; fig. 3, Welbanks et al. 2019), as they can accrete and retain more gas directly from the disc and thus lower their metallicity. The retrieved C/O can inform us about the relative importance of solid accretion to gas accretion during formation, as the solid icy materials in protoplanetary discs tend to be O-rich, while the gas tends to be more C-rich. As mentioned above, since more massive planets such as WASP-43b (~ 2 Jupiter mass) accrete more gas from the disc, their atmospheres should naturally have supersolar C/O ratios. If we combine our retrieved metallicity and C/O ratio, we can broadly

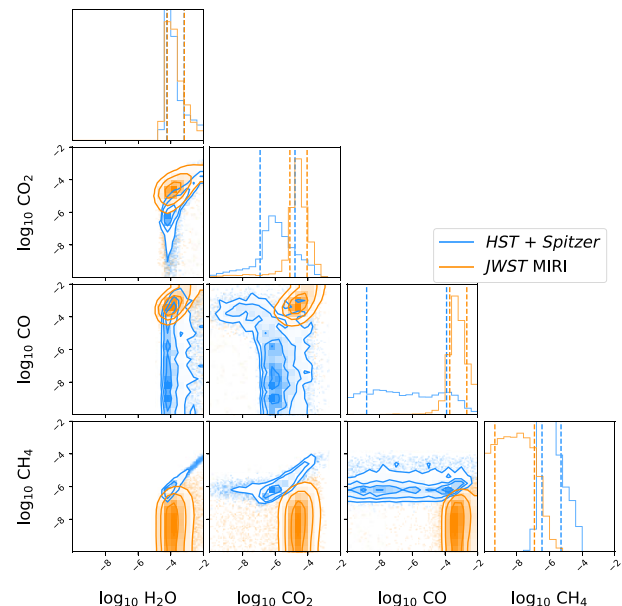


Figure 10. Comparison of posterior distributions from (1) *HST*/WFC3 and *Spitzer*/Infrared Array Camera (IRAC) data set analysed in Yang et al. (2023); (2) *JWST* MIRI/LRS data set analysed in this work.

constrain the formation environment of WASP-43b. As pointed out by Cridland, Eistrup & van Dishoeck (2019a), chemical processing of carbon-rich dust grains within the protoplanetary discs (Anderson et al. 2017; Gail & Tieloff 2017; Klarmann, Ormel & Dominik 2018) is a significant source of gaseous carbon and can impact the C/O ratios of the giant planets that form within. Our retrieved metallicity and C/O ratio favours the ‘ongoing’ model in Cridland et al. (2019a), where ongoing chemical processes continually erode carbon from dust grains in the protoplanetary disc (fig. 6, Cridland et al. 2020). Additionally, according to Madhusudhan et al. (2017), who studied the link between pebble accretion and C/O ratio, our retrieved C/O ratio is more consistent with a planet that formed outside the water ice line and then migrated through the disc, but probably not outside the CO_2 iceline as it would lead to substantially subsolar metallicity. The hypothesis that WASP-43b formed beyond the water iceline is supported by Mordasini et al. (2016), who argued that planets formed within the water iceline should have C/O ratio less than 0.2 assuming a standard disc chemistry model. The analysis presented here is a simplistic one since the impact of non-standard disc chemistry on the atmospheric composition of hot Jupiters can be highly uncertain, and we cannot rule out the possibility that the C/O ratio we retrieved is not representative of the entire atmosphere (see the discussion in 3.2). In addition to the C/O ratio, the N/O ratio has also gained interest as a way of connecting giant planet observations to their formation scenarios (Piso, Pegues & Öberg 2016; Cridland et al. 2020; Ohno & Ueda 2021; Turrini et al. 2021; Notsu et al. 2022), though additional modelling beyond the scope of this work is required to estimate the N/O ratio from our retrieved NH_3 abundance, as the abundance of N_2 may be significant but is not constrained by our model (see Ohno & Fortney 2023a, b).

3.6 1D chemical modelling

We examine if our retrieved dayside hot spot thermal structure in 3.3 can produce our retrieved molecular abundances in 3.3, assuming our tentatively estimated metallicity and C/O ratio in 3.2.

This is motivated by our assumption that the atmospheric chemical abundances are homogenized to the dayside values by circulation in the pressure region probed by our observation (Cooper & Showman 2006; Agúndez et al. 2014; Mendonça et al. 2018b; Venot et al. 2020; Baeyens et al. 2021). We employ a 1D chemical model (Tsai et al. 2017, 2021), which simultaneously takes into account thermochemistry, photochemistry, and vertical mixing. In particular, we want to explore if we can produce the level of NH_3 retrieved in 3.2, which is more than one order of magnitude higher than several model predictions for a WASP-43b like planet (Venot et al. 2020; Baeyens et al. 2021; figs 14 and 21 of Fortney et al. 2020). We extend our retrieved representative dayside TP profile in 3.3 down to 100 bar using the double-grey analytical expression of Heng, Mendonça & Lee (2014) to model the chemistry in the deep atmosphere, assuming a T_{int} of 595 K. We find that the temperature of our retrieved dayside profile T_{day} at 1 bar is so high that the vertical quenching of chemical abundances happens at approximately 1 bar. Therefore, the thermal structure below the 1 bar pressure level does not impact the photospheric composition, and the results of our 1D chemical modelling is insensitive to the particular value of T_{int} that we choose, since we always have to match our retrieved TP profile at around 1 bar. We then test three different metallicities corresponding to the median and the 1σ central credible interval bounds of our estimated metallicity while keeping the input C/O ratio at 0.8. We also experiment with different values of vertical eddy diffusion coefficient K_{zz} (e.g. Parmentier, Showman & Lian 2013) and N/H ratios to see if we can produce our retrieved NH_3 abundance using the 1D chemical model.

We now describe the results of our 1D chemical modelling. We first examine the dayside chemical abundances under thermochemical equilibrium without vertical mixing and photochemistry. As shown in Fig. 11, our retrieved H_2O and CO abundances are consistent with their equilibrium abundances, which are roughly uniform with pressure. Our retrieved CH_4 upper bound is also consistent with the dayside equilibrium abundance, which is lower than 10^{-6} for pressures lower than 10 bar. However, our retrieved abundances for CO_2 and NH_3 are much higher than the equilibrium abundances. We defer the investigation of CO_2 as its detection significance is relatively low and focus on the issue of NH_3 . We include the effect of photochemistry and vertical mixing on the equilibrium abundances in Fig. 12. We show three models that can increase the level of NH_3 compared to the equilibrium model: (1) $200\times$ enhanced N/H, $K_{zz} = 10^{13} \text{ cm}^2 \text{ s}^{-1}$ (dash-dotted lines); (2) $1000\times$ enhanced N/H, $K_{zz} = 10^8 \text{ cm}^2 \text{ s}^{-1}$ (solid lines); (3) $1000\times$ enhanced N/H, $K_{zz} = 10^{10} \text{ cm}^2 \text{ s}^{-1}$ (dashed lines). We illustrate two pathways to increasing the NH_3 abundance in the context of a one-dimensional model: invoking stronger vertical mixing to quench the NH_3 abundance to deeper pressure levels or increasing the overall N abundance in the atmosphere. In the context of our 1D chemical modelling, an improbably strong enhancement of N/H ($> 200\times$ enhanced N/H) is necessary to obtain our retrieved NH_3 abundance, even with unlikely high-vertical mixing. The real picture, with 3D circulation, is much more complex, and we plan to investigate the atmospheric chemistry of WASP-43b more comprehensively in a future paper.

4 LIMITATIONS

Given the constraints in data quality, in particular the lack of spatial resolution, and the gaps in our theoretical understanding of hot Jupiter atmospheres, we have to make a series of assumptions in our modelling. We now discuss the validity of our assumptions and how they impact the conclusions we draw from the data.

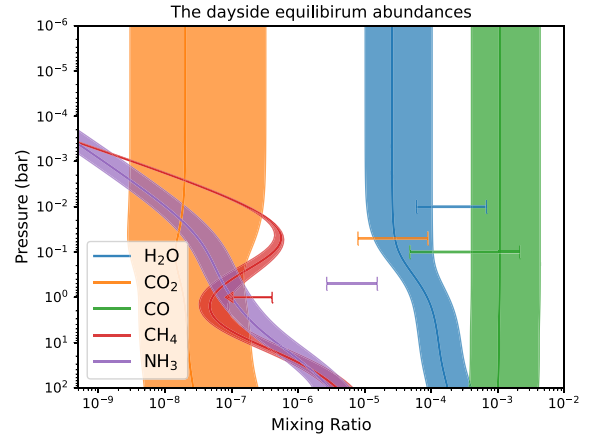


Figure 11. Comparison of dayside thermochemical equilibrium abundances with the chemical abundances retrieved from the MIRI data. The shaded regions represent a range of 1D thermochemical equilibrium models consistent with our estimated metallicity in 3.2. The horizontal error bars show the 1σ credible intervals of our retrieved molecular abundances using the MIRI data. The right edge of the red horizontal arrow shows the 95 per cent upper limit of CH_4 abundance retrieved using the MIRI data. Note that the vertical coordinates of the error bars and the arrow are arbitrary.

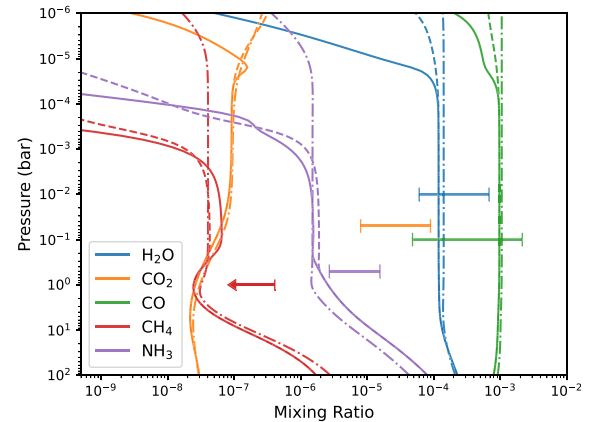


Figure 12. Dayside chemical abundances with vertical mixing and photochemistry. The solid and dashed lines show the model with $1000\times$ enriched N elemental abundance, with the solid lines showing $K_{zz} = 10^8 \text{ cm}^2 \text{ s}^{-1}$ and the dashed lines showing $K_{zz} = 10^{10} \text{ cm}^2 \text{ s}^{-1}$. The dash-dotted lines show the model with $200\times$ enriched N elemental abundance and $K_{zz} = 10^{13} \text{ cm}^2 \text{ s}^{-1}$. Note that the vertical coordinates of the error bars and the arrow are arbitrary.

4.1 Thermal structure

Our 2D parametric temperature model comes with various assumptions. We first assume that a latitudinally averaged structure can effectively model the phase curve data. In our 2D temperature model, temperature varies as a function of longitude and pressure while being held constant as a function of latitude on isobars (see equation 1 and Fig. 1). We then interpret the retrieved temperature map as a meridional mean temperature map, using $\cos(\text{latitude})$ as the weight. We base this interpretation on the fact that Irwin et al. (2020) and Yang et al. (2023) find the temperature structure retrieved from synthetic phase curve data simulated from a GCM closely resembles the latitudinally averaged thermal structure of the GCM using $\cos(\text{latitude})$ as the weight. Our temperature model further

assumes that the thermal structure is symmetric about the equator (North–South symmetry) and about the meridian going through the centre of the dayside region (East–west symmetry) and that the thermal structure on the nightside is more homogeneous than that on the dayside. As discussed in 2.2, this meridional mean structure is inspired by GCM studies, though we cannot explore some of the more complex mean thermal structures seen in GCM simulations, such as pressure-dependent hotspot offset. Our parametric modelling is only one of the first steps in understanding the thermal structure of hot Jupiters through observations, and we expect that as more *JWST* data come in, a more sophisticated model set-up will be necessary. Looking ahead, we can constrain the dayside latitudinal thermal variation by using eclipse mapping data (Rauscher et al. 2007; Challener & Rauscher 2023; Boone, Grant & Hammond 2024), where high-cadence flux measurements are taken as the star gradually eclipses the planet. In future work, we plan to incorporate MIRI eclipse mapping data of WASP-43b (Hammond et al. 2024) in retrievals to constrain the 3D thermal structure of WASP-43b.

4.2 Clouds

We assume that the atmosphere of WASP-43b is free from significant aerosols in the pressure range probed by the observation. The assumption that the dayside region contains no optically thick aerosols in the pressure ranges probed by MIRI is supported by numerous studies: transmission spectroscopy with *HST*/WFC3 (Kreidberg et al. 2014), inference of bond albedo from *HST*/WFC3 + *Spitzer*/IRAC phase curves (Stevenson et al. 2017), reflected light photometry with *HST* WFC3/Ultraviolet and Visible Light channel (UVIS) (Fraine et al. 2021), a joint analysis of *CHEOPS*, *TESS*, and *HST* WFC3/UVIS data (Scandariato et al. 2022), and a retrieval study of VLT/CRIRES + high-resolution dayside emission spectra (Lesjak et al. 2023). Furthermore, a cloud-free dayside photosphere is consistent with the high-dayside temperatures we retrieve in this work. On the other hand, a cloudy nightside has been inferred from large day–night brightness temperature contrast (Kataria et al. 2015; Stevenson et al. 2017; Mendonça et al. 2018a; Morello et al. 2019; Irwin et al. 2020). Bell et al. (2024) also find that cloudy GCMs fit the phase curve at nightside phases better than cloudless GCMs. Notably, both the retrieval analysis in Bell et al. (2024) and the retrieval analysis in this work show that the observed spectra can be adequately fit with cloud-free parametric atmospheric models. As discussed in 3.3, it has been shown that a more isothermal temperature profile can mimic the effects of cloud opacity in low-resolution spectral retrievals (Burningham et al. 2017; Mollière et al. 2020), which explains the ability of simple models to fit the observation without invoking cloud opacity. The omission of clouds from our model means that our retrieved nightside thermal structure would not be reliable if the nightside of WASP-43b is, in fact, cloudy. However, if our assumption that the atmospheric chemical abundances are constant with longitude and pressure holds, then the likely cloud-free dayside still anchors our retrieved chemical abundances. Since we retrieve spectra from four different orbital phases simultaneously, we do not expect our retrieved chemical abundances to be significantly biased. On the other hand, as mentioned in 3.2, even though our abundance constraints may not be biased in the presence of nightside clouds, a significant amount of oxygen may be removed from the gas-phase molecules if the cloud composition is rich in silicates (Woitke et al. 2018). Consequently, based on our gas-phase abundance constraints, we would underestimate the metallicity and overestimate the C/O ratio of WASP-43b.

4.3 Chemistry

We assume that the VMRs of all molecules in our model atmosphere are constant with longitude, latitude, and pressure. This assumption holds well for H_2O and CO , as both thermochemical equilibrium models and models that include photochemistry and mixing predict H_2O and CO abundances to be approximately uniform in the observed pressure region for a WASP-43b like planet (e.g. fig. 5, Venot et al. 2020). The assumption of uniform H_2O abundance is supported by the retrieval analysis in Bell et al. (2024), which showed no significant phase variation in H_2O abundance (see Fig. 5). On the other hand, the spatial distributions of CH_4 , NH_3 , and CO_2 in thermochemical equilibrium vary significantly with longitude and pressure, especially in the case of CH_4 , which would be present at a detectable level in the nightside of WASP-43b under thermochemical equilibrium. However, by coupling a simplified chemical kinetics scheme to a GCM of HD 209458b, Cooper & Showman (2006) showed that the CH_4 abundance should be homogenized between 1 bar and 1 millibar to be representative of the low-dayside abundance when strong zonal circulation is present. The uniform CH_4 prediction was later corroborated by Agúndez et al. (2014), who coupled a more realistic chemical network to a simplified circulation scheme that included both vertical mixing and a uniform zonal wind. Modelling specific to the planetary parameters of WASP-43b showed that dynamical time-scale dominates over chemical time-scale between 1 bar and 1 millibar (Mendonça et al. 2018b; Venot et al. 2020), and we expect no significant longitudinal variation in the abundances of CO_2 , CH_4 , and NH_3 . The assumption of constant-with-longitude CH_4 abundance is supported by the non-detection of CH_4 at all orbital phases in Bell et al. (2024). There is also circumstantial evidence of strong zonal circulation on WASP-43b, for example, the ~ 8 degrees east hotspot shift that is retrieved in this work and in Bell et al. (2024), and a wind speed of $3\text{--}8\text{ km s}^{-1}$ inferred from high-resolution emission spectroscopy (Lesjak et al. 2023). While theoretical models do predict some vertical variation in the abundances of CO_2 , CH_4 , and NH_3 , the variation in CO_2 with pressure is within one order of magnitude and thus likely insignificant compared to the uncertainties in its retrieved abundance. The vertical variations of CH_4 and NH_3 warrant further investigation and a parametrization of vertical quenching may be required if we are sensitive to a wide range of pressures. However, as shown in Figs 6 and 7, the transmission weighting function of our best-fitting model peaks strongly between 1 and 0.1 bar. Given the limited pressure range probed by our observation, we feel the modelling assumption that the abundances of CH_4 and NH_3 are constant with pressure is justified.

We assume a H_2/He dominated model atmosphere and only include five spectrally active molecules: H_2O , CO_2 , CO , CH_4 , and NH_3 . We chose this set of molecules based on chemical modelling of hot Jupiter atmospheres (Agúndez et al. 2014; Venot et al. 2020; Baeyens et al. 2021), taking into account the opacities of molecules, as well as guided by previous retrieval studies of WASP-43b such as Bell et al. (2024) and Yang et al. (2023). While this set of molecules adequately explains the spectra, we could not fit an absorption signature at $8.75\text{ }\mu\text{m}$ that only appears at phase 0.5. The investigation of more chemical species and spatial variation in chemistry is left for future work.

4.4 Radiative transfer

We assume that the correlated- k approximation for radiative transfer calculations (Lacis & Oinas 1991) introduces negligible errors in

modelling exoplanet spectra, as found by Mollière et al. (2015), Irwin et al. (2020), and Leconte (2021). Furthermore, we use channel-averaged k -tables binned to our data's $0.5\ \mu\text{m}$ spectral resolution. We assess the accuracy of this approximation by computing spectra from our best-fitting atmospheric model using k -tables at a resolution of $R=1000$. The results are in good agreement with the spectra calculated using k -tables at the spectral resolution of our data: the differences between the two spectra are found to be well within the measurement uncertainties. The errors introduced by our radiative transfer method are thus unlikely to bias the retrieved physical parameters significantly. Since our Bayesian inference algorithm requires the generation of hundreds of thousands of models to explore the model parameter space, a faster radiative transfer routine is much preferred if there is no significant loss in accuracy. None the less, as data quality advances, we need to more carefully assess the errors introduced in our forward modelling, particularly our radiative transfer calculation method, as well as the errors intrinsic to the opacity data we employ.

4.5 Data reduction

In this work, we are treating the spectra at different orbital phases as independent measurements to boost the signal-to-noise ratio of the retrieval. In reality, these spectra are correlated through the data reduction process, as phase curves are conventionally extracted by fitting a combination of Fourier series and systematics model to the time-series observations, wavelength by wavelength (Bell et al. 2024; Changeat et al. 2024). In effect, we have assumed that the measurement uncertainties of our spectra are Gaussian and independent, when in reality they are non-Gaussian and correlated. Consequently, we expect our posterior distribution to be biased. Given the goodness-of-fit of our model to the data in Fig. 2, we expect our retrieved molecular abundances are still consistent with the data. A full account of this issue is beyond the scope of this study, and we refer the reader to a possible resolution by Changeat et al. (2024), where the atmospheric retrieval is performed simultaneously with data reduction on the time-series observations.

5 CONCLUSION

We find evidence of H_2O (6.5σ), NH_3 (4σ), CO (3.1σ), CO_2 (2.5σ), and no evidence of CH_4 in the atmosphere of the hot Jupiter WASP-43b by retrieving four emission spectra observed at different orbital phases using the MIRI/LRS on the *JWST*. We fit the spectra simultaneously by employing a 2D atmospheric model accounting for a non-uniform temperature profile across the dayside and assuming uniform chemical abundances (Yang et al. 2023). Our approach more confidently and precisely constrains the molecular abundances than the previous retrievals in Bell et al. (2024), which analysed each orbital phase separately. Based on our abundance constraints, we tentatively estimate the metallicity of WASP-43b at $1.6^{+4.9}_{-1.0} \times$ solar and its C/O ratio at $0.8^{+0.1}_{-0.2}$, though we stress the need for further observations as the detection significance of the carbon-bearing molecules in this study are relatively low. In particular, the scheduled *JWST* NIRSpec observation (GTO 1224, PI: S. Birkmann) will likely revise these estimates by better constraining the abundances of CO and CO_2 , giving us an even clearer picture of the bulk chemistry of WASP-43b. Our tentative detection of NH_3 , seen for the first time in the emission spectra of a hot Jupiter, may be verified in the future by reliably reducing the $10.6\text{--}12\ \mu\text{m}$ region of our MIRI/LRS data that is excluded from our retrieval due to residual instrument systematics, or by further observations such as

the scheduled *JWST* NIRSpec observation. The precise abundance constraints in this work showcase the need for multidimensional atmospheric models to retrieve phase curve data and open up exciting avenues for follow-up studies on the atmospheric chemistry and formation history of WASP-43b.

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DATA AVAILABILITY

The data used in this paper are associated with *JWST* DD-ERS programme 1366 (PIs Batalha, Bean, and Stevenson; observation 11) and are publicly available from the Mikulski Archive for Space Telescopes (<https://mast.stsci.edu>). Additional intermediate and final results from Bell et al. (2024) that are analysed in this work will be archived on Zenodo at <https://zenodo.org/doi/10.5281/zenodo.10525170> upon final publication of Bell et al. (2024). All other data underlying this article will be shared on reasonable request to the corresponding author.

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APPENDIX A: FULL POSTERIOR

Fig. A1 shows the full posterior distribution of the retrieval model parameters listed in Table 1.

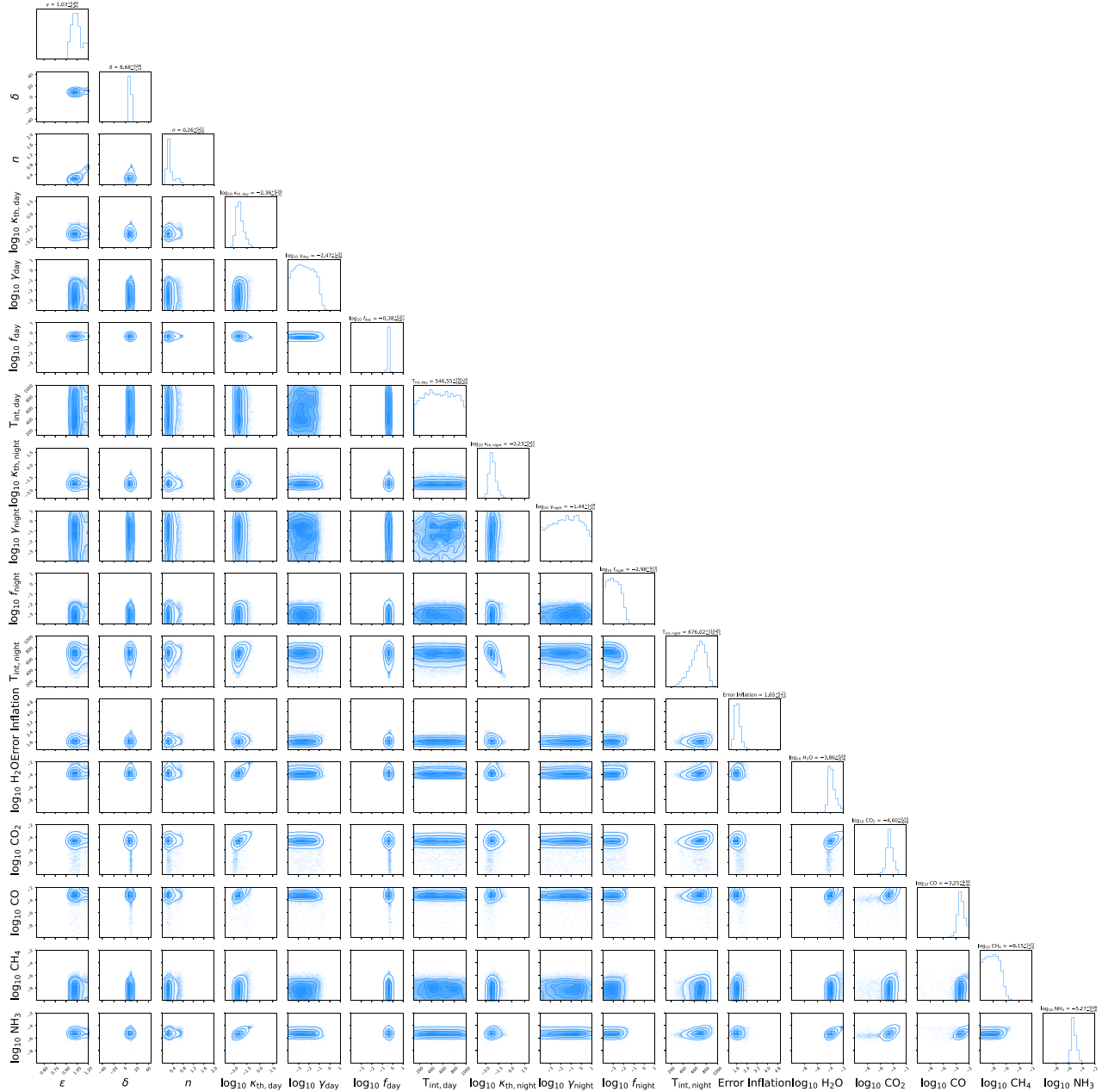


Figure A1. Posterior distribution of all the model parameters listed in Table 1. The diagonal plots are the marginalized posterior distributions, and the labels above show the posterior medians and 1σ central credible intervals.

APPENDIX B: COMPARISON WITH PAST STUDIES

We compare our retrieved H_2O abundance to past studies that retrieved H_2O abundance. Note that apart from Lesjak et al. (2023),

who analysed ground-based high-resolution spectra observed with VLT/CRIRES+, all other studies analysed low-resolution space telescope data. For Bell et al. (2024), we list the phase-by-phase retrieval results from the three retrieval pipelines that included H_2O as a free parameter (see also Fig. 5 in the main text).

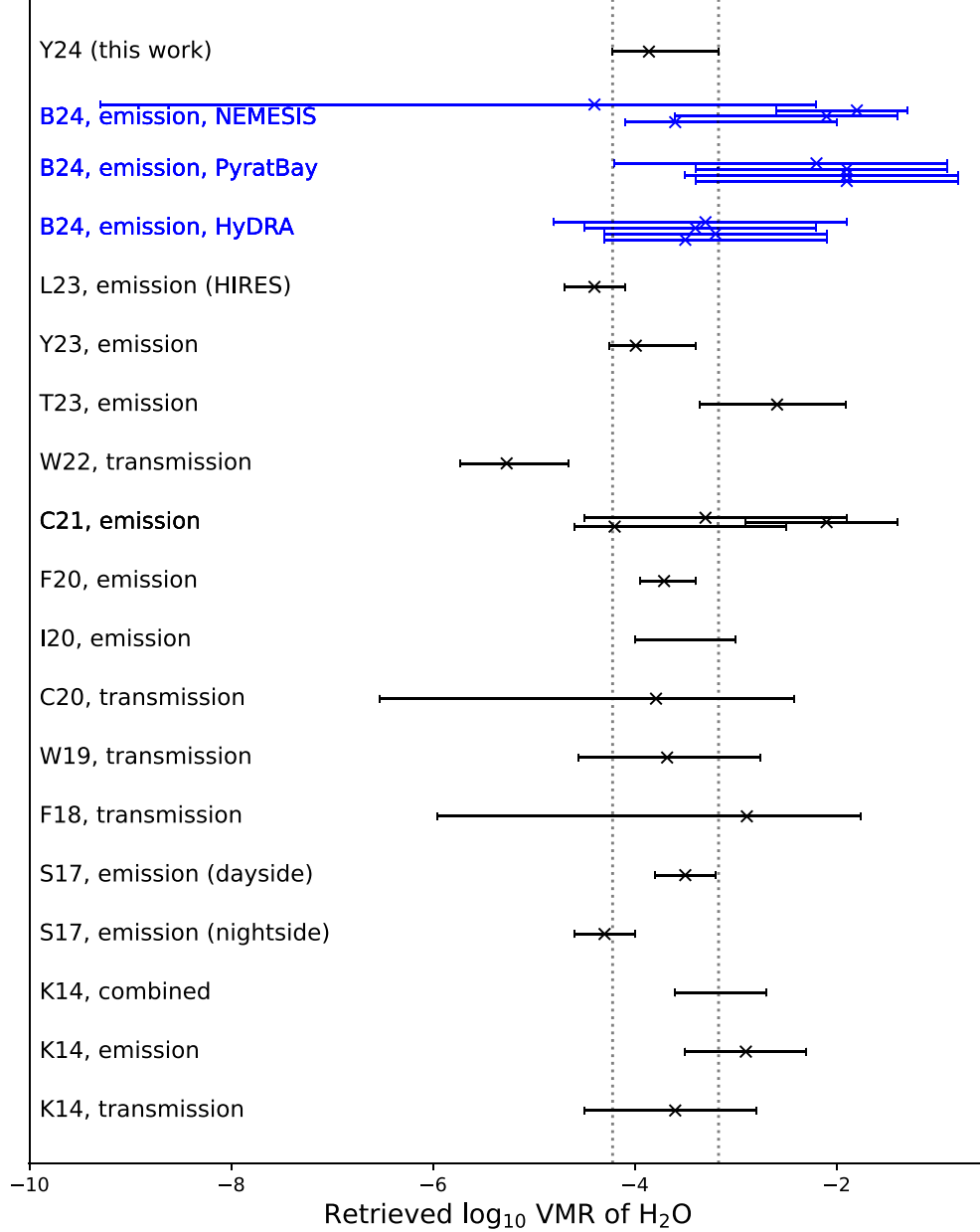


Figure B1. Comparison to past studies of WASP-43b that retrieved H_2O abundance. The error bars represent the 1σ credible intervals, and the crosses mark the posterior medians for studies that quoted such values. The vertical dashed lines mark the 1σ credible interval retrieved in this work. Note that both C21 and B24 performed phase-by-phase retrievals, hence we plot multiple error bars to show the constraints from all phases. The constraints for T23 are taken from the Scat.Cloud model in their Table A6. The list of studies is as follows: K14: Kreidberg et al. (2014); S17: Stevenson et al. (2017); F18: Fisher & Heng (2018); W19: Welbanks et al. (2019); C20: Chubb et al. (2020); I20: Irwin et al. (2020); F20: Feng et al. (2020); C21: Cubillos et al. (2021); W22: Welbanks & Madhusudhan (2022); T23: Taylor & Parmentier (2023); Y23: Yang et al. (2023); L23: Lesjak et al. (2023); B24: Bell et al. (2024); Y24: this work. Note that K14, S17, F18, W19, C20, I20, F20, C21, W22, T23, Y23 analysed subsets of *HST*/WFC3 and *Spitzer*/IRAC data, L23 analysed VLT/CRIRES+ data, and B24 and Y24 analysed *JWST*/MIRI data.

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