

Calibrating the absolute magnitude of type Ia supernovae in nearby galaxies using [O II] and implications for H_0

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ABSTRACT

The present state of cosmology is facing a crisis where there is a fundamental disagreement in measurements of the Hubble constant (H_0), with significant tension between the early and late Universe methods. Type Ia supernovae (SNe Ia) are important to measuring H_0 through the astronomical distance ladder. However, there remains potential to better standardize SN Ia light curves by using known dependencies on host galaxy properties after the standard light curve width and colour corrections have been applied to the peak SN Ia luminosities. To explore this, we use the 5-yr photometrically identified SNe Ia sample obtained by the Dark Energy Survey, along with host galaxy spectra obtained by the Australian Dark Energy Survey. Using host galaxy spectroscopy, we find a significant trend with the equivalent width (EW) of the [O II] $\lambda\lambda$ 3727, 29 doublet, a proxy for specific star formation rate, and Hubble residuals. We find that the correlation with [O II] EW is a powerful alternative to the commonly used mass step after initial light-curve corrections. Applying this [O II] EW correction to 20 SNe Ia in calibrator galaxies observed with WiFeS, we examined the impact on SN Ia absolute magnitudes and H_0 . Our [O II] EW corrections result in H_0 values ranging between 73.04 and 73.51 km s^{−1} Mpc^{−1}, with a combined statistical and systematic uncertainty of ~ 1.31 km s^{−1} Mpc^{−1}. However, even with this additional correction, the impact of host galaxy properties in standardizing SNe Ia appears limited in reducing the current tension ($\sim 5\sigma$) with the cosmic microwave background result for H_0 .

Key words: galaxies: general – cosmology: observations – transients: supernovae.

1 INTRODUCTION

Cosmology is currently facing a crisis known as the Hubble tension, where there is a disagreement between different methods measuring the Hubble constant (H_0). The most significant discrepancy is between the values derived from the cosmic microwave background (CMB; Planck Collaboration VI 2020) and the local distance ladder (Riess et al. 2022). Currently, the Hubble tension is at a significance level of 5 – 6 σ (Di Valentino et al. 2021).

The local and direct approach to measuring H_0 is through the astronomical distance ladder, which relies on the use of standardizable candles, such as Type Ia supernovae (SNe Ia). SNe Ia are excellent cosmological probes in understanding the expansion rate of the Universe. By calibrating the absolute magnitudes of SNe Ia (M_0^B) in nearby galaxies, one can use SNe Ia in the Hubble flow and measure H_0 (Dhawan et al. 2020; Freedman 2021; Khetan et al. 2021; Riess et al. 2022). Using 42 SNe Ia for which host galaxy distances are available using Cepheids, Riess et al. (2022) found $M_0^B = -19.253 \pm 0.027$ mag, and determined $H_0 = 73.04 \pm 1.04$ km s^{−1} Mpc^{−1}.

By increasing the sample of nearby SNe Ia in galaxies that have distance estimates from Cepheids, one can increase the precision in

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the measurement of H_0 . However, the current rate of finding nearby galaxies with SN Ia and Cepheid distances is $\sim 1 \text{ yr}^{-1}$ (Riess et al. 2022). Alternatively, the sample of SN Ia calibrator galaxies can be increased by including hosts that have distances derived using the tip of the red giant branch (TRGB; Freedman 2021) or surface brightness fluctuations (Jensen et al. 2021). In addition to reducing statistical uncertainties, these approaches will help better understand and control systematic uncertainties.

To standardize SNe Ia for use in cosmology, corrections are made to account for the relationship between the peak magnitude and the width of SN Ia light curve (Phillips 1993) and the SN Ia colour (Tripp 1998). To further reduce the intrinsic scatter, many studies have used broad-band photometry and spectroscopy of the host galaxies to correct for trends between the colour and width-corrected luminosities of SNe Ia and the properties of their host galaxies. Correlations with Hubble residuals have been found with host gas-phase metallicity (D’Andrea et al. 2011; Pan et al. 2014; Moreno-Raya et al. 2018), stellar age (Childress et al. 2013; Rose, Garnavich & Berg 2019), specific star formation rate (sSFR; Lampeitl et al. 2010; D’Andrea et al. 2011; Childress et al. 2013; Rigault et al. 2015; Rigault et al. 2020; Briday et al. 2022; Dixon et al. 2022; Galbany et al. 2022; Martin et al. 2024), rest-frame colour (Roman et al. 2018; Kelsey et al. 2021, 2023), and host galaxy dust (Brout & Scolnic 2021; Dixon et al. 2022; Meldorf et al. 2023).

The most common correction is with host galaxy stellar mass (Guy et al. 2010; Kelly et al. 2010; Sullivan et al. 2010; Childress et al. 2013; Uddin et al. 2017; Smith et al. 2020; Kelsey et al. 2021). This is applied as a step correction around $10^{10} M_\odot$, and is known as the ‘mass step’. SNe Ia in low-mass galaxies are fainter after standard light-curve correction compared to SNe Ia in high-mass galaxies.

However, the origin of the mass step remains poorly understood. One possibility is that the mass step is an artefact, driven by the overly simplistic single-valued colour correction that is applied to SN Ia luminosities (Brout & Scolnic 2021; Meldorf et al. 2023; Popovic et al. 2023). There are at least two physically plausible mechanisms for the colour correction. First, more energetic SNe Ia will be brighter, hotter and therefore bluer. This is intrinsic to the SN. Secondly, dust will redden and dim SNe Ia. This is extrinsic to the SN. There is a no priori reason to expect that the relationship between colour and luminosity should be the same for both, yet this is what is assumed. The consequence of this assumption might be the mass step.

Many galaxy properties correlate with the properties of the dust in them. For example, more massive galaxies are dustier (Triani et al. 2021), as are galaxies with higher sSFRs (Orellana et al. 2017), which is the star formation rate (SFR) normalized by stellar mass. These correlations then naturally lead to trends between Hubble residuals and the properties of the hosts if the colour correction is unable to capture both reddening by dust and the intrinsic colour–luminosity at the same time. For example, Dixon et al. (2022) find a correlation between Hubble residuals and the Balmer decrement, a measure of extinction by dust.

Previous work by Dixon et al. (2022) derived host galaxy properties using stacked spectra instead of broad-band photometry to explore what is physically driving the mass step. The most significant correlation uncovered was with the equivalent width (EW) of the [O II] $\lambda\lambda$ 3727, 29 doublet which is an indicator of the sSFR.

In this paper, we build upon this finding by directly measuring the [O II] EW for each SN Ia host galaxy. By analysing these SNe Ia, we are able to derive correlations between their Hubble residuals and [O II] EWs. We find that the [O II] EW trend is more significant than the commonly used mass step correction, highlighting its potential

Table 1. From the DES5YR photometric sample containing 1499 SNe Ia, we apply specific cuts that result in a final sample of 707 OzDES host galaxies for our analysis.

Selection Cut	# SNe Ia
DES5YR SNe Ia sample	1499
OzDES host spectra	1248
Redshift reliability > 95%	1191
Not contaminated by SN Ia light	744
EW [O II] < 0	707

for improving the standardization of SNe Ia. Next, we explore the impact of applying our [O II] EW correction to a sample of 20 SNe Ia hosted in galaxies ($z < 0.012$) with Cepheid-derived distances, which are used in constraining M_0^B . Finally, applying these results to the Pantheon + SNe Ia Hubble flow sample, we update the values M_0^B and H_0 .

2 HUBBLE FLOW GALAXIES

2.1 DES and OzDES

The Dark Energy Survey (DES) ran for six observing seasons from 2013 to 2019 and used the 570 megapixel Dark Energy Camera (DECam; Flaugher et al. 2015) situated on the 4-m Victor M. Blanco Telescope (DES Collaboration 2016). DES has observed hundreds of millions of galaxies and discovered thousands of supernovae (Hartley et al. 2022). The data are passed through the DES Image Processing Pipeline (Morganson et al. 2018), and SN transients are identified with a difference imaging pipeline (Kessler et al. 2015).

The Australian Dark Energy Survey (OzDES) was undertaken over the same observing period as DES, using the 3.9-m Anglo-Australian Telescope (AAT) at Siding Spring Observatory, along with the AAOmega spectrograph and the 2dF fibre positioner (Yuan et al. 2015; Childress et al. 2017). The 2dF fibre positioner allocates fibres, and accommodates up to 8 guide stars and 392 science targets within a 2.1 degree field, which aligns with the DECam imager’s field of view (FoV). The primary aims of OzDES were measuring redshifts of SN hosts, confirming the spectral-type of SNe, and monitoring AGNs over a wide redshift range (Hoormann et al. 2019). The wavelength coverage is between 3700 and 8800 Å, and the faintest objects have an apparent magnitude $r \approx 24$ mag. The second OzDES data release contains 375 000 spectra of 39 000 objects and is described in Lidman et al. (2020).

2.2 Sample selection

We use the DES5YR photometrically identified SNe Ia sample described in Möller et al. (2022) and updated in Vincenzi et al. (2024) for the final cosmology analysis (DES Collaboration 2024). SNe Ia were classified using the SuperNNova classifier (Möller & de Boissière 2020).

Initially, we have 1499 SNe Ia, where the probability of being a SNe Ia is greater than 0.5 (DES Collaboration 2024). We then apply specific selection cuts to our sample of hosts (see Table 1). Each DES transient is associated with its host galaxy, which is identified using the smallest directional light radius from the deep image stacks discussed in Wiseman et al. (2020). Spectroscopic redshifts are obtained from the OzDES global redshift catalogue (Lidman et al. 2020), and cuts are made regarding the redshift reliability for each galaxy. We only consider host redshifts with a quality flag of

3 or higher. This depends on identifying either a prominent single feature or multiple weaker features, resulting in a confidence level exceeding 95%. We also remove galaxies that may have been affected by supernova light contamination. Specifically, we exclude hosts that were observed within two months before or five months after the peak luminosity of the SN Ia explosion. After this cut, there are 744 SNe Ia. We then make a final cut by omitting hosts with positive [O II] EWs (see Section 2.4), resulting in a final sample of 707 SN Ia host galaxies for our analysis.

2.3 Deriving Hubble residuals

Each DES5YR SN Ia light curve is fit using the SALT3 light-curve model (Kenworthy et al. 2021; Taylor et al. 2023), which is based on SALT2 (Guy et al. 2010). We note that the differences between using SALT2 and SALT3 have been found to have a minimal impact on SN Ia cosmology results (Taylor et al. 2023). The standardization parameters consist of the colour (c), stretch (x_1), and peak brightness (m_x). We can then derive the observed distance modulus (μ_{obs}) for each SNe Ia using the modified Tripp equation (Tripp 1998):

$$\mu_{\text{obs},i} = m_{x,i} + \alpha x_{1,i} - \beta c_i + \delta_{\text{host},i} - M - \delta_{\text{bias},i}, \quad (1)$$

where M is the SN Ia absolute magnitude, with $x_1 = 0$ and $c = 0$. δ_{host} represents an additional correction for observed correlations between the SN Ia peak brightness and host galaxy properties. Commonly this is expressed in the form of a ‘mass step’, where $\delta_{\text{host}} = \gamma G_{\text{host}}$.

$$G_{\text{host}} = \begin{cases} +1/2 & \text{for } M_* > M_{\text{step}} \\ -1/2 & \text{otherwise,} \end{cases} \quad (2)$$

where M_* is the galaxy stellar mass, and γ is the size of the mass step. The division point (M_{step}) is commonly taken as $10^{10} M_{\odot}$. The global fitting parameters for the DES5YR analysis are $\alpha = 0.161 \pm 0.001$, $\beta = 3.12 \pm 0.03$, and $\gamma = 0.038 \pm 0.007$ (DES Collaboration 2024). The terms indexed by ‘ i ’ denote parameters specific to individual SNe Ia. δ_{bias} accounts for biases arising from selection effects and light-curve fitting, and is obtained from simulations using ‘Beams with Bias Correction’ (BBC: Kessler & Scolnic 2017). The bias correction depends on redshift, stretch, colour and host galaxy mass (‘BBC4D’: Popovic et al. 2021). The light-curve fitting process is discussed in greater detail in DES Collaboration (2024).

The Hubble residuals ($\Delta\mu$) in our analysis represent the difference between the observed distance modulus (μ_{obs}) and the distance modulus measured by DES using a flat Λ CDM cosmological model (μ_{cosmo}), where $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\Omega_m = 0.352$ (DES Collaboration 2024):

$$\Delta\mu = \mu_{\text{obs}} - \mu_{\text{cosmo}}. \quad (3)$$

We then construct a Hubble diagram for our SNe Ia, as illustrated in Fig. 1.

2.4 SN Ia host galaxy properties

We use penalized pixel-fitting (pPXF), which is a full-spectrum fitting approach to extract stellar population parameters for each galaxy (Cappellari & Emsellem 2004). In building each synthetic spectrum, we utilize the E-MILES spectral library of single stellar populations (Vazdekis et al. 2016), with stars older than 30 Myr, spanning a metallicity range of $[M/H]$ from -1.792 to $+0.26$, and assuming a Salpeter (1955) initial mass function with a slope of 1.30 (Vazdekis et al. 2010). They were obtained at the 2.5-m Isaac Newton Telescope in Spain and cover the wavelength range 3525–7500 Å, with a spectral resolution of 2.5 Å. To account for factors

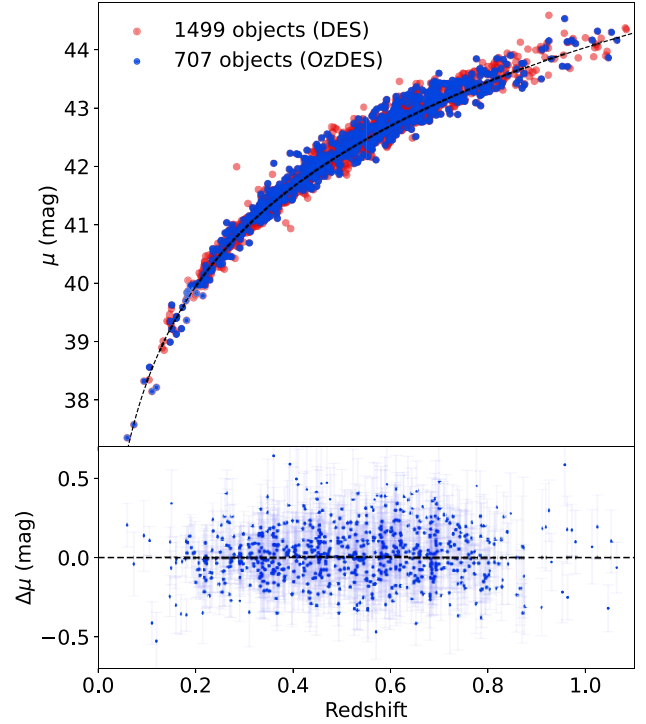


Figure 1. Upper plot: Hubble diagram showing the DES5YR sample containing 1499 SNe. After our selection cuts, we obtain a sample of 707 SNe Ia. Lower plot: Hubble residuals ($\Delta\mu$) for 707 SNe Ia.

such as extinction and calibration errors, a fourth-order multiplicative polynomial was used to warp the spectral continuum. Full details of the pPXF methodology are given in Cappellari (2017).

Building on the investigation in Dixon et al. (2022), we choose to focus on the [O II] emission line, given it has the strongest correlation obtained from the OzDES spectra and is an indicator of sSFR. Importantly, we now measure [O II] EW for each OzDES host galaxy. This is achieved by using pPXF to obtain the stellar continuum and emission-line components. An example host galaxy spectrum, characteristically exhibiting a low signal-to-noise ratio (S/N) of 3, is shown in Fig. 2. Excluded from this analysis are objects with positive EWs. These are likely to be passive galaxies with little ongoing star formation.

We estimate uncertainties by perturbing each point in our spectra, where the magnitude of the perturbation is taken from a Gaussian that has a mean of zero and a variance that is determined from the variance spectrum. We run pPXF on 500 perturbed spectra and then measure the uncertainty in [O II] EW.

Fig. 3 explores the relationship between host galaxy properties derived from photometry, which include stellar mass, $U - R$ rest-frame colour (DES Collaboration 2024; Vincenzi et al. 2024) and our [O II] EW measurements. The plot highlights a strong trend between stellar mass and [O II] EW. Additionally, the comparison with $U - R$ shows that galaxies with higher stellar mass and a lower sSFR tend to be redder.

2.5 Different fitting approaches

Hubble residuals have been found to vary with host galaxy properties, and are typically characterized with a step function, such as the well-known ‘mass step’. However, the physical origin of why such a step relation should exist is unclear. We can explore this systematic by

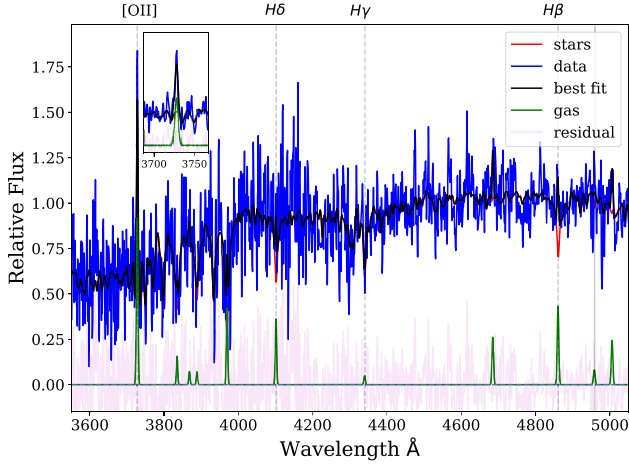


Figure 2. An example of using pPXF (Cappellari 2017) to measure the [O II] EW for an OzDES galaxy. The blue line shows the galaxy spectrum, the black line is the best fit, the red line is the stellar continuum, the green line is the gas component, and the pink line illustrates the residual flux. We zoom in on the [O II] emission line to illustrate that this measurement ($-8.94 \pm 1.63 \text{ \AA}$) is achievable even with a low S/N for the continuum.

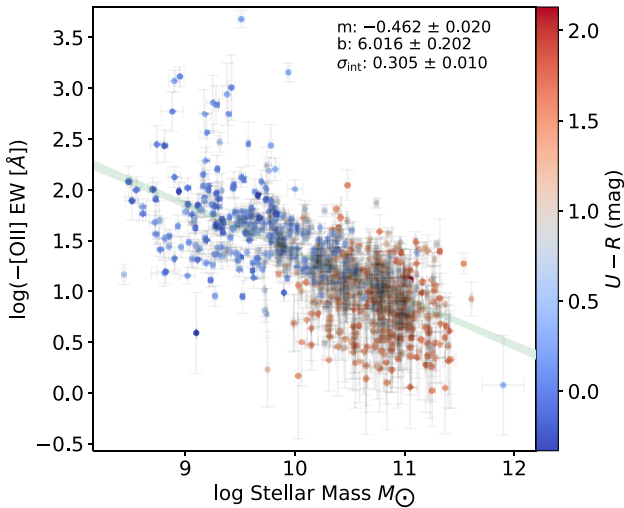


Figure 3. We find a strong correlation between host galaxy stellar mass and the log of the negative EW of [O II] for our sample of 707 galaxies. The slope of the linear fit is -0.462 ± 0.020 and we find that more massive galaxies tend to have lower sSFRs. Additionally, we add a comparison with $U - R$ rest-frame colour, showing more massive, passive galaxies tend to be redder.

using different fitting functions in obtaining our corrections. In this paper, we account for both x and y uncertainties in the fit. This allows us to determine the optimal fits for a linear trend, step function, and smoothed step function (hyperbolic tan).

In Fig. 4, we find a strong trend (4.0σ) between Hubble residuals, $\Delta\mu(x_1, c, \delta_{\text{bias}})$ ¹, and [O II] EW. Fainter objects after light curve and bias corrections tend to reside in galaxies that have higher [O II] EWs and therefore higher sSFRs. Additionally, we find that the smoothed step function prefers a significant amount of smoothing in obtaining the best fit. This converges towards the linear fit and away from

¹Hubble residuals that exclude the mass step correction but include stretch, colour, and bias corrections.

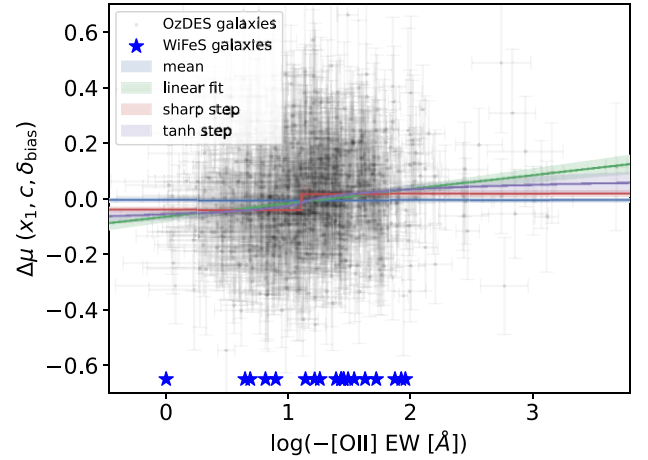


Figure 4. Host galaxy sample containing 707 objects. We plot [O II] EW against Hubble residuals ($\Delta\mu$), derived using stretch (x_1), colour (c) and bias corrections (δ_{bias}) for each SN Ia light curve. Compared to the final DES analysis (DES Collaboration 2024), we omit the mass step correction (δ_{host}) to explore the potential of using an [O II] EW correction instead. A strong trend with a slope of 0.050 ± 0.013 (4.0σ significance) is evident, where fainter SNe Ia tend to be in hosts with larger [O II] EWs. Additionally, we examine the impact of different fitting functions (linear, step, smoothed step/tanh). The 20 SN Ia calibrator galaxies (blue stars) are plotted at a fixed $\Delta\mu = -0.65$ for reference and are discussed in Section 3. We note that most of the SN Ia calibrators reside in galaxies with higher sSFRs.

the step function. Recent studies have found similar trends with Hubble residuals (Rigault et al. 2020; Briday et al. 2022; Dixon et al. 2022; Galbany et al. 2022). However, these studies employ different methods to derive the sSFR for each host galaxy (photometry versus spectroscopy) and utilize different light-curve correction parameters.

2.6 Splitting by SN Ia light-curve properties

We can also split our SNe Ia by colour (Fig. 5), using the dividing line at $c = -0.025$ (Brout & Scolnic 2021). We find a noticeable difference between the red SNe ($c > -0.025$) and blue SNe $c \leq -0.025$, namely that redder SNe have a more significant trend with Hubble residuals and [O II] EW (3.9σ), compared to bluer SNe, where the correlation is weaker (1.8σ). Kelsey et al. (2021) found a similar result when splitting on $U - R$ rest frame colour. Overall, these findings support the notion that galaxies hosting bluer SNe are less impacted by environmental dependencies and are better suited for use in cosmology (Kelsey et al. 2023).

3 SN IA CALIBRATOR GALAXIES

3.1 Observing and data reduction

Our first step is to quantify host stellar population properties, with a focus on the [O II] EW measurements from nearby galaxies for which SNe Ia and Cepheid distances are available. Our sample is built from the recent analysis undertaken by the SH0ES collaboration (SNe, H0, for the Equation of State of Dark Energy) and is described in Riess et al. (2022). We note that the Cepheid calibration used by SH0ES relies on multiple distance anchors, including *Gaia* EDR3 parallaxes, masers in NGC 4258, and detached eclipsing binaries in the Large Magellanic Cloud. Additionally, the SH0ES galaxies

Table 2. SN Ia calibrator galaxies observed using WiFeS. For each galaxy, we show the redshift, RA, DEC, SH0ES SN Ia absolute magnitude ($M_{0,i}^B$), [O II] EW, and host stellar mass. We note that for NGC1448, we take the weighted average of the two SNe Ia hosted by that galaxy (Riess et al. 2022).

Galaxy	z	RA	Dec	$M_{0,i}^B$ (mag)	$\sigma_{M_{0,i}^B}$	[O II] EW (Å)	σ_{OII}	log Stellar mass ($M_{\odot}$)
NGC1015	0.0088	02 ^h 38 ^m 11 ^s .565	−01 ^d 19 ^m 07 ^s .020	−19.220	0.120	−1.44	0.95	9.906
NGC1309	0.0071	03 ^h 22 ^m 06 ^s .556	−15 ^d 23 ^m 59 ^s .794	−19.337	0.102	−20.09	0.09	9.890
NGC1365	0.0055	03 ^h 33 ^m 36 ^s .458	−36 ^d 08 ^m 26 ^s .370	−19.479	0.108	−27.58	2.55	10.732
NGC1448	0.0039	03 ^h 44 ^m 31 ^s .915	−44 ^d 38 ^m 41 ^s .380	−19.199	0.116	−28.08	1.74	11.280
NGC1559	0.0043	04 ^h 17 ^m 35 ^s .750	−62 ^d 47 ^m 01 ^s .225	−19.361	0.106	−31.25	0.53	10.375
NGC2442	0.0049	07 ^h 36 ^m 23 ^s .842	−69 ^d 31 ^m 50 ^s .960	−19.223	0.105	−23.74	3.16	12.198
NGC3370	0.0043	10 ^h 47 ^m 04 ^s .039	+17 ^d 16 ^m 25 ^s .310	−19.186	0.097	−6.80	0.66	10.196
NGC4038	0.0054	12 ^h 01 ^m 31 ^s .770	−18 ^d 50 ^m 41 ^s .300	−19.207	0.158	−44.29	0.90	10.682
NGC4424	0.0015	12 ^h 27 ^m 11 ^s .575	+09 ^d 25 ^m 14 ^s .312	−19.369	0.232	−21.89	0.10	9.633
NGC4536	0.0060	12 ^h 34 ^m 28 ^s .129	+02 ^d 11 ^m 16 ^s .37	−19.287	0.142	−50.76	2.07	9.686
NGC4639	0.0035	12 ^h 42 ^m 52 ^s .378	+13 ^d 15 ^m 26 ^s .713	−19.364	0.150	−6.95	2.57	9.802
Mk1337	0.0085	12 ^h 52 ^m 34 ^s .701	−09 ^d 46 ^m 35 ^s .724	−19.267	0.163	−91.50	3.90	9.554
NGC5468	0.0095	14 ^h 06 ^m 34 ^s .891	−05 ^d 27 ^m 10 ^s .719	−19.127	0.104	−42.62	4.49	10.441
NGC5584	0.0055	14 ^h 22 ^m 23 ^s .811	−00 ^d 23 ^m 14 ^s .820	−18.971	0.095	−19.49	1.96	10.331
NGC5643	0.0040	14 ^h 32 ^m 40 ^s .778	−44 ^d 10 ^m 28 ^s .600	−19.362	0.089	−64.53	1.24	10.530
NGC5861	0.0062	15 ^h 09 ^m 16 ^s .091	−11 ^d 19 ^m 17 ^s .980	−19.287	0.147	−11.84	2.88	10.591
NGC5917	0.0063	15 ^h 21 ^m 32 ^s .550	−07 ^d 22 ^m 37 ^s .523	−19.284	0.154	−157.51	2.19	9.184
NGC7329	0.0109	22 ^h 40 ^m 24 ^s .199	−66 ^d 28 ^m 44 ^s .580	−19.244	0.140	−2.20	0.19	10.501
NGC7541	0.0090	23 ^h 14 ^m 43 ^s .857	+04 ^d 32 ^m 02 ^s .040	−19.086	0.176	−30.28	0.84	10.935
NGC7678	0.0116	23 ^h 28 ^m 27 ^s .860	+22 ^d 25 ^m 16 ^s .573	−19.106	0.181	−11.91	1.71	10.530

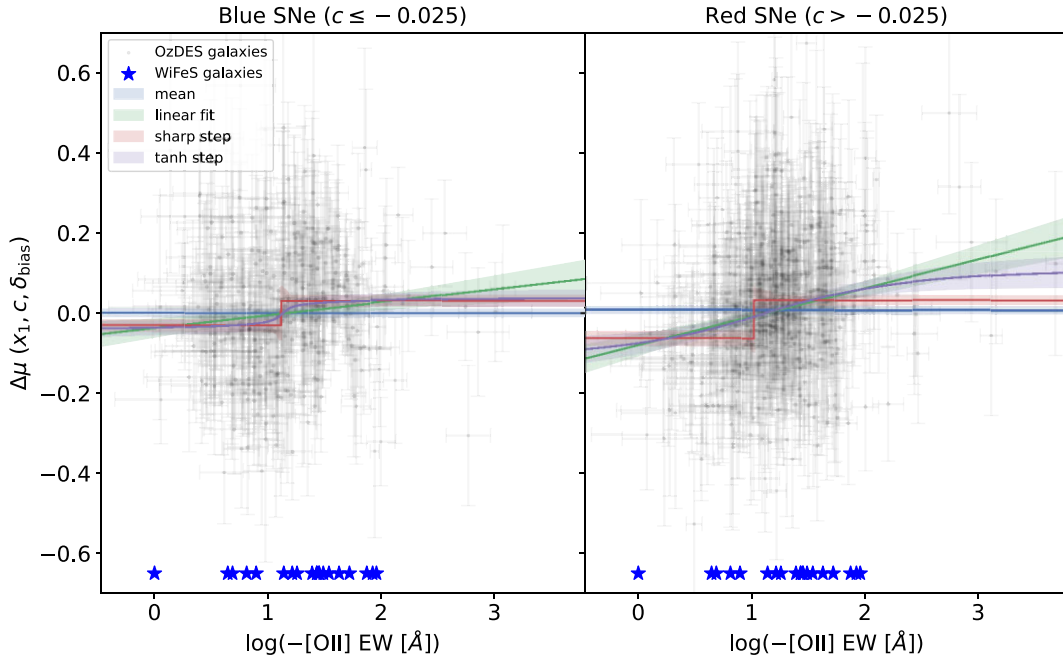


Figure 5. OzDES host galaxies split into subsamples based on SN Ia colour, with 295 blue SNe Ia ($c < -0.025$) and 412 red SNe Ia ($c \geq -0.025$). We find that the trend seen in Fig. 4 is more significant for redder SNe Ia (3.9σ), and is weaker for bluer SNe Ia (1.8σ).

also form part of the SN Ia sample used in the Pantheon + analysis (Brout et al. 2022; Scolnic et al. 2022), encompassing 1701 SN Ia light curves spanning the redshift range of $0.001 < z < 2.26$.

Initially, we selected a subset of 28 galaxies, limiting our selection to a declination ≤ 20 deg for observing objects using the 2.3m ANU telescope at Siding Spring Observatory. We utilize the WiFeS instrument (Dopita et al. 2007, 2010) which is a double-beam, image-slicing, integral-field spectrograph with a 25 by 38 arcsec FoV. We

use the RT560 dichroic beamsplitter along with the R3000/B3000 gratings, covering the wavelength range 3300 – 9200 Å. Our galaxies were observed in nod & shuffle mode, with two 1600s exposures, comprising of 800s on the object and 800s on sky positions, respectively. The median seeing was 2 arcsec. Due to time constraints, we observed 20 galaxies which are listed in Table 2.

Our WiFeS data were reduced using the PyWiFeS data-reduction pipeline (Childress et al. 2014). The wavelength solution was derived

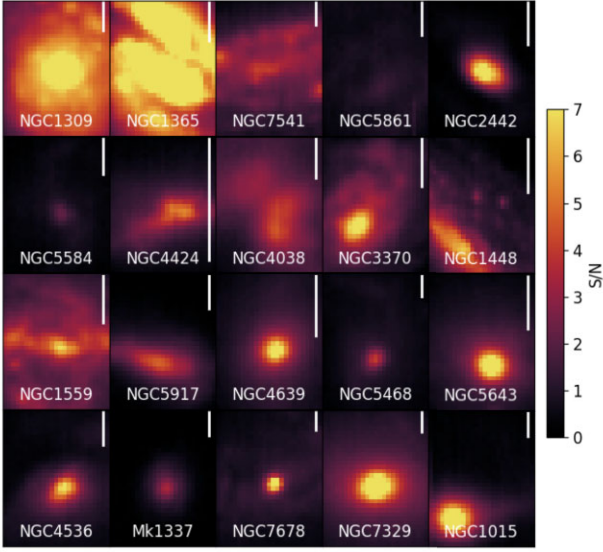


Figure 6. S/N maps for 20 SN Ia calibrator galaxies observed using WiFeS. The white line in each subplot represents a scale of ~ 1 kpc, given each galaxy’s redshift.

using arc lamp exposures. Wire frames are obtained for spatial alignment of the slitlets while flat-field exposures are used to remove pixel-to-pixel sensitivity variations. For each night, a standard star was acquired for flux calibration and to correct for telluric absorption. Typically, we observed a white-dwarf spectrophotometric standard. For each exposure, two data cubes from the blue/red CCD’s are generated, containing spatial and spectral information across the FoV.

3.2 Aperture selection

The next step is to extract a 1D spectrum from each WiFeS cube, with the goal of measuring integrated [O II] EW’s that are comparable to those from OzDES. The OzDES sample spans the redshift range 0.12–1.06, with a median redshift of 0.52. At these distances, the 2 arcsec fibre aperture corresponds to a projected aperture ≈ 5 –16 kpc (but with seeing effects on a similar scale). Recognizing the risk of ‘aperture effects’ creating a systematic difference between the low- z and high- z measurements, we use the largest aperture that the data can support, trying to make sure that we capture as much of the [O II] emission as possible, and being mindful of strong [O II] EW gradients (see Figs 6 and 7). A potential challenge associated with this selection relates to spaxel quality in regions near the aperture edges. To address this, we exclude poor-quality ($S/N < 0.5$) spaxels from our analysis.

We can further explore the systematic uncertainty associated with our aperture selection through three different cases. First, we can select all of the available spaxels to best compare with the OzDES hosts. Secondly, we can make a cut based on the S/N of each spaxel. This effect introduces a bias towards spaxels with stronger lines. However, this bias is mitigated when the spaxels are combined. Thirdly, we then explore the impact of selecting a smaller aperture, and combine the variation in each approach to obtain the systematic involving our selection choice. Overall, for the linear correction, the systematic impact remains consistent regardless of our aperture selection and only differs by 0.005 mag. The variation in step correction is more significant, which is expected for galaxies that fall around the step location. However, the impact is small (~ 0.01

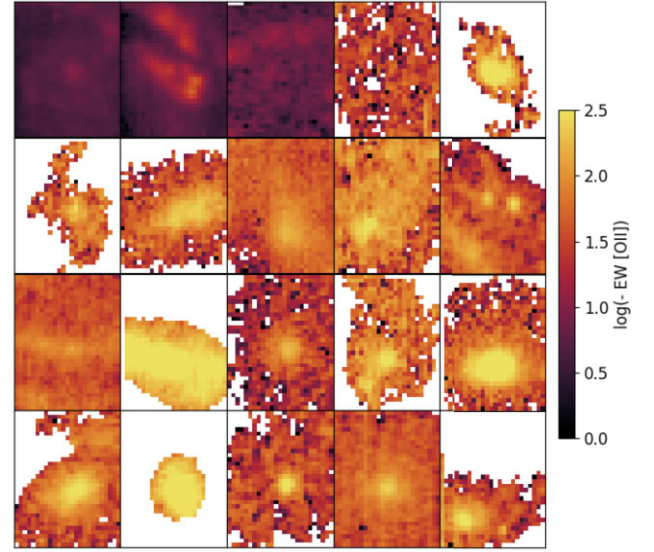


Figure 7. [O II] EW maps for each of the 20 SN Ia calibrator galaxies. A mask is applied to include spaxels with an S/N greater than 0.5. This additionally helps remove spaxels that are more heavily impacted by sky subtraction residuals.

mag). Aperture corrections in spectroscopic surveys can artificially increase the size of the mass step by 0.02 mag and its significance (Galbany et al. 2022). Hence, we adopt this as a systematic in our analysis.

Once extracted, the red and blue spectra are spliced together and we determine the redshift of the combined spectrum using MARZ (Hinton et al. 2016). We note the spectra from the red arm are rebinned to match the spectral resolution of the blue arm.

4 CORRECTING SN IA ABSOLUTE MAGNITUDES

We now explore the impact of correcting SNe Ia for environmental dependence using [O II] EW. Similar to the OzDES hosts, we measure [O II] EW using pPXF for each of the 20 SN Ia calibrator galaxies observed using WiFeS (Fig. 8). SN Ia absolute magnitudes ($M_{0,i}^B$) are obtained from the SH0ES analysis (Riess et al. 2022) and are listed in Table 2. Hereafter, $\widetilde{M}_{0,i}^B$ represents the [O II] EW corrected absolute magnitude for each SN Ia, while M_0^B denotes the weighted average of the 20 SN Ia calibrator galaxies. We note that our M_0^B will differ from SH0ES, given that it is a subsample.

We determine our corrections by analysing the trends between $\Delta\mu$ and host galaxy properties (stellar mass, [O II], $U - R$). We explore these trends with a range of light-curve correction parameters which include x_1 , c , δ_{host} , δ_{bias} , and [O II].

- (i) Case 1 baseline: $\Delta\mu(x_1, c, \delta_{\text{bias}})$
- (ii) Case 1: $\Delta\mu(x_1, c, [\text{O II}], \delta_{\text{bias}})$
- (iii) Case 2 baseline: $\Delta\mu(x_1, c, \delta_{\text{host}}, \delta_{\text{bias}})$
- (iv) Case 2: $\Delta\mu(x_1, c, \delta_{\text{host}}, [\text{O II}], \delta_{\text{bias}})$

These corrections correspond to the rows in Fig. 9, and summarized in Table 3. We discuss these results in more detail in Section 5.3.

Next, we examine two specific cases that apply different [O II] EW corrections derived using the Case 1/2 baselines.

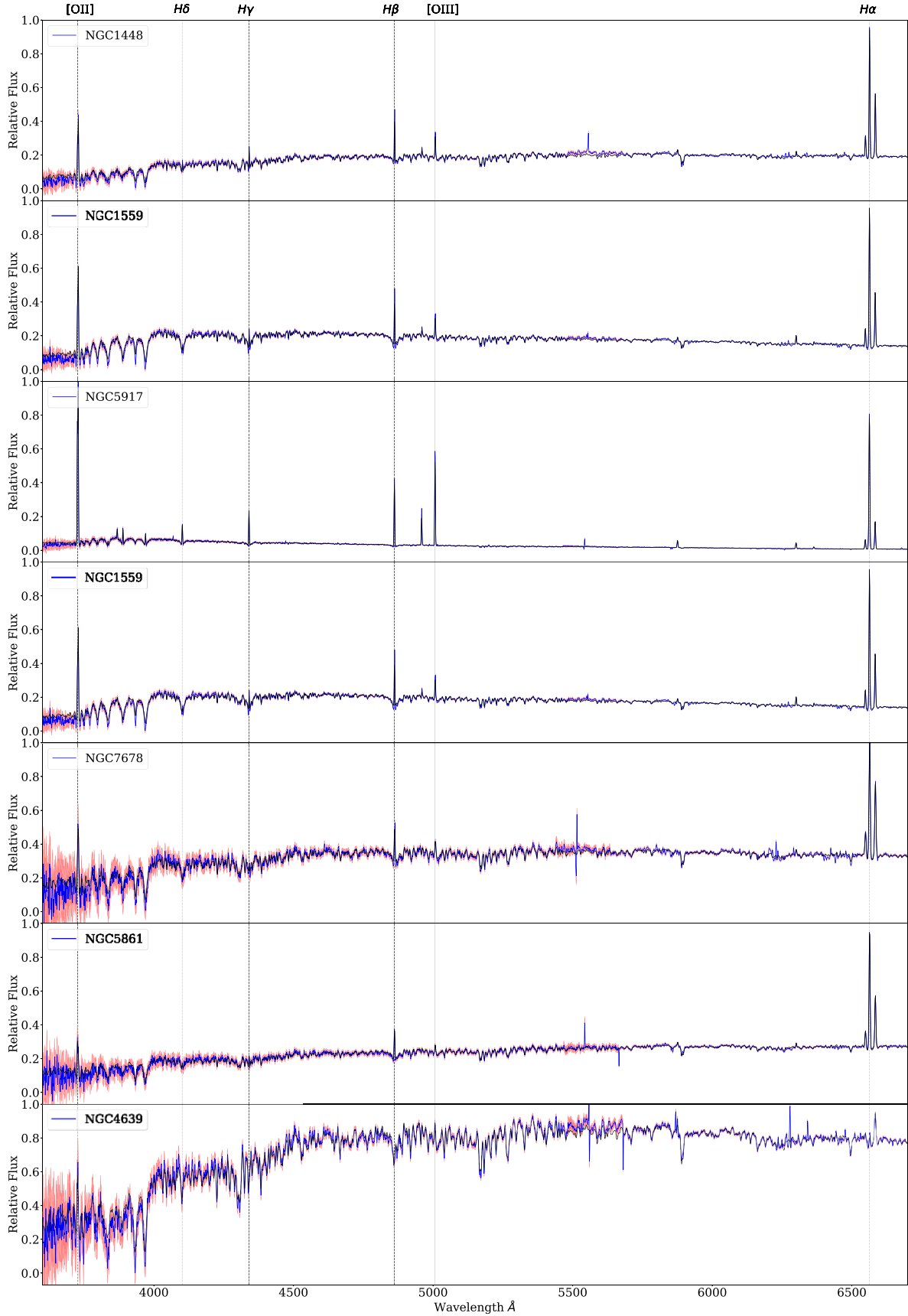


Figure 8. Examples of host galaxy spectra obtained using WiFeS. The original spectrum is in blue, the best fit is shown in black and is obtained using penalized-pixel fitting (pPXF; Cappellari 2016). The red region illustrates the standard deviation of the flux. Prominent spectral features that can be identified are [O II], [O III], and the Balmer lines ($H\alpha$, $H\beta$, $H\gamma$, $H\delta$).

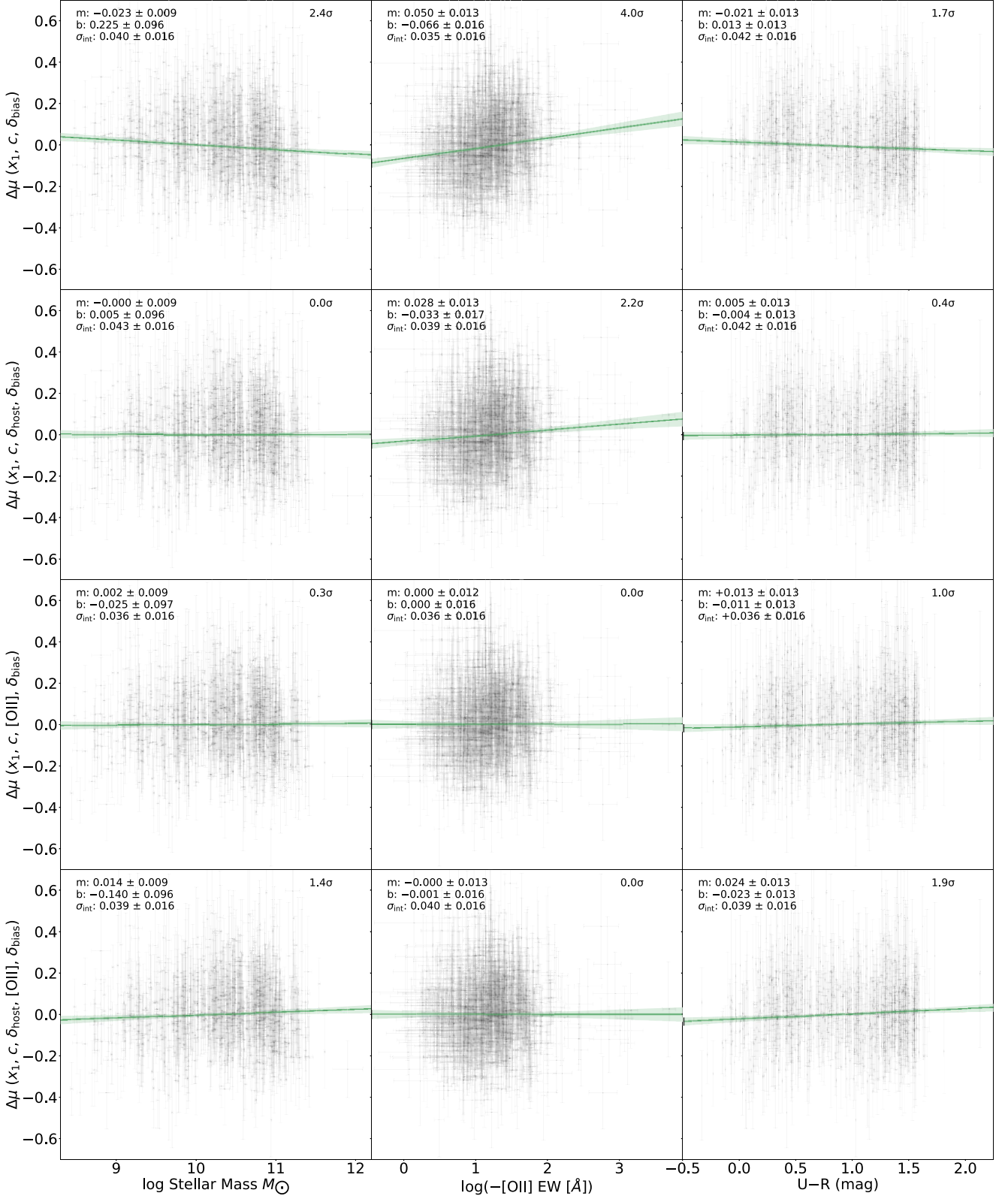


Figure 9. Comparing the effectiveness of host stellar mass, [O II] EW, and $U - R$ rest-frame colour as additional corrections. First row: $\Delta\mu$ are derived using stretch (x_1), colour (c), and bias corrections (δ_{bias}). We find that [O II] EW exhibits the most significant trend (4.0 σ) and smallest intrinsic scatter. Second row: We then introduce a mass step correction δ_{host} to obtain new $\Delta\mu$ values. Interestingly, a trend with [O II] EW persists (2.2 σ), while stellar mass and $U - R$ colour show no significant trends. Third row: Here, we explore the case of replacing δ_{host} with our [O II] EW correction. Fourth row: Lastly, we examine the case of utilizing both a mass step and [O II] EW correction. We find that the best approach for minimizing environmental dependence with $\Delta\mu$ across all three parameters ($\leq 1\sigma$) involves a combination of x_1 , c , [O II] and δ_{bias} .

Table 3. Summary of the trends observed between $\Delta\mu$ and host galaxy properties for a range of light-curve corrections (see Fig. 9). These parameters consist of x_1 , c , δ_{host} , δ_{bias} , and $[\text{O II}]$.

$\Delta\mu$	Slope	Significance (σ)	σ_{int}
Stellar mass			
$x_1, c, \delta_{\text{bias}}$	-0.023 ± 0.009	2.4	0.040
$x_1, c, \delta_{\text{host}}, \delta_{\text{bias}}$	-0.000 ± 0.009	0.0	0.043
$x_1, c, [\text{O II}], \delta_{\text{bias}}$	0.002 ± 0.009	0.3	0.036
$x_1, c, \delta_{\text{host}}, [\text{O II}], \delta_{\text{bias}}$	0.014 ± 0.009	1.4	0.039
[O II]			
$x_1, c, \delta_{\text{bias}}$	0.050 ± 0.013	4.0	0.035
$x_1, c, \delta_{\text{host}}, \delta_{\text{bias}}$	0.028 ± 0.013	2.2	0.039
$x_1, c, [\text{O II}], \delta_{\text{bias}}$	0.000 ± 0.012	0.0	0.036
$x_1, c, \delta_{\text{host}}, [\text{O II}], \delta_{\text{bias}}$	-0.000 ± 0.013	0.1	0.040
U-R			
$x_1, c, \delta_{\text{bias}}$	-0.021 ± 0.013	1.7	0.042
$x_1, c, \delta_{\text{host}}, \delta_{\text{bias}}$	0.005 ± 0.013	0.4	0.042
$x_1, c, [\text{O II}], \delta_{\text{bias}}$	0.013 ± 0.013	1.0	0.036
$x_1, c, \delta_{\text{host}}, [\text{O II}], \delta_{\text{bias}}$	0.024 ± 0.013	1.9	0.039

4.1 Case 1: $\Delta\mu$ ($x_1, c, [\text{O II}], \delta_{\text{bias}}$)

The first case involves replacing the mass step correction (δ_{host}) with our $[\text{O II}]$ EW correction. We achieve this by removing δ_{host} as described in DES Collaboration (2024), where $\gamma = 0.038$ mag, and are left with the Case 1 baseline which can now be applied as a linear correction using $[\text{O II}]$ EW. However, we need to also remove the Pantheon + mass step correction ($\delta_{\text{host}}^{P+}$) for each of the SN Ia calibrators. We accomplish this by adding $\delta_{\text{host}}^{P+}$ to $M_{0,i}^B$, where the size of γ is set to 0.019 mag (Brout et al. 2022) to align with the intrinsic scatter model (P21: Popovic et al. 2021) used in the DES5YR analysis (DES Collaboration 2024). We can then obtain $\widetilde{M}_{0,i}^B$ as is shown below:

$$\widetilde{M}_{0,i}^B = M_{0,i}^B + \delta_{\text{host}}^{P+} + (0.050 \times [\text{O II}]_i - 0.066), \quad (4)$$

where $[\text{O II}] = \log(-[\text{O II}] \text{ EW})$. Each step is shown in Fig. 10, where we first remove $\delta_{\text{host}}^{P+}$, then apply a linear and step $[\text{O II}]$ EW correction, before we calculate $\widetilde{M}_{0,i}^B$ for each SN Ia calibrator galaxy.

After applying a $[\text{O II}]$ EW linear correction, we measure $\widetilde{M}_0^B = -19.253 \pm 0.033$ mag. For a step correction, $\widetilde{M}_0^B = -19.247 \pm 0.033$ mag. We take the linear and step-fitting approaches to be the upper/lower extremes, defining the systematic range depending on our choice of fitting function. However, we do find that the linear fit is the preferred option, with the smallest deviance information criterion (DIC). Additionally, a smoothed step approach tends to converge to a linear trend, rather than a step.

4.2 Case 2: $\Delta\mu$ ($x_1, c, \delta_{\text{host}}, [\text{O II}], \delta_{\text{bias}}$)

In this case, we explore the impact of applying an additional $[\text{O II}]$ EW correction to the mass-corrected DES5YR Hubble residuals. Even after applying δ_{host} as a mass-step correction in the Case 2 baseline, our analysis in Fig. 9 (second row) reveals a 2.2σ trend between the remaining $\Delta\mu$ and $[\text{O II}]$ EW. We can then directly apply this correction to the SN Ia calibrators, and examine the impact on $M_{0,i}^B$:

$$\widetilde{M}_{0,i}^B = M_{0,i}^B + (0.028 \times [\text{O II}]_i - 0.033). \quad (5)$$

Applying this $[\text{O II}]$ EW correction results in a slight shift towards fainter $\widetilde{M}_0^B$ values which range from -19.244 to -19.239 mag (Fig. 11). This approach, which combines the mass step and $[\text{O II}]$ EW, proves to be less effective than Case 1 (fourth row, Fig. 9). We find that replacing the mass step directly with our $[\text{O II}]$ EW correction (Case 1) is the better choice for mitigating trends between Hubble residuals and host galaxy properties (third row, Fig. 9). Further discussion on the effectiveness of these host galaxy parameters is provided in Section 5.3.

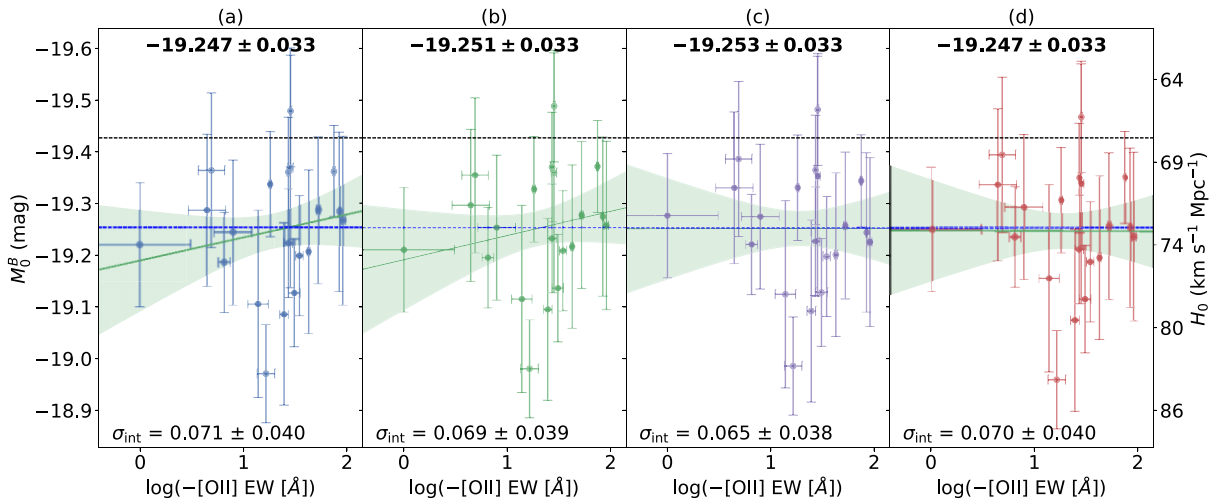


Figure 10. Case 1: Exploring the impact of different $[\text{O II}]$ EW corrections on SN Ia absolute magnitudes. (a) shows $M_{0,i}^B$ for each SN Ia calibrator galaxy with the Pantheon + mass step applied, (b) is with the mass step removed, while (c) and (d) are with linear and step $[\text{O II}]$ corrections respectively. We measure the intrinsic scatter (σ_{int}) in each subplot and report the mean M_0^B , with values ranging from -19.253 to -19.247 mag. The $[\text{O II}]$ correction helps minimize the trend between M_0^B and $[\text{O II}]$, as seen by comparing the linear fits (green). The dashed blue line represents the SH0ES analysis (Riess et al. 2022), and the black dashed line corresponds to the Planck H_0 result (Planck Collaboration VI 2020).

Table 4. Impact of different host galaxy corrections on the SN Ia absolute magnitude and H_0 . We compare the two cases discussed in Section 4. With no additional correction for our galaxies, we measure $M_0^B = -19.247 \pm 0.033$ mag and $H_0 = 73.24 \pm 1.11$ km s⁻¹ Mpc⁻¹. We also show the slope and significance of the trends in each case before and after corrections are implemented.

OzDES Hubble Flow SNe Ia	WiFeS Calibrator SNe Ia	Slope	Sig (σ)	M_0^B (mag)	H_0 (km s ⁻¹ Mpc ⁻¹)
	M_0^B (SH0ES)			-19.253 ± 0.027	73.04 ± 1.01
	M_0^B	-0.044 ± 0.071	0.62	-19.247 ± 0.033	73.24 ± 1.11
Case 1: $\Delta\mu$ (x_1 , c , [O II], δ_{bias})	$M_0^B + \delta_{\text{host}}^{P+}$	-0.048 ± 0.072	0.67	-19.251 ± 0.033	73.11 ± 1.11
	$M_0^B + \delta_{\text{host}}^{P+} + [\text{OII}]_{\text{linear}}$	-0.003 ± 0.070	0.05	-19.253 ± 0.033	73.04 ± 1.11
	$M_0^B + \delta_{\text{host}}^{P+} + [\text{OII}]_{\text{step}}$	0.001 ± 0.072	0.01	-19.247 ± 0.033	73.24 ± 1.11
Case 2: $\Delta\mu$ (x_1 , c , δ_{host} , [O II], δ_{bias})	$M_0^B + [\text{OII}]_{\text{linear}}$	-0.018 ± 0.072	0.25	-19.244 ± 0.034	73.34 ± 1.15
	$M_0^B + [\text{OII}]_{\text{step}}$	-0.011 ± 0.073	0.14	-19.239 ± 0.033	73.51 ± 1.11

Table 5. Summary of the systematic error contributions for Case 1 (see Section 4.1). This includes the selection of fitting functions used to derive [O II] EW correction, uncertainties in our choice of aperture, and the impact of not applying the [O II] EW correction to SNe Ia used for the third rung of the distance ladder. We estimate the total systematic uncertainty by taking the quadrature sum of these contributions.

Case 1	ΔM_0^B (mag)	ΔH_0 (km s ⁻¹ Mpc ⁻¹)
Fitting approach	0.007	0.23
Aperture selection	0.020	0.66
Differential impact	0.009	0.30
Total	0.023	0.76

Table 6. Similar to Table 5 but here we show the systematic budget for Case 2 (see Section 4.2). We note the size of the [O II] EW correlation in Case 2 compared to Case 1, results in a reduced differential impact of 0.005 mag and hence a smaller overall systematic uncertainty.

Case 2	ΔM_0^B (mag)	ΔH_0 (km s ⁻¹ Mpc ⁻¹)
Fitting approach	0.004	0.14
Aperture selection	0.020	0.66
Differential impact	0.005	0.17
Total	0.021	0.70

5 DISCUSSION

5.1 Impact on H_0

We can now use our corrected M_0^B to make a revised estimate of H_0 . We define the luminosity distance (d_L), which is commonly expressed using the distance modulus (μ), the apparent (m) and absolute magnitudes (M_0) as shown below.

$$\mu = m - M_0 = 5 \log_{10}(d_L) + 25. \quad (6)$$

We then use a kinematic expression of our cosmological model which gives the d_L as a function of redshift. The expression below is derived from a Taylor expansion of the Hubble–Lemaître law for a flat Universe, consisting of additional parameters, the cosmic deceleration, $q_0 = -\ddot{a}a^{-2}$ and jerk, $j_0 = -\ddot{a}\dot{a}^{-3}a^2$, where a is the scale factor and the dots are derivatives with respect to cosmic time. Given our current cosmological model, $q_0 = -0.55$ and $j_0 = 1.0$.

We can neglect the higher order terms and obtain the expression:

$$d_L \approx \frac{cz}{H_0} \left(1 + \frac{1}{2}(1 - q_0)z + \frac{1}{6}(1 - q_0 - 3q_0^2 + j_0)z^2 \right) = \frac{cz}{H_0} Q(z). \quad (7)$$

Substituting into equation (7), we obtain a simplified expression to estimate H_0 :

$$\log_{10} H_0 = \frac{M_0^B + 5a_B + 25}{5}, \quad (8)$$

$$a_B = \log_{10} cz + \log_{10} Q(z) - \frac{m}{5}, \quad (9)$$

where M_0^B is the absolute magnitude of the SN Ia calibrators, a_B is the intercept of the distance–redshift relation, and we take $a_B = 0.714158$ which is the baseline value used in the SH0ES analysis (Riess et al. 2022).

Without applying any additional correction, we measure $M_0^B = -19.247 \pm 0.033$ mag and $H_0 = 73.24 \pm 1.11$ km s⁻¹ Mpc⁻¹. Importantly, depending on the chosen fitting approach and application of our [O II] EW correction (detailed in Cases 1 and 2), we find a difference of 0.47 km s⁻¹ Mpc⁻¹ with H_0 values ranging from 73.04 to 73.51 km s⁻¹ Mpc⁻¹. These results are summarized in Table 4.

As discussed in Section 4.2, we find that the more effective approach in reducing the environmental dependence with Hubble residual is to directly replace the mass step correction with a linear [O II] EW correction. Subsequently, we adopt $M_0^B = -19.253 \pm 0.033$ mag, and determine $H_0 = 73.04 \pm 1.11$ km s⁻¹ Mpc⁻¹ as our nominal result. Ultimately, the shift in H_0 (up by 0.37% or down by 0.27%) is small regardless of the applied [O II] EW correction. This suggests that environmental dependence on SN Ia brightness has a limited effect on reconciling the discrepancy with the CMB result for H_0 , and a significant tension ($\sim 5\sigma$) between these measurements persists.

5.2 Systematic uncertainties

An important bias affecting H_0 measurements, as identified by Rigault et al. (2015), arises from systematic differences in distances derived for SNe Ia in passive and star-forming environments. In our analysis, this bias arises after applying an [O II] EW correction derived from the OzDES hosts to SNe Ia in the calibrator sample (second rung) and not to the SNe Ia in the Hubble flow (third rung). The Hubble flow Pantheon + SNe Ia, which has been corrected for the mass step, lacks spectra of all the host galaxies. Although we cannot apply our [O II] EW correction due to the lack of [O II] EW measurements, we can utilize the OzDES hosts ($z < 0.4$) to assess the potential impact of this bias.

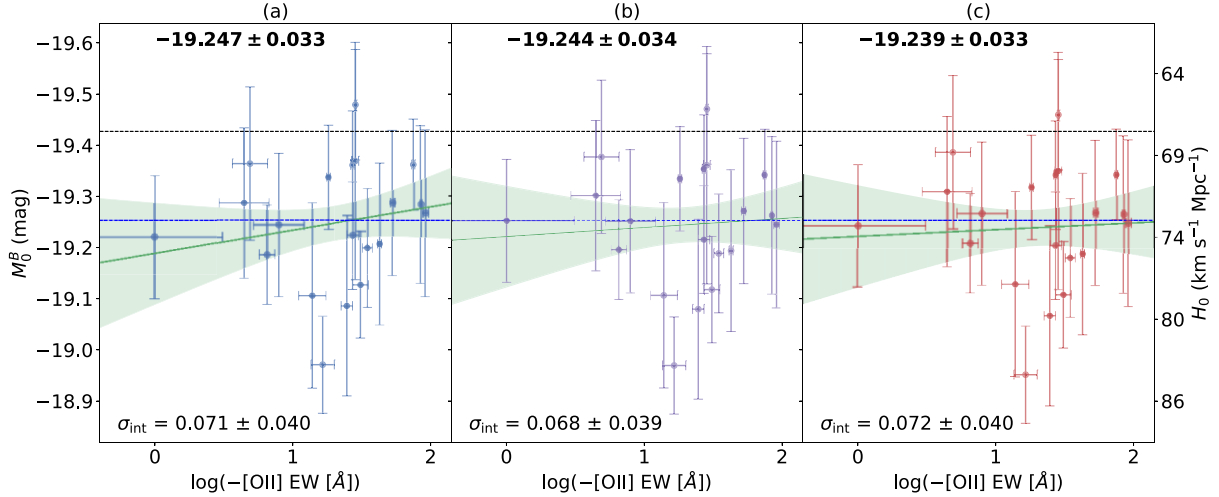


Figure 11. Case 2: Similar to Fig. 10, but here we only apply [O II] corrections to $M_{0,i}^B$, where the mass step has been corrected for. We find mean M_0^B values ranging between -19.247 and -19.239 mag. While the [O II] correction reduces the trend between M_0^B and [O II], its effectiveness is less than in Fig. 10.

To explore the impact of our [O II] EW correction across the second and third rungs, we can determine the differential impact on the variation in SNe Ia luminosities. The bias can be estimated by multiplying the relative differential fraction of host galaxy properties between the second and third rungs by the slope of the correlation with Hubble residuals, e.g. $[\text{O II}]_{\text{slope}} \times ([\text{O II}]_{\text{second}}^{\text{mean}} - [\text{O II}]_{\text{third}}^{\text{mean}})$. The mean $\log(-[\text{O II}] \text{ EW})$ of the WiFeS calibrator galaxies is 1.30, while that of the OzDES hosts ($z < 0.4$) is 1.13. Using the slope measured in Fig. 4, the impact of this bias is ~ 0.009 mag in M_0^B .

To minimize the differential impact and achieve greater consistency with SH0ES, we select the top 80 percent of star-forming OzDES hosts. This selection mirrors the focus on star-forming hosts in the SH0ES analysis, ensuring a more comparable sample. The new OzDES sample has a mean $\log(-[\text{O II}] \text{ EW})$ of 1.36, closer to that of the calibrator galaxies, and gives a reduced impact ~ 0.003 mag in M_0^B . While the smaller size of the applied correction mitigates the differential impact of different stellar populations on H_0 measurements, selecting SN Ia hosts with similar properties across the different rungs can further reduce this impact.

In Tables 5 and 6, we show the systematic errors in our analysis. One of the main differences between Cases 1 and 2, is the differential impact between the different rungs, which is driven by the size of the applied correction. Ultimately, we find that the impact of this systematic in our analysis is minimal.

5.3 Which is more important, host galaxy mass or [O II] EW?

Galaxies are complex systems with many factors driving trends and potentially being intertwined. Ultimately, we want to explore whether stellar mass or [O II] is dominating the correlation with Hubble residuals, or if the two factors are strongly covariant.

In Fig. 9, we compare the effectiveness of host stellar mass and [O II] EW as additional SN Ia light-curve parameters. This analysis is performed by applying different light-curve correction parameters specific to each galaxy property (Table 3). In the first row, we remove the mass step correction (δ_{host}) as defined in DES Collaboration (2024). We find that [O II] EW exhibits the most significant trend (4.0σ) compared with stellar mass and $U - R$ colour. In the second row, a mass step has been applied and along with the ‘BBC4D’ bias correction Popovic et al. (2021).

We find that the slope is flat for both stellar mass and $U - R$. However, we uncover a 2.2σ trend with [O II] EW, which suggests that the correlation has not been entirely removed. In the third row, we instead directly apply our [O II] EW correction as an alternative to implementing a mass step correction. We find this correction is more effective than a mass step, reducing the significance of the trends across all three galaxy parameters to $\leq 1\sigma$. Additionally, we find that this approach results in the smallest intrinsic scatter, further suggesting that [O II] EW does better than the mass step correction. In the fourth row, we also examine the case with all light-curve fitting parameters, but find that this approach is not as effective.

To further probe these trends, we then split our sample by stellar mass or [O II] EW, with roughly the same number of objects in each bin. Then for each of the bins, we plot Hubble residual against the other galaxy parameter ([O II] EW/stellar mass). In Fig. 12, we find that when splitting our sample in bins of [O II] EW, the lines differ. This suggests that the relationship between Hubble residual and stellar mass varies for galaxies with different sSFRs. When instead binning by stellar mass, the lines are more similar and overlap to a greater extent. This suggests that [O II] EW, as an indicator of sSFR, is more prominent in driving the correlation with Hubble residual for our OzDES host galaxies.

5.4 $U - R$ and [O II] EW

As [O II] EW is derived using host galaxy spectroscopy, it would be advantageous to obtain a photometric alternative. One ideal candidate that has been shown to strongly correlate with Hubble residual is $U - R$ rest-frame colour (Kelsey et al. 2021, 2023). In Fig. 13, we fit a linear model to our data and obtain the following relation:

$$\log(-[\text{O II}] \text{ EW}) = -0.620 \times (U - R) + 1.799. \quad (10)$$

Both $U - R$ and the [O II] EW will depend on the stellar population, the amount of dust and its distribution. For example, star-forming regions, may be dustier than regions devoid of star formation.

The bias correction relies on a model describing the residual scatter after standard light-curve corrections are applied. Some of these models (e.g. P21) depend on host properties. The bias corrections could affect the correlations we measure. We find that using a dust-mass bias correction model (Popovic et al. 2021) does result in

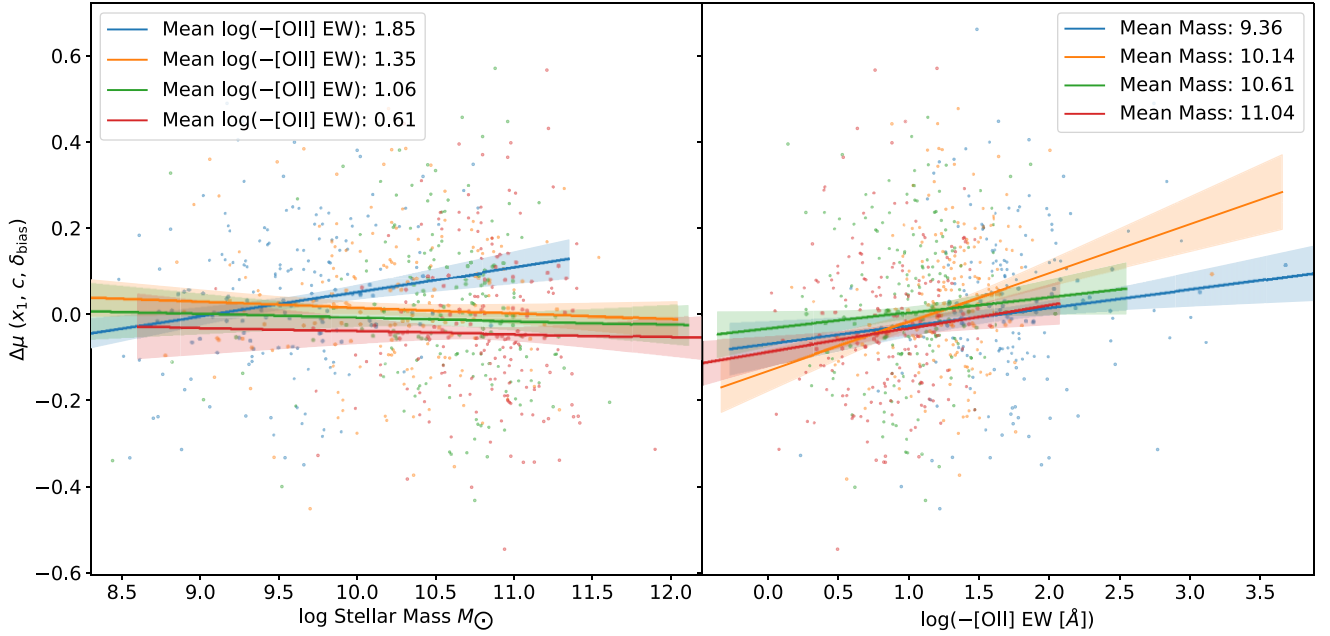


Figure 12. Left: First, we split the OzDES host galaxies into four $[\text{O II}]$ EW bins with roughly the same number of objects. We then plot $\Delta\mu$ against stellar mass for each of the bins. Right: Similar, but instead we split by stellar mass, and then plot $\Delta\mu$ against $[\text{O II}]$ EW.

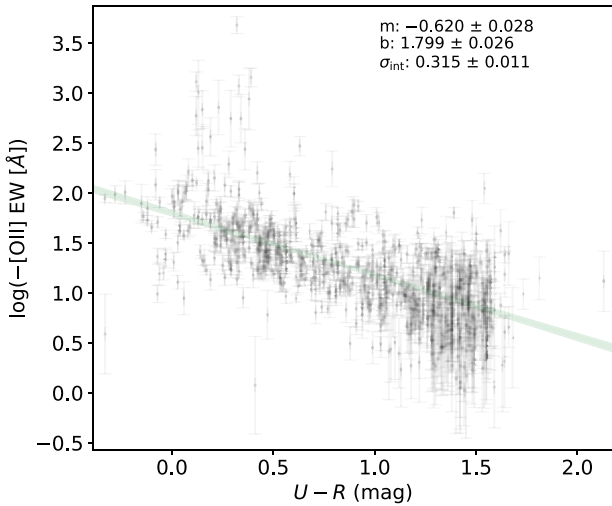


Figure 13. Deriving a photometric proxy for $[\text{O II}]$ EW: We find a very strong correlation between $[\text{O II}]$ EW and $U - R$ rest-frame colour.

reducing the correlation between SN Ia brightness and host galaxy properties, but not entirely. Wiseman et al. (2023) also discuss the potential improvements in having additional galaxy-age light-curve correction parameters, where they find that sSFR is an ideal tracer for exploring this correlation with Hubble residuals.

Ultimately, the commonly used approach to correcting SN Ia light curves using host stellar mass does not fully remove the observed trends. Our analysis finds that the most effective strategy for improving the standardization of SN Ia light curves involves both the ‘BBC4D’ bias correction term and an additional parameter accounting for host sSFR, a tracer of stellar population age, instead of relying solely on stellar mass. We highlight the further impact that using $[\text{O II}]$ EW can in correcting for host galaxy properties such as

stellar mass or $U - R$. Importantly, this can be achieved with one spectroscopic feature.

5.5 Future work

It would be interesting to increase our sample of SN Ia calibrators to cover all of the 37 galaxies in the SHOES analysis. Additionally, we can extend our SN Ia calibrator sample to include TRGB (Freedman 2021) and SBF hosts (Jensen et al. 2021). This would help provide more detail on the variation of SNe Ia in different galactic environments and perhaps a better understanding of the origin of differences in H_0 for different calibrators.

In the coming years, surveys such as TiDES (Swann et al. 2019) will use 4MOST and will be able to observe and obtain spectra for tens of thousands of SN Ia host galaxies. For each galaxy, we can obtain a measurement of $[\text{O II}]$ EW and a significantly larger sample size will help drive down statistical uncertainties.

Additionally, the size of the OzDES sample allows one to explore how the relation changes with redshift. This can be probed with even greater statistical significance with TiDES. Redshift evolution could bias the value of cosmological parameters that are determined from SNe Ia.

The additional benefit of using the $[\text{O II}]$ emission line, is that it can be observed up to $z \sim 1$, unlike the stronger Balmer lines. However, this may bias the sample towards galaxies which are brighter, and have a higher sSFR. With increasing redshift, it becomes increasingly difficult to measure the $[\text{O II}]$ EW. Ultimately, our work highlights the value of $[\text{O II}]$ EW as an effective alternative to stellar mass in standardizing SN Ia luminosities.

6 CONCLUSION

The SN Ia standardizable candle forms an important component of the astronomical distance ladder. However, correlations between Hubble residuals and the properties of the SN Ia hosts impact their

use in measuring distances. Using [O II] EW in our sample of 707 OzDES host galaxies, we obtain a correction that is more significant than the commonly used mass step correction. We then apply our [O II] EW correction to a sample of 20 SN Ia calibrator galaxies observed using WiFeS, and calibrate the SN Ia absolute magnitude. Applying this result to the Pantheon + analysis, we find H_0 values ranging between 73.04 and 73.51 km s⁻¹ Mpc⁻¹, depending on our fitting approach in deriving the [O II] EW correction. The change in the value of H_0 is negligible when using the [O II] EW in place of or in addition to the mass of host galaxies in adjusting SN Ia luminosities. The tension between nearby and distant measurements of H_0 remains.

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DATA AVAILABILITY

The DES-SN photometric SN Ia catalogue will be made available as part of the DES5YR SN cosmology analysis at <https://des.ncsa.illinois.edu/releases/sn>. The OzDES-DR2 spectra used in this paper can be publicly accessed at <https://docs.datacentral.org.au/ozdes/overview/dr2>.

REFERENCES

- Briday M. et al., 2022, *A&A*, 657, A22
- Brout D., Scolnic D., 2021, *ApJ*, 909, 26
- Brout D. et al., 2022, *ApJ*, 938, 110
- Cappellari M., 2016, *MNRAS*, 466, 798

- Cappellari M., 2017, *MNRAS*, 466, 798
- Cappellari M., Emsellem E., 2004, *PASP*, 116, 138
- Childress M. et al., 2013, *ApJ*, 770, 108
- Childress M. J., Vogt F. P. A., Nielsen J., Sharp R. G., 2014, *Ap&SS*, 349, 617
- Childress M. J. et al., 2017, *MNRAS*, 472, 273
- D’Andrea C. B. et al., 2011, *ApJ*, 743, 172
- DES Collaboration, 2016, *MNRAS*, 460, 1270
- DES Collaboration et al., 2024, *ApJ*, 973, L14
- Dhawan S., Brout D., Scolnic D., Goobar A., Riess A. G., Miranda V., 2020, *ApJ*, 894, 54
- Di Valentino E. et al., 2021, *Class. Quantum Gravity*, 38, 153001
- Dixon M. et al., 2022, *MNRAS*, 517, 4291
- Dopita M., Hart J., McGregor P., Oates P., Bloxham G., Jones D., 2007, *Ap&SS*, 310, 255
- Dopita M. et al., 2010, *Ap&SS*, 327, 245
- Flaugher B. et al., 2015, *AJ*, 150, 150
- Freedman W. L., 2021, *ApJ*, 919, 16
- Galbany L. et al., 2022, *A&A*, 659, A89
- Guy J. et al., 2010, *A&A*, 523, A7
- Hartley W. G. et al., 2022, *MNRAS*, 509, 3547
- Hinton S., Davis T. M., Lidman C., Glazebrook K., Lewis G., 2016, *Astron. Comput.*, 15, 61
- Hoormann J. K. et al., 2019, *MNRAS*, 487, 3650
- Jensen J. B. et al., 2021, *ApJS*, 255, 21
- Kelly P. L., Hicken M., Burke D. L., Mandel K. S., Kirshner R. P., 2010, *ApJ*, 715, 743
- Kelsey L. et al., 2021, *MNRAS*, 501, 4861
- Kelsey L. et al., 2023, *MNRAS*, 519, 3046
- Kenworthy W. D. et al., 2021, *AJ*, 923, 265
- Kessler R., Scolnic D., 2017, *ApJ*, 836, 56
- Kessler R. et al., 2015, *AJ*, 150, 172
- Khetan N. et al., 2021, *A&A*, 647, A72
- Lampeitl H. et al., 2010, *ApJ*, 722, 566
- Lidman C. et al., 2020, *MNRAS*, 496, 19
- Martin B., Lidman C., Brout D., Tucker B. E., Dixon M., Armstrong P., 2024, *MNRAS*, 533, 2640
- Meldorf C. et al., 2023, *MNRAS*, 518, 1985
- Möller A., de Boissière T., 2020, *MNRAS*, 491, 4277
- Möller A. et al., 2022, *MNRAS*, 514, 5159
- Moreno-Raya M. E., Galbany L., López-Sánchez Á. R., Mollá M., González-Gaitán S., Vilchez J. M., Carnero A., 2018, *MNRAS*, 476, 307
- Morganson E. et al., 2018, *PASP*, 130, 074501
- Orellana G. et al., 2017, *A&A*, 602, A68
- Pan Y. C. et al., 2014, *MNRAS*, 438, 1391
- Phillips M. M., 1993, *ApJ*, 413, L105
- Planck Collaboration VI, 2020, *A&A*, 641, A6
- Popovic B., Brout D., Kessler R., Scolnic D., Lu L., 2021, *ApJ*, 913, 49
- Popovic B., Brout D., Kessler R., Scolnic D., 2023, *ApJ*, 945, 84
- Riess A. G. et al., 2022, *ApJ*, 934, L7
- Rigault M. et al., 2015, *ApJ*, 802, 20
- Rigault M. et al., 2020, *A&A*, 644, A176
- Roman M. et al., 2018, *A&A*, 615, A68
- Rose B. M., Garnavich P. M., Berg M. A., 2019, *ApJ*, 874, 32
- Salpeter E. E., 1955, *ApJ*, 121, 161
- Scolnic D. et al., 2022, *ApJ*, 938, 113
- Smith M. et al., 2020, *MNRAS*, 494, 4426
- Sullivan M. et al., 2010, *MNRAS*, 406, 782
- Swann E. et al., 2019, *The Messenger*, 175, 58
- Taylor G. et al., 2023, *MNRAS*, 520, 5209
- Triani D. P., Sinha M., Croton D. J., Dwek E., Pacifici C., 2021, *MNRAS*, 503, 1005
- Tripp R., 1998, *A&A*, 331, 815
- Uddin S. A., Mould J., Lidman C., Ruhlmann-Kleider V., Zhang B. R., 2017, *ApJ*, 848, 56
- Vazdekis A., Sánchez-Blázquez P., Falcón-Barroso J., Cenarro A. J., Beasley M. A., Cardiel N., Gorgas J., Peletier R. F., 2010, *MNRAS*, 404, 1639
- Vazdekis A., Koleva M., Ricciardelli E., Röck B., Falcón-Barroso J., 2016, *MNRAS*, 463, 3409
- Vincenzi M. et al., 2024, *ApJ*, 975, 86
- Wiseman P. et al., 2020, *MNRAS*, 495, 4040
- Wiseman P., Sullivan M., Smith M., Popovic B., 2023, *MNRAS*, 520, 6214
- Yuan F. et al., 2015, *MNRAS*, 452, 3047
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