

Impact of flavour coupling on $SO(10)$ -inspired leptogenesis

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Abstract

We discuss the impact of flavour coupling on the predictions of low energy neutrino parameters from $SO(10)$ -inspired leptogenesis (SO10INLEP). The right-handed (RH) neutrino mass spectrum is strongly hierarchical and successful leptogenesis relies on generating the asymmetry from next-to-lightest RH neutrino decays (N_2 -leptogenesis) and circumventing the lightest RH neutrino washout. These two conditions yield distinctive predictions such as a lower bound on the lightest neutrino mass $m_1 \gtrsim 1$ meV. We first review the status of SO10INLEP, noticing how cosmological observations are now testing a particular neutrino mass window, $m_1 \simeq (10\text{--}30)$ meV, where only the first octant is allowed and a large range of values for the Dirac phase is excluded. Including flavour coupling, we find that the lower bound relaxes to $m_1 \gtrsim 0.65$ meV. Moreover, new muon-dominated solutions appear slightly relaxing the upper bound on the atmospheric mixing angle. We also study the impact on strong thermal SO10INLEP (ST-SO10INLEP) scenario where, in addition to successful leptogenesis, one can washout a large pre-existing asymmetry. Contrarily to naive expectations, for which flavour coupling could jeopardise the scenario, allowing a large pre-existing asymmetry to survive unconditionally, we show, and explain analytically, that ST-SO10INLEP is still viable within almost the same allowed region of parameters. There is even a slight relaxation of the m_1 viable window from $(9\text{--}30)$ meV to $(4\text{--}40)$ meV for a 10^{-3} pre-existing asymmetry. The new results from atmospheric neutrinos, mildly favouring normal ordering and first octant, are now in nice agreement with the predictions of ST-SO10INLEP. Intriguingly, the predicted $0\nu\beta\beta$ signal is starting to be within the reach of KamLAND-Zen.

1 Introduction

We are entering an important stage for the prospect of testing high scale thermal leptogenesis [1] within minimal type I seesaw models [2], the minimal scenario of leptogenesis [3]. Colliders have found no evidence of new physics at the TeV scale and below so far, placing strong constraints (and doubts) on low scale leptogenesis scenarios. At the same time, low energy neutrino experiments continue to progress steadily and in the next years they will be able to provide information on the neutrino mixing unknowns and to test the absolute neutrino mass scale in regions of the parameter space that are quite important for high scale leptogenesis scenarios. Moreover, the discovery of gravitational waves [4] is stimulating many new ideas on how to test high scale leptogenesis [5, 6, 7, 8]. For these reasons, the often proclaimed statement for which high scale leptogenesis scenarios are untestable is today outdated.

An example of well motivated, testable, high scale leptogenesis scenario is SO10INLEP. This relies on quite minimal assumptions on the Dirac neutrino mass matrix, so-called $SO(10)$ -inspired conditions [9]. They are typically realised within various grandunified models, in particular within $SO(10)$ models. With these assumptions, and barring the highly fine-tuned compact spectrum case [10, 11, 12], the emerging spectrum of RH neutrinos is strongly hierarchical. In particular the lightest RH neutrino, N_1 , has, typically, a mass $M_1 \sim (10^5\text{--}10^6)$ GeV, well below the lower bound for successful leptogenesis [13, 14]. On the other hand, the next-to-lightest RH neutrino, N_2 , turns out to have a mass with just the correct order-of-magnitude, $M_2 \sim 10^{11}$ GeV, to produce an asymmetry able to explain the observed value, realising successful N_2 -leptogenesis [15]. In order for the CP asymmetries of N_2 to be sufficiently large, the existence of a third heavier RH neutrino species, N_3 , is necessary to produce the needed interference in N_2 decays. This nicely fits within the properties of $SO(10)$ -models that indeed, notoriously, predict the existence of three RH neutrino species, one for each of the three family $SO(10)$ fermionic sixteenth-plet representations. When $SO(10)$ -inspired conditions are imposed, typically one obtains $M_3 \sim (10^{14}\text{--}10^{16})$ GeV. However, notice that for successful thermal leptogenesis it is sufficient that the reheat temperature $T_{\text{RH}} \gtrsim M_2 \sim 10^{11}$ GeV, since only the N_2 's need to thermalise, while N_3 plays a role at the virtual level only, in the interfering loop diagrams for N_2 -decays. Still this is a crucial role, since this interference is the source of the N_2 - CP asymmetries and, ultimately, of the observed baryon asymmetry. For this reason, when N_3 decouples, in the limit $M_3 \gg 10^{16}$ GeV, all flavoured CP asymmetries of N_2 tend to vanish. This yields an upper bound on M_3 and, consequently, from the seesaw formula, a lower bound on the lightest neutrino mass [16]. Although one can have a sizeable asym-

metry generated from N_2 -decays at very high temperatures, $T \sim M_2$, for long time such a strongly hierarchical RH neutrino spectrum has been regarded as an obstacle to reproduce the correct observed baryon asymmetry within SO10INLEP. The reason is that, within an unflavoured description of SO10INLEP, the washout from N_1 at $T \sim M_1$ is necessarily strong and a successful N_2 -leptogenesis scenario cannot be realised for such a strongly hierarchical spectrum. However, when flavour effects are included [17, 18], the calculation of the $B - L$ asymmetry splits into three different contributions, one for each charged lepton flavour [19]. The strong washout from N_1 also splits into three weaker components. In this way an allowed region in the space of low energy neutrino parameters opens up and successful SO10INLEP can be attained [16, 20]. The asymmetry is dominantly produced in the tauon flavour. Subdominant muon-dominated flavour solutions also exist, while electron-dominated solutions just fall short of reproducing the observed asymmetry.¹

It is interesting that successful SO10INLEP is realised only within a region of the entire low energy neutrino parameter space. This yields constraints, and predictions, to be confronted with the experimental results [20]. A first important feature of SO10INLEP predictions is that the allowed region for inverted ordered neutrino masses (IO) is marginal [20, 12], since it necessarily requires the atmospheric neutrino mixing angle to be in the second octant and a lower bound on the lightest neutrino mass. For $\theta_{23} \lesssim 50^\circ$, as suggested by latest neutrino oscillation data global analyses at 3σ [23], one has a lower bound on the lightest neutrino mass $m_1 \gtrsim 20$ meV, corresponding to $\sum_i m_i \gtrsim 130$ meV. This is in tension with the cosmological upper bound on the sum of neutrino masses obtained combining CMB anisotropy, baryon acoustic oscillation observations and supernovae data [24, 25, 26, 27]

$$\sum_i m_i < 120 \text{ meV} \quad (95\% \text{ C.L.}) . \quad (1)$$

For this reason, at 3σ , SO10INLEP is viable only for normal ordering (NO) and this is in nice agreement with the latest results from atmospheric neutrino oscillation experiments, that tend to favour NO [28]. This is also confirmed by the most recent global analyses disfavouring IO at $\sim 2.5\sigma$ [23]. Therefore, in the following we will focus exclusively on NO and we will always refer to the cosmological upper bound on the lightest neutrino mass

$$m_1 < 30 \text{ meV} \quad (95\% \text{ C.L.}) , \quad (2)$$

¹They appear marginally in a supersymmetric version [21]. We should also mention that there is a second way to circumvent the N_1 washout, even when this is strong in all three flavours [21]. If one the lightest Dirac neutrino mass is at least two orders of magnitude lower than the up quark mass, then M_1 gets lower than the sphaleron freeze-out temperature $T_{\text{sph}}^{\text{off}} \simeq 132 \text{ GeV}$ [22]. In this case the washout acts only on the lepton asymmetry but not on the observed baryon asymmetry.

derived from Eq. (1) in the case of NO.²

As mentioned, the most interesting constraint from successful SO10INLEP is the existence of the lower bound on the lightest neutrino mass [16, 20]. This depends in a non-trivial way on the CP -violating Dirac phase δ and the atmospheric mixing angle θ_{23} [29]. There is a large region of the plane δ versus θ_{23} that is already incompatible with SO10INLEP at 95% C.L. since the cosmological upper bound on neutrino masses Eq. (1) is violated. In the semi-hierarchical regime, for $10 \text{ meV} \lesssim m_1 \lesssim 30 \text{ meV}$, most stringent constraints hold within SO10INLEP. In particular, SO10INLEP is incompatible with the atmospheric mixing angle in the second octant. On the other hand, this is also a very interesting neutrino mass regime, since only for m_1 in the (10–30) meV range ST-SO10INLEP can be realised [30]. This is a special scenario where, in addition to the condition of successful leptogenesis, one can also wash out a pre-existing asymmetry as large as 10^{-1} [31]. For a hierarchical RH neutrino spectrum, without even imposing $SO(10)$ -inspired conditions, successful strong thermal leptogenesis can only be realised for quite special conditions: a tauon-dominated N_2 -leptogenesis scenario [31] and for $m_1 \gtrsim 10 \text{ meV}$ [32]. It is then quite non trivial that within successful SO10INLEP, one can also satisfy the conditions for strong thermal leptogenesis for a subset of the solutions. In particular, they can only be satisfied for $m_1 \lesssim 30 \text{ meV}$, just coinciding with the current cosmological upper bound Eq. (2). This is another important point showing how we are now just entering a crucial stage in testing SO10INLEP and its strong thermal case ST-SO10INLEP.

For this reason, it is of great importance to understand how solid and accurate these theoretical predictions are. The constraints on the low energy neutrino parameters have been derived in various papers and with independent codes. They have also been derived analytically, expressing the RH neutrino mass spectrum and mixing matrix in terms of the low energy neutrino parameters [12, 33]. In this way one obtains an analytical expression of the final baryon asymmetry in terms of the low energy neutrino parameters. However, the calculation of the asymmetry has been done within a set of approximations that neglects different effects. Three of them are at the level of field theory description:

- Flavour coupling effects from spectator processes [34, 35, 36].
- Running of parameters from radiative corrections.
- Partial equilibration of spectator processes [37].

²In the IO case the upper bound Eq. (1) would imply an even more stringent upper bound on the lightest neutrino mass, $m_3 < 16 \text{ meV}$ (95% C.L.). The significant difference between the upper bound on the lightest neutrino mass in the NO and IO cases, shows that the upper bound Eq. (1) is essentially ruling out quasi-degenerate neutrinos at 95% C.L. .

Two additional ones concern the kinetic theory description. The asymmetry has been so far calculated solving simple momentum integrated rate equations in the fully flavoured regime, neglecting:

- Momentum dependence (requiring solution of the full Boltzmann equations) [38].
- Decoherence (requiring solution of density matrix equation) [34, 17].

In this paper, encouraged by the current agreement of (ST-)SO10INLEP predictions with neutrino oscillation experiment results and by the fact that absolute neutrino mass scale experiments are starting to test a crucial regime, as discussed above, we start implementing these effects, studying the impact of flavour coupling effects on (ST-)SO10INLEP. This is certainly the most urgent extension of the calculation of the asymmetry within (ST-)SO10INLEP to be considered. In particular, in the case of ST-SO10INLEP one can even legitimately suspect that flavour coupling could jeopardise the whole scenario. This is because a contribution from a large pre-existing asymmetry might survive until the present, leaking from one flavour to another, thus escaping both N_2 and N_1 washout. As we will discuss in detail, and anticipated in the abstract, this does not happen and, ultimately, even in the ST-SO10INLEP case, account of flavour coupling yields just some slight corrections that, however, might prove to be important when atmospheric neutrino mixing angle and Dirac phase will be measured with errors $\sim 1^\circ$ and $\sim 10^\circ$, respectively, by DUNE+T2HK [39]. At the same time we will highlight some new important aspects of (ST-)SO10INLEP predictions. When these are confronted with the latest results from global analyses on neutrino mixing parameters, we show how the current best fit for NO and first octant, nicely agrees with (ST-)SO10INLEP predictions, removing the tension that was existing with previous results [40] favouring second octant for θ_{23} . Moreover, as we mentioned, even without imposing strong thermal leptogenesis, for $m_1 \sim (10\text{--}30)\text{ meV}$ the second octant is not compatible with SO10INLEP constraints. In addition, there are large ranges of δ that are excluded even for $m_1 \lesssim 10\text{ meV}$. Therefore, during next years, neutrino oscillation experiments will either be able to increase the statistical significance of the agreement or of course, will be able to rule out SO10INLEP. It is then important, not to draw incorrect conclusions, to start a systematic study of theoretical uncertainties to reduce them at the level of the expected experimental errors.

This is the main aim of the paper, that is structured in the following way. In Section 2 we review SO10INLEP, discussing the assumptions and the equations to calculate the asymmetry, describing the results that we obtained ignoring flavour coupling (we have re-derived then once more confirming previous analyses and updating the experimental

constraints in the plots) and comparing the predictions to the latest experimental constraints. In Section 3 we discuss flavour coupling and how this modifies the calculation of the final asymmetry. We show the results for some benchmark points and compare them to the case when flavour coupling is ignored. We also point out how the calculation of the asymmetry becomes much more CPU time consuming. For that reason, the derivation of the hyper-surface in the space of parameters, giving the allowed region, becomes much more challenging to derive. This is indeed so convoluted to require at least one million of points for a precise determination. The use of the analytical expressions for the RH neutrino masses and mixing matrix [12, 33] greatly help in this respect. We show the results of the scatter plots giving both the three-dimensional projection of the hypersurface in the space $(\delta, \theta_{23}, m_1)$, as previously done in [29] neglecting flavour coupling, and different two-dimensional projections. We also highlight the different flavour dominance for each point, confirming that, even including flavour coupling effects, the bulk of the solutions, approximately 90%, are tauonic solutions, while muonic solutions represent a subdominant 10% contribution. However, when flavour coupling is included, we show that a very small 0.1% contribution of electronic solutions appear and also new muonic solutions. These new muonic solutions fall in a region of the parameter space that would otherwise be excluded.³ In Section 4 we discuss ST-SO10INLEP, showing how the calculation of the evolution of a large pre-existing asymmetry, and its relic value, is modified by the account of flavour coupling. Here, we show how the allowed region not only survives but even, for some values of δ , allows slightly larger values of θ_{23} . Moreover, the allowed range of m_1 values also slightly enlarges. Finally, in Section 5, we draw some final remarks on the importance of (ST-)SO10INLEP and of our results within the current quest for new physics.

2 $SO(10)$ -inspired leptogenesis

In this section we review general features of SO10INLEP without the inclusion of flavour coupling. We briefly discuss the set of assumptions that define SO10INLEP and then we show how the requirement of successful SO10INLEP produces constraints on the low energy neutrino parameters that we compare with the latest neutrino oscillation experimental results. In this way we can discuss the status of SO10INLEP showing that there is currently a non-trivial compatibility between SO10INLEP predictions and low energy

³A special subset of the muonic solutions that we obtain were also found within a specific Pati-Salam grandunified model with discrete flavour symmetries [41]. However, this special subset is now excluded by the latest neutrino oscillation experimental results since they require $\theta_{23} \simeq 54^\circ$.

neutrino experimental data. We also show how with the new cosmological upper bound on neutrino masses, we are starting to test a very interesting neutrino mass range for SO10INLEP.

2.1 Seesaw mechanism and low energy neutrino data

Inspired by $SO(10)$ models, we extend the SM introducing three RH neutrinos N_{R1}, N_{R2} and N_{R3} with Yukawa couplings h and a Majorana mass term M . In the flavour basis, where both charged lepton mass matrices m_ℓ and M are diagonal, one can write the leptonic mass terms generated after spontaneous symmetry breaking by the Higgs expectation value $v = 174 \text{ GeV}$ as ($\alpha = e, \mu, \tau$ and $I = 1, 2, 3$)

$$-\mathcal{L}_{\ell+\nu}^m = \overline{\alpha_L} D_{m_\ell} \alpha_R + \overline{\nu_{\alpha L}} m_{D\alpha I} N_{RI} + \frac{1}{2} \overline{N_{RI}^c} D_M N_{RI} + \text{h.c.}, \quad (3)$$

where $D_{m_\ell} \equiv \text{diag}(m_e, m_\mu, m_\tau)$ is the diagonal charged lepton mass matrix, $D_M \equiv \text{diag}(M_1, M_2, M_3)$ is the diagonal Majorana mass matrix and $m_D = hv$ is the neutrino Dirac mass matrix. In the seesaw limit, for $M \gg m_D$, the mass spectrum splits into two sets of Majorana neutrino eigenstates, a light set, ν_1, ν_2 and ν_3 , with masses $m_1 \leq m_2 \leq m_3$, given by the seesaw formula [2]

$$D_m = U^\dagger m_D \frac{1}{D_M} m_D^T U^\star, \quad (4)$$

with $D_m = \text{diag}(m_1, m_2, m_3)$, and a heavy set, N_1, N_2 and N_3 , with masses almost coinciding with the three M_I in D_M . The matrix U , diagonalising the light neutrino mass matrix $m_\nu = -m_D M^{-1} m_D^T$ in the weak basis, can be identified with the PMNS lepton mixing matrix. This can be parameterised in terms of the usual three mixing angles θ_{ij} , the Dirac phase δ and the Majorana phases ρ and σ , as

$$U = \begin{pmatrix} c_{12} c_{13} & s_{12} c_{13} & s_{13} e^{-i\delta} \\ -s_{12} c_{23} - c_{12} s_{23} s_{13} e^{i\delta} & c_{12} c_{23} - s_{12} s_{23} s_{13} e^{i\delta} & s_{23} c_{13} \\ s_{12} s_{23} - c_{12} c_{23} s_{13} e^{i\delta} & -c_{12} s_{23} - s_{12} c_{23} s_{13} e^{i\delta} & c_{23} c_{13} \end{pmatrix} \text{diag}(e^{i\rho}, 1, e^{i\sigma}), \quad (5)$$

where $s_{ij} \equiv \sin \theta_{ij}$ and $c_{ij} \equiv \cos \theta_{ij}$. As we said, successful SO10INLEP cannot be realised within IO, compatibly with the latest upper bounds on neutrino masses and atmospheric neutrino mixing angle. For this reason we only consider NO. In this case latest global analyses of neutrino oscillation experiment results, including atmospheric neutrino data from Super-Kamiokande and IceCube collaborations, find for the mixing

angles and leptonic Dirac phase δ the following best fit values, 1σ errors and 3σ intervals [23]:

$$\begin{aligned}\theta_{13} &= 8.56^\circ \pm 0.11^\circ \in [8.19^\circ, 8.89^\circ], \\ \theta_{12} &= 33.68^{+0.73^\circ}_{-0.70^\circ} \in [31.63^\circ, 35.95^\circ], \\ \theta_{23} &= 43.3^{+1.0^\circ}_{-0.8^\circ} \in [41.3^\circ, 49.9^\circ], \\ \delta &= -148^{+26^\circ}_{-41^\circ} \in [-236^\circ, 4^\circ].\end{aligned}\tag{6}$$

As one can notice, there is a 3σ exclusion interval, $\delta \ni [4^\circ, 134^\circ]$, for the Dirac phase that disfavors $\sin \delta > 0$. Neutrino oscillation experiments are also sensitive to squared neutrino mass differences, finding for the solar neutrino mass scale

$$m_{\text{sol}} \equiv \sqrt{m_2^2 - m_1^2} = (8.65 \pm 0.11) \text{ meV},\tag{7}$$

and for the atmospheric neutrino mass scale

$$m_{\text{atm}} \equiv \sqrt{m_3^2 - m_1^2} = (50.1 \pm 0.2) \text{ meV}.\tag{8}$$

The sum of the two scales yields a lower bound on the sum of the neutrino masses:

$$\sum_i m_i \geq m_{\text{sol}} + m_{\text{atm}} = (58.75 \pm 0.25) \text{ meV} \geq 58.25 \text{ meV} \quad (95\% \text{ C.L.}).\tag{9}$$

No neutrinoless double beta ($0\nu\beta\beta$) decay signal has been detected so far. This implies that there are no experimental constraints on the Majorana phases and that experiments can only place an upper bound on the effective $0\nu\beta\beta$ neutrino mass $m_{ee} \equiv |m_{\nu ee}|$. The most stringent one has been set by the KamLAND-Zen collaboration, that found [42]

$$m_{ee} \leq (28\text{--}122) \text{ meV} \quad (90\% \text{ C.L.}),\tag{10}$$

where the range accounts for nuclear matrix element uncertainties. This translates into an upper bound on the lightest neutrino mass $m_1 \leq (84\text{--}353) \text{ meV}$ (90% C.L.). When this is compared with the upper bound Eq. (2) from cosmological observations (assuming Λ CDM), one can clearly see how the latter is much more stringent.⁴ However, a future $0\nu\beta\beta$ positive signal with a measurement of m_{ee} would not just provide information on the absolute neutrino mass scale but also some partial information on the Majorana phases.

Finally, the KATRIN experiment has recently placed the upper bound

$$m_{\nu_e} \lesssim 0.45 \text{ eV} \quad (90\% \text{ C.L.})\tag{11}$$

on the effective electron neutrino mass [43]. Since this falls in the quasi-degenerate limit, it translates into an equal upper bound on m_1 .

⁴It is even more stringent than it looks like, considering that the cosmological upper bound is at 95% C.L. and the upper bound from $0\nu\beta\beta$ at 90% C.L. .

2.2 Combining seesaw and $SO(10)$ -inspired conditions

Let us now briefly review seesaw models when $SO(10)$ -inspired conditions [9] are imposed on the neutrino Dirac mass matrix. This can be diagonalised with the so-called singular value decomposition (sometimes also referred to as bi-unitary parameterisation) as

$$m_D = V_L^\dagger D_{m_D} U_R, \quad (12)$$

where $D_{m_D} \equiv \text{diag}(m_{D1}, m_{D2}, m_{D3})$ and V_L and U_R are two unitary matrices acting respectively on the left-handed (LH) and RH neutrino fields and operating the transformation from the weak basis (where m_ℓ is diagonal) to the Yukawa basis (where m_D is diagonal).

Parameterising the neutrino Dirac masses m_{Di} in terms of the up quark masses,⁵

$$m_{D1} = \alpha_1 m_{\text{up}}, \quad m_{D2} = \alpha_2 m_{\text{charm}}, \quad m_{D3} = \alpha_3 m_{\text{top}}, \quad (13)$$

we impose $SO(10)$ -inspired conditions [9, 10, 16] defined as

$$(i) \quad \alpha_i = \mathcal{O}(1); \quad (ii) \quad I \leq V_L \lesssim V_{CKM}. \quad (14)$$

With the latter we imply that using for the matrix V_L the same parameterisation as for the leptonic mixing matrix U , the three mixing angles θ_{12}^L , θ_{23}^L and θ_{13}^L do not have values much larger than the three mixing angles in the CKM matrix and in particular $\theta_{12}^L \lesssim \theta_c \simeq 13^\circ$, where θ_c is the Cabibbo angle.

Let us now define $M \equiv U_R^\dagger D_M U_R^T$ and $\tilde{m}_\nu \equiv V_L m_\nu V_L^T$. These are, respectively, the Majorana mass matrix and the light neutrino mass matrix in the Yukawa basis. In this way from the seesaw formula Eq. (4), using the singular value decomposed form Eq. (12) for m_D , we obtain an expression for the inverse Majorana mass matrix in terms of low energy neutrino parameters and theory parameters constrained by $SO(10)$ -inspired conditions:

$$M^{-1} \equiv U_R D_M U_R^T = -D_{m_D}^{-1} \tilde{m}_\nu D_{m_D}^{-1}. \quad (15)$$

From this, diagonalising the matrix on the RH side of Eq. (15), one can derive expressions for the three RH neutrino masses and the RH neutrino mixing matrix U_R in terms of m_ν , V_L and the three α_i 's.

From the analytical procedure discussed in [10, 12, 33], one finds simple expressions for the three RH neutrino masses,

$$M_1 \simeq \frac{\alpha_1^2 m_{\text{up}}^2}{|(\tilde{m}_\nu)_{11}|}, \quad M_2 \simeq \frac{\alpha_2^2 m_{\text{charm}}^2}{m_1 m_2 m_3} \frac{|(\tilde{m}_\nu)_{11}|}{|(\tilde{m}_\nu^{-1})_{33}|}, \quad M_3 \simeq \alpha_3^2 m_{\text{top}}^2 |(\tilde{m}_\nu^{-1})_{33}|, \quad (16)$$

⁵For the values of the up-quark masses at the scale of leptogenesis ($\sim 10^{11}$ GeV), we adopt $(m_{\text{up}}, m_{\text{charm}}, m_{\text{top}}) = (1 \text{ MeV}, 400 \text{ MeV}, 100 \text{ GeV})$ [44].

and for the RH neutrino mixing matrix

$$U_R \simeq \begin{pmatrix} 1 & -\frac{m_{D1}}{m_{D2}} \frac{\tilde{m}_{\nu 12}^*}{\tilde{m}_{\nu 11}} & \frac{m_{D1}}{m_{D3}} \frac{(\tilde{m}_{\nu}^{-1})_{13}^*}{(\tilde{m}_{\nu}^{-1})_{33}^*} \\ \frac{m_{D1}}{m_{D2}} \frac{\tilde{m}_{\nu 12}}{\tilde{m}_{\nu 11}} & 1 & \frac{m_{D2}}{m_{D3}} \frac{(\tilde{m}_{\nu}^{-1})_{23}^*}{(\tilde{m}_{\nu}^{-1})_{33}^*} \\ \frac{m_{D1}}{m_{D3}} \frac{\tilde{m}_{\nu 13}}{\tilde{m}_{\nu 11}} & -\frac{m_{D2}}{m_{D3}} \frac{(\tilde{m}_{\nu}^{-1})_{23}}{(\tilde{m}_{\nu}^{-1})_{33}} & 1 \end{pmatrix} D_{\Phi}, \quad (17)$$

with the three phases in $D_{\Phi} \equiv \text{diag}(e^{-i\frac{\Phi_1}{2}}, e^{-i\frac{\Phi_2}{2}}, e^{-i\frac{\Phi_3}{2}})$ given by [33]

$$\Phi_1 = \text{Arg}[-\tilde{m}_{\nu 11}^*], \quad (18)$$

$$\Phi_2 = \text{Arg}\left[\frac{\tilde{m}_{\nu 11}}{(\tilde{m}_{\nu}^{-1})_{33}}\right] - 2(\rho + \sigma) - 2(\rho_L + \sigma_L), \quad (19)$$

$$\Phi_3 = \text{Arg}[-(\tilde{m}_{\nu}^{-1})_{33}]. \quad (20)$$

One can also derive an expression for the orthogonal matrix. Starting from its definition $\Omega = D_m^{-\frac{1}{2}} U^\dagger m_D D_M^{-\frac{1}{2}}$ [45] that, using Eq. (12), becomes [16]

$$\Omega = D_m^{-\frac{1}{2}} U^\dagger V_L^\dagger D_{m_D} U_R D_M^{-\frac{1}{2}} \quad (21)$$

or, in terms of its matrix elements ($\alpha = e, \mu, \tau$; $i = 1, 2, 3$; $I = 1, 2, 3$),⁶

$$\Omega_{iJ} \simeq \frac{1}{\sqrt{m_i M_J}} \sum_{l=1}^3 m_{Dl} U_{\alpha i}^* V_{Ll\alpha}^* U_{RlJ}, \quad (22)$$

one finds [33]

$$\Omega \simeq \begin{pmatrix} i \frac{(\tilde{m}_{\nu} W^*)_{11}}{\sqrt{m_1 \tilde{m}_{\nu 11}}} \sqrt{\frac{m_2 m_3 (\tilde{m}_{\nu}^{-1})_{33}}{\tilde{m}_{\nu 11}}} \left(W_{21}^* - W_{31}^* \frac{(\tilde{m}_{\nu}^{-1})_{23}}{(\tilde{m}_{\nu}^{-1})_{33}} \right) \frac{W_{31}^*}{\sqrt{m_1 (\tilde{m}_{\nu}^{-1})_{33}}} \\ i \frac{(\tilde{m}_{\nu} W^*)_{12}}{\sqrt{m_2 \tilde{m}_{\nu 11}}} \sqrt{\frac{m_1 m_3 (\tilde{m}_{\nu}^{-1})_{33}}{\tilde{m}_{\nu 11}}} \left(W_{22}^* - W_{32}^* \frac{(\tilde{m}_{\nu}^{-1})_{23}}{(\tilde{m}_{\nu}^{-1})_{33}} \right) \frac{W_{32}^*}{\sqrt{m_2 (\tilde{m}_{\nu}^{-1})_{33}}} \\ i \frac{(\tilde{m}_{\nu} W^*)_{13}}{\sqrt{m_3 \tilde{m}_{\nu 11}}} \sqrt{\frac{m_1 m_2 (\tilde{m}_{\nu}^{-1})_{33}}{\tilde{m}_{\nu 11}}} \left(W_{23}^* - W_{33}^* \frac{(\tilde{m}_{\nu}^{-1})_{23}}{(\tilde{m}_{\nu}^{-1})_{33}} \right) \frac{W_{33}^*}{\sqrt{m_3 (\tilde{m}_{\nu}^{-1})_{33}}} \end{pmatrix}, \quad (23)$$

where we defined $W \equiv V_L U$. The validity of these expressions clearly breaks down in the close vicinity of the two level crossings where either $(\tilde{m}_{\nu}^{-1})_{33}$, or $\tilde{m}_{\nu 11}$, or both, vanish. However, as discussed in [12, 33], the first two cases cannot yield successful leptogenesis. The compact spectrum scenario discussed in the introduction correspond to a situation close to the third case. In all three cases the orthogonal matrix entries diverge, a clear indication of the fine tuning that is needed in the seesaw formula to satisfy neutrino oscillation experimental data.

⁶The different labels denote the three different sets of lepton flavours. While the charged lepton neutrino flavours and the neutrino mass eigenstates form two orthonormal bases, the heavy neutrino lepton flavours cannot in order for the CP asymmetries not to all vanish and have successful leptogenesis.

2.3 $SO(10)$ -inspired leptogenesis

Let us now discuss how the matter-antimatter asymmetry of the universe can be calculated within $SO(10)$ -inspired leptogenesis. This can be expressed in terms of the baryon-to-photon number ratio, whose measured value from *Planck* data (including lensing) combined with external data sets [24], is found

$$\eta_B^{\text{exp}} = (6.13 \pm 0.04) \times 10^{-10}. \quad (24)$$

In general, the final asymmetry is given by the sum of two terms,⁷

$$N_{B-L}^{\text{f}} = N_{B-L}^{\text{p,f}} + N_{B-L}^{\text{lep,f}}. \quad (25)$$

The first term is the relic value of a possible pre-existing asymmetry and the second is the asymmetry generated from the decays of the seesaw neutrinos. The baryon-to-photon number ratio is then also, in general, the sum of two contributions, η_B^{p} and η_B^{lep} , respectively. Typically, one assumes that the initial pre-existing asymmetry, generated after or during inflation and prior to leptogenesis, is negligible. In this way, one finds solutions respecting just successful leptogenesis, i.e., for which $N_{B-L}^{\text{lep,f}}$ reproduces the observed asymmetry.

However, some external mechanism might have generated a large value of the initial pre-existing asymmetry, $N_{B-L}^{\text{p,i}}$, between the end of inflation and the onset of leptogenesis. In the absence of any washout, this would translate into a sizeable value of η_B^{p} comparable or greater than η_B^{exp} . In this case, then one should also add a strong thermal leptogenesis condition, requiring that the initial pre-existing asymmetry is efficiently washed out by seesaw neutrino inverse processes. In Section 4 we discuss how a subset of the solutions satisfying successful SO10INLEP can also simultaneously satisfy this strong thermal leptogenesis condition realising ST-SO10INLEP [30]. We also show how this subset is not disrupted by the inclusion of flavour coupling but just slightly modified. For the time being, we will simply assume that the initial pre-existing asymmetry contribution is absent or in any case negligible.

⁷Strictly speaking one should also include a third term, the asymmetry generated after all RH neutrino have decayed, for example by electroweak baryogenesis. However, such a contribution would be controlled by completely independent new physics so one can only assume that a post-leptogenesis production is absent or in any case negligible. From this point of view, in line with what we wrote in the introduction, the lack of new physics at colliders supports a solution of the matter-antimatter asymmetry puzzle at high scale. Of course, one could consider a situation where two baryogenesis mechanisms both contribute to the final asymmetry but such a situation would be quite fine-tuned and not particularly attractive.

The contribution to η_B from leptogenesis, can be calculated as [46]

$$\eta_B^{\text{lep}} = a_{\text{sph}} \frac{N_{B-L}^{\text{lep,f}}}{N_{\gamma}^{\text{rec}}} \simeq 0.96 \times 10^{-2} N_{B-L}^{\text{lep,f}}. \quad (26)$$

This expression accounts for sphaleron conversion [47, 48, 49, 50] and photon dilution. The second numerical expression holds when the abundances N_X are normalised in a portion of comoving volume such that the ultra-relativistic thermal equilibrium abundance of a RH neutrino $N_{N_i}^{\text{eq}}(T \gg M_i) = 1$. Successful leptogenesis requires $\eta_B^{\text{lep}} = \eta_B^{\text{exp}}$, where η_B^{exp} is the measured value given in Eq. (24).

The final $B - L$ asymmetry from leptogenesis can be calculated, neglecting flavour coupling, as the sum of the three contributions, one from each (charged lepton) flavour, explicitly [19, 51, 36, 52]:

$$\begin{aligned} N_{B-L}^{\text{lep,f}} \simeq & \left[\frac{K_{2e}}{K_{2\tau_2^\perp}} \varepsilon_{2\tau_2^\perp} \kappa(K_{2\tau_2^\perp}) + \left(\varepsilon_{2e} - \frac{K_{2e}}{K_{2\tau_2^\perp}} \varepsilon_{2\tau_2^\perp} \right) \kappa(K_{2\tau_2^\perp}/2) \right] e^{-\frac{3\pi}{8} K_{1e}} + \\ & + \left[\frac{K_{2\mu}}{K_{2\tau_2^\perp}} \varepsilon_{2\tau_2^\perp} \kappa(K_{2\tau_2^\perp}) + \left(\varepsilon_{2\mu} - \frac{K_{2\mu}}{K_{2\tau_2^\perp}} \varepsilon_{2\tau_2^\perp} \right) \kappa(K_{2\tau_2^\perp}/2) \right] e^{-\frac{3\pi}{8} K_{1\mu}} + \\ & + \varepsilon_{2\tau} \kappa(K_{2\tau}) e^{-\frac{3\pi}{8} K_{1\tau}}. \end{aligned} \quad (27)$$

In this expression the (six) $K_{I\alpha}$ are the *flavoured decay parameters*, defined as

$$K_{I\alpha} \equiv \frac{\Gamma_{I\alpha} + \bar{\Gamma}_{I\alpha}}{H(T = M_I)} = \frac{|m_{D\alpha I}|^2}{M_I m_\star}, \quad (28)$$

where $\Gamma_{I\alpha} = \Gamma(N_I \rightarrow \phi^\dagger l_\alpha)$ and $\bar{\Gamma}_{I\alpha} = \Gamma(N_I \rightarrow \phi \bar{l}_\alpha)$ are the zero temperature limit of the flavoured decay rates into α leptons and anti-leptons in the three-flavoured regime, $m_\star \equiv 16 \pi^{5/2} \sqrt{g_\star} / (3\sqrt{5}) (v^2/M_{\text{Pl}}) \simeq 1.08 \text{ meV}$ is the equilibrium neutrino mass, $H(T) = \sqrt{g_\star^{SM}} 8 \pi^3 / 90 T^2 / M_{\text{P}}$ is the expansion rate and $g_\star^{SM} = 106.75$ is the number of ultra-relativistic degrees of freedom in the standard model. The quantities

$$\varepsilon_{2\alpha} \equiv -(\Gamma_{2\alpha} - \bar{\Gamma}_{2\alpha}) / (\Gamma_2 + \bar{\Gamma}_2) \quad (29)$$

are the N_2 -flavoured CP asymmetries, with $\Gamma_2 \equiv \sum_\alpha \Gamma_{2\alpha}$ and $\bar{\Gamma}_2 \equiv \sum_\alpha \bar{\Gamma}_{2\alpha}$. Finally, we defined $K_{2\tau_2^\perp} \equiv K_{2e} + K_{2\mu}$ and $\varepsilon_{2\tau_2^\perp} \equiv \varepsilon_{2e} + \varepsilon_{2\mu}$.

Using the bi-unitary parameterisation Eq. (12) for m_D , the flavoured decay parameters

can be expressed as⁸

$$K_{I\alpha} = \frac{\sum_{k,l} m_{Dk} m_{Dl} V_{Lk\alpha} V_{Ll\alpha}^* U_{RkI}^* U_{RlI}}{M_I m_\star}. \quad (31)$$

The flavoured CP asymmetries can be calculated using [53]

$$\varepsilon_{2\alpha} \simeq \bar{\varepsilon}(M_2) \left\{ \mathcal{I}_{23}^\alpha \xi(M_3^2/M_2^2) + \mathcal{J}_{23}^\alpha \frac{2}{3(1 - M_2^2/M_3^2)} \right\}, \quad (32)$$

where

$$\bar{\varepsilon}(M_2) \equiv \frac{3}{16\pi} \frac{M_2 m_{\text{atm}}}{v^2}, \quad \xi(x) = \frac{2}{3}x \left[(1+x) \ln \left(\frac{1+x}{x} \right) - \frac{2-x}{1-x} \right], \quad (33)$$

$$\mathcal{I}_{23}^\alpha \equiv \frac{\text{Im} \left[m_{D\alpha 2}^* m_{D\alpha 3} (m_D^\dagger m_D)_{23} \right]}{M_2 M_3 \tilde{m}_2 m_{\text{atm}}} \quad \text{and} \quad \mathcal{J}_{23}^\alpha \equiv \frac{\text{Im} \left[m_{D\alpha 2}^* m_{D\alpha 3} (m_D^\dagger m_D)_{32} \right]}{M_2 M_3 \tilde{m}_2 m_{\text{atm}}} \frac{M_2}{M_3}, \quad (34)$$

with $\tilde{m}_2 \equiv (m_D^\dagger m_D)_{22}/M_2$. However, the first term is important only when RH neutrino masses are quasi-degenerate, close to crossing level solutions. In our case such solutions are falling in region of parameter spaces that are now excluded by experiments⁹ and one can reliably use the approximate expression [33]

$$\varepsilon_{2\alpha} \simeq \frac{3}{16\pi} \frac{M_2}{M_3} \frac{\text{Im} \left[m_{D\alpha 2}^* m_{D\alpha 3} (m_D^\dagger m_D)_{23} \right]}{v^2 (m_D^\dagger m_D)_{22}}. \quad (35)$$

Again, using the bi-unitary parameterisation, one obtains:

$$\varepsilon_{2\alpha} \simeq \frac{3}{16\pi v^2} \frac{|(\tilde{m}_\nu)_{11}|}{m_1 m_2 m_3} \frac{\sum_{k,l} m_{Dk} m_{Dl} \text{Im}[V_{Lk\alpha} V_{Ll\alpha}^* U_{Rk2}^* U_{Rl3} U_{R32}^* U_{R33}]}{|(\tilde{m}_\nu^{-1})_{33}|^2 + |(\tilde{m}_\nu^{-1})_{23}|^2}. \quad (36)$$

Notice that in general the asymmetry within minimal leptogenesis with three RH neutrinos would depend on 18 quantities: $\varepsilon_{I\alpha}, K_{I\alpha}$ ($I = 1, 2, 3$ and $\alpha = e, \mu, \tau$). However, from

⁸One can also express the decay parameters in terms of the orthogonal matrix:

$$K_{I\alpha} = \left| \sum_j \sqrt{\frac{m_j}{m_\star}} U_{\alpha j} \Omega_{jI} \right|^2, \quad \text{and} \quad K_I = \sum_I \frac{m_j}{m_\star} |\Omega_{jI}|^2. \quad (30)$$

However, in SO10INLEP, the bi-unitary parameterisation clearly provides the primary way to calculate all leptogenesis parameters, the $K_{I\alpha}$'s and the $\varepsilon_{2\alpha}$'s, since, as one can see from Eqs. (16) and (17), one gets directly RH neutrino masses M_I and RH neutrino mixing matrix U_R in terms of low energy neutrino parameters in m_ν , neutrino Dirac masses m_{Di} and parameters in V_L .

⁹For example, a crossing level solution where $\tilde{m}_{\nu 33}^{-1} = 0$ appears at $m_1 \simeq 0.1$ eV for the range of values of θ_{23} we are considering. In any case, in the code we use the full expression for $\varepsilon_{2\alpha}$, also including the contribution from the self-energy diagram to the CP asymmetries.

Eq. (27) one can see that in SO10INLEP, since it realises a N_2 -leptogenesis scenario, the asymmetry depends only on nine quantities, $\varepsilon_{2\alpha} K_{2\alpha}, K_{1\alpha}$ ($\alpha = e, \mu, \tau$), while there is no dependence on $\varepsilon_{1\alpha}, \varepsilon_{3\alpha}, K_{3\alpha}$ ($\alpha = e, \mu, \tau$). These nine quantities can be expressed in terms of the seesaw parameters. As we have seen, using the bi-unitary parameterisation and the seesaw formula, one obtains Eqs. (16) and (17) for the three Majorana masses and the RH neutrino mixing matrix U_R in terms of the six parameters in V_L , the nine parameters in the low energy neutrino matrix m_ν and the three Dirac neutrino masses m_{Di} . Therefore, in this way, the three RH neutrino masses and the six parameters in U_R are traded off with the 9 low energy neutrino parameters. Moreover, the $\varepsilon_{2\alpha}$'s depend, approximately, only on m_{D2} and the flavoured decay parameters do not depend on the Dirac neutrino masses m_{D1}, m_{D2} and m_{D3} . In this way, ultimately, one obtains for the final baryon-to-photon number ratio an expression of the form

$$\eta_B^{\text{lep}} \simeq \eta_B^{\text{lep}}(m_{\text{sol}}, m_{\text{atm}}, \theta_{12}, \theta_{13}; \theta_{23}, \delta, m_1; \rho, \sigma; V_L, \alpha_2). \quad (37)$$

Notice that we have grouped the parameters in the following way:

- The first four low energy neutrino parameters are those measured accurately and precisely enough that one can assume small Gaussian errors. In the scatter plots, these parameters have been Gaussianly randomly generated. However the errors are so small that just fixing them at the best fit values would practically produce the same results;
- We have then the three neutrino unknowns θ_{23} , δ and m_1 . Here we have been very conservative not to contaminate the predictions from SO10INLEP with the partial experimental information we have. We have uniformly randomly generated the atmospheric neutrino mixing angle in the range $\theta_{23} = [38^\circ, 52^\circ]$ and the neutrino oscillation CP violating phase in the full range $\delta = [-\pi, \pi]$. Finally we have uniformly logarithmically generated the lightest neutrinos mass in the range $\log(m_1/\text{eV}) = [-4, 0]$.
- The Majorana phases have been simply randomly uniformly generated in the full range $[0, 2\pi]$.
- The 6 parameters in V_L have been uniformly randomly generated. The three mixing angles have been capped to the CKM matrix values, following $SO(10)$ -inspired conditions, explicitly we adopted: $\theta_{12}^L \leq 13^\circ \simeq \theta_{12}^{CKM} \equiv \theta_c$, $\theta_{23}^L \leq 2.4^\circ \simeq \theta_{23}^{CKM}$, $\theta_{13}^L \leq 0.2^\circ \simeq \theta_{13}^{CKM}$. In any case the results do not depend on a precise choices of these upper bounds [30]. Finally, for the parameter α_2 , we have used throughout the paper a kind of maximum standard value, $\alpha_2 \leq 5$.

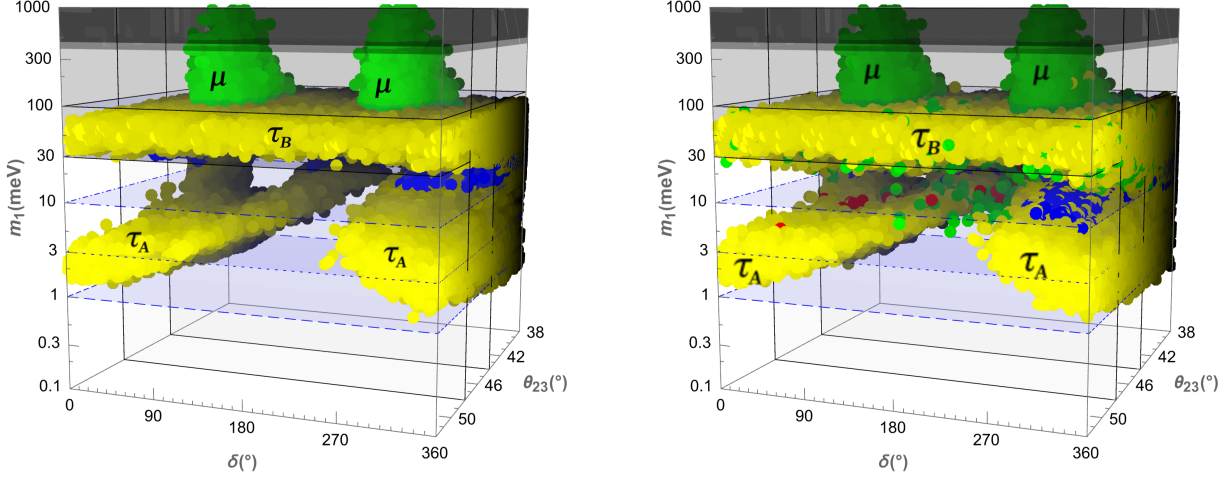


Figure 1: Scatter plot of the solutions obtained imposing successful SO10INLEP neglecting flavour coupling (left panel) and accounting for flavour coupling (right panel). The three grey areas correspond to the excluded regions by the three upper bounds on the absolute neutrino mass scale: Eq. (2) from cosmological observations, Eq. (10) from $0\nu\beta\beta$ and Eq. (11) from tritium beta decay. The three planes in light blue simply help understanding the 3-dim shape. Colour code: tauonic, muonic and strong thermal solutions are, respectively, yellow, green and blue points.

2.4 Successful leptogenesis condition and solutions

The successful leptogenesis condition

$$\eta_B^{\text{lep}}(m_{\text{sol}}, m_{\text{atm}}, \theta_{12}, \theta_{13}; \theta_{23}, \delta, m_1; \rho, \sigma; V_L, \alpha_2) = \eta_B^{\text{exp}} \quad (38)$$

identifies an allowed region in the space of the nine low energy neutrino parameters. In the approximation $V_L = I$ this would correspond to an eight-dimensional hypersurface. Turning on the three angles and three phases in V_L , this hypersurface turns into a layer with some thickness. It is important that it does not fill the whole parameter space and this yields constraints and predictions. In order to visualise this region, we show in Fig. 1 the three-dimensional projection, on the 3-dim space $(\delta, \theta_{23}, m_1)$ of the three unknown neutrino parameters [29], of a scatter plot obtained for $\alpha_2 = 5$ in the 15-dim space of the parameters in U and V_L by generating about 2×10^6 solutions satisfying successful $SO(10)$ -

inspired leptogenesis and all low energy neutrino experimental constraints in Eq. (6).¹⁰ One can immediately notice how solutions are found only for $m_1 \gtrsim 1$ meV, so that there is a clear lower bound on the absolute neutrino mass scale. The three dark grey regions indicate the upper bounds on neutrino masses Eq. (2) from cosmological observations, Eq. (10) from $0\nu\beta\beta$ and Eq. (11) from tritium beta decay.

Each point in the scatter plots is generated imposing the following two conditions:

- $\chi^2(m_{\text{sol}}, m_{\text{atm}}, \theta_{12}, \theta_{13}) \equiv \left(\frac{m_{\text{atm}} - \bar{m}_{\text{atm}}}{\delta m_{\text{atm}}}\right)^2 + \left(\frac{m_{\text{sol}} - \bar{m}_{\text{sol}}}{\delta m_{\text{sol}}}\right)^2 + \left(\frac{\theta_{12} - \bar{\theta}_{12}}{\delta \theta_{12}}\right)^2 + \left(\frac{\theta_{13} - \bar{\theta}_{13}}{\delta \theta_{13}}\right)^2 < \chi_{\text{max}}^2$;
- $\eta_B > \bar{\eta}_B^{\text{exp}} - 3\delta\eta_B^{\text{exp}} = 6.01 \times 10^{-10}$.

We have conservatively used $\chi_{\text{max}}^2 = 25$. However, the obtained allowed regions are not sensitive to a precise value of χ_{max}^2 . For example, a more stringent value $\chi_{\text{max}}^2 = 16$ yields basically the same regions. This is because the errors on the four measured neutrino oscillation parameters are sufficiently small that fluctuations around the mean value produce a very small change of the asymmetry. We have also tried to combine the two conditions including η_B in the calculation of the χ^2 . This just simply slows down the search of the solutions without any change, as expected. The reason is quite simple: for any point where the predicted asymmetry is higher than the experimental value, one can always change the values of the other parameters, those not involved in the projection, to lower the value of the asymmetry. Therefore, these two conditions provide the most efficient (and sensible) way to produce the allowed regions.

2.4.1 Three types of solutions

There are three types of solutions in SO10INLEP [20]. These correspond to three different regions in the parameter space that can be partly distinguished in the three dimensional projection scatter plot in Fig. 1. The different colours indicate the flavour giving the main contribution to the final asymmetry, specifically:

- The yellow points indicate tauon-flavour dominated solutions characterised by $K_{1\tau} \lesssim 1$. One can distinguish two distinct types of tauonic solutions: τ_A and τ_B solutions:
 - The τ_A solutions are characterised by strong washout at the production ($K_{2\tau} \gtrsim 10$). Moreover, one has $K_{1\mu} \gg 1$, so that the muon contribution is many orders of magnitude below the observed value. For current allowed experimental values of θ_{23} as in Eq. (6), they can only be realised for $0.9 \text{ meV} \lesssim m_1 \lesssim 30 \text{ meV}$.
 - The blue points indicate the subset of tauon flavour dominated solutions that

¹⁰The success rate is approximately 0.01% for the parameter ranges we have adopted in the scan.

realise ST-SOINLEP and that we will discuss in Section 5. The τ_A solutions can realise ST-SOINLEP while the τ_B cannot.

- The τ_B solutions are characterised by a mild washout at the production with $2 \lesssim K_{2\tau} \lesssim 10$. In this case one can also have simultaneously $K_{1\mu} \lesssim 1$ and in any case $K_{1\mu} \lesssim 10$. However $K_{2\tau_2^\perp} \gg 10$. The muonic contribution is now only 2-3 orders of magnitudes smaller than the observed value. They can only be realised for $m_1 \simeq (20-60)$ meV.
- The green points indicate solutions with a sizeable muon-flavour component of the asymmetry. The bulk of these solutions is characterised by a dominant muonic contribution. There are some rare solutions where the tauonic contribution can be sizeable though sub-dominant, as large as 10%. In general, we will refer to them as muonic solutions. They are characterised by $K_{1\mu} \lesssim 1$ and $K_{2\tau_2^\perp} \lesssim 10$. Like for τ_B solutions, one can have simultaneously $K_{1\mu}, K_{1\tau} \lesssim 1$ and in any case $K_{1\tau} \lesssim 10$. Therefore, the difference between τ_B and muonic solutions is mainly given by the value of $K_{2\tau_2^\perp}$.

One can notice how the cosmological upper bound on neutrino masses is now excluding almost completely muon-dominated and τ_B solutions. However, the τ_A solutions represent the bulk of the solutions, about 90% of the total.

It is also interesting to notice that for all $SO(10)$ -inspired solutions one has $K_{2e} \ll K_{2\mu} \simeq K_{2\tau_2^\perp}$. This means that the electronic component of the leptons produced from N_2 -decays is negligible. Essentially, all leptons produced from N_2 decays have a flavour composition that lies quite precisely on the muon-tauon plane. The reason can be easily understood analytically. Indeed from Eq. (31) and Eq. (17) one easily derives $K_{2e}/K_{2\mu} \lesssim |V_{L21}|^2 \sim \theta_c^2 \sim 0.05$. One can see from Eq. (27) that this result, combined with the fact that also $\varepsilon_{2e} \ll \varepsilon_{2\mu}$, implies that phantom terms are always negligible.

In Fig. 2 we plot RH neutrino masses, decay parameters and asymmetries versus m_1 for three benchmark points in the parameter space and for $(\alpha_1, \alpha_2, \alpha_3) = (1, 5, 1)$ (the values of α_1 and α_3 only affect M_1 and M_3). For all of them one has $\eta_B^{\text{lep}} \geq \eta_B^{\text{exp}}$ within some range of m_1 values. Therefore, successful SO10INLEP is realised for some special values of m_1 where $\eta_B^{\text{lep}} = \eta_B^{\text{exp}}$. Each of these three benchmark points provides an example of solution belonging to one of the three types: τ_A (left panels), τ_B (central panels) and μ (right panels). At this stage of the discussion, one should focus only on the dashed lines, corresponding to neglect flavour coupling. The values of the different parameters for each solution are given in Table 1. The muonic solution has been selected among those with a sizeable sub-dominant tauonic contribution to the asymmetry. This can be

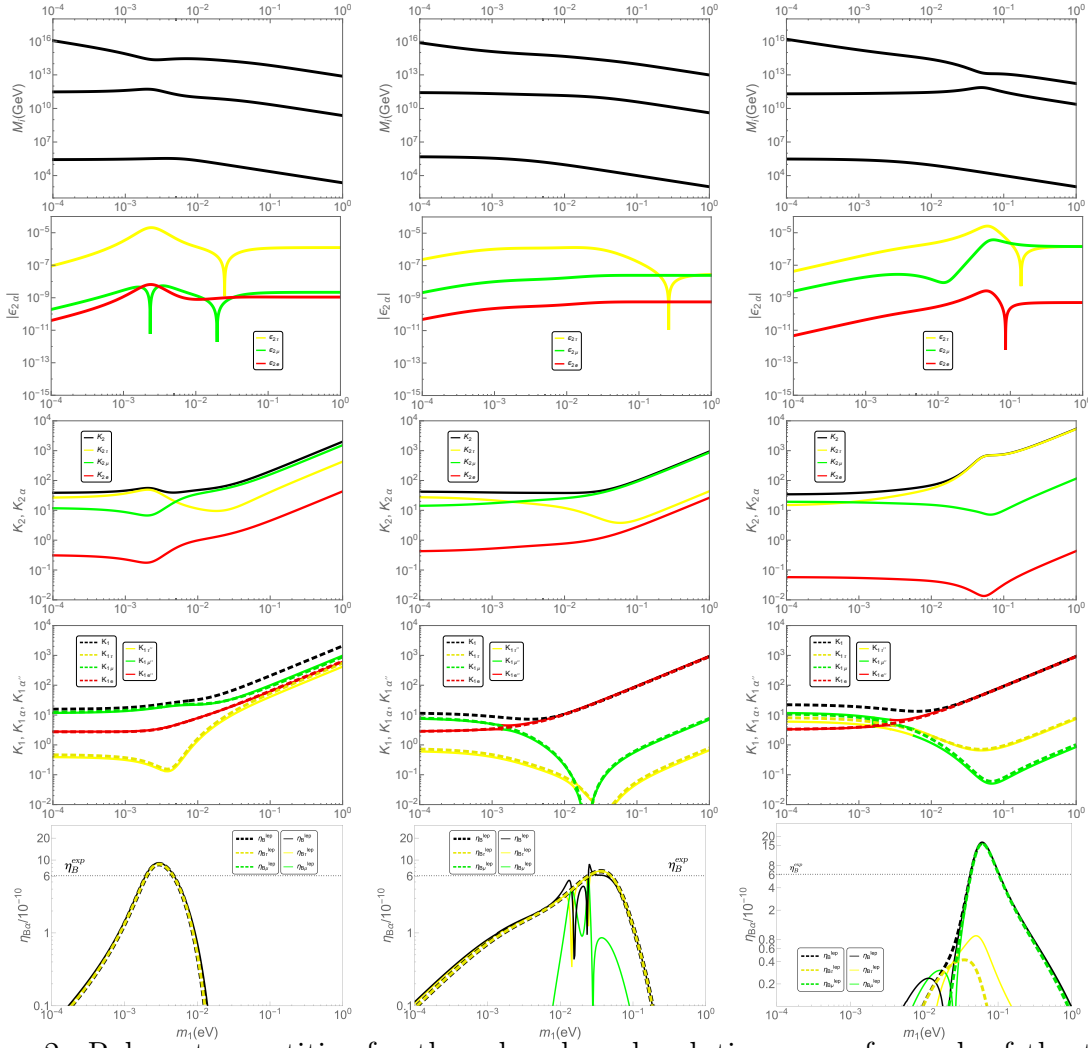


Figure 2: Relevant quantities for three benchmark solutions, one for each of the three types discussed in the main text: τ_A solution (left panels), τ_B solution (central panels), μ solution (right panel). The values of the parameters are shown in Table 1.

clearly seen in the bottom right panel (yellow dashed-line), the tauonic contribution to the total asymmetry is $\sim 6\%$. However, it requires a value $m_1 = 45$ meV in tension with the current cosmological upper bound. Indeed, as we mentioned and one can see from Fig. 1, the current cosmological upper bound Eq. (2) tends to rule out τ_B and μ solutions, so that only the bulk of τ_A solutions is still fully allowed.

2.4.2 Lower bound on m_1

The major feature that can be noticed in the three-dimensional scatter plot in Fig. 1 is the existence of a lower bound on m_1 . It is also clear that this lower bound depends

	θ_{12}	θ_{13}	θ_{23}	δ	ρ/π	σ/π	θ_{12}^L	θ_{13}^L	θ_{23}^L	ρ_L/π	σ_L/π	δ_L/π
τ_A	33.44°	8.61°	43.12°	-62.4°	0.31	0.92	9.46°	0.16°	0.0445°	0.35	0.49	1.59
τ_B	30.9°	8.2°	41.5°	-34.59°	0.952	0.920	9.90°	0.035°	0.86°	0.66	0.99	1.99
μ	33.03°	8.41°	42.81°	103.4°	0.016	0.54	3.23°	0.12°	2.29°	1.01	1.10	1.91

Table 1: Values of the six low energy neutrino parameters in U and six parameters in the unitary matrix V_L for the three benchmark solutions in Fig. 2, as indicated. Best fit values of m_{sol} and m_{atm} are assumed. For each solution, the observed asymmetry is reproduced for two values of m_1 . These three benchmark points are indicated with black stars in all scatter plots, where we choose $m_1 = 2.6 \text{ meV}, 39 \text{ meV}$ and 45 meV for τ_A , τ_B and μ solutions, respectively.

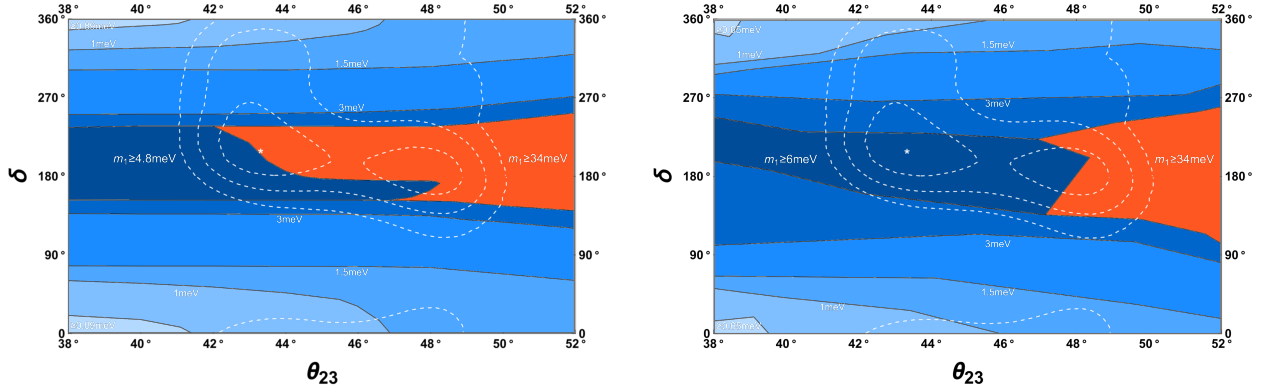


Figure 3: Values of the lower bound on m_1 in the $\theta_{23} - \delta$ plane (isocontour lines) without flavour coupling (left panel) and with flavour coupling. The white dashed lines are the 1σ , 2σ and 3σ experimental constraints from [23]. The orange area is the area excluded by the cosmological upper bound (2). The green area is the portion of the plane that opens up thanks to the muonic solutions from flavour coupling effects.

both on θ_{23} and even more strongly on δ . In the light of an accurate measurement of these two neutrino mixing unknowns at DUNE+T2HK, it is then important to derive the value of the lower bound on m_1 for each point in the plane $\theta_{23}-\delta$. In Fig. 3 we show iso-contour lines of the lower bound on m_1 in the plane δ versus θ_{23} . We also superimpose the two-dimensional experimental constraints from latest global analyses [23]. It can be seen how at 3σ the global lower bound is $m_1 \gtrsim 0.9 \text{ meV}$. An interesting feature is that there is a large region of the plane $\theta_{23}-\delta$, in red colour, that is incompatible with the cosmological upper bound Eq. (1) and that the 1σ experimental regions fall almost entirely just within this region. This clearly shows how SO10INLEP is being strongly tested by low energy neutrino experimental results. Therefore, it is very important to have a clear understanding of theoretical uncertainties to draw firm conclusions from a comparison between theoretical calculations and experimental constraints. As we will see

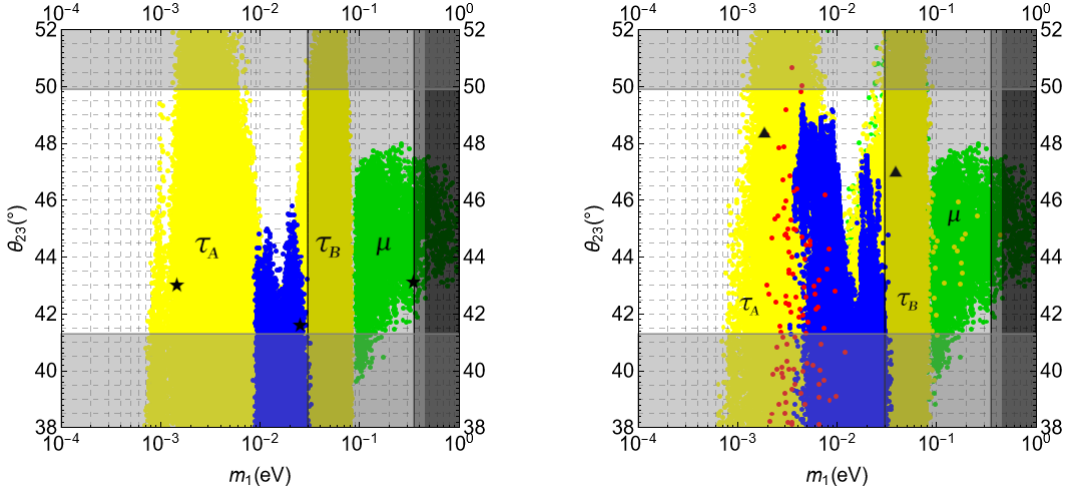


Figure 4: Two-dimensional projection on the plane m_1 - θ_{23} of the scatter plot of the solutions obtained imposing successful SO10INLEP, neglecting flavour coupling effects (left panel) and accounting for flavour coupling effects (right panel).

in the next section, flavour coupling plays an important role in this respect.

2.4.3 Upper bound on θ_{23}

We can further project these solutions from the 3-dim scatter plot on two-dimensional planes, to show better the constraints on the different parameters. In the left panel of Fig 4 we show the projection on the $m_1 - \theta_{23}$ plane. The most interesting feature is that, in the range $m_1 \sim (10-30)$ meV, the second octant is not allowed. This is another important clear test for SO10INLEP and notice that it is independent of the value of δ .

2.4.4 Excluded regions for δ

In the left panel of Fig 5 we also show the projection on the $m_1 - \delta$ plane. One can notice how the lower bound on m_1 is much more stringent in the half $\delta = 180^\circ \pm 90^\circ$ compared to the other half range $\delta \sim 0^\circ \pm 90^\circ$. In particular, the current 1σ experimental range of values for δ falls in the first half that is allowed only for $m_1 \gtrsim 10$ meV. It is also curious how there is a kind of hole for $\delta \sim 30^\circ - 120^\circ$ and $m_1 \sim (5-15)$ meV. This shows again how a future precise determination of δ will represent a crucial test for SO10INLEP.

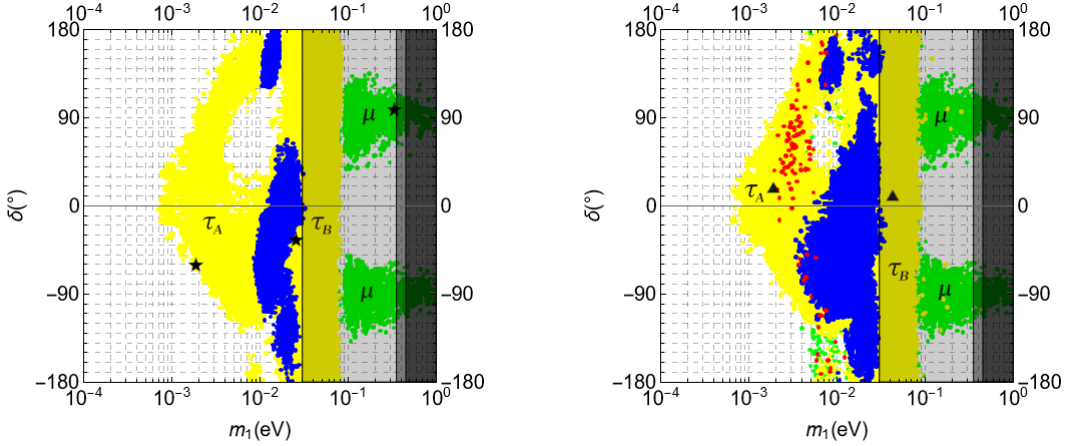


Figure 5: Two-dimensional projection on the plane m_1 - θ_{23} of the scatter plot of the solutions obtained imposing successful SO10INLEP, neglecting flavour coupling effects (left panel) and accounting for flavour coupling effects (right panel). As in Fig. 1, the vertical grey areas denote the excluded regions by the three upper bounds on m_1 : Eq. (2) from cosmological observations, Eq. (10) from $0\nu\beta\beta$ and Eq. (11) from tritium beta decay. The horizontal grey area is the 3σ excluded θ_{23} range of values from Eq. (6).

2.4.5 Majorana phases

In the left panel of Fig 6 we show the projection of the scatter plot solutions on the $\sigma - \rho$ plane. One can notice how large part of the plane is excluded. This is because, as well understood analytically [12, 33], the Majorana phases need to be within certain ranges in order for either $K_{1\tau} \lesssim 1$ or $K_{1\mu} \lesssim 1$ (or both). Notice how in this plane the muon solution allowed regions (green points) are well disconnected by the tauon solution allowed region (yellow points) and lay around the points $(\sigma, \rho) = ((2n + 1)\pi/2, m\pi)$, with $m, n \in \mathbb{Z}$. These are indeed the values that make possible to have $K_{1\mu} \lesssim 1$ [33].

In Fig. 7 we also shown the projection on the plane $m_1 - \sigma$. It can be clearly seen how the values of σ for the muon solutions are different from those for the τ solutions.

2.4.6 $0\nu\beta\beta$ effective neutrino mass

The observation of a $0\nu\beta\beta$ signal would be a crucial discovery supporting, in general, leptogenesis since it would be a clear sign that lepton number is violated at tree level. Specifically, for SO10INLEP, it would provide the measurement of an additional low energy neutrino observable depending on all nine low energy neutrino parameters. This would represent another important experimental testing experimental constraint. In Fig. 8 we show the allowed region in the traditional plane m_1 - m_{ee} . The vertical grey area is the

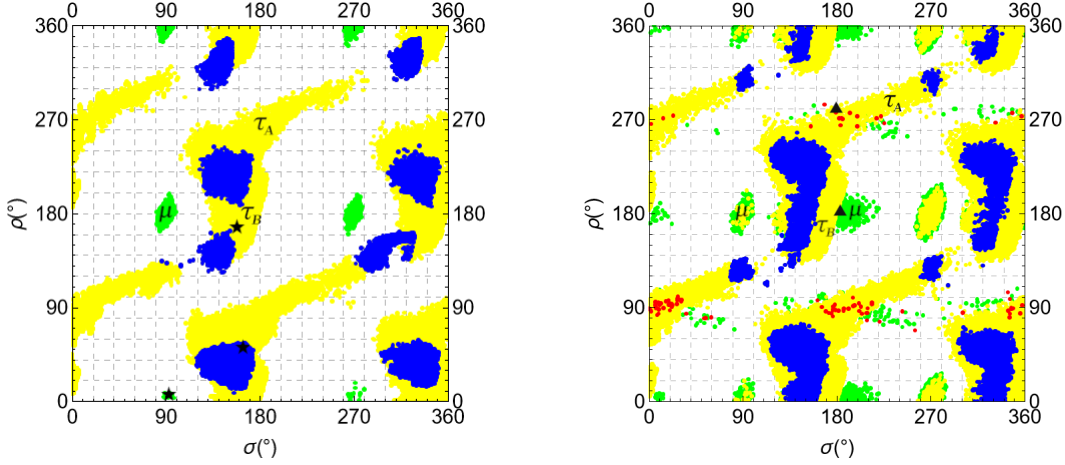


Figure 6: Two-dimensional projection on the plane σ - ρ of the scatter plot of the solutions obtained imposing successful SO10INLEP, neglecting flavour coupling effects (left panel) and accounting for flavour coupling effects (right panel).

conservative excluded region from KamLAND-Zen upper bound given in Eq. (10), while the dashed line indicates the most stringent upper bound also given in Eq. (10). The thin black lines indicate the general allowed region compatible with the experimental constraints (6) (i.e., no successful SO10INLEP condition is imposed). One can see how SO10INLEP strongly restricts this region. Therefore, the simultaneous experimental determination of m_1 and m_{ee} would provide a very strong test of SO10INLEP.

3 Including flavour coupling

The expression in Eq. (27) for the final $B-L$ asymmetry assumes that the evolution of the three flavoured asymmetries $\Delta_\alpha \equiv B/3 - L_\alpha$ ($\alpha = e, \mu, \tau$), where the L_α 's are the lepton asymmetries in lepton doublets, proceeds independently of each other, in an uncoupled way. Sphaleron process conserve the Δ_α 's and the total $B-L$ asymmetry is the sum of the three. In this way one can solve the set of uncoupled differential equations for the Δ_α 's and at the end $N_{B-L}^f = \sum_\alpha N_{\Delta_\alpha}^f$ and the result is the expression in Eq. (27) [19, 51, 36, 52]. The assumption of uncoupled evolution of the flavoured asymmetries relies on a simple picture where: N_2 -decays source initially the asymmetry in the flavour direction ℓ_2 ; decoherence of leptons states results into an independent washout from N_2 -inverse processes for the two flavoured asymmetries Δ_τ and $\Delta_e + \Delta_\mu$ at $T \sim M_2$ (two flavour regime) and into an independent washout from N_1 -inverse processes for all three Δ_α at $T \sim M_1$; finally, it assumes that the washout is just depending on each Δ_α separately.

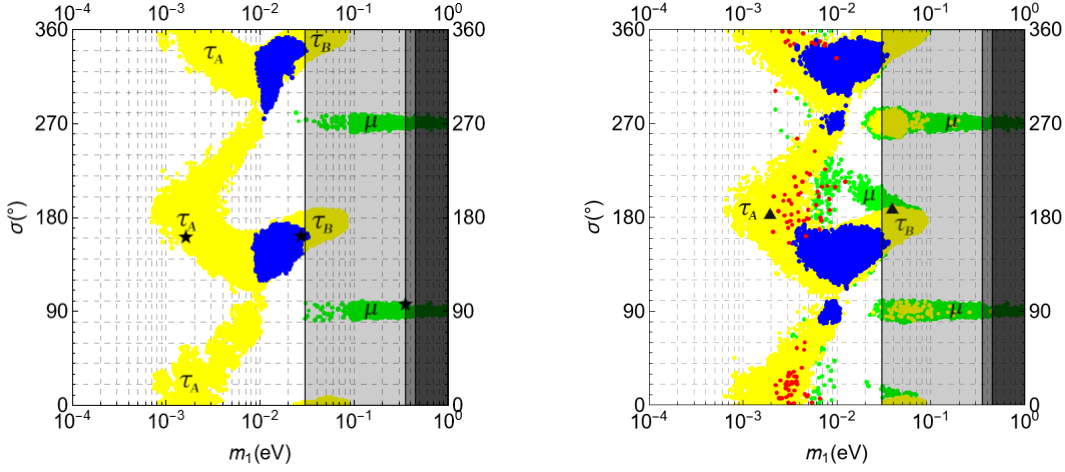


Figure 7: Two-dimensional projection on the plane m_1 - θ_{23} of the scatter plot of the solutions obtained imposing successful SO10INLEP, neglecting flavour coupling effects (left panel) and accounting for flavour coupling effects (right panel).

Within this simple picture, though decoherence splits the evolution of the asymmetry first in two and then in three flavoured asymmetries, these still develop independently of each other because the washout in a certain flavour is always depending, linearly in very good approximation, just on the asymmetry in that flavour. However, this picture is incomplete because it neglects spectator processes of different nature [34, 35, 17, 51]: (i) the generation of an asymmetry from N_2 -decays into lepton doublets is also accompanied by the generation of a hypercharge asymmetry into the Higgs bosons and (ii) by a baryonic asymmetry into squarks via sphaleron processes; (iii) the lepton asymmetry from lepton doublets is also redistributed to RH charged particles. In this way the washout of a Δ_α asymmetry also depends on the asymmetries stored into the Higgs doublets (this is the primary source of flavour coupling) and into quarks (this a secondary source of flavour coupling). To be more clear, since the Higgs doublet asymmetry is flavour blind, if a Δ_α asymmetry has been generated in a certain flavour, necessarily at the washout the asymmetry into the Higgs doublets will induce the generation of a $\Delta_{\beta \neq \alpha}$ asymmetry in the other flavour(s) β : therefore, inverse processes do in this case generate a flavoured asymmetry rather than wash it out. This means that the washout strictly speaking refers to the Higgs asymmetry and the N_2 -decay produced Δ_α asymmetry but not to $\Delta_{\beta \neq \alpha}$. This is the primary source of flavour coupling. Secondly, the redistribution of the asymmetries into quarks and RH charged particles also couples the Δ_α asymmetries.

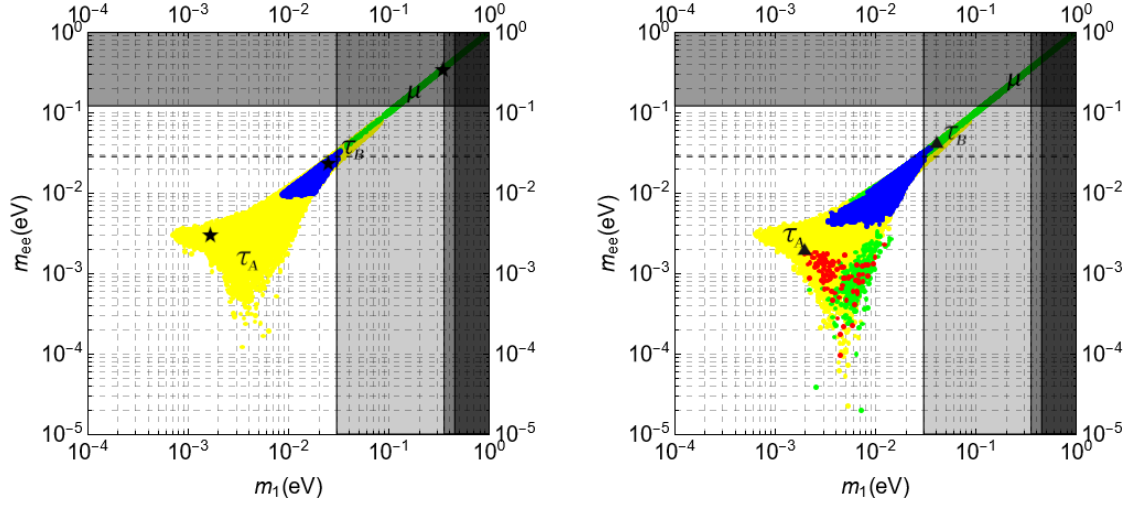


Figure 8: Projection in the plane $m_1 - m_{ee}$ of the scatter plot of the solutions obtained imposing successful SO10INLEP, neglecting flavour coupling effects (left panel) and accounting for flavour coupling effects (right panel).

3.1 A more tangled expression for the final asymmetry

Solving the system of coupled differential equations at the production, for $T \sim M_2$, and at the N_1 -washout, for $T \sim M_1$, leads to the following analytical expression for each final Δ_α asymmetry [36]:

$$N_{\Delta_\alpha}^f = \sum_{\alpha''} V_{\alpha\alpha''}^{-1} \left[N_{\alpha''}^{T \sim T_L} e^{-\frac{3\pi}{8} K_{1\alpha''}} \right], \quad (39)$$

where the three $N_{\alpha''}^{T \sim T_L}$ ($\alpha'' = e'', \mu'', \tau''$) are the asymmetries produced by N_2 -decays at $T = T_L \sim M_2$ in the (rotated) flavours α'' and are given by

$$N_{\alpha''}^{T \sim T_L} = \sum_{\beta=e,\mu,\tau} V_{\alpha''\beta} N_{\Delta_\beta}^{T \sim T_L}. \quad (40)$$

Before explaining in detail the meaning and definition of all involved quantities in this general expression, notice that now the final $B - L$ asymmetry, though still given by $N_{B-L}^f = \sum_{\alpha=e,\mu,\tau} N_{\Delta_\alpha}^f$, is the sum of 27 terms instead of just 3. The matrix

$$V \equiv \begin{pmatrix} V_{e''e} & V_{e''\mu} & V_{e''\tau} \\ V_{\mu''e} & V_{\mu''\mu} & V_{\mu''\tau} \\ V_{\tau''e} & V_{\tau''\mu} & V_{\tau''\tau} \end{pmatrix} \quad (41)$$

is the matrix that diagonalizes P_1^0 , i.e. $V P_1^0 V^{-1} = P_{1''}^0 \equiv \text{diag}(P_{1e''}^0, P_{1\mu''}^0, P_{1\tau''}^0)$, where the matrix P_1^0 is defined as

$$P_1^0 \equiv \begin{pmatrix} P_{1e}^0 C_{ee}^{(3)} & P_{1e}^0 C_{e\mu}^{(3)} & P_{1e}^0 C_{e\tau}^{(3)} \\ P_{1\mu}^0 C_{\mu e}^{(3)} & P_{1\mu}^0 C_{\mu\mu}^{(3)} & P_{1\mu}^0 C_{\mu\tau}^{(3)} \\ P_{1\tau}^0 C_{\tau e}^{(3)} & P_{1\tau}^0 C_{\tau\mu}^{(3)} & P_{1\tau}^0 C_{\tau\tau}^{(3)} \end{pmatrix}, \quad (42)$$

and the three-flavour coupling matrix is given by [54]

$$C^{(3)} \equiv \begin{pmatrix} C_{ee}^{(3)} & C_{e\mu}^{(3)} & C_{e\tau}^{(3)} \\ C_{\mu e}^{(3)} & C_{\mu\mu}^{(3)} & C_{\mu\tau}^{(3)} \\ C_{\tau e}^{(3)} & C_{\tau\mu}^{(3)} & C_{\tau\tau}^{(3)} \end{pmatrix} = \begin{pmatrix} 188/179 & 32/179 & 32/179 \\ 49/358 & 500/537 & 142/537 \\ 49/358 & 142/537 & 500/537 \end{pmatrix}, \quad (43)$$

and we defined $P_{1\alpha}^0 \equiv K_{1\alpha}/K_1$. The rotated flavoured decay parameters are defined as $K_{1\alpha''} \equiv P_{1\alpha''}^0 K_1$.

The three flavour asymmetries $N_{\Delta_\beta}^{T \sim T_L}$ in the standard flavours $\beta = e, \mu, \tau$, at the production, are given by

$$\begin{aligned} N_{\Delta_\tau}^{T \sim T_L} &= U_{\tau\tau_2^\perp}^{-1} \left[U_{\tau_2^\perp \tau_2^\perp} \varepsilon_{2\tau_2^\perp} + U_{\tau_2^\perp \tau} \varepsilon_{2\tau} \right] \kappa(K_{2\tau_2^\perp}) + U_{\tau\tau'}^{-1} \left[U_{\tau'\tau_2^\perp} \varepsilon_{2\tau_2^\perp} + U_{\tau'\tau} \varepsilon_{2\tau} \right] \kappa(K_{2\tau}), \\ N_{\Delta_e}^{T \sim T_L} &= \left[\frac{K_{2e}}{K_{2\tau_2^\perp}} N_{\Delta_{\tau_2^\perp}}^{T \sim T_L} + \left(\varepsilon_{2e} - \frac{K_{2e}}{K_{2\tau_2^\perp}} \varepsilon_{2\tau_2^\perp} \right) \kappa(K_{2\tau_2^\perp}/2) \right], \\ N_{\Delta_\mu}^{T \sim T_L} &= \left[\frac{K_{2\mu}}{K_{2\tau_2^\perp}} N_{\Delta_{\tau_2^\perp}}^{T \sim T_L} + \left(\varepsilon_{2\mu} - \frac{K_{2\mu}}{K_{2\tau_2^\perp}} \varepsilon_{2\tau_2^\perp} \right) \kappa(K_{2\tau_2^\perp}/2) \right], \end{aligned} \quad (44)$$

where

$$N_{\Delta_{\tau_2^\perp}}^{T \sim T_L} = U_{\tau_2^\perp \tau_2^\perp}^{-1} \left[U_{\tau_2^\perp \tau_2^\perp} \varepsilon_{2\tau_2^\perp} + U_{\tau_2^\perp \tau} \varepsilon_{2\tau} \right] \kappa(K_{2\tau_2^\perp}) + U_{\tau_2^\perp \tau'}^{-1} \left[U_{\tau'\tau_2^\perp} \varepsilon_{2\tau_2^\perp} + U_{\tau'\tau} \varepsilon_{2\tau} \right] \kappa(K_{2\tau}).$$

The matrix

$$U \equiv \begin{pmatrix} U_{\tau_2^\perp \tau_2^\perp} & U_{\tau_2^\perp \tau} \\ U_{\tau' \tau_2^\perp} & U_{\tau' \tau} \end{pmatrix} \quad (45)$$

is the analogous of V , accounting for flavour coupling in the two-flavour regime. It is defined as the matrix that diagonalises

$$P_2^0 \equiv \begin{pmatrix} P_{2\tau^\perp}^0 C_{\tau_2^\perp \tau_2^\perp}^{(2)} & P_{2\tau^\perp}^0 C_{\tau_2^\perp \tau}^{(2)} \\ P_{2\tau}^0 C_{\tau \tau_2^\perp}^{(2)} & P_{2\tau}^0 C_{\tau \tau}^{(2)} \end{pmatrix}, \quad (46)$$

i.e. $U P_2^0 U^{-1} = \text{diag}(P_{2\tau^\perp}^0, P_{2\tau}^0)$, where the two-flavour coupling matrix is given by [54]

$$C^{(2)} \equiv \begin{pmatrix} C_{\tau_2^\perp \tau_2^\perp}^{(2)} & C_{\tau_2^\perp \tau}^{(2)} \\ C_{\tau \tau_2^\perp}^{(2)} & C_{\tau \tau}^{(2)} \end{pmatrix} = \begin{pmatrix} 581/589 & 104/589 \\ 194/589 & 614/589 \end{pmatrix} \quad (47)$$

and we defined $P_{2\alpha}^0 \equiv K_{2\alpha}/K_2$. We have now fully defined all quantities entering the expression for the flavoured asymmetries in Eq. (39). As we said, the final $B - L$ asymmetry is still obtained as the sum of the three flavoured asymmetries and it can be checked that, for $U = V = I$, one recovers the expression Eq. (27) valid when flavour coupling is neglected. We can also unpack the first sum in Eq. (39) and write explicitly:

$$\begin{aligned}
N_{\Delta_\alpha}^f &= V_{\alpha e''}^{-1} \left[\sum_{\beta} V_{e''\beta} N_{\Delta_\beta}^{T \sim T_L} \right] e^{-\frac{3\pi}{8} K_{1e''}} \\
&+ V_{\alpha \mu''}^{-1} \left[\sum_{\beta} V_{\mu''\beta} N_{\Delta_\beta}^{T \sim T_L} \right] e^{-\frac{3\pi}{8} K_{1\mu''}} \\
&+ V_{\alpha \tau''}^{-1} \left[\sum_{\beta} V_{\tau''\beta} N_{\Delta_\beta}^{T \sim T_L} \right] e^{-\frac{3\pi}{8} K_{1\tau''}} .
\end{aligned} \tag{48}$$

This expression shows how now each flavoured asymmetry $N_{\Delta_\alpha}^f$ is not simply given by one term containing a N_1 washout exponential suppression term described by $e^{-3\pi K_{1\alpha}/8}$ but it also contains terms that are washed out by exponentials $e^{-3\pi K_{1\delta''}/8}$ with $\delta \neq \alpha$. In this way, even though $K_{1\alpha} \gg 1$, there can still be unsuppressed contributions to $N_{\Delta_\alpha}^f$ from terms with $K_{1\delta} \ll 1$. Even though these terms are weighted by factors $V_{\alpha\delta}^{-1}$, containing off-diagonal terms of the $C^{(3)}$ matrix that are $\mathcal{O}(0.1)$, they can be dominant in some cases and therefore, in general, they have to be accounted for.

3.2 Successful leptogenesis solutions

Despite the much higher intricacy of the expression Eq. (39), accounting for flavour coupling, compared to Eq. (27), neglecting flavour coupling, the discussion we made in 2.3 on the parameter dependence and the consequent successful leptogenesis condition Eq. (38) still holds. Flavour coupling does not introduce any new parameter. It only introduces many new terms in the expression for the final asymmetry that are proportional to the off-diagonal terms in the flavour coupling matrix. These are usually sub-dominant and give a small correction but for particular values of the parameters they can yield new solutions in regions of parameter space that would otherwise be inaccessible when flavour coupling is neglected.

The search of these solutions is the main objective of including flavour coupling. The result is shown in the right panel of Fig. 1. Here we show again the three-dimensional projection of a scatter plot of a about 2×10^6 solutions obtained for $\alpha_2 = 5$ and scanning over the 15 parameters in U and V_L , respecting the experimental constraints in Eq. (6).

	θ_{12}	θ_{13}	θ_{23}	δ	ρ/π	σ/π	θ_{12}^L	θ_{13}^L	θ_{23}^L	ρ_L/π	σ_L/π	δ_L/π
$\tau \rightarrow \mu$	35.18°	9.11°	46.73°	9.97°	1.033	1.064	2.172°	0.109°	0.842°	1.05	1.13	0.23
$\tau \rightarrow e$	30.04°	8.325°	48.24°	16.56°	1.47	1.021	4.859°	0.155°	1.238°	1.72	1.25	0.31

Table 2: Values of the six low energy neutrino parameters in U and six parameters in the unitary matrix V_L for the two benchmark solutions shown in Fig. 9. Best fit values of m_{sol} and m_{atm} are assumed. For each solution, the observed asymmetry is reproduced for two values of m_1 . For $m_1 = 42 \text{ meV}$ (left) and $m_1 = 2 \text{ meV}$ (right) one has a muon-dominated and electron-dominated solution, respectively. These two solution are indicated with triangles in the scatter plots.

The solutions have been again obtained imposing $\chi^2 < \chi_{\text{max}}^2 = 25$ and $\eta_B > 6.01 \times 10^{-10}$. The first thing one can notice is that the main structure of the allowed region is not changed by flavour coupling and, in particular, one still has a lower bound on m_1 . On the other hand, it can be noticed how there are new muonic solution appearing in a region that was before excluded, at values of $m_1 \sim 10 \text{ meV}$. Moreover, electronic solutions, indicated by red points, now appear in the region of τ_A solutions.

3.2.1 Two new types of solutions

Let us have a closer look at these new solutions. As we did in the last section for the three types of old solutions, in Fig. 9 we plot different quantities ($M_I, \varepsilon_\alpha, K_{1\alpha}, K_{2\tau}, K_{2\tau_2^\perp}$) versus m_1 for two benchmark new solutions, one muonic and one electronic, with values of the parameters given in Table 2. As one can see, successful leptogenesis is satisfied for $m_1 = 42 \text{ meV}$ and $m_1 = 2.1 \text{ meV}$ for the muon-dominated and the electron-dominated solution, respectively. One can also see how the dominant CP asymmetry is the tauonic CP asymmetry while both the muonic and the electronic ones would be too small, in the absence of flavour coupling, to yield the observed baryon asymmetry. Therefore, the existence of these two new types of solutions crucially relies on the account of flavour coupling (without flavour coupling the muon and electron asymmetries, respectively, would be orders of magnitude below the observed value). We should add, however, that these new electronic solutions are very marginal ones. This can be somehow understood by the fact that the peak of the asymmetry in the shown benchmark case is just above the observed value. In the 2×10^6 point scatter plot we generated, we obtained only less than 200 electronic solution (i.e., less than 0.1%). The new muonic solutions are less special and we obtained about 2k of them (i.e., $\sim 1\%$).

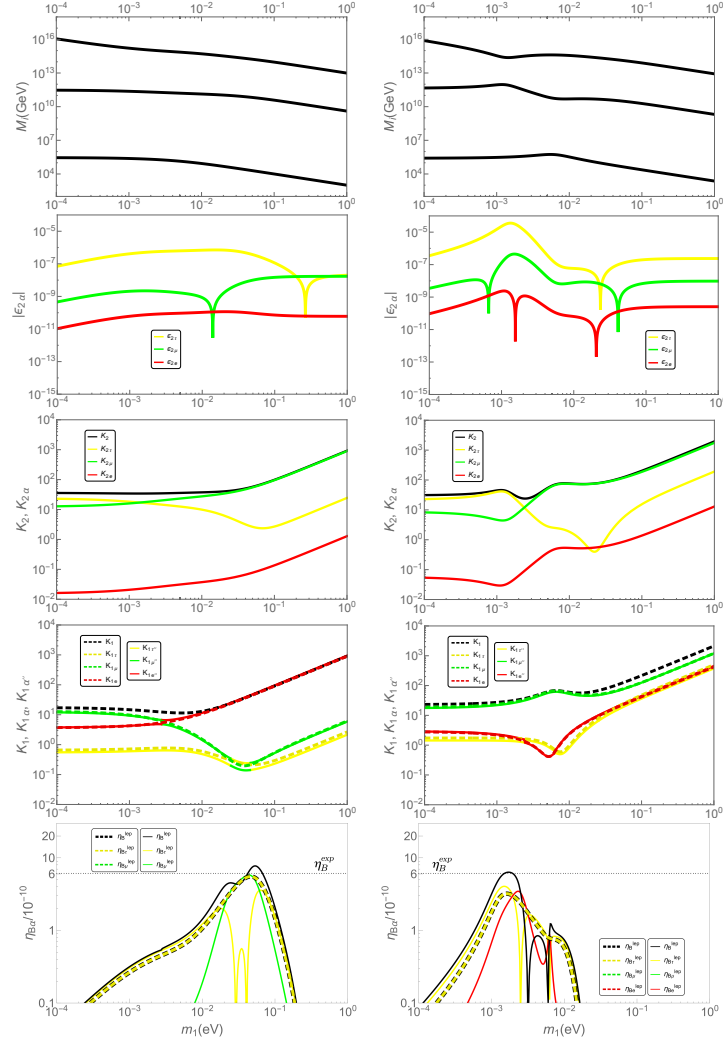


Figure 9: Relevant quantities for two benchmark new type of solutions: one new μ -solution (left panels), one new e -solution (right panels). The values of the parameters for these solutions are shown in Table 2.

3.2.2 Lower bound on m_1

In the right panel of Fig. 3 we again show the isocontour lines for the m_1 lower bound in the δ versus θ_{23} plane. One can see that they are just slightly modified by the inclusion of flavour coupling and, for example, the absolute minimum value gets slightly relaxed from $m_1 \geq 0.89 \text{ meV}$ to $m_1 \geq 65 \text{ meV}$. The most remarkable difference is that now the excluded region by the cosmological upper bound Eq. (2), in red in Fig. 3, gets reduced and first octant values of θ_{23} are now possible also for values $\delta \sim \pi$. For these values of δ , one has that values of the atmospheric mixing angle $\theta_{23} \gtrsim 47^\circ$ are still excluded even when flavour coupling is included. In particular, one can see how the best fit point falls

now in the green region that is the one opened by the muonic solutions induced by the account of flavour coupling.

3.2.3 Upper bound on θ_{23}

In the right panel of Fig. 4, we show how flavour coupling affects the upper bound on θ_{23} versus m_1 . One can see that the new muonic solutions tend to cover, very marginally, the region for $m_1 \simeq 15$ meV and $\theta_{23} \simeq 44^\circ$ – 45° that would be excluded without an account of flavour coupling. However, all second octant values of θ_{23} still remain excluded for $m_1 = (10\text{--}30)$ meV.

3.2.4 Excluded regions for δ

In the right panel of Fig. 5 we show the constraints in the plane δ versus m_1 . One can see how the new muonic solutions tend to fill some existing gaps for $m_1 \simeq 10$ meV and δ in the range between $\pi/2$ and $3\pi/2$. However, the hole at $\delta \simeq \pi/2$, though shrinking, still survives. For m_1 much below 10 meV the constraints are practically unchanged.

3.2.5 Majorana phases

In the right panel of Fig. 6 we show the constraints in the plane ρ versus σ . This plot is quite interesting because it clearly shows how the new solutions, both muonic and electronic, stem from the tauonic CP asymmetry. The regions they fill are completely disconnected from the primary muonic solutions that are present also in the absence of flavour coupling and for which the asymmetry is proportional to the muon CP asymmetry. One can also notice how the new solutions extend to higher values of σ for $\rho \sim \pi$. The same conclusions can be also drawn looking at the right panel of Fig. 7 where we show the constraints in the plane σ versus m_1 .

3.2.6 $0\nu\beta\beta$ effective neutrino mass

Finally, one can see from the right panel of Fig. 8 that flavour coupling does not really change the constraints in the plane m_{ee} versus m_1 . The new solutions are confined within the region that was anyway accessible also without accounting for flavour coupling.

4 Strong thermal $SO(10)$ -inspired leptogenesis

In this section, we finally discuss ST-SO10INLEP [30]. It is intriguing that the rather special conditions that are needed to realise strong thermal leptogenesis [31, 32], are naturally satisfied by a subset of the solutions realising successful SO10INLEP. In this Section we first re-derive, and update, the predictions on low energy neutrino parameters from ST-SO10INLEP neglecting flavour coupling, and then we discuss the impact of flavour coupling.

4.1 Neglecting flavour coupling

Let us assume that some external mechanism has created an initial pre-existing B-L asymmetry $N_{B-L}^{p,i}$ prior to leptogenesis and prior to any washout from RH neutrinos. We also assume that the N_3 mass is larger than the reheat temperature, so that the N_3 -washout is absent.¹¹ Because of the RH neutrino inverse processes, the pre-existing asymmetry will undergo a dynamical evolution that has to take into account also the different flavour regimes. Finally, the relic value of the pre-existing asymmetry can be also expressed as the sum of three contributions from each flavour i.e., $N_{B-L}^{p,f} = \sum_{\alpha} N_{\Delta_{\alpha}}^{p,f}$. Neglecting flavour coupling, the expressions for the $N_{\Delta_{\alpha}}^{p,f}$ are given by [30, 40]

$$\begin{aligned} N_{\Delta_{\tau}}^{p,f} &= (p_{p\tau}^0 + \Delta p_{p\tau}) e^{-\frac{3\pi}{8}(K_{1\tau}+K_{2\tau})} N_{B-L}^{p,i}, \\ N_{\Delta_{\mu}}^{p,f} &= \left\{ (1 - p_{p\tau}^0) \left[p_{\mu\tau_2^{\perp}}^0 p_{p\tau_2^{\perp}}^0 e^{-\frac{3\pi}{8}(K_{2e}+K_{2\mu})} + (1 - p_{\mu\tau_2^{\perp}}^0) (1 - p_{p\tau_2^{\perp}}^0) \right] + \Delta p_{p\mu} \right\} e^{-\frac{3\pi}{8}K_{1\mu}} N_{B-L}^{p,i}, \\ N_{\Delta_e}^{p,f} &= \left\{ (1 - p_{p\tau}^0) \left[p_{e\tau_2^{\perp}}^0 p_{p\tau_2^{\perp}}^0 e^{-\frac{3\pi}{8}(K_{2e}+K_{2\mu})} + (1 - p_{e\tau_2^{\perp}}^0) (1 - p_{p\tau_2^{\perp}}^0) \right] + \Delta p_{pe} \right\} e^{-\frac{3\pi}{8}K_{1e}} N_{B-L}^{p,i}. \end{aligned} \quad (49)$$

In this expression the quantities $p_{p\tau}^0$ and $p_{p\tau_2^{\perp}}^0$ denote the fractions of the initial pre-existing asymmetry in the tauon flavour and in the flavour $\tau_2^{\perp}$, where, as already defined, $\tau_2^{\perp}$ is the electron and muon flavour superposition component in the leptons produced by N_2 -decays (or equivalently the flavour component that is washed-out in the inverse processes producing N_2).¹² We also introduced the two quantities $p_{\alpha\tau_2^{\perp}}^0 \equiv K_{2\alpha}/(K_{2e} + K_{2\mu})$ ($\alpha = e, \mu$). These give the fractions of the α flavour in the $\tau_2^{\perp}$ component of the lepton state $|\ell_2\rangle$ produced by N_2 -decays. The terms $\Delta p_{p\tau}, \Delta p_{pe}$ and $\Delta p_{p\mu}$ are the so-called phantom terms and such that $\Delta p_{pe} + \Delta p_{p\mu} + \Delta p_{p\tau} = 0$. They are phantom terms taking into account the possibility that the mechanism that created the pre-existing asymmetry produced

¹¹This assumption is conservative but ultimately it is irrelevant whether N_3 -washout is included or not, the predictions on the low energy neutrino parameters change just slightly [30].

¹²Notice that in general $p_{p\tau}^0 + p_{p\tau_2^{\perp}}^0 \neq 1$ since in general the projection of the pre-existing asymmetry flavour direction on the $\tau^{\perp}$ plane, is not parallel to $\tau_2^{\perp}$.

lepton states that are not exactly the CP conjugated of the anti-lepton states. It can be crossed-checked that in Eq. (49), if one switches-off the washout terms, then the sum of the three final flavoured pre-existing asymmetries is equal to the total initial pre-existing asymmetry, as it should be.

The strong thermal leptogenesis condition requires that the relic value of the pre-existing asymmetry is negligible with respect to the final value of the asymmetry produced by N_2 -decays. For definiteness, we will impose a condition $N_{B-L}^{\text{p,f}} < 0.1 N_{B-L}^{\text{lep,f}}$.

The only possibility to have successful strong thermal leptogenesis is just the N_2 tauon-dominated leptogenesis scenario [31] that is also representing the bulk of solutions in SO10INLEP, as we discussed. This is quite an intriguing coincidence. The reason why only tauon-dominated solutions can realise strong thermal leptogenesis can be understood directly from Eq. (49). One can indeed have in this case $K_{2\tau}, K_{1\mu}, K_{1e} \gg 1$ and $K_{1\tau} \lesssim 1$ in a way to wash out at $T \sim M_2$ the tauonic component of the pre-existing asymmetry and at $T \sim M_1$ also the electronic and muonic components. At the same time, having $K_{1\tau} \lesssim 1$, the tauonic asymmetry produced from leptogenesis is the only that survives the N_1 -washout.

There is only a very special caveat [40]. If one imposes $(1 - p_{\mu\tau_2}^0) = K_{2e}/(K_{2e} + K_{2\mu}) \lesssim 0.1 N_{B-L}^{\text{lep,f}}/N_{B-L}^{\text{p,i}} \lll 1$ and $K_{1\mu} \lesssim 1$, instead of $K_{1\mu} \gg 1$, one can have a strong thermal muon-dominated solution. We noticed that $K_{2e}/(K_{2e} + K_{2\mu}) \sim \theta_c^2 \sim 0.05$ is indeed small but in this case, for a meaningfully large value of pre-existing asymmetry $\sim 10^{-3}$, as we will assume, one needs this ratio at least four orders of magnitude smaller. These are incredibly special solutions, unless they can be justified within a particular model. In [40] these solutions were found with a tiny $\sim 10^{-9}$ rate. They have to be considered then more like oddities rather than real ST-SO10INLEP solutions.

By imposing a strong thermal leptogenesis condition in addition to successful SO10INLEP, one singles out a special subset of solutions within those we obtained just imposing successful SO10INLEP. We have assumed as benchmark value $N_{B-L}^{\text{p,i}} = 10^{-3}$. In the scatter plots, in Fig. 1 and Figs. 3–9, these solutions are denoted by blue points. The success rate for these solutions in our scan is 10^{-6} . The results fully confirm previous results [30, 12, 33, 40]. It is still useful to notice some important aspects:

- As one can see they are all tauon-dominated solutions obtained with a 10^{-6} success rate, except for a couple of special muon-dominated solutions (success rate 10^{-9}) we mentioned.
- From Fig. 6 one can see how one the requested values of the Majorana phases for ST-SO10INLEP become so constrained that ST-SO10INLEP would be a way to

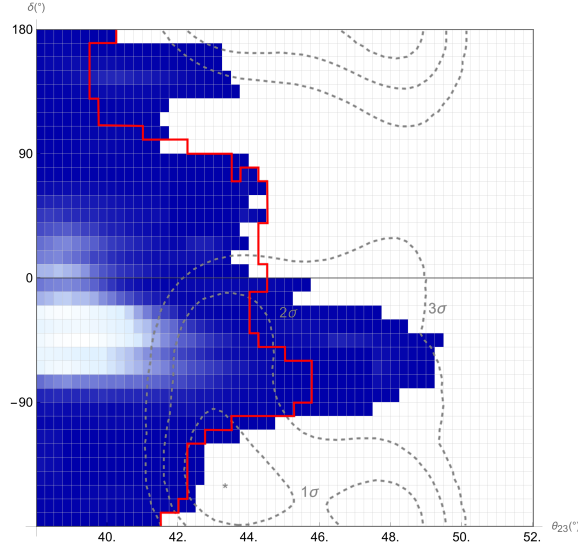


Figure 10: ST-SO10INLEP allowed region in the plane δ versus θ_{23} . The red line delimits the region when flavour coupling is neglected, the coloured region is the result in the case when flavour coupling is included. The colour code denotes point density: lighter blue means higher density of solutions. The thin dashed lines are the 1σ , 2σ and 3σ experimental constraints from [23].

determine their values.

- The absolute neutrino mass scale range is very restricted. For the benchmark value we assumed, $N_{B-L}^{\text{p,i}} = 10^{-3}$, one has that successful ST-SO10INLEP requires $m_1 = (10\text{--}30)$ meV and, from Fig. 8, one can see a corresponding similar range for m_{ee} . It is interesting that both cosmological observations and $0\nu\beta\beta$ experiments are now starting to test this range.
- The bulk of solutions realising ST-SO10INLEP are of τ_A type. The τ_B type solutions have the problem that $K_{1\mu}$ tends to be too small to washout the muonic component of the pre-existing asymmetry. For example for $V_L = I$ one finds $K_{1\mu} \lesssim 4$ [12] and in this case they would be absolutely unable to yield strong thermal solutions. However, turning on angles in V_L values as large as $K_{1\mu} \sim 10$ are possible and for this reason some τ_B strong thermal solutions do exist, though they are only $\sim 1\%$ of the τ_B solutions.

Finally, we have focused on an accurate determination of the allowed region in the plane δ versus θ_{23} for strong thermal leptogenesis. The results is shown with a red line in Fig. 10. This well reproduces the result in [40]. The small differences can be ascribed

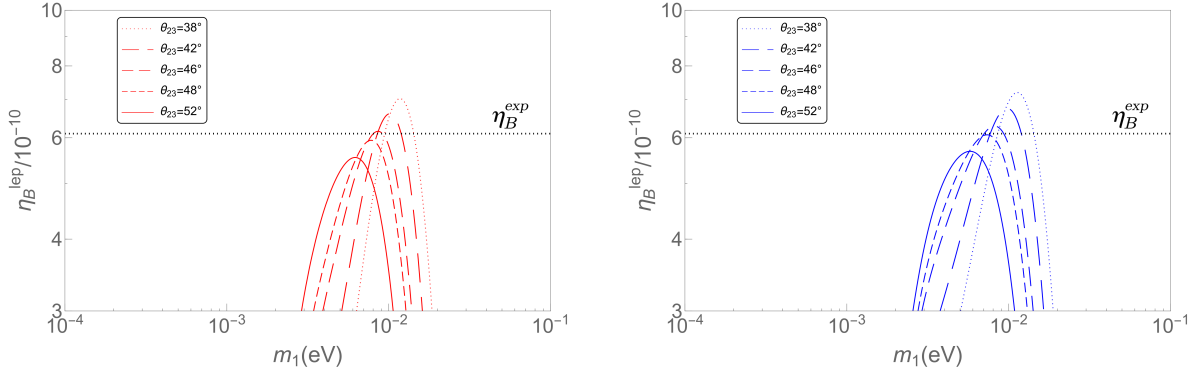


Figure 11: Plots of the final value of η_B^{lep} versus m_1 for different values of θ_{23} , for fixed $\delta = -50^\circ$ and for values of the other parameters maximising the asymmetry shown in Table 3. The right (left) panel is for the case with (no) flavour coupling.

	θ_{12}	θ_{13}	θ_{23}	δ	ρ/π	σ/π	θ_{12}^L	θ_{13}^L	θ_{23}^L	ρ_L/π	σ_L/π	δ_L/π
No F.C.	33.62°	9.09°	See Figure	-49°	0.312	0.834	4.11°	0.183°	2.41°	1.92	1.24	0.18
With F.C.	33.62°	9.09°	See Figure	-49°	0.312	0.834	4.11°	0.183°	2.41°	1.92	1.24	0.18

Table 3: Values of the low energy neutrino parameters in U (except θ_{23} given in Fig. 11) and six parameters in the unitary matrix V_L for the plots shown in Fig. 11.

to the updated experimental ranges for neutrino oscillation parameters in Eq. (6) we are adopting. The thin dashed lines are the 1σ , 2σ and 3σ experimental constraints from [23]. One can see how ST-SO10INLEP is in nice agreement with current experimental results within 1σ . The fact that currently the best fit is in the first octant releases the tension that was found with old experimental data in [40]. The best fit δ value is actually too low but this is very sensitive to the employed data sets in the analysis, mainly because of an existing tension between T2K and NOVA data. In the left panel of Fig. 11 we show plots of η_B^{lep} versus m_1 for the indicated values of θ_{23} and for fixed $\delta = -50^\circ$, corresponding to the value allowing the maximum value of θ_{23} . All other parameters are fixed to values that maximise the asymmetry and are given in Table 3. One can see how for increasing values of θ_{23} the peak value of η_B^{lep} decreases. This is well understood since the asymmetry is approximately $\propto 1/\sin^4 \theta_{23}$ [12, 33]. One can see how the maximum value of θ_{23} we obtain is 46° . This is consistent with the value obtained in [40] and it clearly shows how second octant is incompatible with ST-SO10INLEP.

4.2 Including flavour coupling

We have now finally to discuss the impact of flavour coupling on ST-SO10INLEP. This is one of the main motivations of the paper. A naive, but legitimate, expectation is that

flavour coupling might completely disrupt ST-SO10INLEP. The reason is that a large pre-existing asymmetry, thanks to flavour coupling, can find its way to survive into the tauon flavour where the washout is weak. As we are going to explain, this expectation is incorrect.

We have first of all to generalise the expression (49) for the relic value of the pre-existing asymmetry. If we start with an initial pre-existing asymmetry $N_{B-L}^{\text{p,i}}$, that we assume to be generated at some temperature above $\sim 10^{12}$ GeV, below 10^{12} GeV the tauon component of lepton states is measured by tauon charged lepton interactions and the pre-existing asymmetry will be the sum of a tauon component, given by $p_{\text{p}\tau}^0 N_{B-L}^{\text{p,i}}$, and a component along a flavour direction given by the projection of the flavour composition of pre-existing leptons on the e - μ flavour plane given by $(1 - p_{\text{p}\tau}^0) N_{B-L}^{\text{p,i}}$. In addition, in general, there should be also the phantom terms that however we neglect just to shorten the notation (in any case again results are not sensitive to them). Similarly to the solution for the asymmetry generated by leptogenesis, at $T \sim M_2$ the N_2 -washout, in the presence of flavour coupling, uncouples in a rotated flavour basis τ' and $\tau_2^\perp$. The difference is that now there is no source term and, therefore, at the end of the washout processes the washout is described by exponential factors and, rotating back to the standard flavour basis, one finds

$$\begin{aligned}
N_{\Delta_\tau}^{\text{p},T \lesssim T_L} &= U_{\tau\tau_2^\perp}^{-1} \left[U_{\tau_2^\perp \tau_2^\perp} (1 - p_{\text{p}\tau}^0) + U_{\tau_2^\perp \tau} p_{\text{p}\tau}^0 \right] N_{B-L}^{\text{p,i}} e^{-\frac{3\pi}{8} (K_{2e} + K_{2\mu})} \\
&\quad + U_{\tau\tau'}^{-1} \left[U_{\tau' \tau_2^\perp} (1 - p_{\text{p}\tau}^0) + U_{\tau' \tau} p_{\text{p}\tau}^0 \right] N_{B-L}^{\text{p,i}} e^{-\frac{3\pi}{8} K_{2\tau}}, \\
N_{\Delta_{\tau_2^\perp}}^{\text{p},T \lesssim T_L} &= U_{\tau_2^\perp \tau_2^\perp}^{-1} \left[U_{\tau_2^\perp \tau_2^\perp} (1 - p_{\text{p}\tau}^0) + U_{\tau_2^\perp \tau} p_{\text{p}\tau}^0 \right] N_{B-L}^{\text{p,i}} e^{-\frac{3\pi}{8} (K_{2e} + K_{2\mu})} \\
&\quad + U_{\tau_2^\perp \tau'}^{-1} \left[U_{\tau' \tau_2^\perp} (1 - p_{\text{p}\tau}^0) + U_{\tau' \tau} p_{\text{p}\tau}^0 \right] N_{B-L}^{\text{p,i}} e^{-\frac{3\pi}{8} K_{2\tau}}
\end{aligned} \tag{50}$$

$$\begin{aligned}
N_{\Delta_{\tau_2^\perp}}^{\text{p},T \lesssim T_L} &= (1 - p_{\text{p}\tau}^0) (1 - p_{\text{p}\tau_2^\perp}^0) N_{B-L}^{\text{p,i}}.
\end{aligned} \tag{51}$$

where the matrix U is the same we introduced in Eq. (45) when we introduced flavour coupling for the asymmetry produced by N_2 -decays. It can be noticed that, for $U = I$, one recovers the result in the absence of flavour coupling and N_1 -washout in Eq. (49). Also notice that the component of the pre-existing asymmetry along the flavour $\tau_2^\perp$, orthogonal to τ and $\tau_2^\perp$, escapes the washout at $T \sim M_2$ [34, 55].

We have now still to evolve the asymmetry down to temperatures below M_1 . First of all, at temperatures $T \sim 10^9$ GeV also RH muon interactions become effective and break the coherence of $\tau^\perp$ -flavour leptons. The asymmetries in the electron and muon flavour,

prior to N_1 -washout, can then be written as

$$N_{\Delta_e}^{\text{p}, T \gtrsim M_1} = p_{e\tau_2^\perp}^0 N_{\Delta_{\tau_2^\perp}}^{\text{p}, T \lesssim T_L} + (1 - p_{e\tau_2^\perp}^0) N_{\Delta_{\tau_2^\perp}}^{\text{p}, T \lesssim T_L}, \quad (52)$$

$$N_{\Delta_\mu}^{\text{p}, T \gtrsim M_1} = p_{\mu\tau_2^\perp}^0 N_{\Delta_{\tau_2^\perp}}^{\text{p}, T \lesssim T_L} + (1 - p_{\mu\tau_2^\perp}^0) N_{\Delta_{\tau_2^\perp}}^{\text{p}, T \lesssim T_L}. \quad (53)$$

Finally, the N_1 -washout stage can be taken into account exactly as for the asymmetry produced from leptogenesis, using an equation similar to Eq. (39) and Eq. (40), obtaining

$$N_{\Delta_\alpha}^{\text{p}, \text{f}} = \sum_{\alpha''} V_{\alpha\alpha''}^{-1} \left[\sum_{\beta=e,\mu,\tau} V_{\alpha''\beta} N_{\Delta_\beta}^{\text{p}, T \gtrsim M_1} e^{-\frac{3\pi}{8} K_{1\alpha''}} \right]. \quad (54)$$

Using now this much more intricate expression, we have obtained ST-SO10INLEP solutions shown, again with blue points, in the right panels of Fig. 1 and Figs. 2–8. The first thing to notice is that the region survives and in fact it even slightly increases. In particular, the most noticeable change is that the lower bound on m_1 gets relaxed and from $m_1 \gtrsim 8$ meV, obtained neglecting flavour coupling, we have now found $m_1 \gtrsim 3$ meV. One can also see from Fig. 8 that, corresponding, the lower bound on m_{ee} gets relaxed from $m_{ee} \gtrsim 9$ meV to $m_{ee} \gtrsim 4$ meV. One can also see from Fig. 6 that the two disconnected subregions now get linked by the appearance of the new solutions.

This implies, that in fact the washout of a pre-existing asymmetry, gets even easier when flavour coupling is taken into account, contrarily to naive expectations. The reasons why that happens were already partially envisaged in [33]. Even though flavour coupling induces part of the pre-existing asymmetry to be transferred in the tauon flavour where the washout is absent ($K_{1\tau} \ll 1$) to let the asymmetry produced from leptogenesis to survive, the off-diagonal term that is responsible for the coupling is proportional to $K_{1\tau}$ itself and so intrinsically small. That means that the same conditions that are necessary to have strong thermal leptogenesis are also sufficient to protect it from the potential disruptive effect of flavour coupling. In addition flavour coupling makes in a way that the flavoured decay parameters responsible for the washout of the pre-existing electron and muon asymmetries are not the standard K_{1e} and $K_{1\mu}$ but these are replaced by K_{1e}'' and $K_{1\mu}''$. The origin of the lower bound on m_1 was studied in detail in [32, 12] and it was due to the fact that $K_{1e} \simeq m_{ee}/m_\star \simeq 0.8 m_1/m_{ee}$. For an initial electronic pre-existing asymmetry $\sim 10^{-4}$, the strong thermal conditions implies $K_{1e} \gtrsim 8$ that translates immediately into $m_{ee} \gtrsim 8$ meV and $m_1 \gtrsim 9$ meV. However, when flavour coupling is included the washout of the electron asymmetry is described by K_{1e}'' that is a linear combination primarily of K_{1e} but with a small contribution from $K_{1\mu} \simeq (m_{\text{sol}} s_{12}^2 + m_{\text{atm}} s_{13} c_{23})^2 / m_\star m_{\text{sol}} s_{12}^2 \simeq 25$. This contamination from a large $K_{1\mu}$ makes possible to have smaller values of K_{1e} and, therefore, of m_{ee} and m_1 .

5 Final remarks

We could benefit of the analytical description of SO10INLEP given in [33] to re-derive the allowed region in the space of low energy neutrino parameters including flavour coupling. The results confirm the broad picture already discussed in previous papers but we noticed how we are currently entering a new stage for the prospect of testing SO10INLEP. There are also interesting new aspects introduced by flavour coupling. Let us briefly summarise our results.

- The cosmological upper bound Eq. (2) definitively rules out quasi-degenerate neutrinos and the latest upper bounds from $0\nu\beta\beta$ and tritium decay experiments confirm this so far. Absolute neutrino mass scale experiments are then starting to test the most interesting absolute neutrino mass range for the opportunities to test SO10INLEP, since for $10\text{ meV} \lesssim m_1 \lesssim 30\text{ meV}$ there is a stringent upper bound on the atmospheric mixing angle that excludes second octant.
- An interesting new feature introduced by flavour coupling is that new muon-dominated solutions appear in a portion of the space of parameters that would be otherwise excluded if flavour coupling is not taken into account. There are also some completely new electron-dominated solutions appearing (in already allowed region of parameter space) but these are quite marginal and do not seem to have a particular interest.
- We have confirmed the existence of ST-SO10INLEP solutions even when flavour coupling is taken into account. The allowed region gets even larger and, importantly, the lower bound on the absolute neutrino mass scale gets more relaxed. The washout of a large initial pre-existing asymmetry $N_{B-L}^{\text{p,i}} \sim 10^{-3}$ requires $m_{ee}(m_1) \gtrsim 4(3)\text{ meV}$. This is still large enough that future generation $0\nu\beta\beta$ experiments will be able to test it. The fact that ST-SO10INLEP is protected by flavour coupling is somehow a built-in feature: since $K_{1\tau} \lesssim 1$, this necessarily limits the possibility to have a too large leak of the pre-existing asymmetry from the e and μ flavours into the τ flavour so that the pre-existing asymmetry can be still efficiently washed-out for a subset of the SO10INLEP solutions. This is an important result of our paper since it shows that the ST-SO10INLEP scenario is stable under perturbations introduced by a more sophisticated calculation of the asymmetry.
- Our procedure, providing an approach toward a map of $SO(10)$ -inspired models, can also be useful within the quest for a realistic grandunified model since it gives a complementary analytical insight and a global picture of the solutions. Indeed

it has provided quite a useful insight in connection with a recent identification of a realistic fit of fermion parameters within a minimal $SO(10)$ scenario that is also able to reproduce the observed baryon asymmetry with N_2 -leptogenesis [56]. This falls within the category of SO10INLEP scenarios but extended to allow for a large, in fact maximal, value of θ_{23}^L .¹³

In conclusion, our results represent the first step toward a systematic understanding of theoretical uncertainties within SO10INLEP that is particularly important in view of future experimental expected progress both in long baseline neutrino oscillation experiments and absolute neutrino mass scale experiments. It is a very exciting time for the possibility of testing SO10INLEP. Both long baseline and absolute neutrino mass scale experiments will gradually test its predictions and constraints. In particular, $0\nu\beta\beta$ experiments are finally entering the mass range that is predicted by the strong thermal version of SO10INLEP. An account of flavour coupling has confirmed and even strengthened the picture, making the predictions more robust. At the same time, it also provides useful information for the optimization of successful strategy toward the identification of a grandunified realistic model.

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¹³Recently, it was also emphasized the role that a triplet Higgs might have in $SO(10)$ -leptogenesis and in the search of realistic minimal $SO(10)$ -models [57]. For a recent mini-review on studies embedding high scale thermal leptogenesis within grandunified models see [58].

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